

COSMOLOGICAL CONSTRAINTS FROM THE REDSHIFT DEPENDENCE OF THE VOLUME EFFECT: GALAXY ANGULAR 2-POINT CORRELATION FUNCTION

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ABSTRACT

We propose to use the redshift dependence of the galaxy angular 2-point correlation function (2pCF) to constrain cosmological parameters. When an incorrect set of parameters is assumed for the coordinate transformation from redshift space to comoving space, it produces *redshift-dependent mis-scaling* in the constructed galaxy distribution. As a result, we observe a redshift-dependent rescaling of the measured galaxy angular 2pCF. Using mock galaxy catalogues (containing millions of galaxies at redshifts 0, 0.5, 1, 1.5, 2) from the Horizon Run 4 N-body simulation, we find that the redshift evolution of the angular 2pCF shape, as a function of scale, is sensitive to the adopted cosmological parameters. However, it is relatively insensitive to the gravitational growth of structure, the galaxy bias, and the effect of redshift space distortions. Analyzing the 2pCF shape on scales $5h^{-1}\text{Mpc} \leq r_{\perp} \leq 40h^{-1}\text{Mpc}$, we derive tight constraints on Ω_m and w with 68.3% CL uncertainties of 0.015 and 0.05. The method could be applicable to future large scale structure surveys, especially photometric surveys such as DES, LSST, to derive tight cosmological constraints. This work is a continuation of our previous works (Li et al. 2014, 2015, 2016) as a strategy to constrain cosmological parameters using redshift-invariant physical quantities; they are all powerful in constrain cosmology and have a lot of merits, e.g., are less affected by the systematic effects introduced by non-linear clustering or RSD, does not require accurate modeling of galaxy clustering, and are applicable on relative small clustering scales.

Keywords: large-scale structure of Universe — dark energy — cosmological parameters

1. INTRODUCTION

The discovery of cosmic acceleration (Riess et al. 1998; Perlmutter et al. 1999) implies the existence of a “dark energy” component in the Universe or the breakdown of Einstein’s gravity theory on cosmological scales. The theoretical explanation and observational probe of cosmic acceleration has attracted tremendous effort in the last two decades and are still far from well understood or accurately measured (Weinberg 1989; Li et al. 2011; Weinberg et al. 2013).

In an effort to probe the cosmic expansion history the large scale structure (LSS) surveys enable measurements of two geometrical quantities, the angular diameter distance D_A and the Hubble factor, H . If they were precisely measured as a function of redshift, tight constraints can be placed on cosmological parameters, e.g. the matter density Ω_m and the equation of state (EoS) of dark

energy w .

When an incorrect cosmological model is assumed for the coordinate transformation from redshift space to comoving space, geometric distortions are produced in the constructed galaxy distribution. These distortions include the “volume effect” which is the mis-scaling of the size of structures, and the Alcock-Paczynski (AP) effect, the shape distortion induced by the fact that distances along and perpendicular to the line of sight are fundamentally different.

Various statistical methods have been proposed to measure these effects. The AP effect was first suggested to be measured through the anisotropic clustering of galaxies (Ballinger Peacock & Heavens 1996; Matsubara & Suto 1996). Later, with the advent of large observational spectroscopic programs, this was applied to a series of LSS surveys (Outram et al. 2004; Blake et al. 2011; Chuang & Wang 2012; Reid et al. 2012; Beutler et al. 2013; Linder et al. 2014; Song et al. 2014; Jeong et al.

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2014; Sutter et al. 2014; López-Corredoira 2014; Alam et al. 2016; Beutler et al. 2016; Sanchez et al. 2016). Although there have also been alternative approaches using the symmetry properties of galaxy pairs (Marinoni & Buzzi 2010; Jennings et al. 2011; Bueno Belloso et al. 2012) and cosmic voids (Ryden 1995; Lavaux & Wandelt 2012; Sutter et al. 2014; Mao et al. 2016). Methods proposed to measure the volume effect include the number counting of galaxy clusters (Press & Schechter 1974; Viana & Liddle 1996), the topology of LSS (Park & Kim 2010), the BAO scale (Eisenstein et al. 1998; Blake & Glazebrook 2003; Seo & Eisenstein 2003), and the shape of 2pCF and power spectrum (Sánchez et al. 2006; Sánchez et al. 2009).

In Li et al. (2014, 2015, 2016) we developed a novel strategy to probe D_A and H from the LSS data. We found that the AP or volume effect in general introduces a geometric distortion that evolves significantly with redshift; this phenomenon can be utilized to distinguish the AP effect from the other effects, such as the redshift-space distortion (RSD). The radial distances of galaxies are inferred from the measured redshifts, which are distorted by the peculiar motion of galaxies, leading to apparent distortion in the redshift space galaxy distribution (Jackson 1972; Kaiser 1987; Ballinger Peacock & Heavens 1996). It is the major systematic effects limiting our ability to probe the geometry of LSS from the galaxy distribution. However, we found that the influence of RSD is significantly reduced when we consider only the redshift dependence of the distortion.

Li et al. (2014) attempted to use the *galaxy density gradient field* to characterize the redshift evolution of the LSS distortion. We tested the idea on the Horizon Run 3 (HR3) N-body simulations (Kim et al. 2011), showing that the redshift evolution of the gradient field can be used to infer unbiased constraints on cosmological parameters and that it is, largely, insensitive to RSD. The same topic was revisited in Li et al. (2015), using a slightly modified version of the usual anisotropic galaxy two-point correlation function (2pCF), $\xi(\mu)$, where $\mu \equiv \cos\theta$ and θ is the angle between the connection of the galaxy pairs and the line of sight direction. When an incorrect cosmological model is adopted, the shape of $\xi(\mu)$ is altered due to the anisotropic distortion of the LSS (the AP effect), and the amplitude is shifted due to the change in the comoving volume elements; both effects have significant redshift dependence. Although the RSD effect significantly distorts $\xi(\mu)$, it exhibits much less redshift evolution. The method was applied to BOSS DR12 galaxies (Li et al. 2016) to probe the redshift dependence of the AP effect, and we arrive at tight cosmological constraints competitive with mainstream probes of SNIa, BAO, CMB.

In this paper we continue our exploration of this strategy and present a new method using the redshift evolution of the shape of the *galaxy angular 2pCF* as a function of radial scale, $\omega(r_\perp)$, to constrain cosmology. A wrongly adopted cosmology result in *redshift-dependent mis-scaling* of the galaxy distribution, which leads to redshift-evolution of the shape of $\omega(r_\perp)$. The signal is sensitive to the cosmological parameters and relatively insensitive to the other effects such as the gravitational growth of structure, the galaxy bias, and the RSD effect. We apply our method to HR4 mock galaxies, using hundreds of millions of galaxies at redshifts $0 \leq z \leq 2$ and develop the statistical procedure to constrain cosmological parameters.

