

MEASURING SOLAR DIFFERENTIAL ROTATION WITH AN ITERATIVE PHASE CORRELATION METHOD

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ABSTRACT

A reliable inference of the differential rotation rate of the solar photosphere is essential for models of the solar interior. The work presented here is based on a novel iterative phase correlation technique, which relies on the measurement of the local shift, at central meridian, between two images separated by a given time interval. Consequently, it does not require any specific reference features, such as sunspots or supergranules, nor extended observations spanning several months. The reliability of the method is demonstrated by applying it to high spatial and temporal resolution continuum images of the solar photosphere, at 1700 Å, acquired by the Solar Dynamics Observatory (SDO) Helioseismic and Magnetic Imager (HMI) over one complete Carrington rotation. The data selected covers the time period of 2020 January 1 to February 2. The method was applied to one day, and to the full time interval. The differential rotation rate derived using this feature-independent technique yields values that fall in the middle of the range of those published to date. Most importantly, the method is suited for the production of detailed rotation maps of the solar photosphere. It also enables the visual and quantitative identification of the North-South asymmetry in the solar differential rotation rate, when present.

Keywords: solar rotation - differential rotation - image registration - phase correlation - subpixel accuracy

1. INTRODUCTION

The identification of sunspots with features on the solar surface led to the discovery of solar rotation in the 17th century, and the variation of the rotation speed with latitude, or the differential rotation rate, by Scheiner (1630). Advances in observational tools over the centuries, led to the discovery of the oscillation modes of the Sun (Leighton et al. 1962), which established the field of helioseismology (see review by Leibacher et al. 1985). Helioseismology further revealed that the differential rotation rate changes with depth in the convection zone (Thompson et al. 1996). Given that the differential rotation rate is now established as a necessary mechanism for generating solar magnetic fields, i.e. the solar dynamo (Elsasser 1950; Babcock & Babcock 1955; Parker 1955), the accurate determination of the differential rotation rate, starting from the solar surface, becomes even more compelling.

Recently Beck (2000) compared the differential rotation rates derived from the most widely used techniques, namely the Doppler shifts of photospheric spectral lines, the Doppler feature tracking, the magnetic feature tracking and the p-mode splittings. He found a

large disparity in the values inferred from these different techniques. The difficulties hampering the accuracy of the measurement are directly linked to the complexity of observed features on the solar surface itself. Additional limitations due to foreshortening effects close to the poles limit the viability of these techniques to low latitudes. Sunspot tracking adds to the limitation, since they are rarely observed above 30° latitude. Furthermore, since all tracked features, such as sunspots, supergranules and granules, for example, are subject to intrinsic temporal variability, they will inevitably bias any measurement of rotation rate.

The goal of this paper is to show how the iterative phase correlation technique, recently developed by Hrazdíra et al. (2020) can be applied to continuum images of the photosphere to derive a reliable, feature-independent, solar differential rotation rate. This method belongs to the class of tracer measurement methods. In the past, the cross-correlation method applied to images of the solar photosphere had already been used for that purpose (Beck 2000). The phase correlation technique was first introduced by Kuglin (1975) for the alignment of displaced images taken with different imaging devices or under varying conditions

of illumination. It has been successfully implemented by Druckmüller (2009) for the alignment of total solar eclipse images taken with different exposure times and different optics, and has rendered images with unprecedented detail (see, e.g. Druckmüller et al. 2014).

The best data available at present for the implementation of this technique are the 1700 Å data from the Helioseismic and Magnetic Imager (HMI) instrument on the Solar Dynamics Observatory (SDO) (Lemen et al. 2011), which represent emission from the photosphere (see Lemen et al. 2011, Table 1). The technique is described in Section 2. The results emerging from the application of this technique to one Carrington Rotation of SDO/HMI data are presented in Section 3, where the differential rotation rate thus derived is compared with inferences from a number of different techniques. Also demonstrated in that section is how the technique recovers the now well-established North-South asymmetry of the solar differential rotation rate. The results are summarized in Section 4 together with some concluding remarks.

