Understanding the Impact of Nanoflare Heating Frequency on the Observed Emission Measure Distribution

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Introduction

- Nanoflare model of Parker [7]: corona heated by impulsive ($\ll \tau_{cool}$), low-energy (10^{24} erg) events produced by twisting, braiding of field lines rooted in the photosphere
- "Smoking gun" of nanoflare heating is the faint, high-temperature component of the emission measure distribution [3, 5]
- Fundamental question: what is the frequency of energy release in the solar corona? Two extreme cases:
  - Low-frequency heating: Time between successive events is much greater than a typical loop cooling time (i.e. approaches the single pulse case)
  - High-frequency heating: Time between successive events is much smaller than a typical loop cooling time (i.e. approaches the steady heating case)
- **Goal:** Use hydrodynamic loop models to determine the influence of heating parameters, including the heating frequency, on the emission measure distribution and associated observables.

Hydrodynamic Modeling

- Zero-dimensional Enthalpy-based Thermal Evolution of Loops (EBTEL) model of Klimchuk et al. [6], Cargill et al. [4] allows for efficient modeling of many thousands of loops
- We use a modified form of the EBTEL model to treat the electron and ion populations separately [for more details, see 1, submitted]
- Applying the "EBTEL method" to the two-fluid hydrodynamic equations [as given in 2], the modified two-fluid EBTEL equations are,

$$\frac{d}{dt}\bar{p}_e = \frac{\gamma-1}{L}\left[\psi_{TR} - (\mathcal{R}_{TR} + \mathcal{R}_C)\right] + k_B\bar{n}\nu_{ei}(\bar{T}_i - \bar{T}_e) + (\gamma-1)\bar{Q}_e, \tag{1}$$

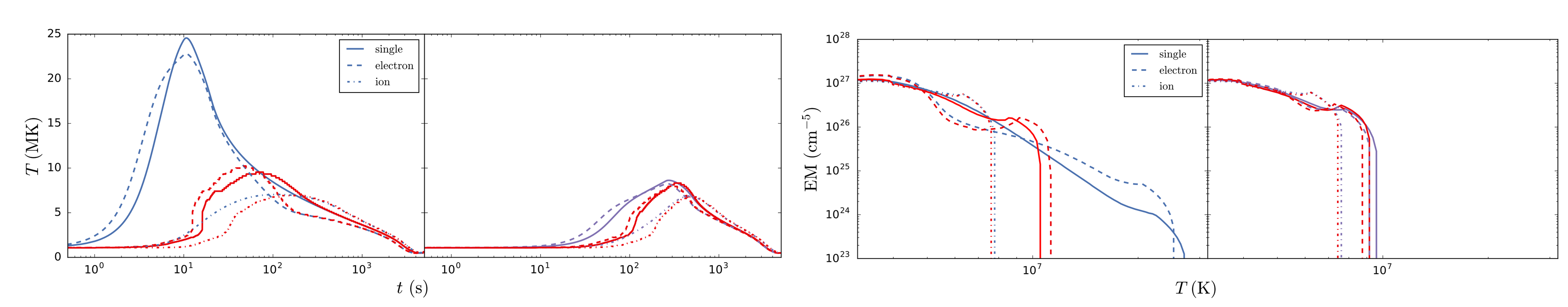
$$\frac{d}{dt}\bar{p}_i = -\frac{\gamma-1}{L}\psi_{TR} + k_B\bar{n}\nu_{ei}(\bar{T}_e - \bar{T}_i) + (\gamma-1)\bar{Q}_i, \tag{2}$$

$$\frac{d}{dt}\bar{n} = \frac{c_2(\gamma-1)}{c_3\gamma L k_B \bar{T}_e}\left(\psi_{TR} - F_{ce,0} - \mathcal{R}_{TR}\right), \tag{3}$$

where ψ_{TR} is a term included to maintain charge and current neutrality and ν_{ei} is the electron-ion binary Coulomb collision frequency
- Assume quasi-neutrality, $n_e = n_i = n$, and closed by equations of state for both the electrons and ions: $p_e = k_B n T_e$ and $p_i = k_B n T_i$