The outline of this paper proceeds as follows. In §2 we briefly review the nature and consequences of the LSS geometric distortion when performing coordinate transforms in a cosmological context. In §3 we describe the N-body mock galaxies used to test our method. The methodology is presented in §4. We conclude in §5.

2. VOLUME EFFECT IN A NUTSHELL

In this section we briefly introduce the scaling effect caused by incorrectly assumed cosmological parameters. A more detailed description has been provided in Li et al. (2014, 2015, 2016).

Suppose that we are probing the size of some objects in the Universe. We measure their redshift span Δz and angular size $\Delta\theta$, then compute their sizes in the radial and transverse directions using the following formulas

$$\Delta r_\parallel = \frac{c}{H(z)} \Delta z, \quad \Delta r_\perp = (1+z) D_A(z) \Delta\theta, \quad (1)$$

where H is the Hubble parameter and D_A is the angular diameter distance. In the particular case of a flat universe composed of a cold dark matter component and a constant EoS dark energy, they take the forms

$$H(z) = H_0 \sqrt{\Omega_m a^{-3} + (1 - \Omega_m) a^{-3(1+w)}},$$

$$D_A(z) = \frac{1}{1+z} r(z) = \frac{1}{1+z} \int_0^z \frac{dz'}{H(z')}, \quad (2)$$

where $a = 1/(1+z)$ is the cosmic scale factor, H_0 is the present value of Hubble parameter and $r(z)$ is the comoving distance.

When incorrect values of Ω_m and w are adopted, the inferred Δr_\parallel and Δr_\perp are also wrong, resulting in wrong estimations of the object's shape (AP effect) and size (volume effect). These effects and their cosmological consequences have been studied in Li et al. (2014, 2015, 2016).

In this paper we focus on the mis-estimation of the angular size Δr_\perp . The ratio of mis-estimation is

$$\alpha_\perp \equiv \frac{\Delta r_{\perp, \text{wrong}}}{\Delta r_{\perp, \text{true}}} = \frac{D_{A, \text{wrong}}}{D_{A, \text{true}}}, \quad (3)$$

where “true” and “wrong” denote the values in the true cosmology and wrongly assumed cosmologies respectively.

An illustration is provided in the left panels of Figure 1. Suppose that the true cosmology is a flat Λ CDM with present matter ratio $\Omega_m = 0.26$ and standard dark energy EoS $w = -1$. If we were to distribute a series of perfect squares at distances ranging from 500 Mpc/h to 3000 Mpc/h, and an observer located at the origin were to measure their redshifts and infer the sizes of the squares using the distance-redshift relations of four incorrect cosmologies

- (i). $\Omega_m = 0.40$, $w = -1.0$,
- (ii). $\Omega_m = 0.15$, $w = -1.0$,
- (iii). $\Omega_m = 0.26$, $w = -0.5$,
- (iv). $\Omega_m = 0.26$, $w = -1.5$,

then the shapes of the squares appear distorted (AP effect), and their sizes are wrongly estimated (volume effect). Cosmological models (i,iii) result in compressed volume elements (in both angular and LOS direction), and the degree of compression increases with increasing

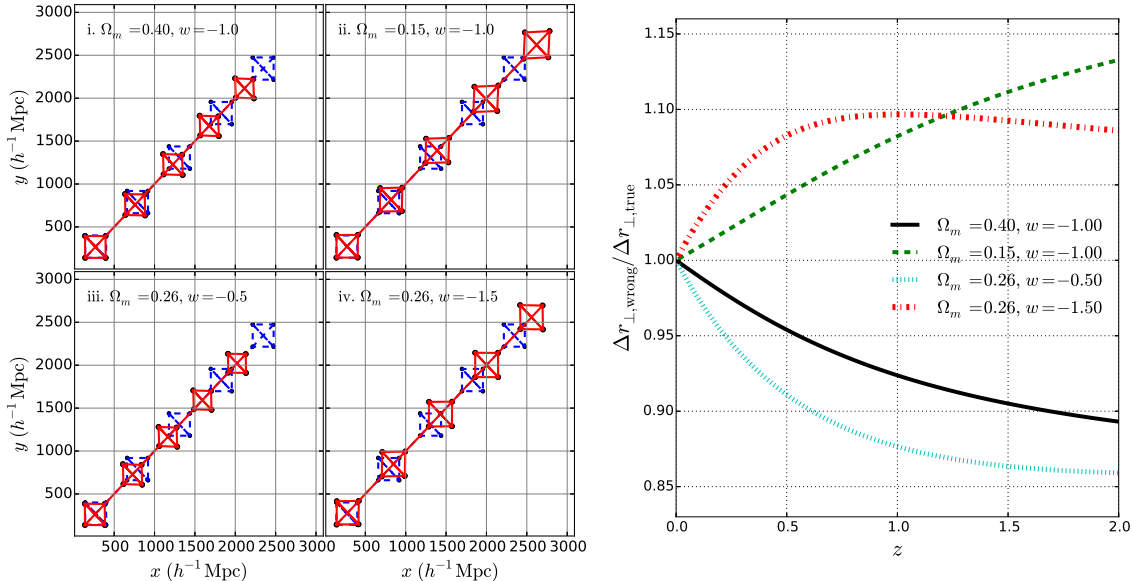


Figure 1. A description of how incorrectly assumed cosmologies can distort the interpretation of observation, assuming $\Omega_m = 0.26, w = -1$ is the true cosmology. Left panel shows a series of five regular squares, measured by an observer at the origin. Their true positions and shapes are plotted in blue dashed lines. When the observer adopts the wrong cosmologies to compute the distance from measured redshifts and infer the positions and shapes of the squares, he/she obtains distorted squares shown as the red solid lines. Right panel shows the redshift evolution of the wrongly estimated angular diameter distance (divided by the correct value).

distance; the situation is opposite for the other two cosmologies.

The mis-estimation of angular size, $\Delta r_{\perp, \text{wrong}} / \Delta r_{\perp, \text{true}}$, is displayed in the right panel of Figure 1. In all cosmologies, $\Delta r_{\perp, \text{wrong}} / \Delta r_{\perp, \text{true}}$ evolves significantly in the redshift range $0 \leq z \leq 2$. As an example, when adopting the quintessence cosmology $\Omega_m = 0.26, w = -0.5$, the angular size is underestimated by 8.9%, 12.3%, 13.6%, 14.1% at $z = 0.5, 1, 1.5, 2$.

In sum, as a consequence of incorrect cosmologies the size of objects is mis-estimated and the magnitude of mis-estimation depends on the redshift. We use the galaxy angular 2pCF to probe the mis-estimation of angular size, $\Delta r_{\perp, \text{wrong}} / \Delta r_{\perp, \text{true}}$.

3. THE SIMULATION DATA

We test the method using the mock galaxy samples produced by the Horizon Run 4 (HR4) N-body simulation (Kim et al. 2015; Hong et al. 2016).