2. MEASUREMENT TECHNIQUE

2.1. Basics of the technique

In this work, the inference of the solar rotational speed at different latitudes is achieved by measuring the local horizontal and vertical image shifts Δx , Δy between pairs of *Solar Dynamics Observatory* (SDO) *Helioseismic and Magnetic Imager* (HMI) continuum full disk 4096x4096 images by means of image registration (see Fig. 1). Measured image shifts are then projected onto the solar surface, allowing the corresponding angular shifts $\Delta\varphi_x$, $\Delta\varphi_y$ to be found. The average angular velocities ω_x and ω_y are then calculated based on the time difference Δt between the two images following the relations:

$$\omega_x = \frac{\Delta\varphi_x}{\Delta t}, \quad (1)$$

$$\omega_y = \frac{\Delta\varphi_y}{\Delta t}, \quad (2)$$

where the time difference Δt is determined by the SDO HMI continuum instrument temporal cadence Δ_t and the difference between the image indices Δ_p , namely

$$\Delta t = \Delta_p \Delta_t. \quad (3)$$

The SDO/HMI temporal cadence is constant ($\Delta_t = 45\text{s}$). Pairs of consecutive images ($\Delta_p = 1$) are used to ensure the best signal-to-noise ratio for the image registration and thus to achieve a high precision measurement of the image shifts Δx and Δy .

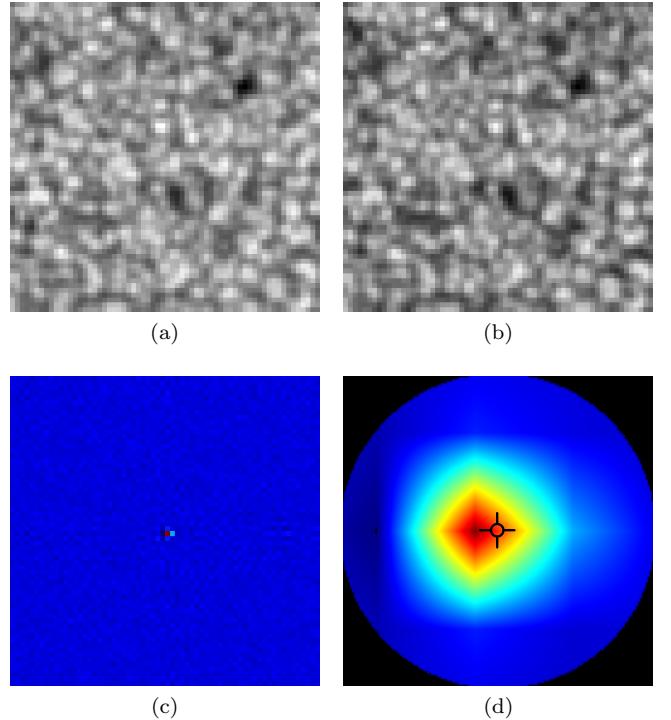


Figure 1. High precision subpixel image shift estimation via iterative phase correlation - first input 64x64px SDO/HMI image subregion (a), second input 64x64px SDO/HMI image subregion (b), heatmap of the phase correlation landscape computed from the two input images (with higher values corresponding to warmer colors) (c), circularly cropped heatmap of the bilinearly upscaled neighborhood around the location of maximal phase correlation with marked center of gravity (centroid) location (d).

2.2. The image registration algorithm

The local image shifts were calculated for a chosen set of pixel coordinates (and their neighborhoods) by the iterative phase correlation (IPC) high precision subpixel image registration algorithm (Hrazdíra et al. 2020). As the vast majority of horizontal shifts between pairs of consecutive ($\Delta t = 45\text{s}$) images lie in the interval $[0.1, 0.2]$ px, the high subpixel precision of the IPC algorithm justifies the choice of the method. Given that the IPC method was recently developed by the authors of this article, its main advantages and limitations were well known and were taken into account during the entire differential rotation estimation process.

For each of the registered pixels, two equal-sized square neighborhoods (subregions) are selected from each of the two SDO/HMI images at equal locations relative to the image reference pixel $[x_{mid}, y_{mid}]$, defined in the header of each SDO/HMI *Flexible Image Transport System - FITS* file. This way, changes in the image coordinates of the solar center caused by movements of

the SDO are taken into account and do not degrade the results. Since the time interval between the correlated images is only 45 seconds, all such changes are negligible in comparison with the measured image shift. Median absolute shift of the solar center was measured at approximately 0.005px in both directions (both according to the FITS file header and according to measurements by the IPC method). The corresponding image subregion pairs are then iteratively phase correlated, yielding the pixel shift estimates Δx and Δy for each of the selected locations.

