

Searching 21-cm Absorption Systems in Chinese Radio Telescopes

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Abstract Neutral hydrogen clouds are known to exist in the Universe, however their spatial distributions and physical properties are poorly understood. Such missing information can be studied by the Chinese new generation radio telescopes by a blind searching of 21-cm absorption systems. We forecast the abilities of surveys of 21-cm absorption systems by two representative radio telescopes in China – Five-hundred-meter Aperture Spherical radio Telescope (FAST) and Tianlai. The result shows that, in a few of years term, these telescopes with either high sensitivity (FAST) or wide field of view (Tianlai) can discover orders of magnitudes more 21-cm absorption systems, than the cumulative discoveries in the past 50 years.

Key words:

1 INTRODUCTION

Neutral hydrogen (HI) clouds are known to exist in the Universe, however only few of them are discovered in the past half a century, and we poorly understand their spatial distribution and physical properties (Wolfe et al. 2005). In damped Lyman- α absorption (DLA) systems, the radio spectrum is also substantially absorbed by the HI hyperfine structure, whose rest frame wavelength is approximately 21 cm (~ 1420 MHz in frequency). These systems are defined to have at least $2 \times 10^{20} \text{ cm}^{-2}$ HI column density, thus being able to house cold 21-cm absorbing gas in the cold neutral medium (CNM).

The 21-cm absorption systems are important in the study of the distribution, location, temperature and structure of neutral gas, and the evolution of neutral gas systems and galaxies over cosmic time scale. Due to the narrow intrinsic line width, the 21-cm absorption systems are proposed to be used to directly measure the cosmic acceleration (Darling 2012; Yu et al. 2014), via the Sandage-Loeb (SL) effect (Sandage 1962; Loeb 1998). The uncertainties are from the poorly understood neutral hydrogen clouds. Path finder surveys, taking less than a year, are needed to improve our understandings of spatial

distribution and physical properties of neutral hydrogen clouds. This can be done by the Chinese new generation radio telescopes – Five-hundred-meter Aperture Spherical radio Telescope (FAST) (Li & Pan 2016) and Tianlai (Chen 2012). In the current and near future, these are the two symbolic modern radio telescopes in China. The single-dish FAST, like Arecibo¹, uses its gigantic single dish to achieve ultimate sensitivity, whereas Tianlai, designed like CHIME², has ultra wide field of view (FoV) per day, by using its cylindrical reflectors and arrays of receivers to scan the northern hemisphere of the sky as the Earth turns.

According to their different design and observing strategies, we forecast their abilities of blind searching of 21-cm absorption systems. In section 2 we show the sensitivity estimation and all the aspects needed to be taken into account, and according to the configurations of FAST and Tianlai, we show their forecast in section 3 and section 4 respectively. Conclusions are made in section 5.

2 SENSITIVITY

To find possible 21-cm absorption systems, radio telescopes could be devised to scan the radio sources in NRAO VLA Sky Survey (NVSS)³. If we assume the distributions of radio sources n_R and HI clouds n_{HI} are uniform over the sky, then the redshift distribution of the 21-cm absorption systems n_{system} would be

$$n_{\text{system}}(z) = n_{\text{HI}}(z) \int_z^\infty n_R(z') dz' \quad (1)$$

regardless of the observation sensitivity, where $n_R(z)$ is given by equation (26) of de Zotti et al. (2010). Spacial distribution of HI is poorly understood and is to be studied in upcoming surveys. Recent studies show that the number density of absorbers to be $dn_{\text{HI}}/dz \sim 0.045$ per line of sight (Wolfe et al. 2005; Zwaan et al. 2007), so we apply this value in the following forecast.

Observation sensitivity depends on the background fluxes and foreground absorptions (current section), as well as telescope configurations and observation strategies (section 3,4).

NVSS contains 2 million sources (with declination $\delta > -40^\circ$, cover 82% of the sky) stronger than 2.5 mJy at $\nu = 1.4$ GHz (Condon et al. 1998) and their flux distribution $n_R(F)$ is given by Condon (1984). Redshifted HI clouds absorb lower frequency bands where the radio sources are typically brighter by $\nu^{-0.7}$ (Condon et al. 1998), or $(1 + z_{\text{HI}})^{0.7}$. The error of the measurement ΔF is given by

$$\Delta F = \text{SEFD} / \sqrt{n_{\text{pol}} \Delta \nu \Delta t}, \quad (2)$$

where SEFD is the system equivalent flux density, n_{pol} is the number of polarizations, $\Delta \nu$ is the line width and Δt is the integration time. $\Delta \nu = (u_{\text{width}}/c)\nu_{\text{obs}}$, where u_{width} is the equivalent line width of the HI absorber. The properties of 21-cm absorbers are primarily derived from followup studies of optical absorbers. The discovery rate from the cross correlation is a lower bound on the expect number of absorbers, since high column density systems in the CNM may systematically obscure potential background optical sources (Yu et al. 2014). There are only 3 blind radio detections, and a survey may discover more systems which are optically obscured. They would likely be cold and at high column density. For sensitivity purposes, we treat all sources as $u_{\text{width}} = 2 \text{ km s}^{-1}$ (Wolfe et al. 1982, 2005). ν_{obs} is the observation frequency. For redshifted 21-cm absorption systems, $\nu_{\text{obs}} = 1420 \text{ MHz}/(1 + z_{\text{HI}})$. The integration time $\Delta t = (n_{\text{obs}} \times 24 \text{ h})\tau$, where n_{obs} is the number of days an object is scanned, and $\tau = \lambda_{\text{obs}}/2\pi D \cos \delta$ is the fractional time the object transits the FoV ($\lambda_{\text{obs}} \ll D$ or $\delta \rightarrow \pi/2$). For 21-cm absorption systems $\lambda_{\text{obs}} = 21 \text{ cm} \times (1 + z_{\text{obs}})$. ΔF is redshift- (frequency-) independent, as $(1 + z_{\text{obs}})$ terms from integration time and line width cancel out in the equation (2).