HR4 was made within a cube of volume $(3.15 h^{-1} \text{Gpc})^3$ using 6300^3 particles with mass $m_p \simeq 9 \times 10^9 h^{-1} \text{M}_{\odot}$. The simulation adopted the second order Lagrangian perturbation theory (2LPT) initial conditions at $z_i = 100$ and a WMAP5 cosmology $(\Omega_b, \Omega_m, \Omega_{\Lambda}, h, \sigma_8, n_s) = (0.044, 0.26, 0.74, 0.72, 0.79, 0.96)$ (Komatsu et al. 2011).

Mock galaxies are produced from the simulation based on a modified one-to-one correspondence scheme (Hong et al. 2016). The most bound member particles (MBPs) of simulated halos are adopted as tracers of galaxies. The merger timescale is computed to obtain the lifetime of merged halos. Merger trees of halos are constructed by tracking their MBPs from $z = 12$ to 0; when a merger event occurs, the merger timescale is computed using the formula of Jiang et al. (2008) to determine when the satellite galaxy is completely disrupted.

The resulting mock galaxies were found to reproduce the 2pCF of the SDSS DR7 volume-limited galaxy sam-

ple (Zehavi et al. 2011) very well. The mock galaxies shows a similar finger of god (FOG) feature (Jackson 1972) as the observation. The projected 2pCF of the mock and observational samples agree within 1σ CL on scales greater than $1 h^{-1} \text{Mpc}$.

We take the snapshot data of the mock galaxies at $z = 0, 0.5, 1, 1.5, 2$. Setting a minimal halo mass of $3 \times 10^{11} h^{-1} \text{M}_{\odot}$, we select 457, 406, 352, 306 and 228 million mock galaxies at the five redshifts, corresponding to a number density of 1.46, 1.30, 1.13, 0.98 and 0.73 in units of $10^{-2} h^3 \text{Mpc}^{-3}$ respectively. Applying a uniform mass cut, at higher redshift we obtained a smaller number of galaxies with *larger bias*.

As an illustration, Figure 2 displays the 2D distribution of a subsample of mock galaxies in five redshifts, with x, y coordinates computed using the “correct” cosmological parameters $\Omega_m = 0.26, w = -1$ (left panels) and a set of wrong cosmological parameters $\Omega_m = 0.05, w = -1.5$ (right panels), respectively.

When the correct cosmology is adopted the cosmic scale is correctly estimated, the main factor leads to redshift evolution of galaxy distribution is the gravitational growth of structure. With decreasing redshift, we clearly see the clusters and filaments form and grow. On the other hand, when the wrong cosmology is adopted, there is an artificial scaling of distances in the constructed map. The separations among galaxies are over-estimated by 25.6%, 47.3%, 62.2%, 71.7% at redshifts of 0.5, 1.0, 1.5, 2.5, leading to a clear redshift evolution of sizes of structures.

The growth of structure increases the clustering and results in more compact structures, while the volume effect maintains the clustering pattern and uniformly rescales structures on all scales. The imprint of each of these effects on the large scale structure is different and should be distinguishable. In the following section we show that they affect the angular correlation function

measurements in different ways, and thus can be easily separated.

4. METHODOLOGY

We use the angular 2pCF as a statistical tool to probe the volume effect. The galaxy 2pCF as a function of galaxy separation in the angular direction, $\omega(r_\perp)$, is computed for snapshot data of mock galaxies at redshifts 0.5, 1, 1.5 and 2. We adopt the Landy-Szalay estimator (Landy & Szalay 1993),

$$\omega(r_\perp) = \frac{DD - 2DR + RR}{RR}, \quad (4)$$

where DD is the number of galaxy–galaxy pairs, DR the number of galaxy–random pairs, and RR is the number of random–random pairs, all separated by a distance defined by $r_\perp \pm \Delta_{r_\perp}$ where we choose $\Delta_{r_\perp} = 1h^{-1}\text{Mpc}$. The random catalogue consists of unclustered points uniformly distributed in the same space as the simulated data. In an effort to reduce the statistical variance of the estimator, we use 50 times as many random points as we have galaxies.

Considering the large number of galaxies and random points the 2pCF is computed part by part in subsamples with size of $1575 h^{-1}\text{Mpc} \times 1575 h^{-1}\text{Mpc} \times 105 h^{-1}\text{Mpc}$. The ‘sheet’-like shape of the subsample is similar to the shape of redshift shells in the real observational case. The Z direction with thickness $105 h^{-1}\text{Mpc}$ is treated as the radial direction (the r direction) and the X-Y directions are the angular plane. For our simulation box size we have 120 such subsamples. The average of the measurements in all subsamples is taken as the 2pCF of the whole sample, while the covariances of them are taken to be the covariance matrix (after multiplying by a factor of $1/\sqrt{119}$).

4.1. Galaxy angular 2pCF: at different redshifts

The upper-left panel of Figure 3 displays the angular 2pCF measured from HR4 mock galaxies. We multiply ω by the separation r_\perp to have similar size of statistical uncertainty on all scales. For convenience we will use the denotation

$$\hat{\omega}_{r_\perp} \equiv r_\perp \omega. \quad (5)$$

Among the different redshift bins there is large variation in the amplitude of $\omega(r_\perp)$. The amplitude is proportional to the clustering strength, which is affected by the gravitational growth of structures and the galaxy bias. The amplitude is highest at $z = 0$ where the structures experienced the most growth. $\omega(r_\perp)$ increases at $z > 1$ with increasing redshift as a reason of increasing galaxy bias.

Although there is large variation in amplitude between the different redshifts bins, the shape of ω_{r_\perp} remains similar at all redshifts; in general it peaks at $r \sim 9h^{-1}\text{Mpc}$ and monotonically drops or increases at larger or smaller scales. The only exception is the small enhancement at $r \lesssim 2h^{-1}\text{Mpc}$, which is caused by the non-linear growth of structures and is much more significant at lower redshift.

In order to directly compare the shape of the 2pCF at different redshifts, in the middle panel of Figure 3 we show the $r_\perp \omega$ normalized by the overall amplitude within $5h^{-1}\text{Mpc} < r < 40h^{-1}\text{Mpc}$ (here after $\hat{\omega}_{r_\perp}$):

$$\hat{\omega}_{r_\perp} \equiv \frac{r_\perp \omega(r_\perp)}{\int_{r_{\perp,\min}}^{r_{\perp,\max}} r_\perp \omega(r_\perp) dr_\perp / (r_{\perp,\max} - r_{\perp,\min})}, \quad (6)$$

where we choose $r_{\perp,\min} = 5h^{-1}\text{Mpc}$, $r_{\perp,\max} = 40h^{-1}\text{Mpc}$ in this analysis. Below $5h^{-1}\text{Mpc}$ the non-linear growth of structure leads to systematic redshift evolution which could be difficult to be reliably accounted; On scales larger than $40h^{-1}\text{Mpc}$, the analytical modeling of the shape of the angular 2pCF is relatively well understood; one can just fit the 2pCF with the theoretical predictions (Bernardeau et al. 2002; Salvador et al. 2014, 2016) instead of using our method (although our method should also be applicable).