Single-nanoflare Results

- Single nanoflare is the most extreme low-frequency case, loop allowed to undergo complete heating and cooling cycle
- In Barnes et al. [1, submitted], we investigated the effect of pulse duration (τ), heat flux limiting, electron versus ion heating, and non-equilibrium ionization (NEI) on the resulting emission measure distribution, $\mathrm{EM}(T)$
- We found that,
  - While very short pulses ($\tau = 20, 40$ s) lead to significant emission above 10 MK, comparisons with field-aligned models show that EBTEL gives an artifically fast rise in density and thus an excess of hot emission for these very short heating times; longer pulses ($\tau = 200, 500$ s) show a cutoff near 10 MK.
  - Compared to pure Spitzer thermal conduction, heat flux limiting (using $f = 1/6$) extends $\mathrm{EM}(T)$ to > 10 MK; extreme values of f (e.g. $f = 1/30$) lead to significant emission > 20 MK.
  - In the case in which the ions are heated, no emission is visible above 8 MK, independent of the pulse duration
  - Calculating T_{eff} due to NEI shows that, even for very short pulses, there is little to no emission visible above 10 MK, for the single-fluid, electron heating, and ion heating cases
- Conclusion: $\mathrm{EM}(T)$ signature of loop plasma heated by a single nanoflare is most likely found in the temperature range $T_m < T < 10^7$ K, where the temperature of maximum emission $T_m \approx 4$ MK for active region cores [8]



(a) (b)

Figure 1: Equilibrium and non-equilibrium (red) ionization results for a single nanoflare lasting 20 s (blue) and 500 s (purple) in the single-fluid case (solid) as well as the case in which the only the elctrons (dashed) or only the ions (dot-dashed) are heated. Fig. 1(a) shows the electron temperature as a function of time for a 20 s pulse (left) and a 500 s pulse (right). Fig. 1(b) shows the resulting $\mathrm{EM}(T)$ for the equilibrium (left, blue and right, purple) and NEI (red) cases.

