

Briefing Space Weather

2022/05/16

1 Sun

1.1 Responsible: José Cecatto

05/09 – No fast wind stream; 5 CME c.h.c. toward the Earth;
05/10 – 1 X1 flare; No fast wind stream; 6 CME c.h.c. toward the Earth;
05/11 – 3 M5- flares; No fast wind stream; 12 CME c.h.c. toward the Earth;
05/12 – 1 M1 flare; No fast wind stream; 8 CME c.h.c. toward the Earth;
05/13 – No fast wind stream; 11 CME c.h.c. toward the Earth;
05/14 – No fast wind stream; 2 CME c.h.c. toward the Earth;
05/15 – 1 M2 flare; Fast wind stream (< 600 km/s); 1 CME c.h.c. toward the Earth;
05/16 – 1 M2 flare; Fast wind stream (<= 550 km/s); No CME toward the Earth;
Prev.: Fast wind stream up to May 18; for the next 2 days relatively low (35% M, 10% X) probability of M / X flares;
also, occasionally other CME can present component toward the Earth.
c.h.c. – can have a component

2 Sun

2.1 Responsible: Douglas Silva

- WSA-ENLIL (CME 2022-05-06T17:12Z, 2022-05-06T21:24Z)
 - The simulation indicates that the Coronal Mass Ejections' flanks will reach the DSCOVR mission between 2022-05-11T04:00Z and 2022-05-11T18:00Z.
- WSA-ENLIL (CMEs 2022-05-10T15:12Z and 2022-05-10T14:48Z)
 - The simulation results indicate that the flanks of combined Coronal Mass Ejections will reach the DSCOVR mission between 2022-05-13T16:00Z and 2022-05-14T06:00Z.
- WSA-ENLIL (CME 2022-05-10T14:48Z)
 - The simulation results indicate that the flank of CME will reach the DSCOVR mission between 2022-05-13T21:00Z and 2022-05-14T11:00Z.

Buracos coronais (SPoCA : Spatial Possibilistic Clustering Algorithm):

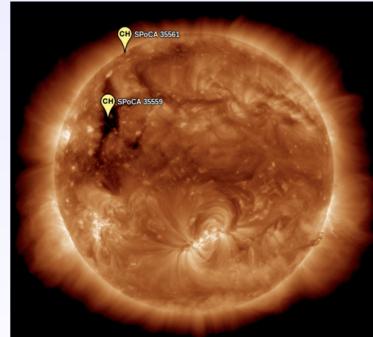
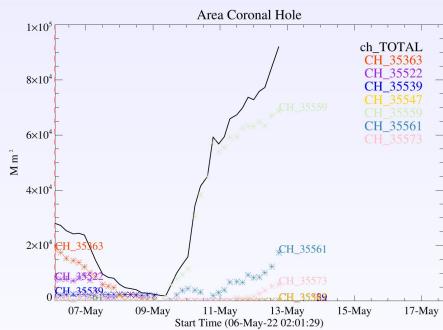
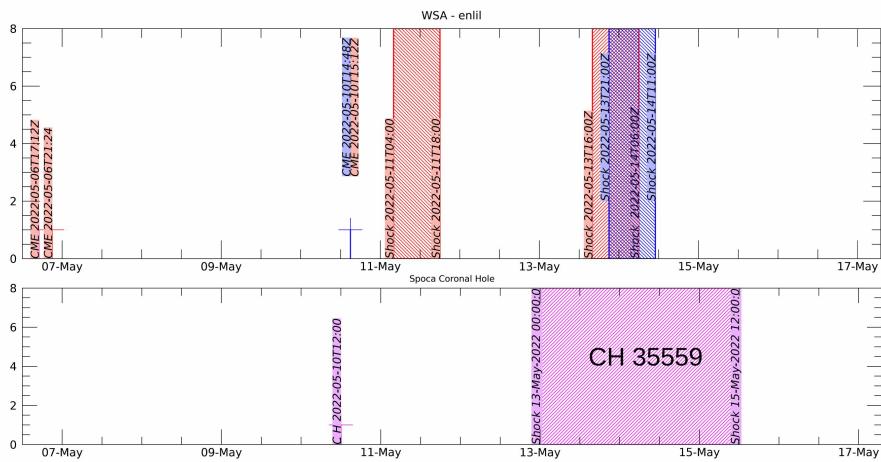


Figura: A linha em preto mostra o resultado da soma das áreas para cada intervalo da detecção realizado pelo SPOCA entre os dias 06 e 13 de maio de 2022

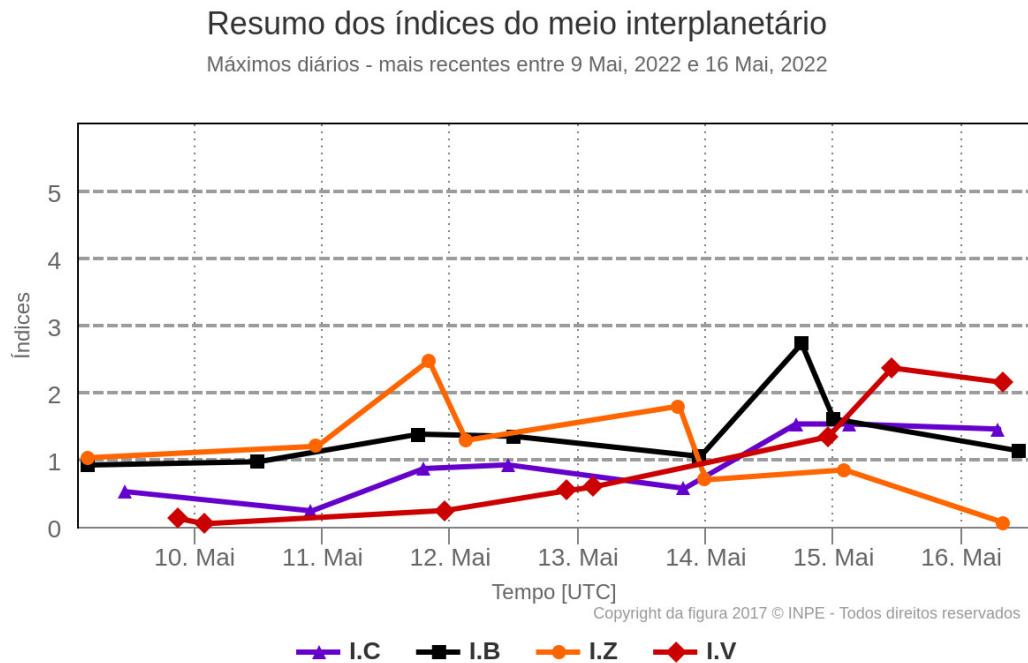
Sobre a imagem em 193 Å do Sol estão destacados os Buracos coronais observados pelo SPOCA por volta das 21:00 UT do dia 10 de maio de 2022.