In the upper-middle panel, the overlapping of $\hat{\omega}_{r_\perp}$ clearly show the minimal redshift evolution of the shape. In the upper-right panel, we further show the residual evolution at high redshifts with respect to $z = 0$. There is only 1-4% enhancement at $r < 10h^{-1}\text{Mpc}$, and < 1.5% suppression at $r > 25h^{-1}\text{Mpc}$. The trend is monotonic with redshift.

There is one detail in the analysis that should be emphasized here. In real observations we may be selecting *different* types of galaxies at different redshifts. Considering this fact in the comparison of $\hat{\omega}_{r_\perp}$ we compare subsamples of galaxies at *different* locations. As an example, if at $z = 0$ we take $\hat{\omega}_{r_\perp}$ measured within $0h^{-1}\text{Mpc} < Z < 105h^{-1}\text{Mpc}$, at higher redshifts we then adopt measurement within $105h^{-1}\text{Mpc} < Z < 210h^{-1}\text{Mpc}$ for a comparison. If we simply compare the 2pCF of the *same* subsample of galaxies at different redshifts, one would significantly underestimate the statistical uncertainty of $\delta\hat{\omega}_{r_\perp}$ by ignoring cosmic variance.

4.2. Galaxy angular 2pCF: cosmological effect

The angular 2pCF of galaxy positions constructed in a wrongly assumed cosmology $\Omega_m = 0.05$, $w = -1.5$ is displayed in the lower panels of Figure 3. The mis-scaling uniformly shifts the clustering pattern on all scales, leading to a biased $\hat{\omega}_{r_\perp}$, which is related with the true $\hat{\omega}_{r_\perp}$ as

$$\hat{\omega}_{r_\perp,\text{wrong}}(r) = \hat{\omega}_{r_\perp,\text{correct}}(\alpha_\perp r), \quad (7)$$

a simple consequence of the fact that the clustering pattern at scale r is rescaled to $\alpha_\perp r$.

The redshift evolution of α_\perp leads to redshift evolution of $\hat{\omega}_{r_\perp,\text{correct}}$, which is displayed in the lower-middle panel. The upscaling of comoving distance leads to a stretched shape. When increasing redshift, the peak location is shifted from $\sim 9h^{-1}\text{Mpc}$ to $\sim 15h^{-1}\text{Mpc}$ at $z = 2$, as a result of the fact that the angular separation is upscaled by 71.7%. Correspondingly, the right panel shows that, compared with $\hat{\omega}_{r_\perp}$ at $z = 0$ there is a 20-40% change at higher redshifts.

The mis-scaling not only changes the shape of the 2pCF but also modifies its amplitude. As is shown in the left panel, due to the stretch of scale the amplitude is enhanced at higher redshifts; the higher the redshift, the greater the enhancement.

Although the change of amplitude could be a more significant cosmological consequence than the alteration of shape, it is mixed with the other effects; the growth of structure and a larger galaxy bias can also lead to a stronger clustering and thus an enhanced amplitude. In order to reliably extract the cosmological information, we just utilize the redshift evolution of the 2pCF shape, which is less affected by other factors.

4.3. Systematic effects

In Sec. 4.1 we showed that $\hat{\omega}_{r_\perp}$ measured from the constant mass cut samples at different redshifts show good