2.3. Locations of the image shift measurements

To reduce the effects of distortions arising due to the perspective projection of the 3D solar surface onto the 2D image, only pixels lying on the central meridian (and their small neighborhoods) were considered. This also allowed for the construction of a central meridian flow speed map, with solar latitude θ on the vertical axis and time t (assuming the image pairs are ordered chronologically) on the horizontal axis. For each of the considered picture pairs, the central subregion pixels were distributed evenly across a portion of the central meridian, symmetrically around the central pixel $[x_{mid}, y_{mid}]$. The entire meridian (from $\theta = -90^\circ$ to $\theta = 90^\circ$) was not considered, given the noticeable distortions caused by the perspective projection of the solar surface near the geographic poles. We label the vertical pixel size of the limited meridional “field of view” as Θ_y , the number of individual registered subregions on the meridian of each picture pair as M and index each of the subregions from north to south with the ascending index μ (starting from zero). We further assume that the standard image coordinate system has the x coordinate (image column) increasing from left to right, and the y coordinate (image row) increasing from top to bottom. To only measure the shifts on the central meridian, the central horizontal pixel coordinates of each subregion $x(\mu)$ need to be fixed for all images at

$$x(\mu) = x_{mid}, \quad (4)$$

while the central vertical pixel coordinates of each subregion $y(\mu)$ (satisfying the equidistant and symmetry conditions mentioned above) can be set as

$$y(\mu) = y_{mid} + \Delta_y \left(\mu - \left\lfloor \frac{M}{2} \right\rfloor \right), \quad (5)$$

where Δ_y is the vertical spacing between the equidistant subregion centers and can be determined by

$$\Delta_y = \frac{\Theta_y}{M-1}. \quad (6)$$

The solar latitude $\theta(\mu)$ can be trigonometrically derived for every meridional subregion index μ as

$$\theta(\mu) = \arcsin \left(\frac{\Delta_y \left(\left\lfloor \frac{M}{2} \right\rfloor - \mu \right)}{R} \right) + \theta_0, \quad (7)$$

where the two newly introduced variables R and θ_0 denote the size of the solar radius in pixels and the latitude of the observer (instrument), and are both described in the SDO/HMI FITS file header. For the differential rotation measurement, the vertical field of view was chosen as $\Theta_y = 3400$ px and the number of subregions as $M = 851$ (an odd number of subregions, assuring the existence of a central element). According to equation 6, this results in a $\Delta_y = 4$ px vertical spacing between the equidistant meridional central subregion pixels.

2.4. Orthographic projection of the image shift

Since the SDO/HMI image is essentially a 2D perspective projection of the 3D solar surface, the measured meridional image shifts $\Delta x(\mu)$ and $\Delta y(\mu)$ (calculated by iteratively phase correlating pairs of image subregions with centers at pixel coordinates $[x(\mu), y(\mu)]$ given by equations 4 and 5) need to be projected back onto the solar surface and transformed into the angular shifts $\Delta\varphi_x(\mu)$ and $\Delta\varphi_y(\mu)$. Since the distance between the Solar Dynamics Observatory and the Sun is $\sim 200R_s$, the image projection was assumed orthographic (parallel projection lines orthogonal to the projection plane - point of perspective at infinite distance). As only the central meridian subregion shifts are calculated for every image pair, the following image shift backprojection equations depend only on the latitude information $\theta(\mu)$ (given in equation 7) and can also be derived trigonometrically. These orthographic backprojections transform the image pixel shifts $\Delta x(\mu)$, $\Delta y(\mu)$ into the angular shifts $\Delta\varphi_x(\theta(\mu))$ and $\Delta\varphi_y(\theta(\mu))$ given by

$$\Delta\varphi_x(\theta(\mu)) = \arcsin \left(\frac{\Delta x(\mu)}{R \cos(\theta(\mu))} \right), \quad (8)$$

$$\Delta\varphi_y(\theta(\mu)) = \theta(\mu) - \arcsin \left(\sin(\theta(\mu)) - \frac{\Delta y(\mu)}{R} \right). \quad (9)$$

2.5. Equations for the image shift-based solar angular velocity

If we now combine six of the previously described equations - the general equations 1 and 2 for the angular velocities $\omega_x(\theta(\mu))$ and $\omega_y(\theta(\mu))$, equation 3 for the temporal difference between the two images Δt , equation 7 for the solar latitude $\theta(\mu)$ and the orthographic projection equations 8 and 9 for the angular

