

Manual of plotastrodata

Yusuke Aso

1 Read Data

The data class `AstroData` can take a fits file.

```
from plotastrodata.analysis_utils import AstroData  
  
d = AstroData(fitsimage='file_name.fits')
```

`AstroData` can take more arguments, such as `Tb` and `sigma`. `Tb=True` means the data values will be converted from flux densities in the unit of Jy beam^{-1} to brightness temperatures in the unit of K. `sigma` specifies how to measure the noise level of the data values: for example, the default option of 'hist' means to use the histogram of the data values.

The data class `AstroFrame` is necessary to form the `AstroData` instance in a useful format. `AstroFrame` can take quantities related to coordinate ranges.

```
from plotastrodata.analysis_utils import AstroFrame  
  
f = AstroFrame(vmin=-5.0, vmax=5.0, vsys=2.0,  
               center='00h00m00s 00d00m00s', rmax=3.0)
```

`vmin`, `vmax`, and `vsys` are in the unit of km s^{-1} . `rmax` is in the unit of arcsec. Instead of `rmax`, other arguments (`xmin`, `xmax`, `ymin`, `ymax`, `xoff`, `yoff`) can be used to adjust the x and y ranges in more detail. When an argument of `fitsimage` is given, the central coordinates `center` is read from the given fits file; the rest frequency is also read from the fits file, which is used for the frequency-velocity conversion. Moreover, an argument of `dist` can specify the distance used to change the unit of spatial coordinates from arcsec to au. An argument of `swapxy` is used to swap the x and y coordinates. An argument of `pv` must be `True` when the `AstroData` instance to be read is a position-velocity (PV) diagram, i.e., the first and second axes are the spatial and velocity coordinates. When `quadrants=True`, the first and third (or second and fourth) quadrants of the PV diagram will be averaged. `xflip` and `yflip` will be used to determine the plotting direction of the x and y axes, respectively; these have a meaning only when a figure is made from the data set.

Forming the `AstroData` instance needs the following command.

```
f.read(d)
```

After this command, the `AstroData` instance has useful attributes in the format of numpy array: a 1D array of `d.x` (as well as `d.y` and `d.v`), a 2D or 3D array of `d.data`, a 1D array of `d.beam`, and a float of `d.sigma`. `d.x` and `d.y` are the relative coordinates in the unit of arcsec from the given `center`. Similarly, `d.v` is the relative coordinate in the unit of km s^{-1} from the given `vsys`. `d.beam` is the beam components `array([bmaj, bmin, bpa])`, where `bmaj` and `bmin` are in the unit

of arcsec, while **bpa** is in the unit of degree. When **Tb=True** above, the data values are converted to the brightness temperature and stored as **d.data**. **d.sigma** is the noise level measured in the way specified above in the same unit as **d.data**. AstroData can take two more arguments, **restfreq** and **cfactor**. **restfreq** is used to explicitly set the rest frequency in the unit of Hz, which is used to convert frequency to velocity and Jy beam⁻¹ to K. When this argument is zero (**restfreq=0**), the frequency axis is not converted to velocity. When the unit of the frequency/velocity axis in the input fits header is m/s, **f.read(d)** divides the coordinate by 1000. The argument **cfactor** specifies a constant factor that **d.data** is multiplied by. This will be useful when one wants to change the intensity unit from Jy beam⁻¹ to mJy beam⁻¹.

In addition to the input attributes, **f.read(d)** adds three attributes to the AstroData instance **d**: **fitsimage_org**, **sigma_org**, and **fitsheader**. **fitsimage_org** saves the input **fitsimage**, while **fitsimage** is updated to None after **f.read(d)**. Similarly, **sigma_org** saves the input **sigma**, which may be a string, while **sigma** is updated to the calculated value. **fitsheader** is the fits header in the format of `astropy.io.fits.open(fitsimage)[0]`.

2 Analyze Data

AstroData also has handy methods to analyze the 2D/3D data.

The **binning** method rebins the 2D/3D data with a given width for each coordinate.

```
d.binning(width=[5, 4, 2])
```

This command takes an average over 5 channels in the velocity direction, 4 pixels in the y direction, and 2 pixels in the x direction. The number of channels and pixels are decreased accordingly after this method. If the length of **width** is 2, the two values are regarded as the widths for the y and x directions.

The **centering** method sets the coordinates so that the spatial center and the systemic velocity have the exact zero coordinates by interpolation.

```
d.centering(includexy=True, includev=False)
```

This command adjusts the x and y coordinates but does not adjust the velocity coordinate. This method will be useful when one wants to quickly get a radial profile along a line passing the center or a PV diagram (using the **rotate** method together).

The **circularbeam** method makes the beam shape circular by additional 2D Gaussian convolution. This method takes no argument.

```
d.circularbeam()
```

The new beam has the major and minor axes same as the old major axis. This method also updates the attribute of **d.beam** accordingly.

The **deproject** method deprojects the 2D/3D data with a given position angle (P.A.) and inclination angle.

```
d.deproject(pa=45, incl=45)
```

pa is the position angle from the north to the east in the unit of degree. **incl** is the inclination angle; **incl=0** means the face-on configuration and thus no deprojection. This command replaces **d.data** and **d.beam** with the deprojected data and the deprojected beam, respectively.