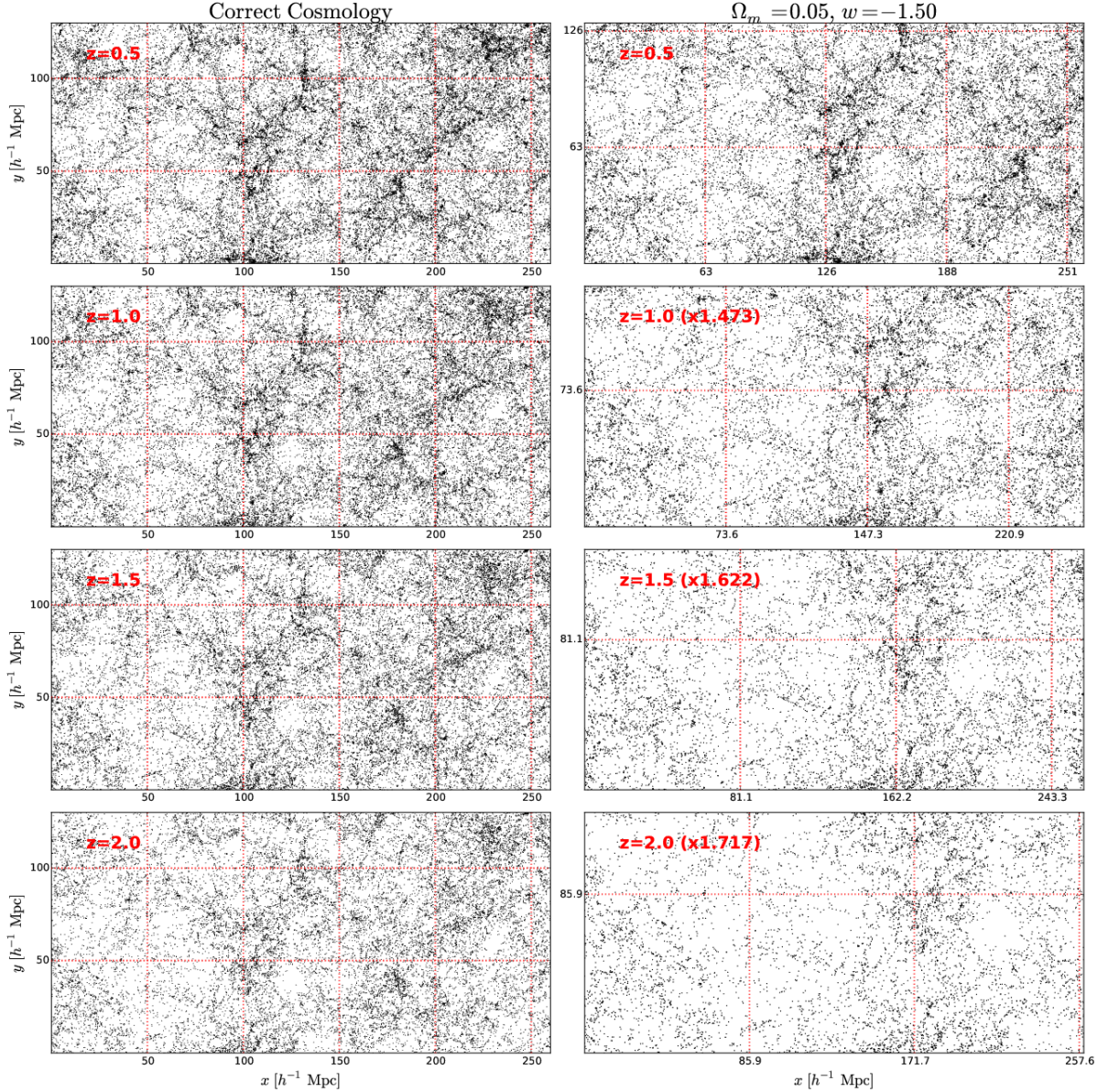


Figure 2. The 2D distribution of a $260 \times 130 \times 105 (h^{-1} \text{Mpc})^3$ subsample of HR4 mock galaxies, at redshifts 0.5, 1, 1.5 and 2. Left and right panels show the x, y coordinates computed using the “correct” parameters $\Omega_m = 0.26$, $w = -1$ and a set of wrong cosmological parameters $\Omega_m = 0.05$, $w = -1.5$, respectively. The growth of structure strengthens the clustering and make structures more compact at lower redshift. When the wrong cosmology is adopted, the comoving distances are mis-scaled and over-estimated by 25.6%, 47.3%, 62.2%, 71.7% at redshifts of 0.5, 1.0, 1.5, 2.5, respectively, leading to a clear redshift evolution of the sizes of structures.

agreement. The gravitational growth of structure will not have a large impact on $\hat{\omega}_{r\perp}$ on scales $\gtrsim 5h^{-1} \text{Mpc}$, yet there are many other factors that may induce redshift evolution including galaxy bias, RSD, redshift error, and the redshift evolution of galaxies properties such as mass, morphology, color, concentration. Here we test two “major” systematical effects; the RSD effect and galaxy bias.

4.3.1. Redshift Space Distortion

The galaxy peculiar velocity contaminates the observed redshift and distorts the inferred galaxy radial position. It is the major systematic in our previous works (Li et al. 2014, 2015, 2016) of statistical analysis on the 3D galaxy distribution.

The effect of RSD is much milder in this work. The angular positions of galaxies are not shifted by RSD at all; the only effect of RSD enter in the procedure of splitting

galaxies into redshift shells. The galaxies observed in a survey are split into shells of subsamples, with different redshift ranges. This allows us to obtain 2pCF measurement at various redshifts. RSD distorts the galaxy redshift and as a result some galaxies (especially those close to the boundaries of shells) are assigned to the wrong redshift shells.

We shift the radial coordinates of galaxies according to the relation

$$\Delta z = (1+z) \frac{v_Z}{c}, \quad (8)$$

where v_Z (Z is the third coordinate of the galaxy in the box, treated as the radial direction in this analysis) is the galaxy peculiar velocity in the LOS direction. This lead to *misclassification* of some galaxies when we split the box into slices.

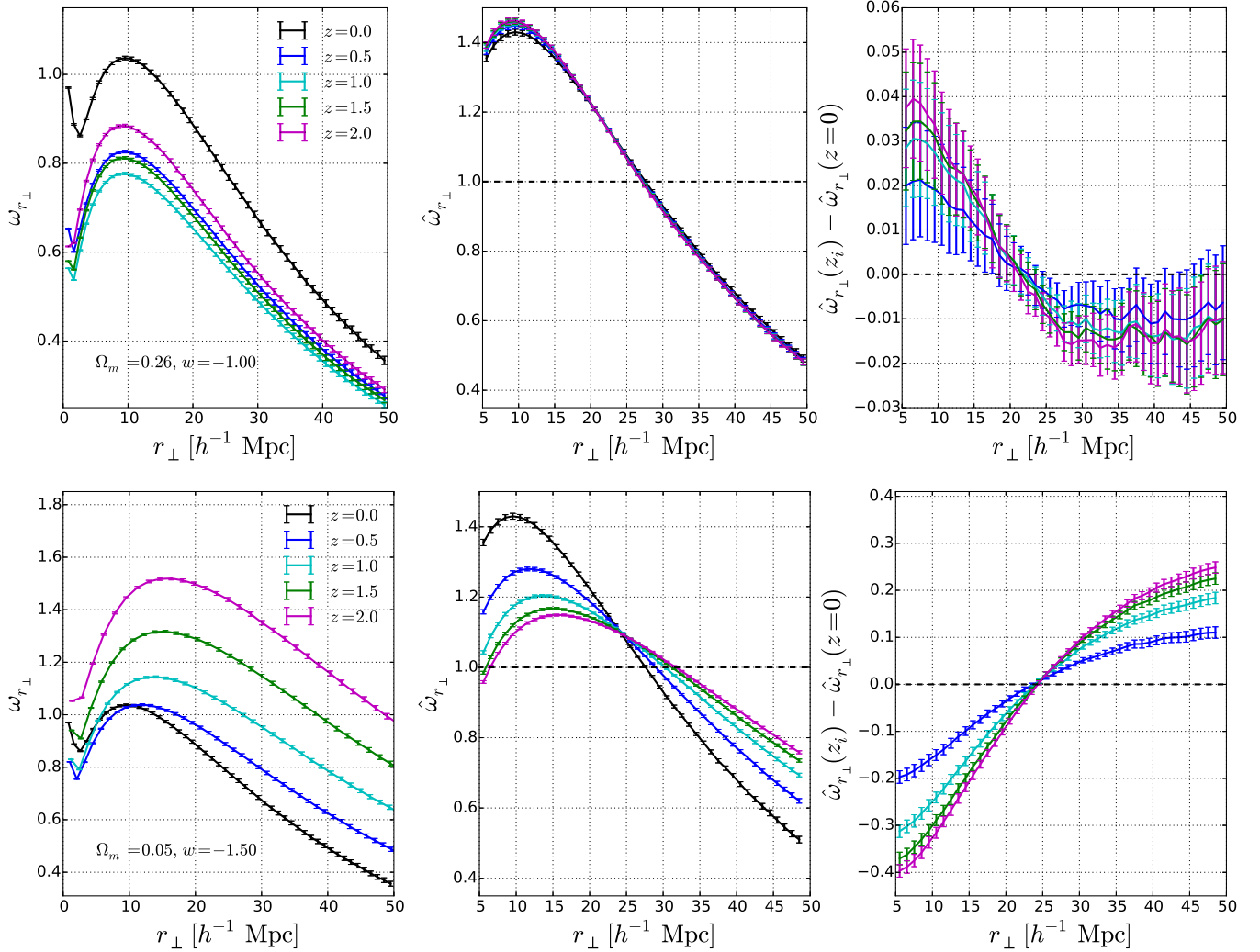


Figure 3. The angular 2pCF ($\omega_{r\perp}$, left panels), the shape of the angular 2pCF ($\hat{\omega}_{r\perp}$, middle panels), and the evolution with respect to the $z = 0$ value ($\hat{\omega}_{r\perp}(z = z_i) - \hat{\omega}_{r\perp}(z = 0)$, right panels). Upper panels show that in the correct cosmology the redshift evolution of the shape of 2pCF is rather small regardless of the large redshift evolution of the amplitude, whereas the lower panels show that, in a wrong cosmology $\Omega_m = 0.05, w = -1.5$, the redshift evolution of $\hat{\omega}_{r\perp}$ is very significant.