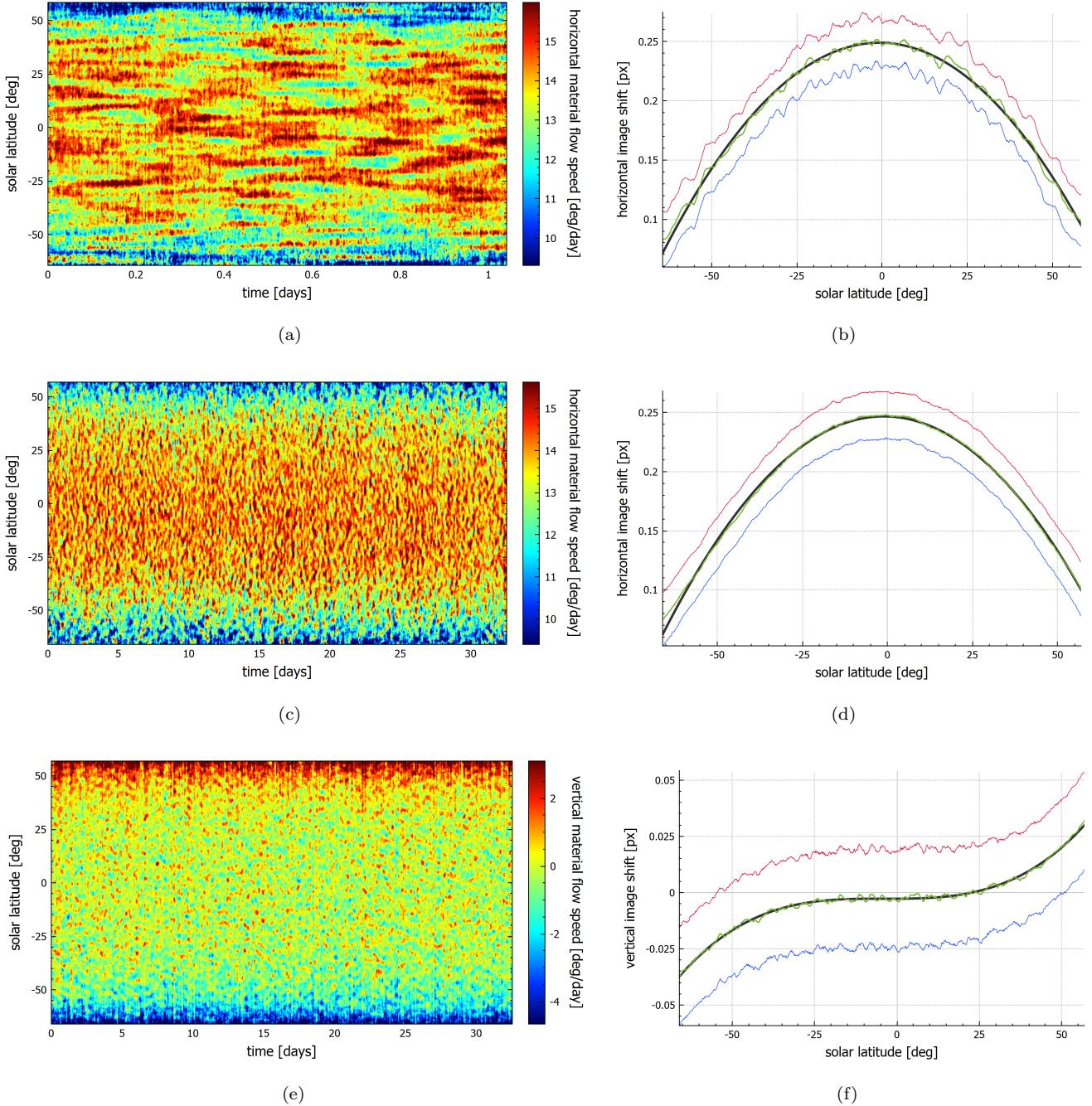


Figure 2. Solar horizontal and vertical flow speed maps (left column) obtained from orthographically backprojecting image shifts measured on 2000 SDO/HMI continuum full-disk images back onto the solar surface with the corresponding plots of the average image shift profiles (right column). Image shift calculations were done on 1999 uniformly distributed image pairs with $\Delta t = 45\text{s}$ intervals between individual images in each pair. First row: horizontal flow speed map (a) and corresponding average horizontal image shift profile $\Delta x(\theta)$ (b) (green) obtained from images spanning approximately a one day period from 1.1.2020 to 2.1.2020, together with its best least squares second degree polynomial fit curve (black) and one sample standard deviation band (blue, red). Second row: same as the first row for a one month period from 1.1.2020 to 2.2.2020. Third row: vertical flow speed map (e) and corresponding average vertical image shift profile $\Delta y(\theta)$ (f) (green) obtained from images spanning approximately a one month period from 1.1.2020 to 2.2.2020, together with its best least squares third degree polynomial fit curve (black) and one sample standard deviation band (blue, red).