WSA - ENLIL SPOCA



3 Interplanetary Medium

3.1 Responsible: Paulo Jauer



- The interplanetary medium region in the last week showed a low/moderate level of plasma perturbations due to the possible interaction of CME and HSS-like structures identified by the DISCOVERY satellite in the interplanetary medium.
- T
- The modulus of the interplanetary magnetic field component showed 1 maximum peak : 14/May at 18:30 of 18 nT.
- The BxBy components showed variations in the analyzed period, both remaining oscillating within the [+10, -10] nT interval, with a change of sector on May 12 at 13:30.
- The bz field component showed fluctuations with a positive value of 12 nT on May 16 at 02:30 and a negative value of -9.19 nT at 19:30 UT on May 11. On average, the Bz component oscillated mostly negative. Conditions favorable to the emergence of geomagnetic disturbances.
- The solar wind density oscillated mostly below 10 p/cm³ during the analyzed period with a maximum peak on May 012 at 1:30 pm from 23 p/cm³, and another peak on May 14 at 5:30 pm from 27 p. /cm³.
- The solar wind speed had oscillated mostly below 400 km/s until the 14th of May at 20:30, changing to higher values with a peak on the 15th of May at 566 km/s.
- The magnetopause position was oscillating on average at the typical 10 Re position. It showed two significant compression on May 14th and 15th of 8.6 Re at 17:30 and 03:30 UT respectively.

4 Radiation Belts

4.1 Responsible: Ligia Alves da Silva

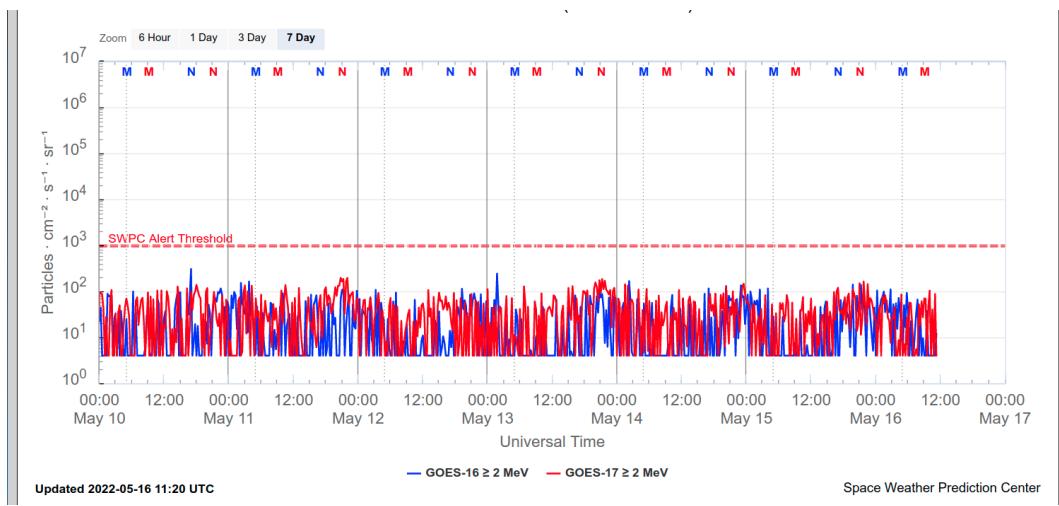


Figura 1: High-energy electron flux (≥ 2 MeV) obtained from GOES-16 and GOES-17 satellite. Source

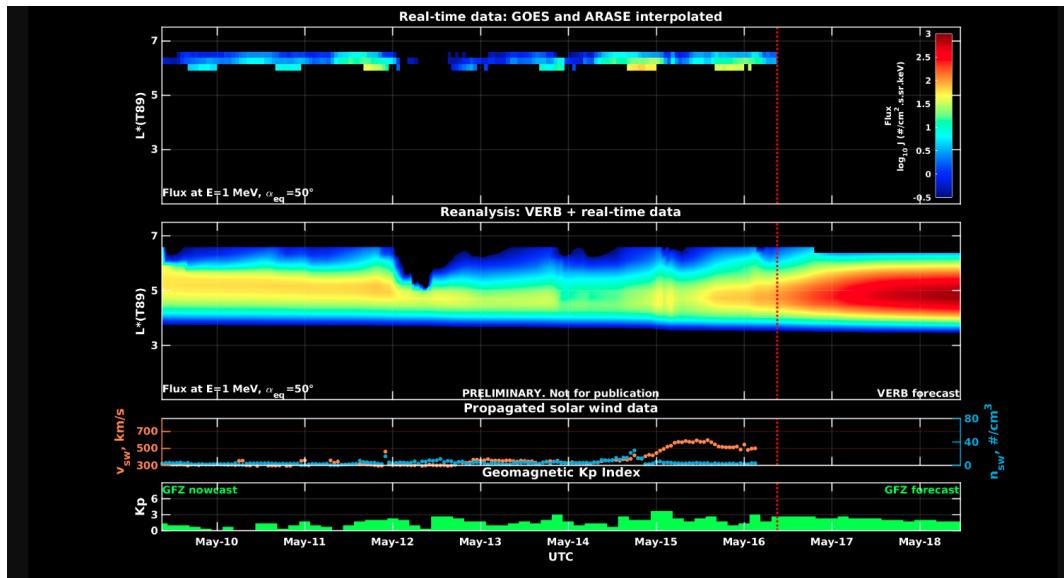


Figura 2: high-energy electron flux data (real-time and interpolated) obtained from ARASE, GOES-16, GOES-17 satellites. Reanalysis's data from VERB code and interpolated electron flux. Solar wind velocity and proton density data from ACE satellite. Source

High-energy electron flux (≥ 2 MeV) in the outer boundary of the outer radiation belt obtained from geostationary satellite data GOES-16 and GOES-17 (Figure 1) is below 102 particles/($\text{cm}^2 \text{ s sr}$) throughout the week of analysis. A significant electron flux decrease is observed at the beginning of May 12th, removing almost completely the high-energy electron from the outer boundary of the outer radiation belt. The outer boundary returns to the initial values of 102 particles/($\text{cm}^2 \text{ s sr}$) from 12:00 UT.

The GOES-16, GOES-17, and Arase satellite data are analyzed and interpolated to observe the high-energy electron flux variability (1 MeV) in the outer radiation belt (Figure 2). Additionally, the VERB code rebuilds this electron considering the Ultra Low Frequency (ULF) waves' radial diffusion. An electron flux increase between 4.5 \pm L-shell \pm 5.5 is observed from 12:00 UT on May 15th. These electron flux variability occurred concomitantly with the arrival of solar wind structures and ULF wave

activities. However, it is important to point out that the data from the ARASE satellite are not available for the week under analysis to confirm the L-shell level of these electron flux variabilities.

5 ULF Waves

5.1 Responsible: José Paulo Marchezi

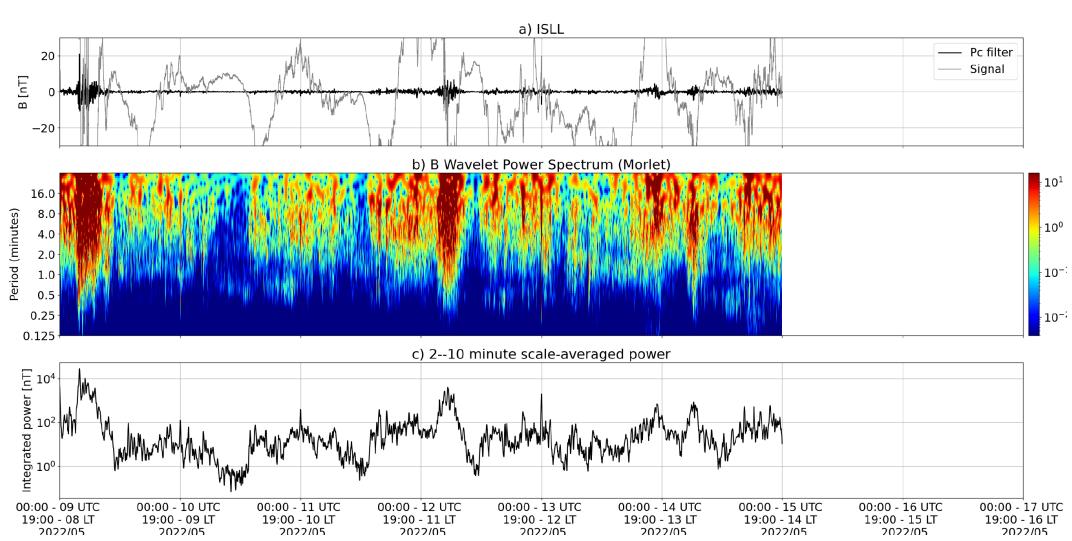


Figura 3: a) signal of the total magnetic field measured in the ISLL Station of the CARISMA network in gray, together with the fluctuation in the range of $Pc5$ in black. b) Wavelet power spectrum of the filtered signal. c) Average spectral power in the ranges from 2 to 10 minutes (ULF waves).