The upper panel of Figure 4 shows how it affects on the 2pCF. Comparing this plot with Figure 3 one can clearly see the effect of RSD. The amplitude of the measured 2pCF is enhanced by $\sim 10\%$ when considering the RSD effect and the slope of $\hat{\omega}_{r\perp}$ is suppressed. However, the redshift evolution of $\hat{\omega}_{r\perp}$ remains small, $\lesssim 4\%$ at $0.5 \leq z \leq 2$. Therefore RSD should not significantly affect the application of our method, as long as its redshift evolution is small; the small effect of RSD can be precisely modelled by simulations and corrected.

Similar to RSD, redshift errors of galaxies also results in fuzzy boundaries of redshift shells. Thankfully, there are some methods for mitigating this effect in photometric surveys using pair-center binning (Nock et al. 2010).

4.3.2. Galaxy Bias

Galaxies are biased tracers of the dark matter field; more massive galaxies reside in regions with higher density contrast and exhibit stronger clustering (i.e. larger bias).

We vary the mass cut and check the effect of galaxy bias on $\hat{\omega}_{r\perp}$. The second row of Figure 4 displays the

2pCF measured from galaxies within a $3175 \times 3175 \times 105$ ($h^{-1}\text{Mpc}$)³ volume, taken from the $z = 0$ snapshot. Four values of minimal mass limits, $3 \times 10^{11} M_\odot$, $1 \times 10^{12} M_\odot$, $4 \times 10^{12} M_\odot$ and $1 \times 10^{13} M_\odot$, are imposed to create samples with different galaxy biases. The measured angular 2pCF are displayed.

Different mass cuts result in large variation in the amplitude of the 2pCF (left panel); but the shape of the normalised 2pCF remains less affected (middle panel). Compared with the sample of $M > 3 \times 10^{11} M_\odot$, samples with mass limits of $1 \times 10^{12} M_\odot$, $4 \times 10^{12} M_\odot$ and $1 \times 10^{13} M_\odot$ have amplitude of 2pCFs enhanced by 10%, 50% and 100%, while the change in the shape is only 0.5%, 2% and 4%, respectively. The effect of galaxy bias becomes much less significant by utilizing the shape of the 2pCF rather than the amplitude.

The combined effect of galaxy bias and RSD results in larger systematics. Galaxies that are more biased predominantly reside within nonlinear, over-dense regions and thus have larger peculiar velocities. The third row of Figure 4 shows that, the difference of $\hat{\omega}_{r\perp}$ among

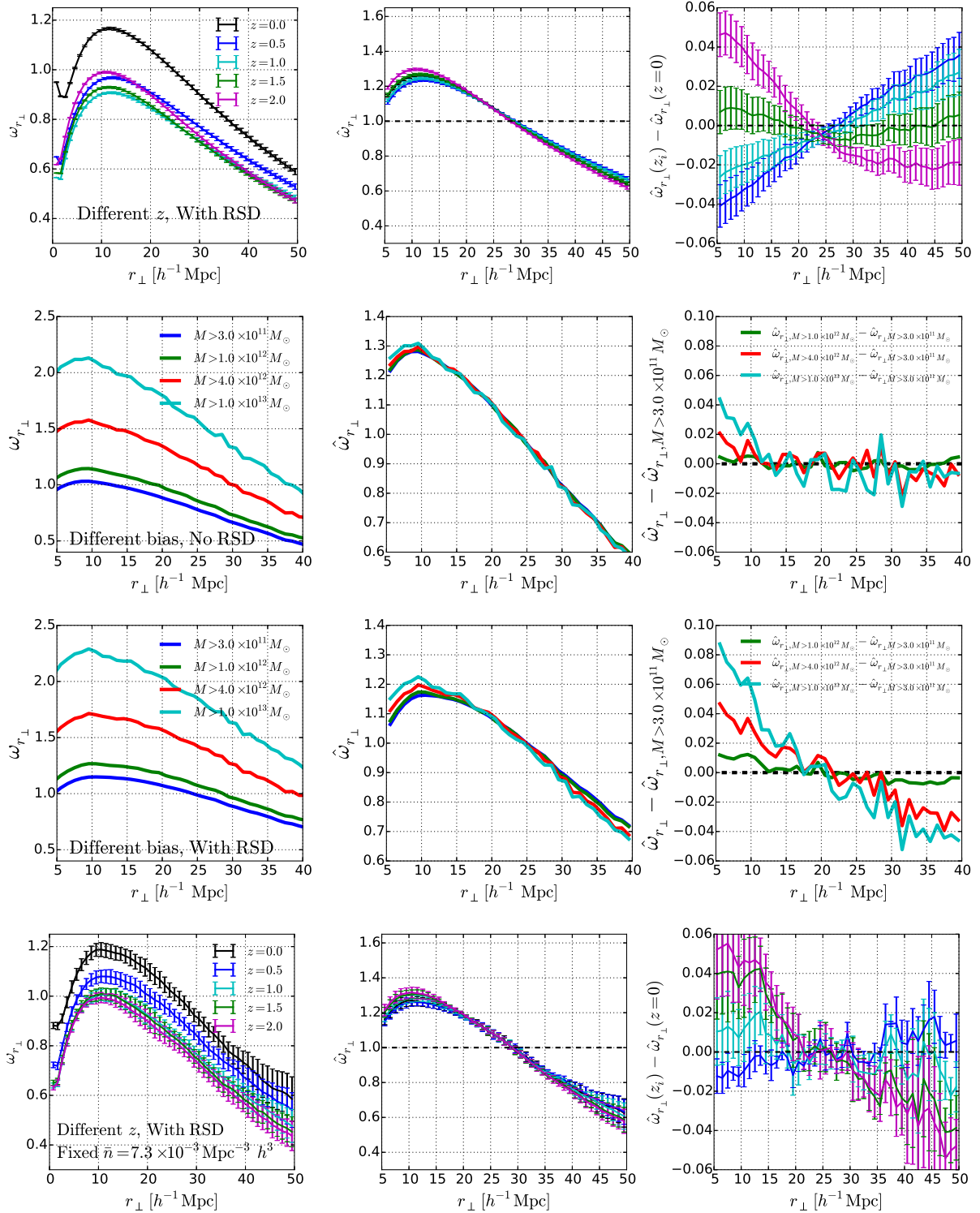


Figure 4. First row: ω_{r_\perp} (left), $\dot{\omega}_{r_\perp}$ (middle), and $\dot{\omega}_{r_\perp}(z = z_i) - \dot{\omega}_{r_\perp}(z = 0)$ (right) calculated with the RSD effect included. Second row: ω_{r_\perp} (left) and $\dot{\omega}_{r_\perp}$ (middle) are shown for four different halo-mass-cuts, 3×10^{11} , 1×10^{12} , 4×10^{12} , & $1 \times 10^{13} M_\odot$, below which we remove from 2pCF calculation. The right panel shows the difference in $\dot{\omega}_{r_\perp}$ between the $3 \times 10^{11} M_\odot$ mass-cut case and the other mass-cut cases. Third row: The same as the second row panels, except that the RSD effect is considered. Fourth row: The redshift evolution in case of using *samples with constant number density* $\bar{n} = 7.3 \times 10^{-3} \text{ Mpc}^{-3} h^3$, and the RSD effect is considered (the error bar is a bit larger since for an illustration purpose we only do the computation for 1/30 of the whole sample).