Energy Budget and Heating Statistics

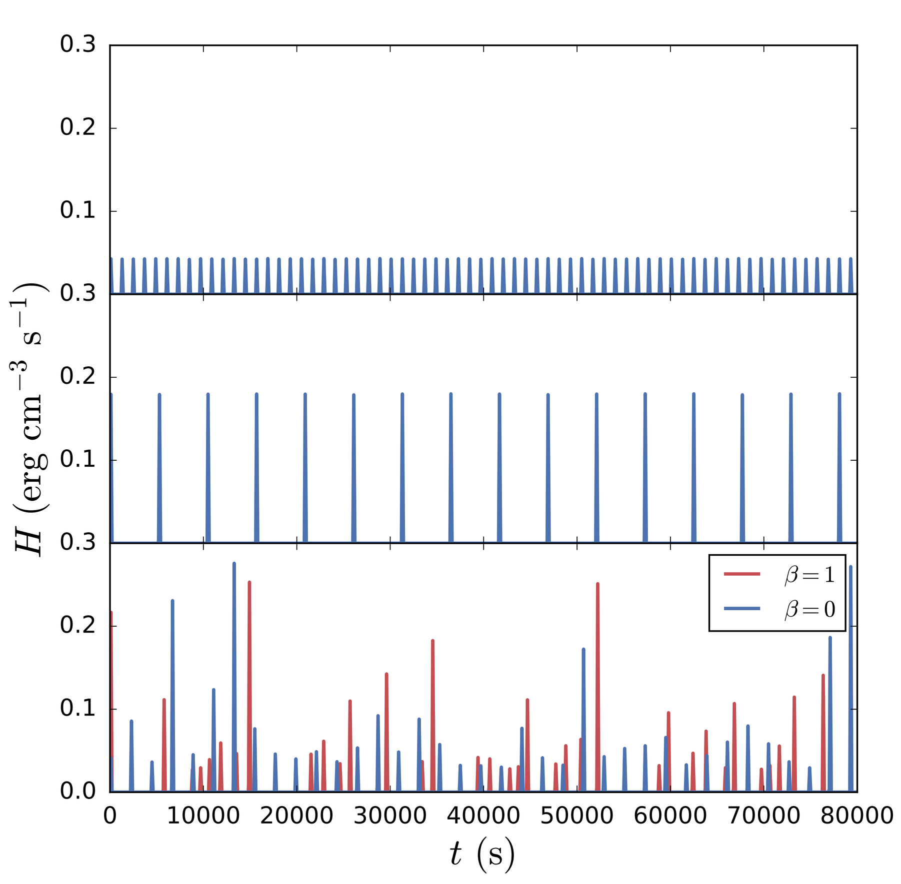


Figure 2: Top (Middle) panel shows Uniform heating amplitudes for $t_N = 1000$ ($t_N = 5000$) s. Bottom panel shows Heating amplitudes drawn from a power-law distribution with $\alpha = -1.5$ and mean wait time $t_N = 2000$ s; the events shown in red (blue) have wait times that depend on the previous event energy (uniform wait times).

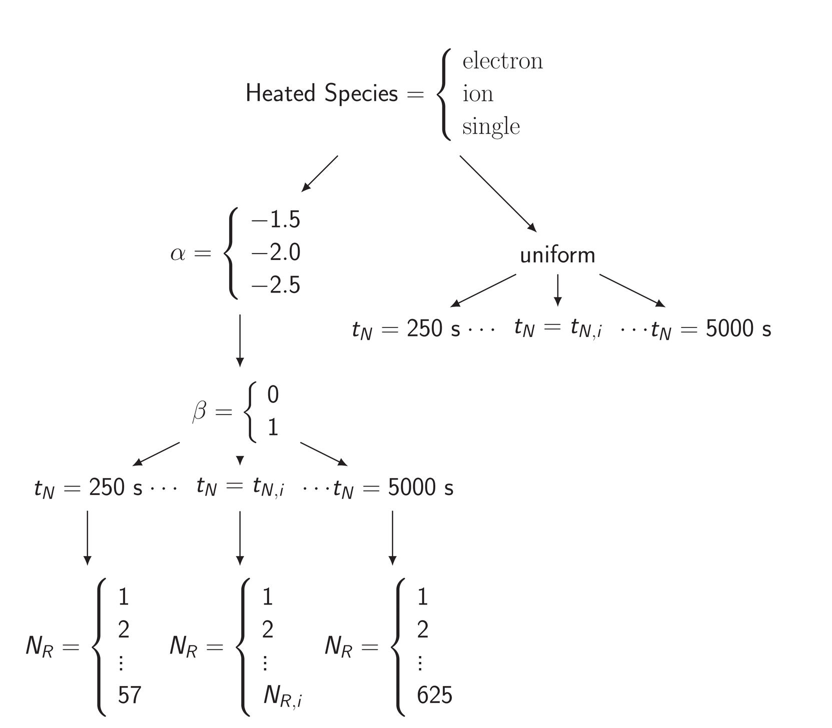


Figure 3: Heating function parameter space. We consider a range of waiting times $250 < t_N < 5000$ s, in increments of 250 s. In the power-law case, a sufficiently large number of runs, N_R is required to sample the distribution. For example, when $t_N = 5000$ s, $N_R = 625$ such that for each $(\alpha, \beta, t_N = 5000)$, we run the model 625 times.

- Loop of half-length $L = 40$ Mm heated by N triangular pulses with duration $\tau = 200$ s over $t_{total} = 8 \times 10^4$ s.
- Each event has maximum heating rate H_i and followed by a waiting time of $t_{N,i}$; static background heating $H_{bg} = 3.5 \times 10^{-5}$ erg cm^{-3} s^{-1}
- H_i can either be uniform such that $H_i = H_0$ for all i or chosen from a power-law distribution with $\alpha = -1.5, -2.0, -2.5$
- The total energy injected into the loop is constrained by

$$H_{eq} = \frac{1}{t_{total}}\sum_{i=1}^{N}\int_{t_i}^{t_i+\tau}\mathrm{d}t\ Q(t) = \frac{\tau}{2t_{total}}\sum_{i=1}^{N}H_i, \tag{4}$$