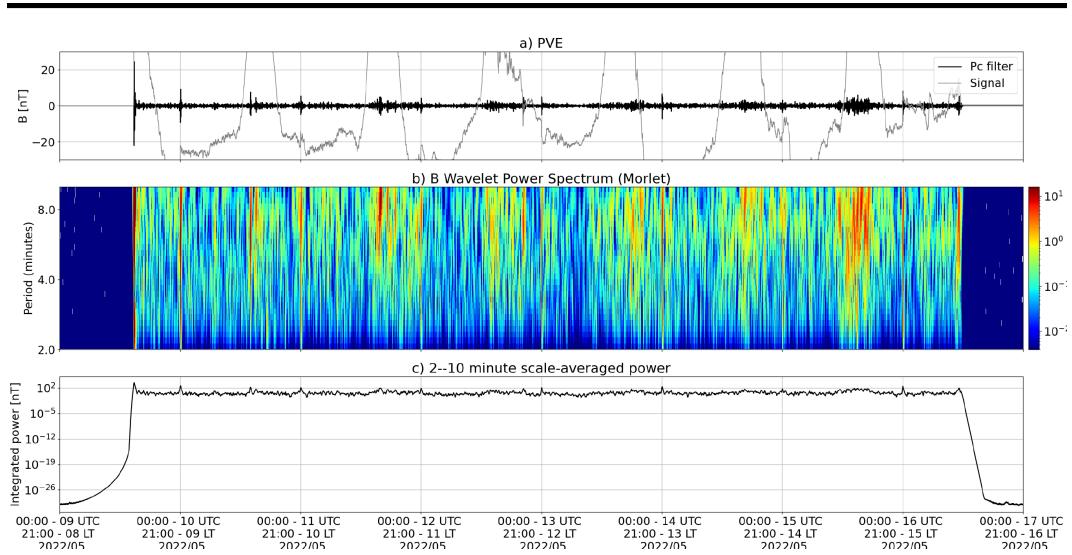


Figura 4: a) signal of the total magnetic field measured in the EMBRACE network in gray, together with the fluctuation in the range of $Pc5$ in black. b) Wavelet power spectrum of the filtered signal. c) Average spectral power in the ranges from 2 to 10 minutes (ULF waves).

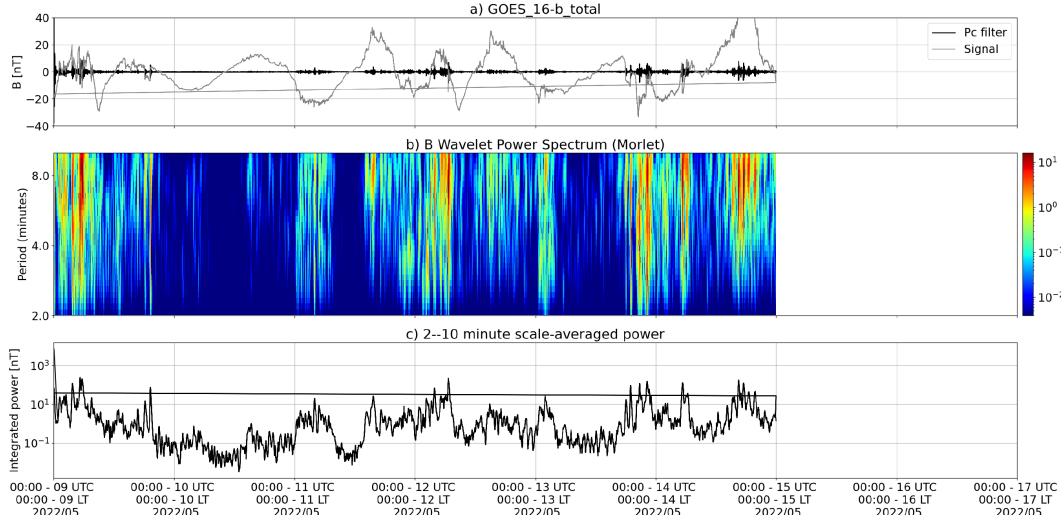
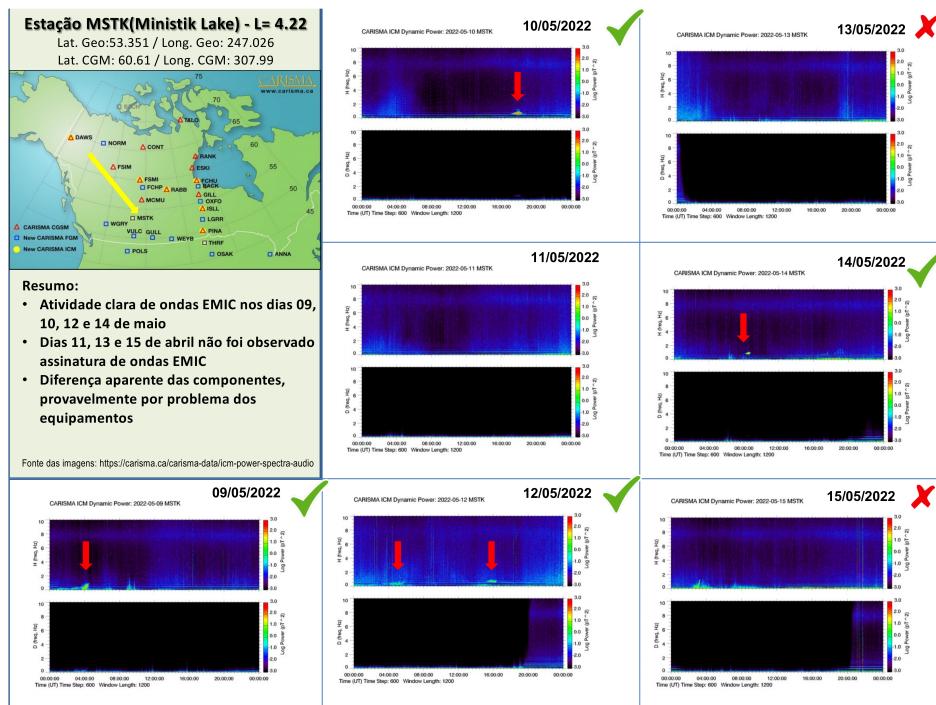
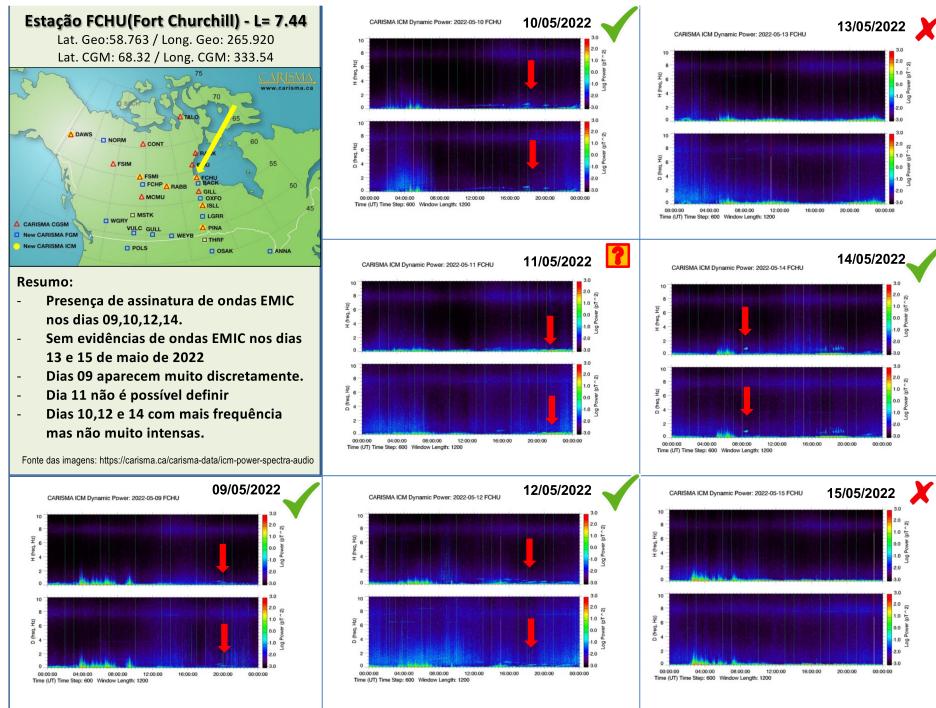