shifts $\Delta\varphi_x(\theta(\mu))$ and $\Delta\varphi_y(\theta(\mu))$, we get the final equations for the estimated angular velocities $\omega_x(\theta(\mu))$ and $\omega_y(\theta(\mu))$ of a meridional pixel with index μ (with pixel indexing as described in section 2.3) in the form of

$$\omega_x(\mu) = \frac{1}{\Delta_p \Delta_t} \operatorname{arccsc} \left(\frac{R}{\Delta x(\mu)} \cos \left(\arcsin \left(\frac{\Theta_y}{M-1} \frac{1}{R} \left(\left[\frac{M}{2} \right] - \mu \right) \right) + \theta_0 \right) \right) \quad (10)$$

for the horizontal solar angular velocity, and

$$\begin{aligned} \omega_y(\mu) = & \frac{1}{\Delta_p \Delta_t} \left(\arcsin \left(\frac{\Delta_y \left(\left[\frac{M}{2} \right] - \mu \right)}{R} \right) + \theta_0 - \right. \\ & \left. \arcsin \left(\sin \left(\arcsin \left(\frac{\Delta_y \left(\left[\frac{M}{2} \right] - \mu \right)}{R} \right) + \theta_0 \right) - \frac{\Delta y(\mu)}{R} \right) \right) \end{aligned} \quad (11)$$

for the vertical solar angular velocity.

3. RESULTS

3.1. Data sets used for the measurement

The horizontal and vertical solar angular velocities $\omega_x(\theta)$ and $\omega_y(\theta)$ were measured for two sets of $P = 1999$ pairs of consecutive images. The first data set uniformly spans the time interval from 1.1.2020 to 2.2.2020. This period approximately equals to one full solar revolution at the equator (where the material flows faster than average) and 80% of a revolution near the poles (where the material flows slower than average).

The second data set uniformly spans the time interval from 1.1.2020 to 2.1.2020. This period approximately equals to 4% of a solar revolution at the equator and 3% of a revolution near the poles.

The duration of both data set intervals (~ 1 day and ~ 1 month) were chosen deliberately to measure both the potentially interesting short-lived behaviour and also the overall rotational speed profile. The specific initial date of both data sets (1.1.2020) was chosen arbitrarily and is not in any way special - both image data sets show a very “calm” behaviour during this period, with almost no clearly discernible large scale features, such as sunspots, present.

The size of the central meridional subregions, from which the image shifts were calculated, was set to 64x64px. This size proved to be large enough for good image registration quality, and at the same time small enough to provide good spatial flow speed resolution.

3.2. Flow speed maps

Since M subregions were equidistantly distributed among different latitudes on the central meridian of every picture and measurements for P picture pairs were made, very detailed 2D horizontal and vertical central meridian flow speed maps with $P \times M$ resolution could be constructed for each data set. These maps can be seen in Fig. 2 (a), (c) and (e).

Although the measured angular velocity clearly varies with time in a given latitude band, the temporally coherent angular velocity values in the flow speed maps demonstrate the good image registration quality. Measured from horizontally (temporally) adjacent image shifts, although being computed from entirely different pairs of images, they are very similar at each latitude, which is demonstrated by the existence of the clearly distinguishable coherent faster than average and slower than average flows in Fig. 2 (a).

The horizontal and vertical image shift $\Delta x(\theta)$, $\Delta y(\theta)$ and angular velocity $\omega_x(\theta)$, $\omega_y(\theta)$ values from all the P image pairs were also averaged at each latitude, resulting in picture-pair averages $\overline{\Delta x(\theta)}$, $\overline{\Delta y(\theta)}$ (see Sections 3.3 and 3.4) and $\overline{\omega_x(\theta)}$, $\overline{\omega_y(\theta)}$ (see Section 3.5). The rather smooth characteristics of the average horizontal and vertical image shift profiles $\overline{\Delta x(\theta)}$ and $\overline{\Delta y(\theta)}$ (see Fig. 2 (b), (d) and (f)) can serve as good indicators of the entire rotational speed measurement quality.