- $H_{eq} \approx 3.6 \times 10^{-3}$ erg cm^{-3} s^{-1} is the time-averaged heating rate such that $T_{peak} \approx 4$ MK, consistent with AR core observations [8].
- Treat $t_{N,i}$ as time needed for the field to "unwind", consistent with the Parker [7] nanoflare picture
  - $\beta = 0$: $t_{N,i} = t_N$ for all i, no dependence on H_i
  - $\beta = 1$: $\varepsilon = LA\tau H_i/2 \propto t_{N,i}$
- Total number of events dependent on t_N, $N = t_{total}/(t_N + \tau)$ such that $N = 16$ when $t_N = 5000$ s
- For the power-law cases, require $NN_R \sim 1 \times 10^4$, where N_R is the number of runs for each unique point in the parameter space, (α, β, t_N)

Emission Measure Distribution

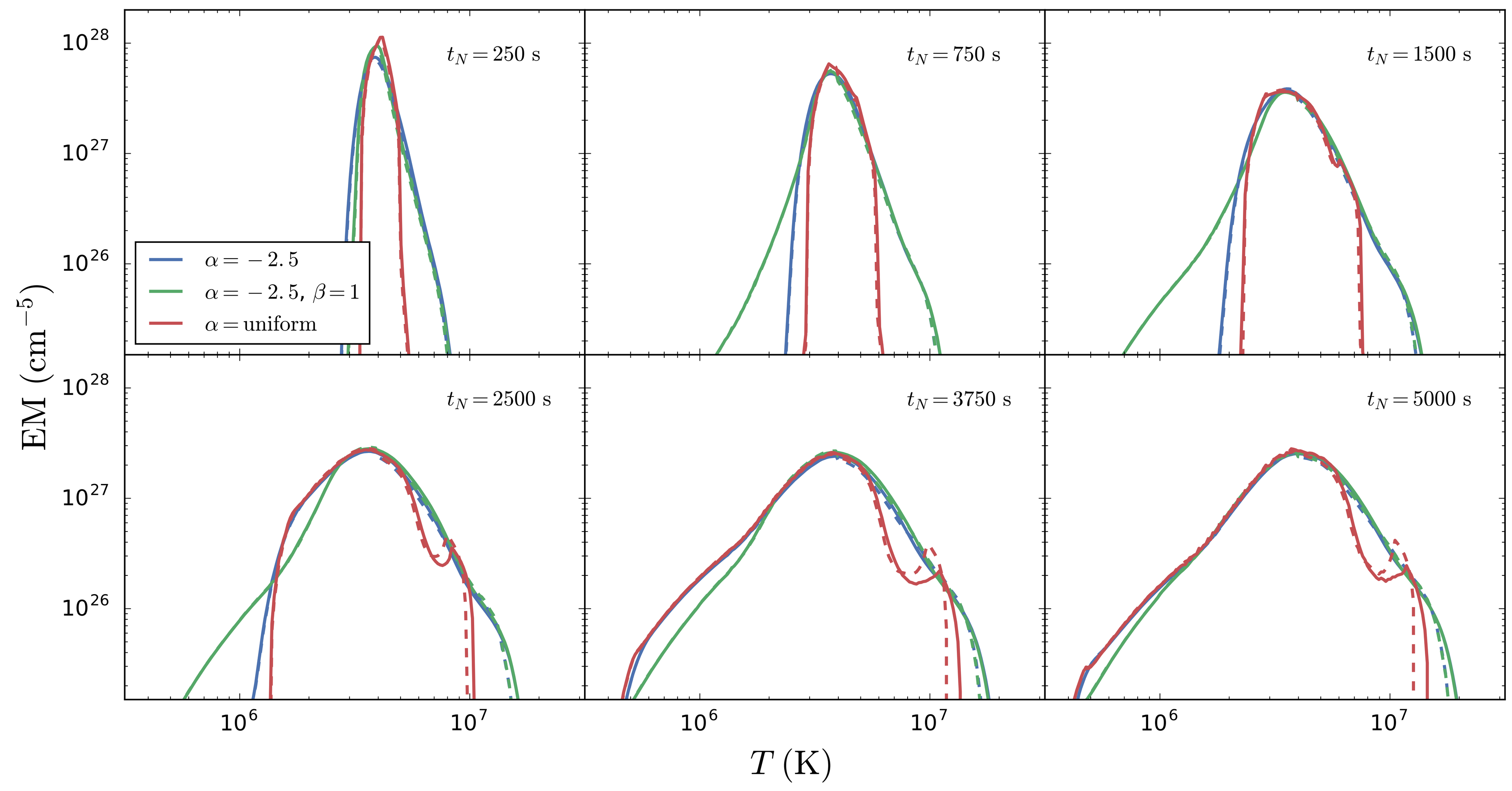


Figure 4: Lorem ipsum

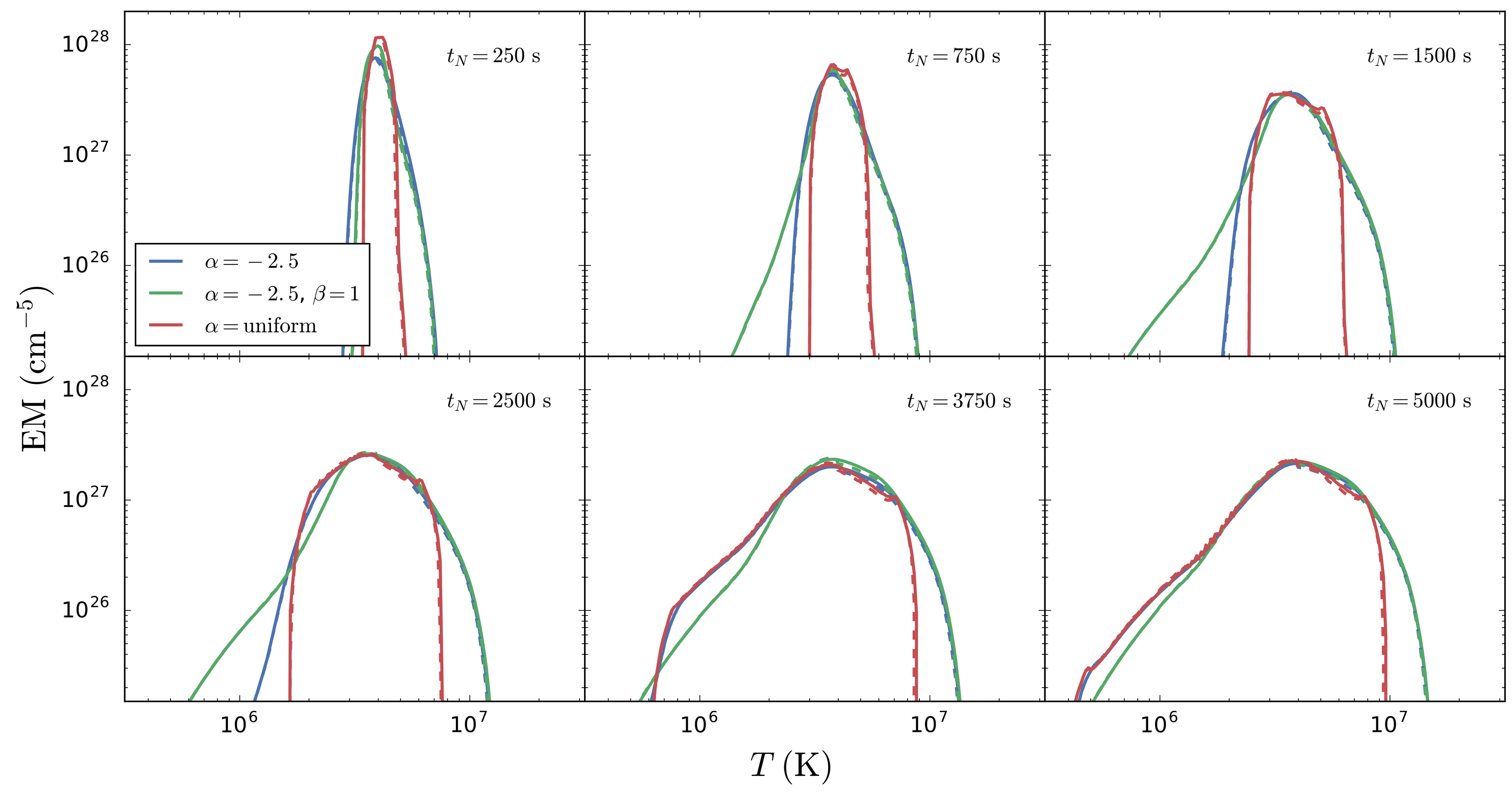
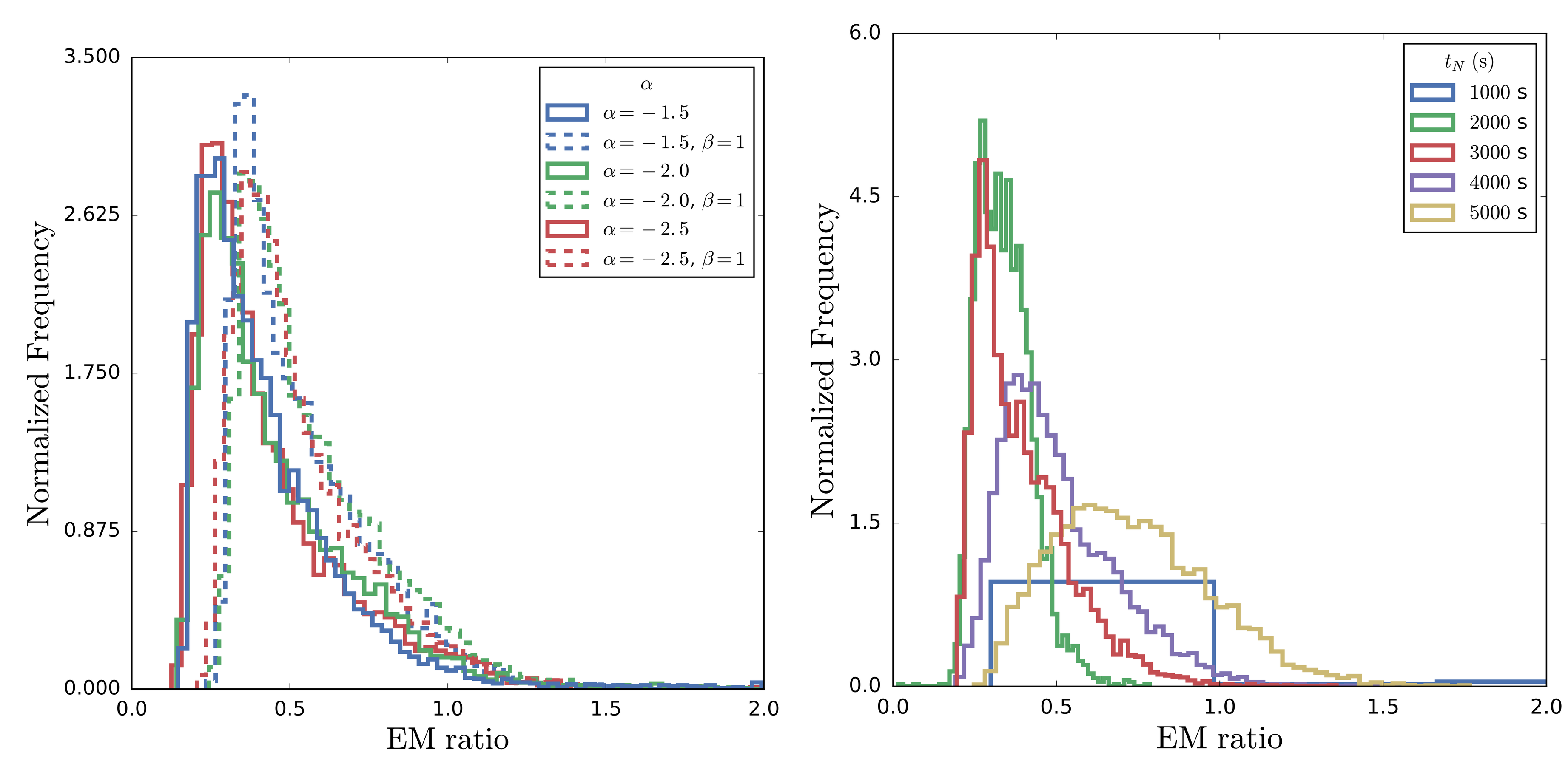


