

# LAB 9

## L<sup>A</sup>T<sub>E</sub>X

The goals of this lab are to learn how to produce a document in L<sup>A</sup>T<sub>E</sub>X, and to practice structured scientific writing. In order to proceed, you will need to first have a L<sup>A</sup>T<sub>E</sub>X distribution installed or use a free on-line service like Overleaf.

### 12.1 Symbols and equations

Below are three Nobel prize-winning excerpts from the physics literature (from Riess et al. 1998; Higgs 1964; Einstein 1905). Write up a L<sup>A</sup>T<sub>E</sub>X document that reproduces them **exactly**, including the equations.

You may choose a template or style file to work from. You will need the **amsmath** package for some of the equation formatting. Sometimes fractions and integral signs turn out unpleasantly squashed, which you can improve by prefixing the commands with `\displaystyle` (or for fractions, replace `\frac` with `\dfrac`).

To avoid typing in everything, you may want to cut and paste from the PDF document, at least for the plain-text components.

**1.** *High-redshift SNe Ia are observed to be dimmer than expected in an empty universe (i.e.,  $\Omega_M = 0$ ) with no cosmological constant.* A cosmological explanation for this observation is that a positive vacuum energy density accelerates the expansion. Mass density in the universe exacerbates this problem, requiring even more vacuum energy. For a universe with  $\Omega_M = 0.2$ , the MLCS and template-fitting distances to the well-observed SNe are 0.25 and 0.28 mag farther on average than the prediction from  $\Omega_\Lambda = 0$ . The average MLCS and template-fitting distances are still 0.18 and 0.23 mag farther than required for a 68.3%(1  $\sigma$ ) consistency for a universe with  $\Omega_M = 0.2$  and without a cosmological constant.

**2.** The simplest theory which exhibits this behavior is a gauge-invariant version of a model used by Goldstone himself: Two real<sup>1</sup> scalar fields  $\varphi_1, \varphi_2$  and a real vector field  $A_\mu$  interact through the Lagrangian density

$$L = -\frac{1}{2}(\nabla\varphi_1)^2 - \frac{1}{2}(\nabla\varphi_2)^2 - V(\varphi_1^2 + \varphi_2^2) - \frac{1}{4}F_{\mu\nu}F^{\mu\nu}, \quad (12.1)$$

where

$$\begin{aligned}\nabla_\mu\varphi_1 &= \partial_\mu\varphi_1 - eA_\mu\varphi_2, \\ \nabla_\mu\varphi_2 &= \partial_\mu\varphi_2 + eA_\mu\varphi_1, \\ F_{\mu\nu} &= \partial_\mu A_\nu - \partial_\nu A_\mu,\end{aligned}$$

$e$  is a dimensionless coupling constant, and the metric is taken as  $-+++$ .

**3.** Letztere ist  $\frac{3}{2}(R/N)T$ , während man für die mittlere Größe des Energiequantums unter Zugrundelegung der Wienschen Formel erhält:

$$\frac{\int_0^\infty \alpha \nu^3 e^{-\frac{\beta\nu}{T}} d\nu}{\int_0^\infty \frac{N}{R\beta\nu} \alpha \nu^3 e^{-\frac{\beta\nu}{T}} d\nu} = 3\frac{R}{N}T.$$

**Submit two files to Canvas:** your final PDF, and your L<sup>A</sup>T<sub>E</sub>X source file (which you can obtain from Overleaf by selecting Project, then Download as ZIP, then extracting the `.tex` file from the downloaded zip file).

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<sup>1</sup>In the present note the model is discussed mainly in classical terms; nothing is proved about the quantized theory...