Figura 5: a) signal of the total magnetic field measured by the GOES 16 satellite, together with the fluctuation in the range of $Pc5$ in black. b) Wavelet power spectrum of the filtered signal. c) Average spectral power in the ranges from 2 to 10 minutes (ULF waves).

On the 9th of May there is an increase in the power of the ULF waves lasting about 12 hours, with continuous characteristics. ULF wave activity decreases on the 10th and 11th, and only increases again in the first half of the 12th, with short-lived oscillations covering the frequency range of the $Pc3-Pc5$ geomagnetic pulsations. The data from the Porto Velho station (PVE) show considerable fluctuations throughout the period, possibly due to local characteristics, under the influence of the equatorial electrojet. In São Martinho da Serra there is an increase in power from the 14th, characteristic of an interaction with an interplanetary shock, followed by a possible High-Speed Stream. The y component of the magnetic field measured by the GOES satellite is the one with the highest wave power, on the 9th, 12th and 14th of May. Summary 10/10 On the 9th of May there is an increase in the power of the ULF waves lasting about 12 hours, with continuous characteristics. ULF wave activity decreases on the 10th and 11th, and only increases again in the first half of the 12th, with short-lived oscillations covering the frequency range of the $Pc3-Pc5$ geomagnetic pulsations. The data from the Porto Velho station (PVE) show considerable fluctuations throughout the period, possibly due to local characteristics, under the influence of the equatorial electrojet. In São Martinho da Serra there is an increase in power from the 14th, characteristic of an interaction with an interplanetary shock, followed by a possible High-Speed Stream. The y component of the magnetic field measured by the GOES satellite is the one with the highest wave power, on the 9th, 12th and 14th of May.

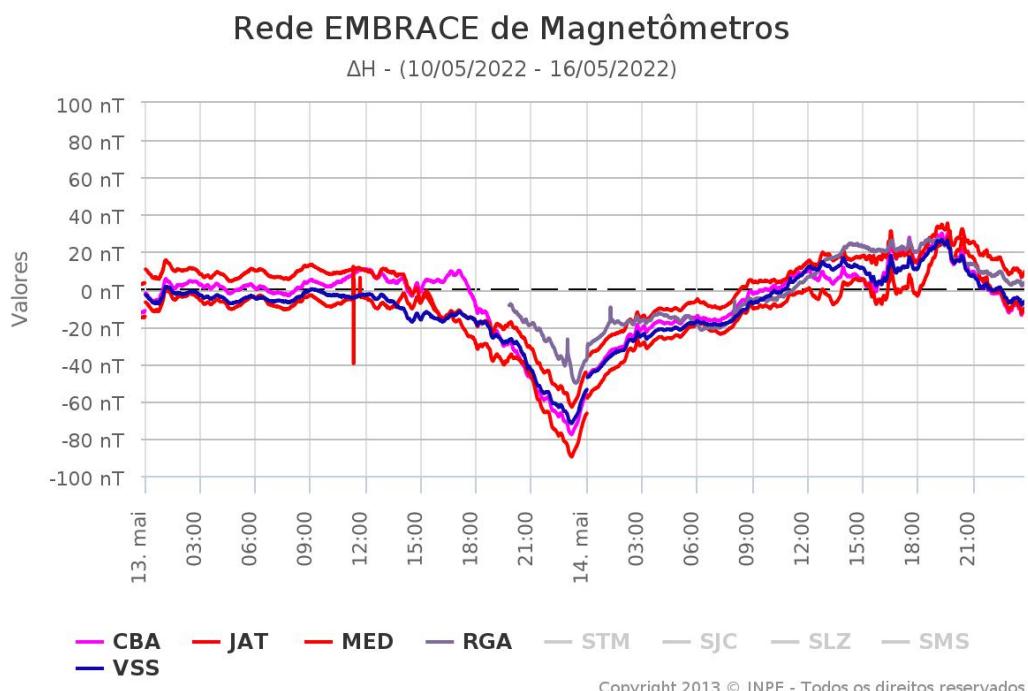
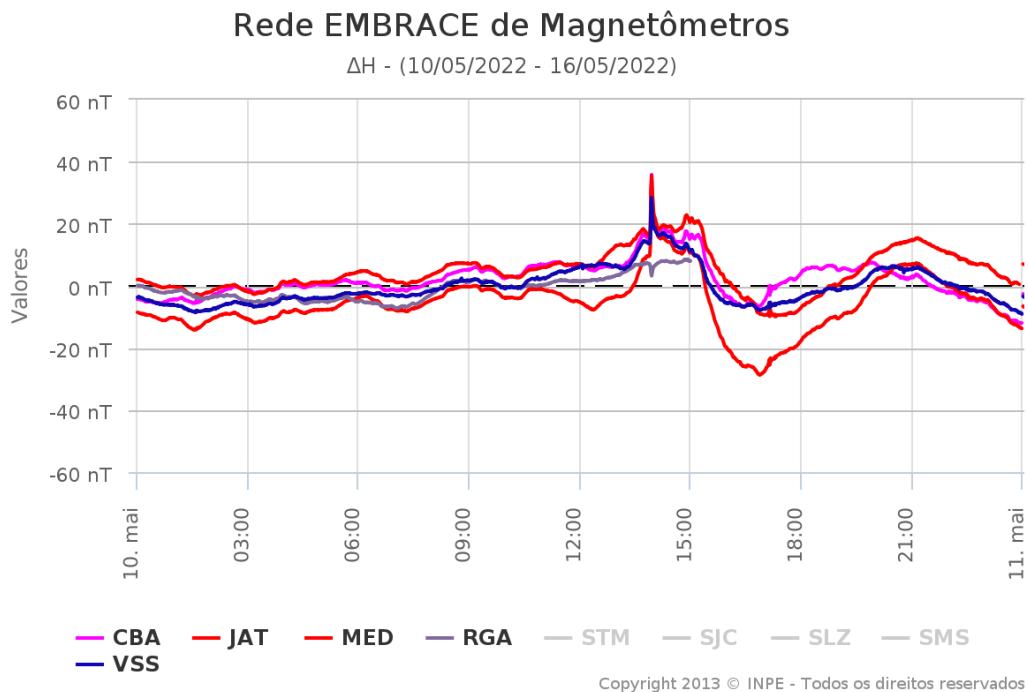
6 Ondas EMIC

