

The Largest Sample of Microlensing Events and the Structure of the Galactic Bulge from the OGLE-III Survey ¹

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ABSTRACT

We present and study the largest and the most comprehensive catalog of microlensing events ever constructed. The sample of standard microlensing events comprises 3718 unique events from years 2001–2009, with 1409 not detected before in real-time by the Early Warning System of the Optical Gravitational Lensing Experiment (OGLE). The search pipeline makes use of Machine Learning algorithms in order to help find rare phenomena among 150 million objects and derive the detection efficiency. Applications of the catalog can be numerous, from analyzing individual events to large statistical studies for the Galactic mass and kinematics distributions and planetary abundances.

We derive the maps of the mean Einstein ring crossing time of events spanning 31 sq. deg. toward of the Galactic Center and compare the observed distributions with the most recent models. We find good agreement within the observed region and we see the signature of the tilt of the bar in the microlensing data. However, the asymmetry of the mean time-scales seems to rise more steeply than predictions, indicating either a somewhat different orientation of the bar or a larger bar width. The map for the events with sources in the Galactic bulge shows a dependence of the mean time-scale on the Galactic latitude, signaling an increasing contribution from disk lenses closer to the plane, related with the height of the disk. Our data present a perfect set for comparing and enhancing new models of the central parts of the Milky Way and creating the 3D picture of the Galaxy.

1. Introduction

Galactic gravitational microlensing is an astrophysical phenomenon, which originates from the fact that in the curved space-time around each sufficiently massive body (e.g., stars) light travels along bent (hence converging) paths. The foundations of the theory lie in the General Theory of Relativity of Einstein (Einstein 1936) and the practical use within the scale of our own Galaxy were developed by Paczyński (Paczynski 1991) and Griest (Griest et al. 1991). This unique astrophysical tool has numerous and very interesting applications, because, unlike other astrophysical phenomena, microlensing is sensitive to the mass of the object passing in front of a background star. This means that, in principle, the lens can be completely dark and it will still cause a microlensing event as long as it aligns with a background star and the relative motion between the source, lens and the observer is large enough for the event to last within the human time-scale. During such an event the background star appears brighter, typically by a factor of few, occasionally reaching hundreds and very rarely thousands, in the case of a perfect alignment when the images form an Einstein ring.

The fact that the microlensing is sensitive to the mass of lensing objects turns this phenomenon into a powerful probe of the mass distribution in the Galaxy, allowing for studying the structure of the Milky Way. Kiraga & Paczynski (1994) defined the microlensing optical depth as a measure of the lensing probability toward the Galactic Centre and showed it as a convenient indicator of the total mass of the lensing populations along the line of sight. The first measurements of the optical depth obtained from the OGLE data (Udalski et al. 1994) and MACHO (Alcock et al. 1997) were significantly higher than expected from theoretical models of the Galaxy (e.g., Dwek et al. (1995), Han & Gould (1995), Binney et al. (2000), Freudreich (1998), Evans & Belokurov (2002)). The discrepancy was partially explained by Paczynski et al. (1994) who introduced a bar to the models, an elongated solidly revolving structure placed at

about a 20 deg angle toward the Sun. In subsequent years the measurements of the optical depth were limited to events in which the sources were the Red Clump giants (RC) from the Galactic Bulge (e.g., Afonso et al. 2003, Popowski et al. 2005). This trick significantly moved the observed values closer to the theoretical predictions, mainly because it mostly probed a single population of sources from the bulge and reduced the problem of blending (crowding) of stars. Blending is a demanding issue to deal with in the extremely crowded sky regions like the Galactic Centre. As shown in Smith et al. (2007), blending can not be completely neglected in the optical depth determinations, even in case of the lensing of the brightest Red Clump stars.

Nevertheless, despite the fact that there were many hundreds of events being detected in real-time every year by the OGLE (Udalski 2003) and MOA (Yock 1998) collaborations, the number of microlensing events used for the studies of the inner parts of the Galaxy were typically much smaller. For instance, the optical depth measured by the EROS group relied on 16 (Afonso et al. 2003) and 120 events (Hamadache et al. 2006), MACHO used 99 events (Alcock et al. 2000) and 42 RC events (Popowski et al. 2005). There were only 9 events used for computing the first ever value of the optical depth from the OGLE-I data (Udalski et al. 1994), then from the OGLE-II project 32 events were used(Sumi et al. 2006). The largest samples of standard single lens events used so far for the optical depth determination were 610 events from the OGLE-III years 2001-2004 (Wyrzykowski 2005) and 474 events found in MOA-II in years 2006-2007 (Sumi et al. 2013). Moreover, Calchi Novati et al. (2008) constrained the Initial Mass Function (IMF) based on just 42 events from MACHO and obtained the slope for Main Sequence stars of $\alpha_{\text{MS}} = 1.7$ and $\alpha_{\text{BD}} = 1.6$ for Brown Dwarfs populations in the Bulge, in agreement with most theoretical predictions. However, a detailed mass spectrum can be obtained when using significantly larger samples of microlensing events. Increasing the number of good quality and robust standard microlensing events is essential for more detailed comparison of the observations to the predictions, also over a range of Galactic coordinates (Mao 2012).