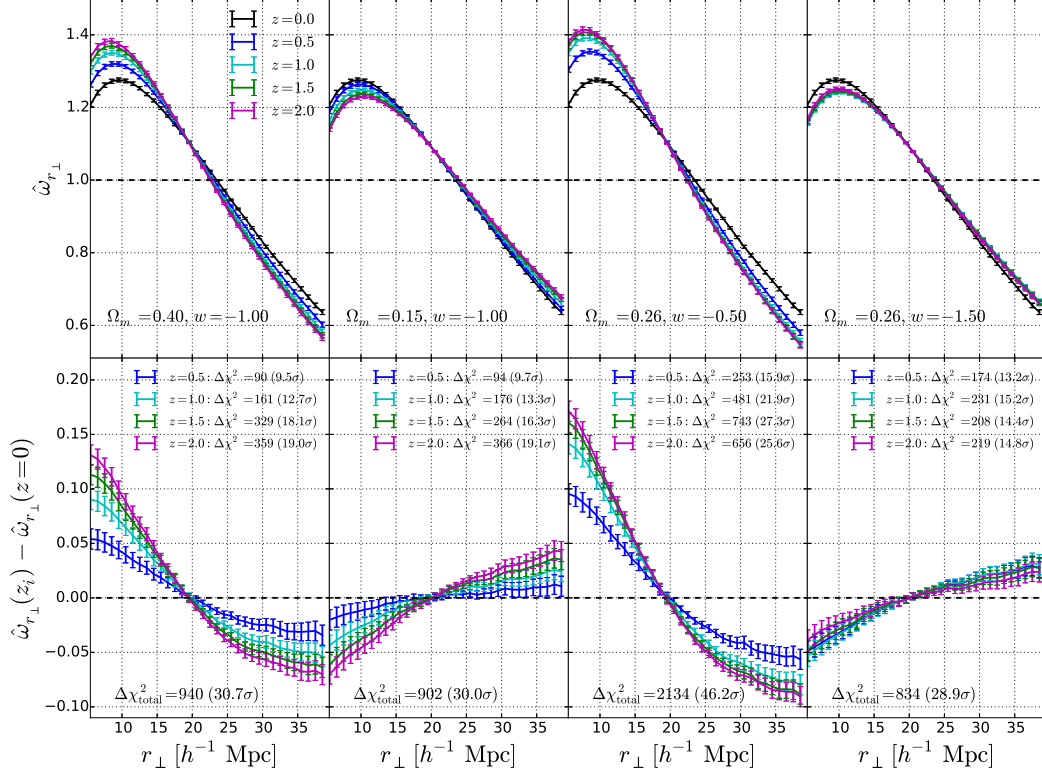


Figure 5. $\hat{\omega}_{r_\perp}$ (upper panels) and $\hat{\omega}_{r_\perp}(z=z_i) - \hat{\omega}_{r_\perp}(z=0)$ (lower panels) at several redshifts, for four wrong cosmologies. The redshift evolution of the angular 2pCF shape in these wrong cosmologies are detected at high CL.

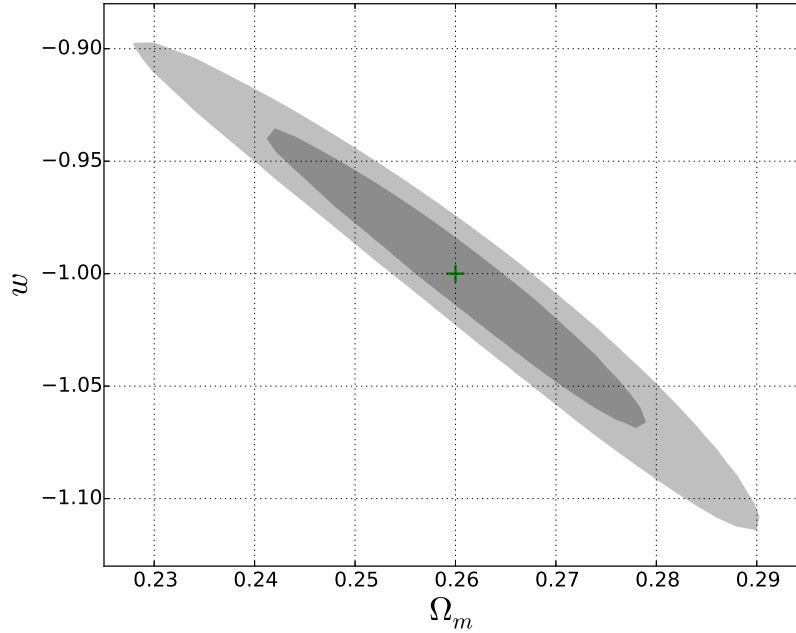


Figure 6. Likelihood contours (68.3%, 95.4%) in the $\Omega_m - w$ plane from our method, based on the redshift evolution of the angular 2pCF shape $\hat{\omega}_{r_\perp}$. Utilizing the HR4 snapshots data of mock galaxies at five redshifts, the method can lead to very tight cosmological constraint.

different mass cut samples reaches 1.5%, 4% and 8% in the case when including the RSD effect.

Finally, the fourth row of the Figure displays the redshift evolution in case of using samples with constant number density $\bar{n} = 7.3 \times 10^{-3} \text{Mpc}^{-3} h^3$, with the RSD effect considered. We see the redshift evolution of $\hat{\omega}_{r_\perp}$ is still as small as $\lesssim 5\%$, similar to the case of using constant mass cut sample.

In summary, the effect of RSD and galaxy bias on the shape of galaxy angular correlation function is, not large, but it is still non-negligible. When dealing with real observational data, one should construct simulation mocks in which these effects are reliably modelled, so that the systematics are corrected and will not lead to significant bias in the derived cosmological constraints.