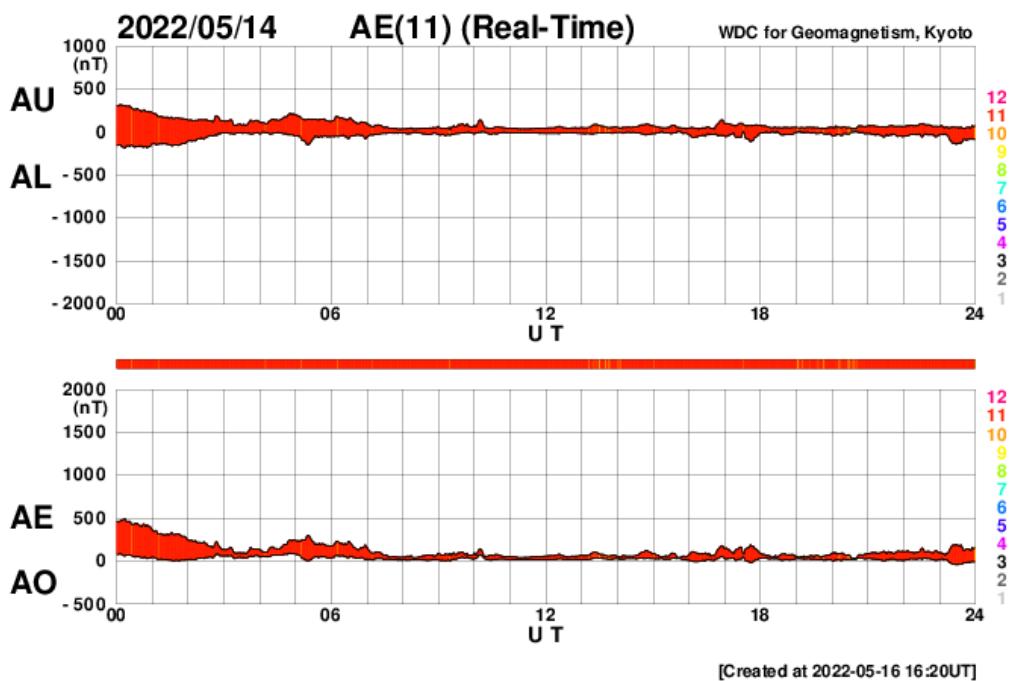
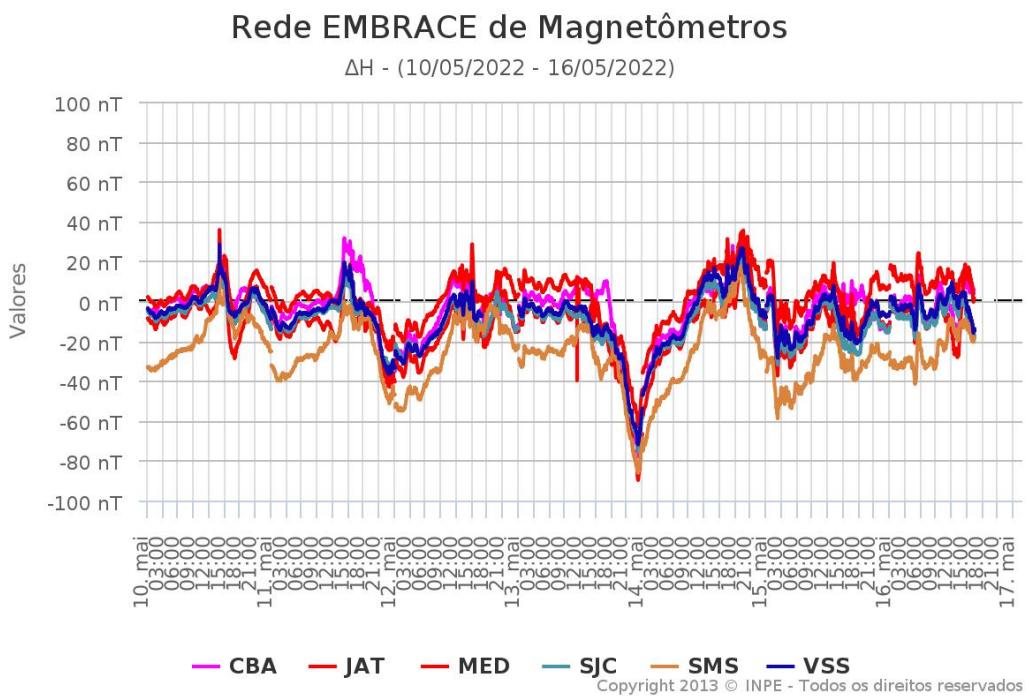
6.1 Responsável: Claudia Medeiros