In the forecast we count $> 10\sigma$ detections, i.e., the absorption line depth is at least $10\Delta F$. The optical depth of HI absorbers are also poorly understood. Recent discoveries of 21-cm absorption systems

¹ <https://www.naic.edu>

² <http://chime.phas.ubc.ca>

³ <http://www.cv.nrao.edu/nvss/>

have about $r = 20\%$ fractional depth (Allison et al. 2015; Zwaan et al. 2015), and we apply this value. Thus, all systems with $rF > 10\Delta F$ are counted.

3 FAST ESTIMATION

FAST has a fixed 500 m dish, and a circle of 300 m diameter is held in the correct parabolic shape, being used at any one time. This enables its zenith angle to be 40° and it can scan objects in declination range $-14^\circ 12' < \delta < 65^\circ 48'$. The 19-beam feed-horn array at L band will be the primary survey instrument for FAST. For a single day it can scan 0.762° in declination δ . FAST survey strategy scans each object only once ($n_{\text{obs}} = 1$), and with its high sensitivity, even faint sources are identified in a single day. We assume a one-month survey by FAST around the celestial equator for simplicity. FAST observes 1.2 to 1.4 GHz (redshift $z_{\text{HI}} < 0.2$), so objects on celestial equator have average transit time $\Delta t \simeq 10$ sec. Taking $\Delta\nu \simeq 9$ kHz, $n_{\text{pol}} = 2$, and SEFD = 1 Jy for FAST. we have $\Delta F \simeq 2.3$ mJy. Taking into account the $\nu^{-0.7}$ flux boost at lower frequencies, $rF > 10\Delta F$ requires $F \gtrsim 0.11$ Jy on 1.4 GHz.

One-month-scan ($n_{\text{obs}} = 30$) around celestial equator by FAST will have $N_R \simeq 10^4$ sources whose $F \gtrsim 0.11$ Jy on 1.4 GHz. We further assume the independency between $n_R(F)$, $n_R(z)$ and $n_{\text{HI}}(z)$, from $n_{\text{system}}(z)$ by integrating equation (1), there will be $N_{\text{system}} \simeq 90/\text{month}$ absorption systems found in redshift range $0 < z < 0.2$.

4 TIANLAI ESTIMATION

Tianlai is currently 30×12 m area with SEFD $\simeq 300$ Jy. Different from FAST, however, Tianlai scans all sources in the northern hemisphere of the sky everyday, and one needs longer periods integration to identify absorbers with fainter background sources. We forecast its one-year-survey ability of searching 21-cm absorption systems in the northern hemisphere.

Because Tianlai locates at latitude $\phi_{\text{site}} = +44^\circ$ and its cylinders are fixed, only objects at the zenith ($\delta = \phi_{\text{site}} = 44^\circ$) fully utilize the SEFD. On the other hand, higher declination objects have longer integration time per day. Thus, ΔF is a function of δ ,

$$\Delta F(\delta) = \frac{\text{SEFD}}{\cos(\delta - \phi_{\text{site}})} \sqrt{\frac{\cos \delta}{n_{\text{pol}} \Delta\nu \Delta t(\delta = 0)}}, \quad (3)$$

where $\Delta t(\delta = 0) \simeq 1.17 \times 10^5$ sec is the total integration time (1-year survey) for objects on celestial equator, for Tianlai frequency range 800 to 900 MHz ($0.58 < z < 0.78$). The effective N_R is given by

$$N_R = \int_0^{\pi/2} 2\pi \cos \delta \left(\int_{10\Delta F(\delta)/r}^{+\infty} n_R(F') dF' \right) d\delta. \quad (4)$$

Note that the inner flux integration converges as $F' \rightarrow +\infty$ because $n_R(F) \sim F^{-2.5}$. Additionally, even if $\tau = \lambda_{\text{obs}}/2\pi D \cos \delta$ fails at $\delta \rightarrow \pi/2$ (the north pole $\tau(\delta = \pi/2) = 1$ is always in FoV), the outer declination integration also converges as $\delta \rightarrow \pi/2$, because the pole area is tiny and neglectable. Equation (4) gives $N_R \simeq 1.3 \times 10^4$, and applying it to equation (1) again we get $N_{\text{system}} \simeq 80/\text{year}$ absorption systems found in redshift range $0.58 < z < 0.78$.

5 CONCLUSION

Acknowledgements HRY acknowledges General Financial Grant No.2015M570884 and Special Financial Grant No. 2016T90009 from the China Postdoctoral Science Foundation. HRY and ULP acknowledge the support of the National Science and Engineering Research Council of Canada.

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