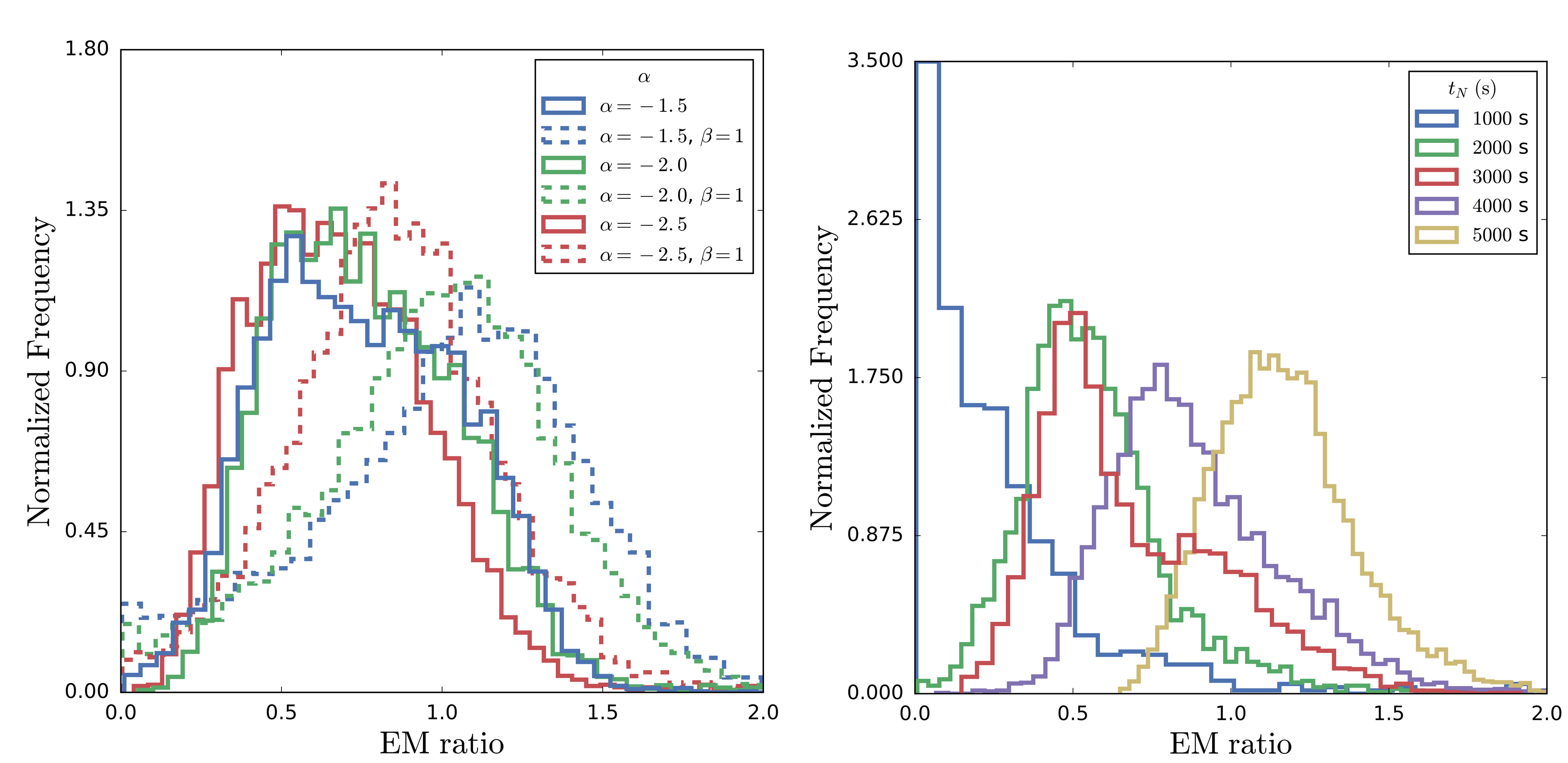
Figure 5: Lorem ipsum

Hot Plasma Diagnostics



(a) (b)

Figure 6: Lorem ipsum



(a) (b)

Figure 7: Lorem ipsum

Conclusions

References

[1] Barnes, W. T., Cargill, P. J., & Bradshaw, S. J. 2016, submitted
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[6] Klimchuk, J. A., Patsourakos, S., & Cargill, P. J. 2008, The Astrophysical Journal, 682, 1351

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