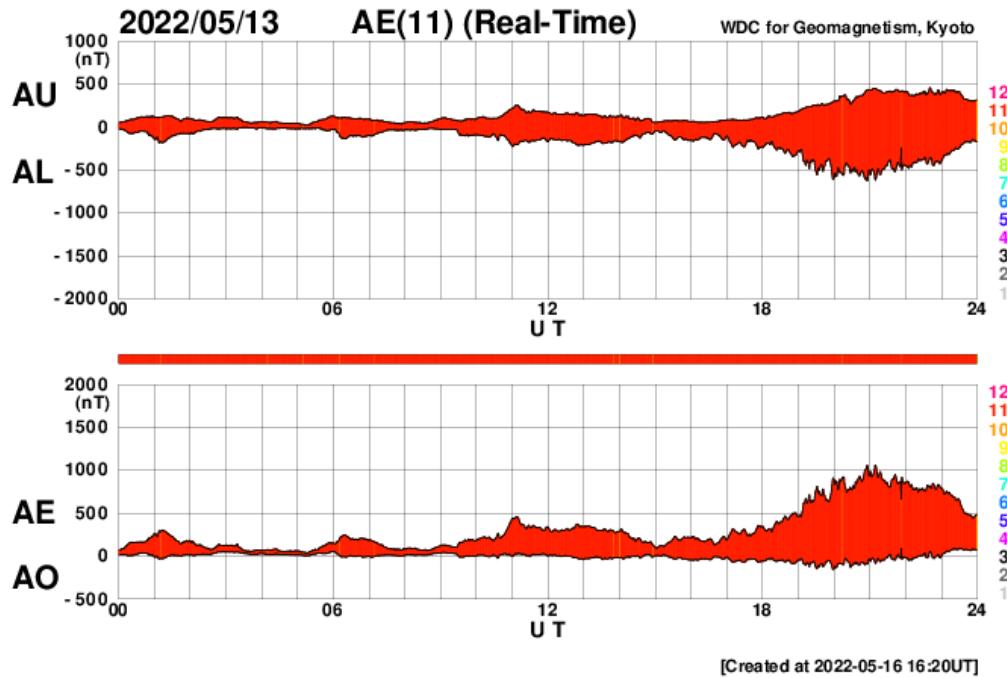
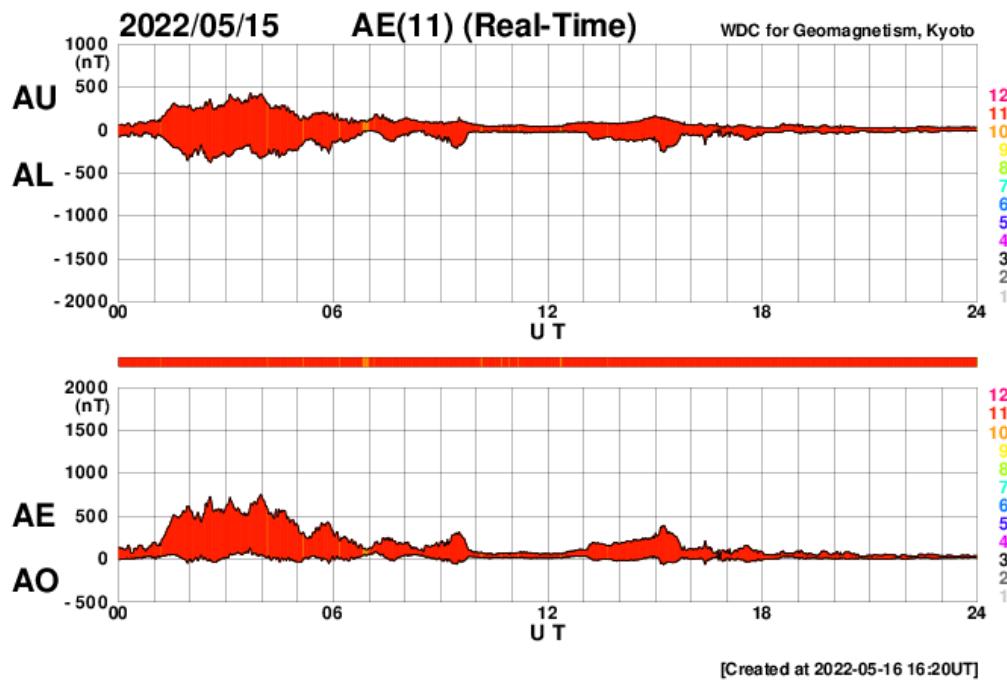


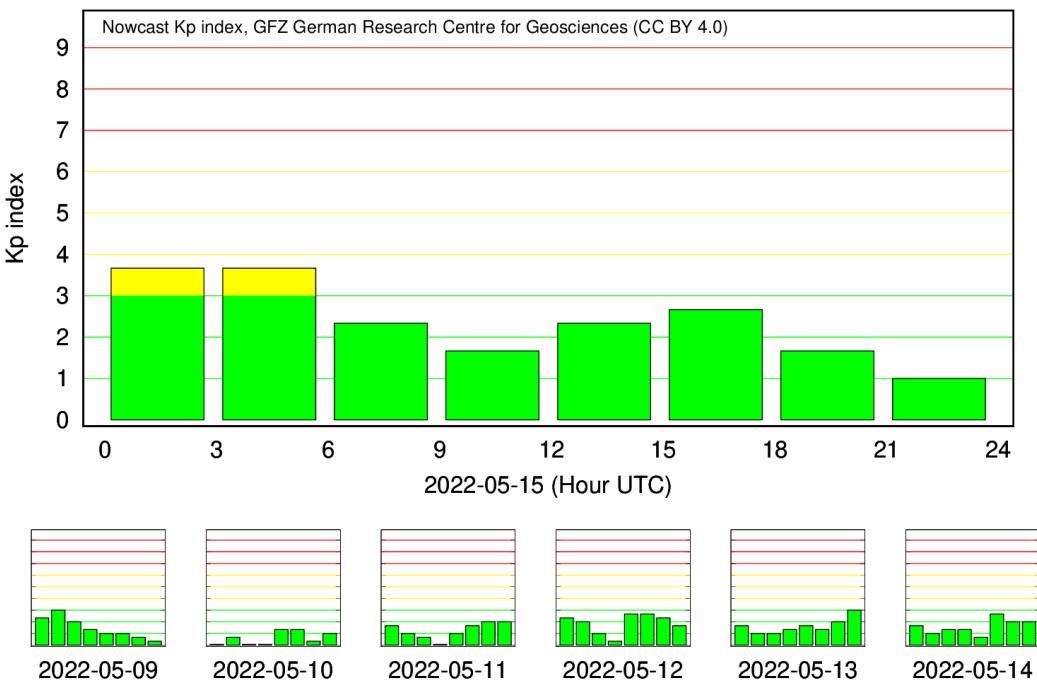
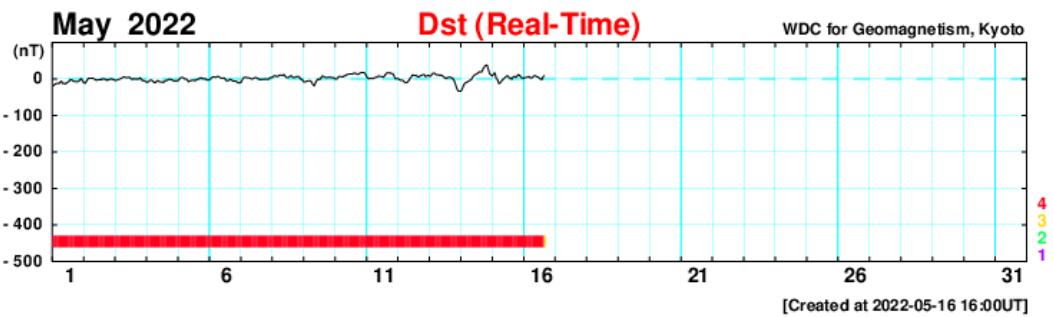
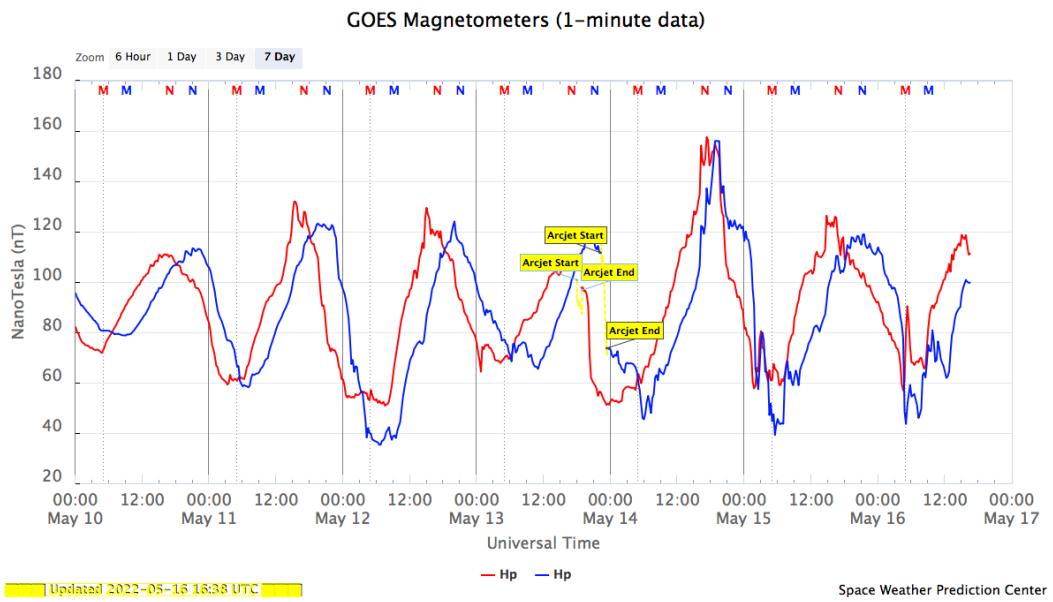
7 Geomagnetism