3.3. Horizontal image shift profile

Upon closer inspection of the average horizontal image shift profiles obtained from both data sets, very coherent curves with very little noise can be seen. The profile obtained from the one day data set (Fig. 2 (b)) is somewhat more chaotic than that of the one month data set (Fig. 2 (d)), which is expected, as the profile captures local material flows and events happening on time scale comparable to the entire data set time span. These events can be seen very clearly as reddish and blueish streaks in the flow speed maps of both data sets, showing regions of faster than average or slower than average material flow, lasting approximately 8 hours. The shorter time span data set flow speed map shows individual flows in more detail, while the longer time span flow speed map demonstrates the practically constant duration and presence of these flows. The profile obtained from the longer time span data set is almost perfectly smooth, displaying the overall global average horizontal rotational speed profile.

3.4. Vertical image shift profile

Although being roughly one order of magnitude less significant than its horizontal counterpart ($\sim 0.01\text{px}$ compared to $\sim 0.2\text{px}$), the measurement result of the average vertical image shift profile $\overline{\Delta y(\theta)}$ is also quite interesting. The vertical profile clearly shows more noisy behaviour, which can be caused by being near the accuracy limit of the image registration precision. In the vertical part of the registration, the iterative phase correlation method is estimating vertical shifts around 0.01 pixels between two 64x64px images, and such shifts are very hard to estimate with higher precision. However, a very clear overall trend can still be seen in the average profile (see Fig. 2 (f)).

In the northern solar hemisphere, the material seems to be flowing southwards, in addition to the natural ($\sim 20\times$ faster) eastwards rotational flow. This southwards part of the flow gets more significant closer to the north pole, roughly following a cubic (or arcsine) trend. The vertical image shift profile appears to be very close to being central-symmetric, meaning that the material appears to be flowing southwards in the northern hemisphere and northwards in the southern hemisphere, resulting in an overall flow towards the equator in both hemispheres, which speeds up significantly near the poles.

3.5. The final flow speed profiles

According to multiple studies (Zhang et al. 2015; Brajša et al. 2000), the southern solar hemisphere seems to be currently (as of 2020) rotating (slightly, but noticeably) faster, than the northern hemisphere. The usual horizontal flow trigonometric fit (denoted $\omega^\dagger(\theta)$) of the form

$$\omega^\dagger(\theta) = A + B \sin^2(\theta) + C \sin^4(\theta) \quad (12)$$

is always zero-symmetric, since even powers of the function $\sin(\theta)$ are always even, and thus cannot capture such north-south asymmetric behaviour. This results in unnecessarily large curve fitting errors if the flow speed profile is not separated into northern and southern parts. On the contrary, polynomial fit of a second degree (denoted $\omega_x^*(\theta)$) in the simple form of

$$\omega_x^*(\theta) = A + B\theta + C\theta^2 \quad (13)$$

can capture any northern-southern hemisphere asymmetries with much less error and with the same amount of fitting parameters. This fit is thus included in all the image shift and flow speed profile figures, where the flow speed is not separated into northern and southern parts.

As the average vertical image shift profile clearly trends towards a cubic, it is only natural to fit the resulting vertical flow speeds with a polynomial of a third

Flow speed curve type	A	B	C	D
HB (1 d.) - trig.	14.16	-0.37	-5.02	-
HB (1 mo.) - trig.	14.10	-0.99	-4.11	-
HN (1 d.) - trig.	14.07	0.88	-8.40	-
HN (1 mo.) - trig.	14.00	-0.71	-5.74	-
HS (1 d.) - trig.	14.23	-1.32	-1.65	-
HS (1 mo.) - trig.	14.14	-0.52	-3.14	-
VN (1 d.) - trig.	-0.15	-0.69	6.34	-
VN (1 mo.) - trig.	-0.06	-1.25	7.81	-
VS (1 d.) - trig.	-0.24	0.28	-3.81	-
VS (1 mo.) - trig.	-0.14	-0.06	-2.58	-
HB (1 d.) - 2 nd deg. poly.	13.36	-0.31	-2.87	-
HB (1 mo.) - 2 nd deg. poly.	14.25	-0.34	-2.93	-
VB (1 d.) - 3 rd deg. poly.	-0.17	-0.46	0.007	3.33
VB (1 mo.) - 3 rd deg. poly.	-0.14	-0.51	0.30	3.37

Table 1. Best (least squares) fitting coefficients A, B, C, D for different types of horizontal (H) and vertical (V) flow speed profile fits for northern (N), southern (S) or both (B) hemispheres, obtained from both one day and one month data sets according to Equations 12, 13 and 14.

degree in the simple form of

$$\omega_y^*(\theta) = A + B\theta + C\theta^2 + D\theta^3. \quad (14)$$

The values of the fitting parameters $A, B, C(D)$ for all mentioned flow profiles in degrees per day for northern, southern and both hemispheres together are provided in Table 1. The final average horizontal and vertical flow speed profiles, together with their corresponding polynomial fits (and together with profiles inferred by different techniques, for comparison, in the horizontal case) can be seen in Fig. 3 (a) and (c).

