

## COSMOLOGY L III PART I

CMB discovery: 1965

Killed steady state proposals

proved hot big bang cosmology beyond any reasonable doubt

Penzias & Wilson

much bigger than expected noise  
isotropic  $\Rightarrow$  cosmic origin

Another group predicted  $\approx 3\text{K}$  background radiation.  
They couldn't measure the spectrum properly  
in 1965.

COBE satellite 1990: measured spectrum.

Found very nice black body radiation curve.  
an estimate temperature from this

$$T \approx 2.7\text{K}$$

Black body radiation:

$$S_{\gamma} c^2 = \left[ \frac{\text{TOTAL}}{\text{PER SURF}} \right] \int_0^\infty \frac{x^3 dx}{e^x - 1} \quad \boxed{\Gamma(4) \zeta(4) = \frac{\pi^4}{15}}$$

Riemann  
Zeta  
function

early times were in thermal equilibrium then

$$\boxed{S_{\gamma} c^2 = a T_g^4} \quad a = \frac{8\pi^5 g_*^4}{15 h^3 c^3}$$

black body radiation exerts isotropic pressure:

$$P_g = \frac{1}{3} S_{\gamma} c^2$$

$$\text{Energy conservation: } \frac{d(SR^3)}{dt} = -3PR^2$$

$$\Rightarrow S_\gamma \propto 1/R^4 \quad S_\gamma C^2 = aT_\gamma^4$$

note that for dust,  
 $S \propto 1/t^3$ , unlike  
radiation which is  
considered here

$$T_\gamma \propto 1/R$$

$$S_\gamma(\nu, t) d\nu = \frac{1}{C^3} \frac{8\pi h\nu^3 d\nu}{\exp\left(\frac{h\nu}{kT_\gamma}\right) - 1}$$

$$T \propto 1/R, \nu \propto 1/R \text{ (redshift)} \Rightarrow$$

radiation  
maintains  
b. b. form  
as univ.  
expands.

number density of BG photons

$$n_\gamma(\nu) d\nu = \frac{8\pi}{C^3} \frac{\nu^2 d\nu}{\exp\left(\frac{h\nu}{kT_\gamma}\right) - 1}$$

integrate up:

$$N_\gamma = 16\pi S(3) \left(\frac{kT_\gamma}{hc}\right)^3 = 410.6 \left(\frac{T_\gamma}{2.726 \text{ K}}\right)^3 \text{ cm}^{-3}$$

critical density is  $\sim 10^8 \text{ /m}^3$

$400^{10^8}$  photons per  $\text{cm}^3$

$$\eta = \frac{n_b}{n_\gamma} = 2.7 \times 10^{-8} S T_\gamma h^2$$

need many more photons  
than baryons. The reason of this  
is not yet well understood.

must have been an asymmetry in matter-antimatter, so after  $\pi^+ \pi^-$  annihilation we have left with  ~~$\pi^+$~~   $\pi^+$  a lot.

$$S_\pi(0) - \frac{S_\pi}{S_C} = 2 \cdot 9 \cdot 10^{-5} \text{ cm}^{-2}$$

↳ bubble parameter

CMB plays negligible control to Edensity of univ. present day.

$$S_{\text{matter}} \propto T^{-3} \quad \& \quad S_\gamma \propto T^{-4}$$

$\Rightarrow$  Edensity of radiation dominates early times.

Furthermore:  $T_\gamma \propto T^{-1} \Rightarrow$  early uni must have been hot & very dense.

~~inflation~~  
i.e. CMB proves hot big bang cosmology

- ~~After~~ Times after CMB discovery
  - is it BB  $\rightarrow$  COBE
  - is it anisotropic?  
i.e. different directions do have different temperatures or not

Discovery of Dipole Anisotropy

## Dipole Anisotropy

if  $v \ll c$ :

$$T'(\theta') \approx T_0 [1 + \frac{v}{c} \cos \theta']$$

i.e. motion of observer through the CMB induces dipole

Put radiometer to U2 spy plane, fly over Shire. anisotropy.

This was measured by Planck satellite as well.

Solar dipole, correct for motion of sun,  
motion compared to local group centre.

Anisotropy still present.

Questions arising:

(CMB defines preferred reference frame. SR violation?  
no, preffered ref frame in FRL metric. (?)

What causes this motion?

inhomogeneity of matter around us

# Cosmology L III P II

de Sitter space

recall FRW: max. by sym space

characterized by 1 number  $k$ .