7.1 Responsible: Livia Riveiro Alves









- In the week of 05/10 to 05/16, the following events related to geomagnetic activity stand out:

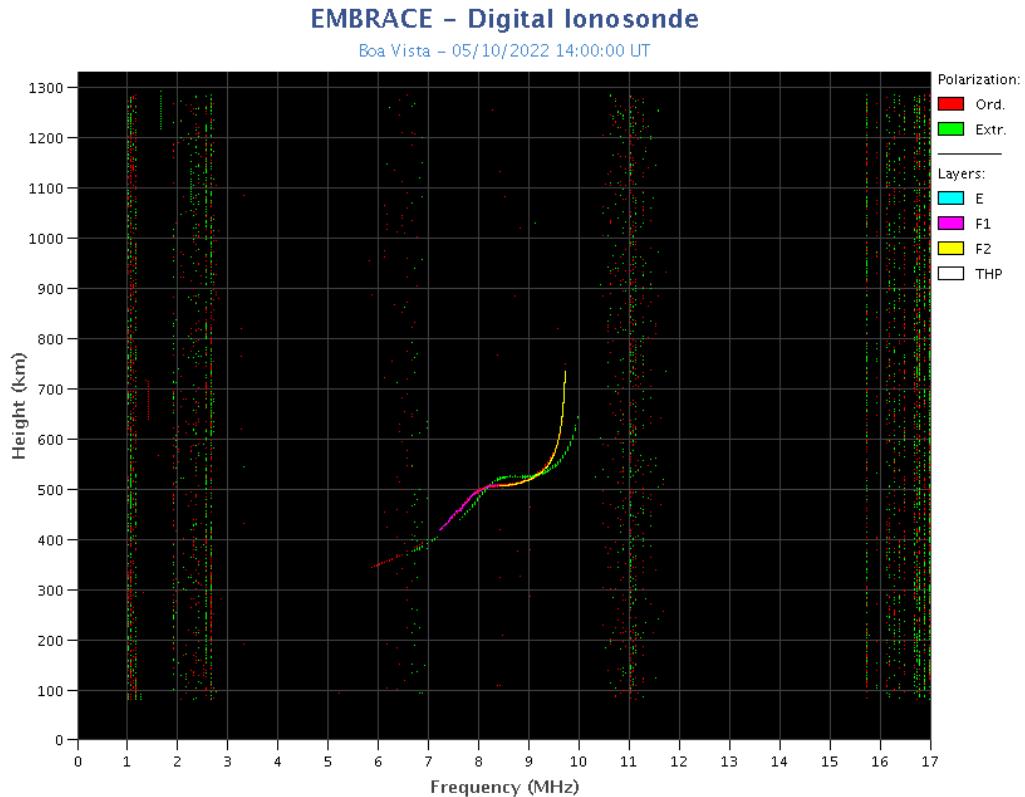
- Data from the Embrace magnetometer network showed instabilities throughout the period, with some events highlighted:
- The biggest disturbances in the H component were recorded on the 11th, 13th, and 14th of May.
- Geomagnetic activity was unstable throughout the AE index, with the Dst index oscillating around zero. Highest K_p of the week was 3+
- The auroral activity was slightly intensified on the 13th, 14th and 15th of May.

8 Ionosphere

8.1 Responsible: Laysa Resende

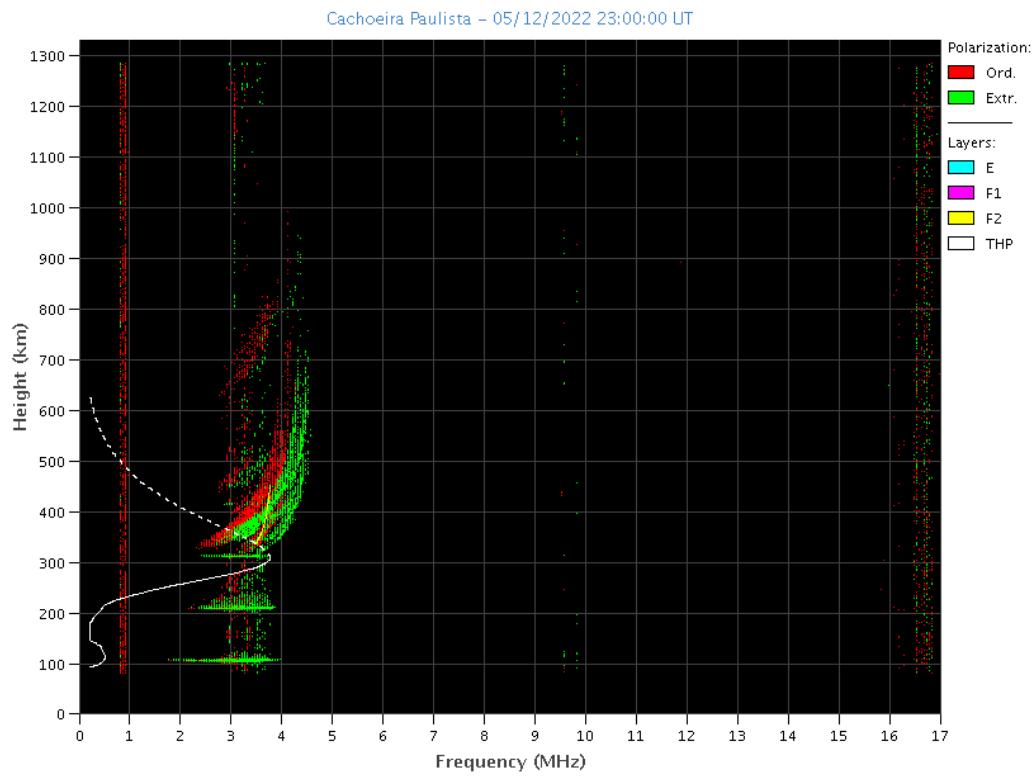
Boa Vista:

- There were spread F during all days in this week.
- The Es layers reached scale 3 on days 12 and 13.
- There was partial blackout on day 10.