Furthermore, for a better visualization of the North-South asymmetry of the resulting horizontal and vertical flows, Fig. 3 (b) and (d) show the average flow profiles and their corresponding best least squares trigonometric fits (according to Equation 12) separately for each hemisphere, with absolute latitude on the horizontal axis. The parameters for all plotted curves can be found in Table 1.

4. SUMMARY AND CONCLUSIONS

The novel iterative phase correlation technique of solar differential rotation estimation described and illustrated in this paper offers a new method for the empirical determination of the differential rotation rate of the solar photosphere. It is based on the measurement of locally varying image shifts between consecutive ($\Delta t = 45\text{s}$) solar images. The technique was applied to SDO/HMI data and roughly covered one Carrington rotation of

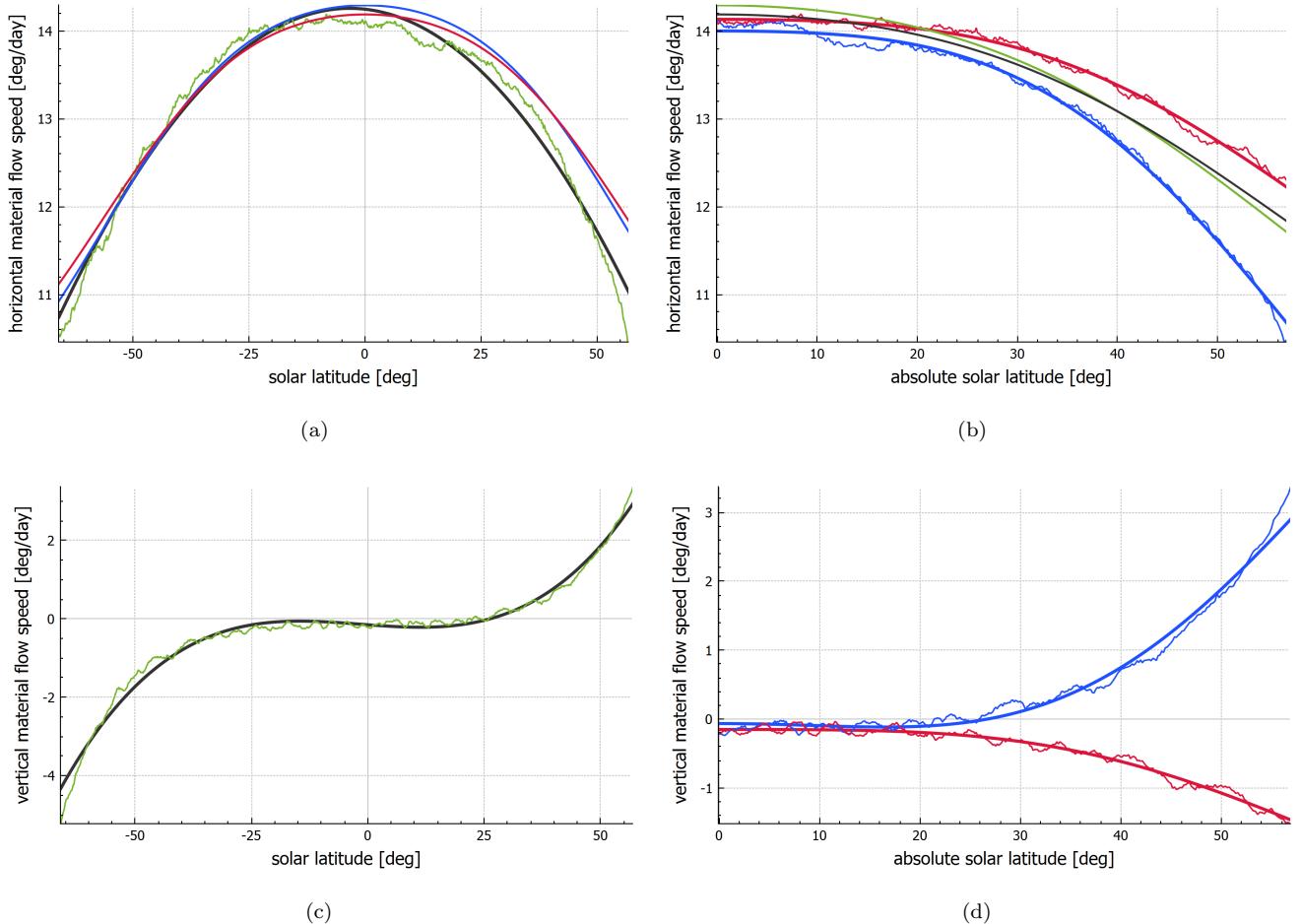


Figure 3. Average measured horizontal (first row) and vertical (second row) flow speed profiles $\overline{\omega_x(\theta)}$ and $\overline{\omega_y(\theta)}$ at different latitudes in degrees per day. Measured from a one month period from 1.1.2020 to 2.2.2020. The flow speed profile was obtained by temporally averaging values at each latitude in the corresponding flow speed map (Fig. 2 (c)). Figures (a) and (c) contain the entire average horizontal and vertical profiles (green) with their corresponding best least squares second and third degree polynomial fits $\omega_x^*(\theta)$, $\omega_y^*(\theta)$ (black). Furthermore, two sample horizontal trigonometric flow speed profiles inferred by different techniques (Lamb 2017; Howard et al. 1983) are also included (blue, red) in (a). Figures (b) and (d) contain the average measured horizontal and vertical flow speed profiles with their corresponding best least squares trigonometric fit curves $\omega_x^\dagger(\theta)$, $\omega_y^\dagger(\theta)$ for northern (blue) and southern (red) hemisphere separately at different absolute latitudes, together with the two sample horizontal trigonometric flow speed profiles inferred by different techniques (green, black) in (b).

continuous observations from 1.1.2020 to 2.2.2020. The technique has a number of outstanding properties which make it superior to other correlation-based techniques used in the past.

The local shift measurement between two images is not limited to contrasty features like sunspots, as it can be applied even to very low contrast structures, such as granulation and/or faculae (see Fig. 1). Hence, the differential rotation rate thus determined is independent of any a priori selection of solar features. Therefore the method can be successfully applied to any location in the photosphere and does not require long observations spanning months or years (see the results obtained by measurements from just one day in Fig. 2 (a) and (b)).