The Einstein ring crossing time, i.e., the time-

¹Based on observations obtained with the 1.3 m Warsaw telescope at the Las Campanas Observatory of the Carnegie Institution of Washington.

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scale of an event, is the only parameter from the standard microlensing model which has a physical meaning. Still, its value is a composition of multiple parameters:

$$t_E = \frac{\sqrt{\kappa M \pi_{\text{rel}}}}{\mu_{\text{rel}}}, \quad (1)$$

where $\kappa = \frac{4G}{c^2 \text{AU}} \approx 8.144 \frac{\text{mas}}{M_\odot}$, M is the mass of the lens, $\pi_{\text{rel}} = 1/D_d - 1/D_s$, D_d and D_s are the distances to the deflector (lens) and the source, respectively, and μ_{rel} is the relative proper motion between the source and the lens. Because of that degeneracy, the time-scales can only be studied statistically in large numbers, as all parameters involved in the t_E creation follow some distributions, which can be modeled for different populations within the Milky Way. As seen in Eq. 1, the strongest influence on the t_E value is by μ_{rel} , therefore the distributions and dispersions of proper motions of different Galactic populations will play the crucial role in interpreting the observed time-scales. Measured distributions of the time-scales of microlensing events have been used in the past as one of the ways of verifying different scenarios for the composition and kinematics of the inner parts of the Galaxy, e.g., Evans & Belokurov (2002), Bissantz & Gerhard (2002), Wood & Mao (2005), Calchi Novati et al. (2008), Kerins et al. (2009), Sumi et al. (2013).

In this paper we analyze the final and complete photometric data set gathered during the OGLE-III project in years 2001-2009. We search for high quality standard microlensing events using the optimized search criteria, supported by the Machine Learning method (Random Forest classifier), and use them to investigate the structure of the inner parts of the Milky Way.

The paper is organized as follows. First we describe the OGLE-III data. In Sec. 3 we describe in detail the procedure of the search for events, divided into several sub-steps. Then, we report on the results of the search and present the catalog, compare it to the EWS and discuss the detection efficiency of events. This is followed by the discussion of the results, in which we describe the properties of the events as a whole, derive the distributions of the time-scales of events over the sky and compare it to the Galaxy models. We summarize and conclude the paper in Sec. 6.

2. Data

The data used in this work was the photometry of 150 million objects toward over 31 sq. deg. of the Galactic Bulge observed on almost 74,000 frames, i.e., about 11,000 billion data points. We selected 91 fields out of all 177 ever observed by the Optical Gravitational Lensing Experiment (OGLE) Udalski et al. (2008) in its third phase from July 2001 until May 2009, which had at least 250 observations. During the OGLE-III phase the Warsaw Telescope, located at the Las Campanas Observatory, Chile, operated by the Carnegie Institution of Washington (now Carnegie Institution for Science), was equipped with a mosaic CCD camera with eight $2\text{k} \times 4\text{k}$ pixels chips covering in total 0.34 square degrees. The typical exposure time in the fields toward the Galactic Bulge was 120s allowing reaching down to nearly 21 mag in the Johnson-Cousin I -band, the filter in which vast majority of the observations were carried out.

Figure 1 shows the positions of OGLE-III fields toward the Galactic Bulge. Each field was observed on average once per three nights, however, from 2005 season the observing strategy changed slightly so that the most dense fields at about $b \sim 2$ deg were observed with higher cadence than the rest, typically once per night. The number of observations collected in the I -band over 8 years per field varied from 251 for BLG344 to 2540 for one of the central fields, BLG102.

There were typically between 1 and 35 V -band data points available for the fields investigated here. Those data were only used to obtain the averaged color of the baseline of the objects. The calibrated color was taken from the OGLE-III Bulge photometric maps (Szymański et al. 2011).

During the operation of the OGLE-III the data collected each night were reduced on-the-fly with the state-of-the-art difference imaging technique (DIA, Woźniak 2000) and preliminary photometry was produced within couple of hours. This was the basis for the Early Warning System (EWS, Udalski 2003), which was designed to look for new microlensing events in the real-time. As a result of seven years of functioning (2002-2009), the OGLE-III EWS reported about 4000 candidates for microlensing events.

Toward the end of the OGLE-III phase, the entire observational material was re-reduced again