3D

de Sitter space:

max. symmetric 4D spacetime.

curvature tensor

$$R_{K250} = k(g_{20}g_{K5} - g_{52}g_{05})$$

Ricci tensor

$$R_{20} = g^{05} R_{K250} = k(N-1)g_{20}$$

Ricci scalar

$$\tau = kN(N-1)$$

$N=4$  for 4D.

Rewrite Einstein field equation:

$$R_{\mu\nu} - \lambda g_{\mu\nu} = -8\pi G(T_{\mu\nu} - \frac{1}{2}g_{\mu\nu}T^0_0)$$

for empty universe:

$$R_{\mu\nu} = \lambda g_{\mu\nu}$$

de Sitter geometry satisfies:  $\Rightarrow$

$$R_{\mu\nu} = 3k g_{\mu\nu}$$

de Sitter geometry is  
sol of EFE if  
there is no matter

$$T_{\mu\nu} = 0 \& \Lambda \neq 0$$

$$\Rightarrow k = \frac{1}{3}\lambda \quad \text{i.e. curvature of spacetime is related to cosm. const.}$$

Geometry of de Sitter space?

consider constant curvature hyperboloid embedded in 2+1 dimensional Minkowski space.

$$-z^2 + x^2 + y^2 = H^{-2} \quad (\text{eq of hyperboloid})$$

similar to FRW: here we embedded sphere

metric:

$$ds^2 = dz^2 - dx^2 - dy^2$$

eliminate  $dz$ :

$$ds^2 = \frac{(x dx + y dy)^2}{x^2 + y^2 - H^2} - dx^2 - dy^2$$

make coordinate transform:

$$x = H^{-1} \cos \theta (\sin \epsilon) \begin{cases} \cos \alpha \\ \sin \alpha \end{cases}$$

$$y =$$

metric reduces to:

$$ds^2 = d\theta^2 - \frac{1}{H^2} \cos^2(\theta \epsilon) dx^2$$

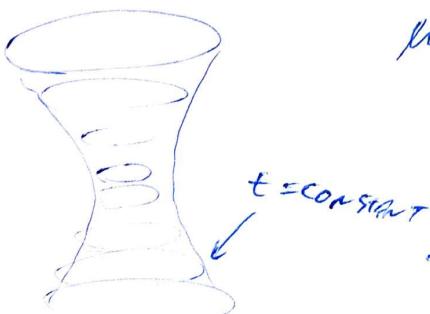
recognize form as FRW with 1 spatial coordinate

corresponds to closed universe

entire surface covered by:

$$-\pi < \theta < \pi,$$

$$0 \leq x \leq 2\pi$$



This universe has minimum size reached at  $t=0$ .

lets make another coord. transform.

[ALGEBRA]

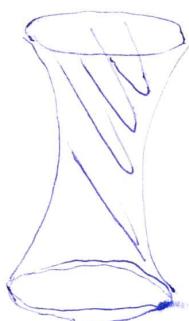
arrive to:

$$\begin{aligned}z &= z(t, \xi) \\y &= y(t, \xi) \\x &= x(t, \xi)\end{aligned}$$

$$ds^2 = dt^2 - H^{-2} e^{(2Ht)} d\xi^2$$

spatially flat universe  
expanding exponentially

spatially flat, just a line



"time-sliced" hyperboloid in a different way

Now we are equipped to  
do embedding in 5D Minkowski space.

loads of algebra.

FRW metric with closed spatial sections:

$$ds^2 = dt^2 - H^2 \cosh^2(Ht) [dx^2 + \sin^2(d\theta^2 + \sin^2 d\phi^2)]$$

analogue to contracting-expanding universe.

closed spatial sections

can choose

time coord.,

change time

$$\text{slicing: } ds^2 = dt^2 - R(t)^2 [dr^2 + r^2 d\theta^2 + r^2 \sin^2 d\phi^2]$$

analogous to top of this page.

$$R \propto \exp(Ht) \quad H = \sqrt{\frac{1}{3}} \quad k = H^2$$

KEY POINTS

if we can timeslice resulting in

closed open flat sections

what does it mean to say that our Universe is approaching the de Sitter solution with flat spatial geometries?

Keep in mind:

de Sitter sol is a vacuum solution

$\Rightarrow$  no unique way of timeslicing  
(why?)  
(existence of matter)

fixes timeslicing -

no arbitrary choice,

matter distortions defines natural timeslicing

de Sitter sol is a max. by symmetric spacetime (ie 4D)  
with curvature  $K$ .

Think of it as a thing in 4D.

Because de Sitter is important for early & late time universe.