The very high temporal resolution of the presented method enables the study of short-period photospheric phenomena and the comparison of measurements on different days separated by an event that may have a global impact on the solar magnetic field, such as a solar flare or a coronal mass ejection. Furthermore, given that the technique does not rely on visible features, it can be used throughout a full solar cycle independently of the presence of specific features in the photosphere.

While methods based on solar markers yield a limited number of data points, the iterative phase correlation technique is applicable to every pixel close to central meridian within a given latitude band. This yields a

unique opportunity to study the rotation rate of the Sun locally.

The method is also able to estimate the locally varying vertical shifts between consecutive solar images. Although being very small and hard to measure ($\sim 0.01\text{px}$ - see Fig. 2 (f)), the vertical image shifts obtained by the novel image registration method show coherent results, displaying a consistent global trend (see Fig. 3 (c)). The reliability of the vertical image shift measurement allows further studies of the general (not only rotational) photospheric movements with great detail. The vertical flow speed results show an overall flow towards the solar equator, with speeds becoming more significant closer to the geographic poles. Both the horizontal and vertical flow speed profiles were fitted with polynomial and trigonometric curves, whose best (least squares) parameters can be seen for both one day and one month observations in Table 1.

In conclusion, the reliability of the approach is underscored by the fact that the average sidereal differential rotation curve obtained is in the middle of other published results (see Fig. 3 (a) and (b)). At the same time it also clearly shows the now well established N-S rotational speed asymmetry. Furthermore, the quality of the results is further supported by the smoothness of the average horizontal image shift profile obtained from a single Carrington rotation of SDO/HMI image data (see Fig. 2 (d)) and also by the coherence of the measured flow speed values inferred from temporally adjacent images. As an added bonus, the novel technique described in this work allows the study of short-lived and previously undetected phenomena associated with the solar magnetic field.

The work of Z. Hrazdíra and M. Druckmüller was supported by the Grant Agency of Brno University of Technology, project No. FSI-S-14-2290.

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