The `histogram` method returns the bins and the histogram in the bins of the attribute of `d.data` by using `numpy.histogram()`.

```
hbin, hist = d.histogram(bins=10)
```

The arguments are the same as those for `numpy.histogram()`. The bins (`hbin`) have the same length as the histogram (`hist`), which is different from the original `numpy.histogram`.

The `gaussfit2d` method performs the 2D Gaussian fitting to a 2D `d.data` or a channel of a 3D `d.data`.

```
res = d.gaussfit2d(chan=12)
```

This command performs the fitting to the channel of 12 in `d.data`. For 2D data, the `chan` argument can be omitted. The output (`res` here) is a dictionary having keys of `popt`, `pcov`, `model`, and `residual`. `popt` and `pcov` are the optimized parameters (peak intensity, central x, central y, major FWHM, minor FWHM, and P.A.) and their covariance. The `model` and `residual` are 2D arrays with the same shape of the fitted 2D array.

The `mask` method puts `numpy.nan` on pixels that satisfies a given condition.

```
d.mask(dataformask=d2.data, includepix=[1e-3, 100], excludepix=[50, 200])
```

This command puts `numpy.nan` on pixels of `d.data` where `d2.data` is outside `[1e-3, 100]` or inside `[50, 200]`, where `d2` is an `AstroData` instance different from `d`. This method will be useful when one wants to put a mask on a moment 1 map using the values of a moment 0 map.

The `profile` method makes line profiles at given spatial positions.

```
v, f, g = d.profile(xlist=[-0.5, 0.2], ylist=[1, 0.8], ellipse=[0.3, 0.2, 45])
```

This command makes line profiles at $(x,y)=(-0.5, 1)$ and $(0.2, 0.8)$ in the unit of arcsec. The intensity of each profile is an average over a boxcar ellipse with the major axis of 0.3 arcsec, the minor axis of 0.2 arcsec, and the P.A. of 45 degrees. Instead of `xlist` and `ylist`, coordinate strings can be input using an argument of `coords`: `coords=['00h00m00.0s 00d00m00.0s', '11h11m11.1s 11d11m11.1s']`. When `ellipse` is omitted, the profile is made by picking up the values at the closest pixel to the given position. When the ellipse size is not so large compared to the pixel size, the integer argument `ninterp` may be useful; this makes the pixel size `ninterp` times finer by interpolation. When the boolean argument `flux` is True, the output profile has the unit of Jy. Additionally, when the boolean argument `gaussfit` is True, the profiles will be fitted with a 1D Gaussian function. The output `v` above is `d.v`. The output `f` above is a list of 1D-array profiles, i.e., 2D array. The output `g` is a list of dictionaries; each dictionary has keys of `best` and `error` (square root of the diagonal components of the covariance).

The `rotate` method rotates the attribute `d.data` by the given angle in the unit of degrees by interpolation.

```
d.rotate(pa=45)
```

When a pixel refers to a position outside the original image, this pixel has `numpy.nan` after this method.

The `slice` method makes a radial profile for a 2D image or a PV diagram for a 3D image along a given direction and length by interpolation.

```
r, f = d.slice(length=3, pa=45, dx=0.2)
```

This commands makes a radial profile (or PV diagram) along a cut with a length of 3 arcsec at P.A.=45 degrees with a separation of 0.2 arcsec. The output **r** and **f** are both 1D arrays of the positional offsets and the intensity at the positions. When the separation **dx** is omitted, the absolute value of the x pixel size is adopted.

The **todict** method returns the attributes as a dictionary.

```
a = d.todict()
```

This output includes keys of **data**, **x**, **y**, **v**, **fitsimage**, **beam**, **Tb**, **restfreq**, **cfactor**, **sigma**, and **center**. This dictionary can be input to methods of **PlotAstroData** as ****a**.

The **writetofits** method exports the instance as a fits file.

```
d.writetofits(fitsimage='new_file_name.fits')
```

The output fits file reuses the header components of the fits file used to make the **AstroData** instance ('**file_name.fits**' above); some header components are updated properly after the above-mentioned methods for analysis, such as **CDEL1**.

3 Plot Data

The class **PlotAstroData** can take an **AstroData** instance through the method of **d.todict()**.

```
from plotastrodata.plot_utils import PlotAstroData

p = PlotAstroData(rmax=3.0)
p.add_color(**d.todict())
p.add_scalebar(length=50 / 140, label='50 au')
p.set_axis()
p.savefig('figure_name.png')
```

These commands make a color map using the **AstroData** instance **d**. **PlotAstroData** can take the same arguments as **AstroFrame** to define the plotting ranges; particularly **rmax** is necessary. The method **p.add_color()** can take a fits file directly instead of the **AstroData** instance, as **p.add_color(fitsimage='file_name.fits')**. The command **p.add_scalebar()** can be omitted if the map does not need to show a scale bar. The command **p.set_axis()** (or **p.set_axis_radec()**) is necessary even without any argument. More detailed usage can be found in the **example.py** file and <https://plotastrodata.readthedocs.io/en/latest/#>.