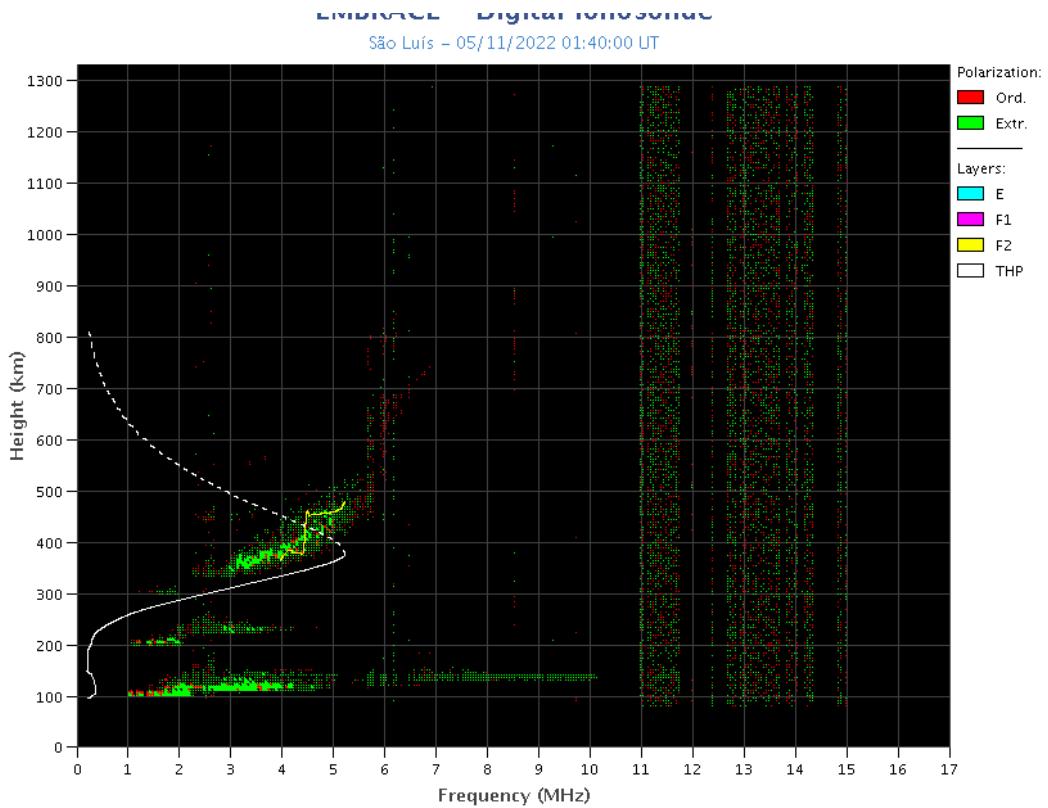
Cachoeira Paulista:

- There were spread F on day 12.
- The Es layers reached scale 3 on days 09, and 12.



São Luís:

- There were spread F during all days in this week.
- The Es layers reached scale 5 on day 11.
- There was partial blackout on day 10.



9 Scintillation

9.1 Responsible: Siomel Savio Odriozola

In this report on the S4 scintillation index, data from SLMA in São Luiz/MA, STSN in Sinop/MG, UFBA in Bahía/BA and SJCE in São José dos Campos/SP are presented. The S4 index tracks the presence of irregularities in the ionosphere having a spatial scale ~360 m. The four analyzed stations did not present relevant values of the S4 index throughout the week. Figure 1 shows the values for SLMA (top panel) and SJCE (bottom panel) stations

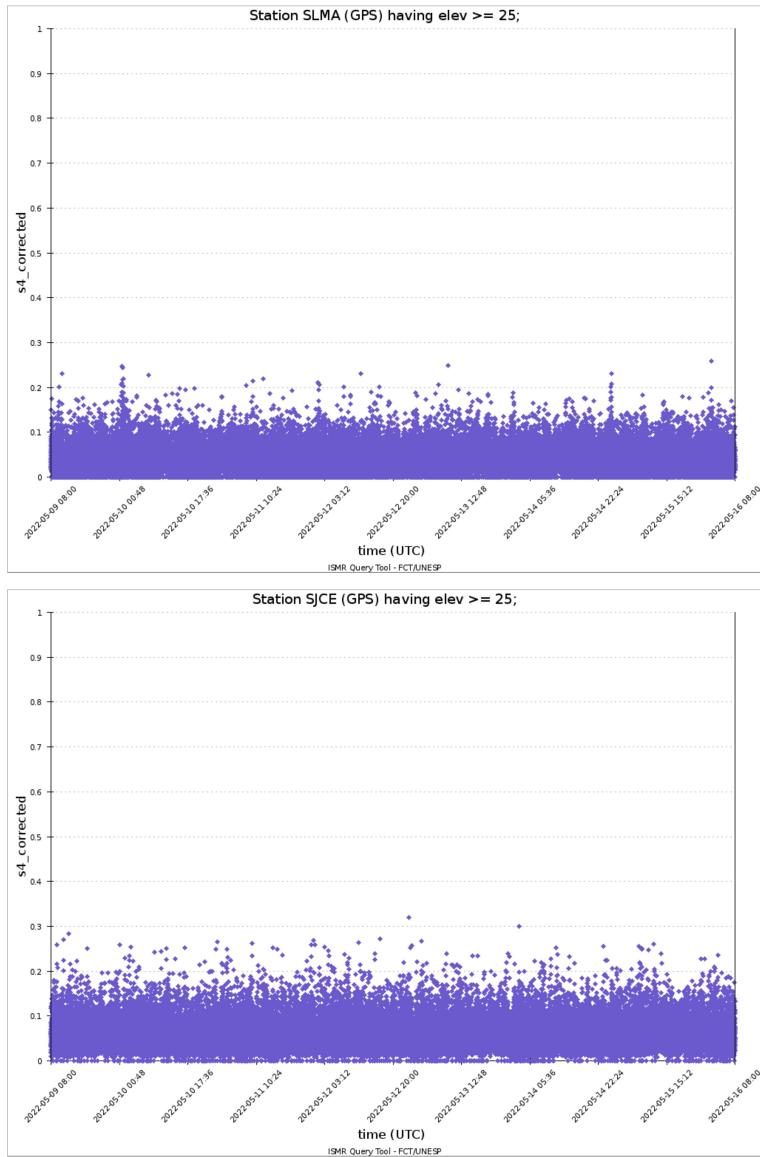


Figura 1: Valores do índice S4 para a constelação GPS para a estação SLMA (painel superior) e a estação SJCE (painel inferior) durante a semana 9—16/05 .

10 All-Sky Imager

10.1 Responsible: LUME

All-Sky Imager EPBs Observation || May 01 - May 07, 2022
Observações das EPBs por meio do imageador All-Sky -
|| 01 - 07 maio, 2022

Observatory Observatório	May 09 maio 09	May 10 maio 10	May 11 maio 11	May 12 maio 12	May 13 maio 13	May 14 maio 14	May 15 maio 15
CA	✓◐	✓◐	✓◐	✓●	✓●	✓●	✓●
BJL	✗	✗	✗	✗	✗	✗	✗
CP	✓◐	✗●	✓●	✓●	✓●	✓●	✗
SMS	✓◐	✗	✗	✗	✗	✗	✓
CA	São João do Cariri						
BJL	Bom Jesus da Lapa						
CP	Cachoeira Paulista						
SMS	São Martinho da Serra						
✓	Observation - Observação						
✗	No Observation - Sem Observação						
○	Clear sky - Céu limpo						
◐	Partly Cloudy - Parcialmente Nublado						
●	Cloudy - Nublado						
✳	Blur image - Desfocar Imagem						

- At the São João do Cariri observatory there was no observation due to technical problems.
- At the Bom de Jesus da Lapa observatory there was no observation due to technical problems.
- At the Cachoeira Paulista no plasma bubbles were observed during the entire week even though there were observations.
- Finally, at the observatory of São Martinho da Serra, no plasma bubble structures were observed during the entire week. 2

TEC

- No plasma bubbles were observed during the entire period. As bubble sea- sonality is at the end, plasma bubbles have small spatial dimensions and they are difficult to observe on TEC maps. 3