4.4. Likelihood Analysis

Adopting the correct cosmological model will resulting in minimal evolution of the normalised 2pCF. Thus we proceed to construct a quantitative likelihood estimator that reflects this property. The procedure here is very similar to Li et al. (2014, 2015, 2016). We compare the high redshift $\hat{\omega}_{r_\perp}$ and the lowest redshift measurement,

$$\delta\hat{\omega}_{r_\perp}(z_i, z_1) \equiv \hat{\omega}_{r_\perp}(z_i) - \hat{\omega}_{r_\perp}(z_1), \quad (9)$$

and design the likelihood function to quantify the redshift evolution of $\hat{\omega}_{r_\perp}$:

$$\chi^2 \equiv \sum_{i=2}^{n_z} \sum_{j_1=1}^{n_r} \sum_{j_2=1}^{n_r} \mathbf{p}(z_i, r_{j_1}) (\mathbf{Cov}_i^{-1})_{j_1, j_2} \mathbf{p}(z_i, r_{j_2}), \quad (10)$$

where n_z is the number of redshifts which is 5 in this analysis, n_r is number of binning in $\hat{\omega}_{r_\perp}(r_\perp)$, which is 35 since we have r_\perp bins with width $1h^{-1}\text{Mpc}$ in a range of $5h^{-1}\text{Mpc} \leq r_\perp \leq 40h^{-1}\text{Mpc}$. $\mathbf{p}(z_i, r_j)$ is the redshift evolution of the angular correlation function shape with systematic effects subtracted

$$\mathbf{p}(z_i, r_j) \equiv \delta\hat{\omega}_\perp(z_i, z_1, r_j) - \delta\hat{\omega}_{r_\perp, \text{sys}}(z_i, z_1, r_j) \quad (11)$$

\mathbf{Cov}_i is the covariance matrix estimated from the $\mathbf{p}(z_i, r_j)$ measured from 120 subsamples. As mentioned previously, for a robust estimation of the covariance matrix, we always compare slices at different locations to include the cosmic variance.

The covariance matrix inferred from a finite number of samples is always a biased estimate of the true matrix (Hartlap et al. 2006). This can be corrected by rescaling the inverse covariance matrix as

$$\mathbf{Cov}_{ij, \text{Hartlap}}^{-1} = \frac{N_s - n_r - 2}{N_s - 1} \mathbf{Cov}_{ij}^{-1}, \quad (12)$$

where $N_s = 120$ is the number of mocks used in covariance estimation. For this analysis the rescaling is as large as 1.43. In the case that one has 2000 mocks the rescaling is less than 1.02.

Figure 5 displays the 2pCF and the likelihood values when adopting the four incorrect cosmologies used in Figure 1. For the cosmologies $\Omega_m = 0.4, w = -1$ and $\Omega_m = 0.26, w = -0.5$, the compression shifts the clustering patterns to smaller scales at higher redshift and make the 2pCFs steeper. For the cosmologies $\Omega_m = 0.15, w = -1$ and $\Omega_m = 0.26, w = -1.5$, the stretch of structure leads to a shallower slope of $\hat{\omega}_{r_\perp}$ at higher redshift. In all cosmologies there is significant detection of redshift

evolution. We compute the χ^2 values according to Equation 10, and find that these cosmologies are disfavored at $\gtrsim 30\sigma$ CL.

5. COSMOLOGICAL CONSTRAINT

We constrain Ω_m and w through a Bayesian analysis (Christensen et al. (2001); also see Lewis & Bridle (2002); Li et al. (2016)). We assume that the likelihood takes the form

$$\mathcal{L} \propto \exp \left[-\frac{\chi^2}{2} \right], \quad (13)$$

and scan the parameter space in the $\Omega_m - w$ plane to obtain the 68.3% and 95.4% CL regions. The result is displayed in Figure 6.

We obtain tight constraints on the two parameters. The 2σ contour lies within the region $0.23 < \Omega_m < 0.29$, $-1.1 < w < -0.9$. The two parameters are very strongly degenerated with each other, and the area of constrained region is quite small. The thin shape of the contour implies that, when combining with other observational data with a different direction of degeneracy (e.g. CMB, BAO), a very tight combined constraint can be obtained.

In the case of fixing one parameter at its best-fit value and constraining the other, one obtains very small 1σ uncertainties of $\delta\Omega_m \approx 0.003, \delta w \approx 0.01$.

6. CONCLUDING REMARKS

We have developed a method for constraining certain cosmological parameters by measuring the redshift evolution of the shape of the galaxy angular correlation function, $\hat{\omega}_{r_\perp}$. We found that a wrongly adopted cosmology results in a redshift-dependent mis-scaling in the constructed galaxy distribution, which in turn leads to a redshift-dependent rescaling of $\hat{\omega}_{r_\perp}(r_\perp)$. The redshift dependent effect is sensitive to cosmology while being relatively insensitive to the gravitational growth of structure, the galaxy bias, and the RSD effect. We tested our method on the HR4 mock galaxies having 457, 406, 352, 306 and 228 million galaxies at redshifts 0, 0.5, 1, 1.5, 2. Analyzing the redshift evolution of $\hat{\omega}_{r_\perp}(r_\perp)$ on scales $5h^{-1}\text{Mpc} \leq r_\perp \leq 40h^{-1}\text{Mpc}$, we derive tight constraints on Ω_m and w .

In this analysis we restrict our focus to the rescaling in the angular direction. One can also use the 2D 2pCF $\xi(s, \mu)$ as a function of both angular and radial scale to fully probe the scaling effect in the 3D galaxy distribution.

There are already works using the angular 2pCF to constrain cosmology by directly comparing the theoretical 2pCF and the observed one (Salvador et al. 2014, 2016). Our method is simpler and just uses the fact that in the correct cosmology the angular 2pCF shape does not exhibit significant redshift evolution. Our method is complementary to these works in that it does not require accurate analytic modeling, and could be applicable on smaller scales.

Li et al. (2015, 2016) shows that using the redshift dependence of the AP effect, one can derive very tight cosmological constraints. The method presented in this paper studies the redshift evolution of the volume effect. These two methods probe different information encoded in the 2pCF and can be combined to give a complete geometric probe of the LSS galaxy distribution.

This method is designed to be applied to current and future LSS surveys, with a particular emphasis on *photometric surveys*. The galaxy angular 2pCF in redshift

shells is much less affected by the large redshift errors compared with the 3D 2pCF. DES will probe hundreds of millions of galaxies at $z \lesssim 1.5$, and LSST will probe tens of billions of galaxies at $z \lesssim 2$. We expect very tight cosmological constraints from these surveys.

Our series of works Li et al. (2014, 2015, 2016), together with this analysis, and also Park & Kim (2010); Morandi & Sun (2016), are introducing a new strategy to constrain cosmology from galaxy clustering. We utilize a statistical pattern which remains constant at all redshifts; this invariance is broken when incorrect cosmologies are adopted to construct the galaxy distribution and introduces redshift evolution of geometric distortion. The systematic effects can be corrected by performing a N-body simulation in a cosmology with values of parameters close enough to the underlying correct one. This avoids the difficulty in the analytical modeling the galaxy

clustering, galaxy bias and RSD, and enables reliable cosmological constraint on relatively small clustering scales. We hope these methods play an important role in deriving cosmological constraints from future LSS surveys.

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APPENDIX

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