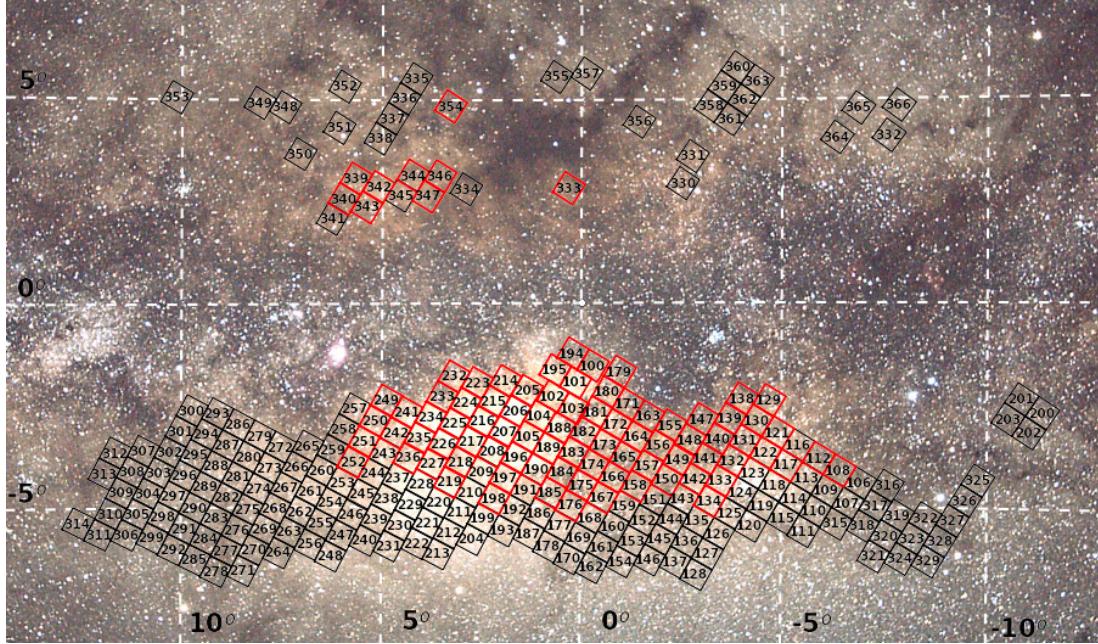


Fig. 1.— Map of the OGLE-III bulge fields in galactic coordinates. Red squares mark fields which were used in the search for microlensing events as they had at least 250 observations over 2001-2009 period. Remaining fields were observed less frequently. Each square contains the number of the field. The background image was taken by Krzysztof Ulaczyk.

with DIA using a new and better set of images for the composite template images (Udalski et al. 2008), yielding significantly better quality output photometry. Figure 2 compares the quality of the photometry of a microlensing event OGLE-2005-BLG-069 reported by the EWS, obtained in the original and new reductions. The quality improvement is primarily due to somewhat better resolution of the new template images.

We searched for microlensing events in this final and complete dataset of the OGLE-III observations of the Galactic bulge.

3. Search procedure

Extraction of rare light curves like microlensing events from vast databases is not an easy task. In the past, the data comprised a relatively small number of objects, making it feasible to visually inspect a small few thousands of light curves, after applying some basic selection criteria. Nowadays, and in case of the OGLE-III data, it is necessary to seek help from automated methods of microlensing events selection. Moreover, a fully au-

tomatized search pipeline allows for obtaining the detection efficiency of events depending on their parameters. Machine Learning algorithms were already successfully used in time-domain astronomy, e.g., Belokurov et al. (2003), Wyrzykowski et al. (2003), Debosscher et al. (2007), Richards et al. (2011), Pawlak et al. (2013), for automatization of the discovery and classification in large data sets.

In our previous searches for microlensing events in the OGLE data in the LMC and SMC fields (Wyrzykowski et al. 2009, 2010, 2011a, 2011b), as well as in the preliminary search in the OGLE-III Bulge data (Wyrzykowski 2005), we relied on a number of cuts applied to a small number of computed features, primarily microlensing model parameters. In those searches the most powerful discriminator was typically the goodness of fit of the model, however, it worked fine only for the events with well observed light curves, with corrected photometric error-bars and with well understood noise. Therefore, such a cut could remove plausible standard microlensing events, the photometry of which could have been affected by

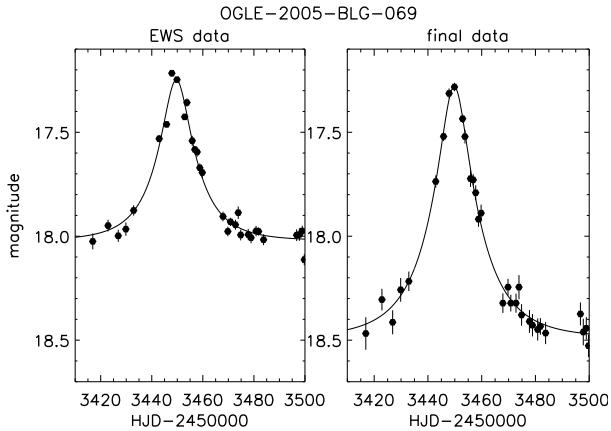


Fig. 2.— Comparison of the EWS data for a microlensing event detected in real-time and the reduced data for the same event along with the best fitted microlensing models (solid lines). The scatter in the light curve decreased around the peak and the scaling factor have changed significantly in the new data. The measured time-scale changed from 16.3 to 19.1 days, for old and new, respectively, and the blending parameter changed from 0.48 to 0.71, suggesting that more stars were resolved on the new reference image.

some observational or instrumental problems, generally very common in such crowded areas as the Galactic bulge, e.g., blending with variable stars. As shown in Wyrzykowski et al. (2006) events with variability in the baseline can be useful in constraining at least some of the physical parameters of typically severely degenerated microlensing models.

Here, we developed and applied a new method for selecting microlensing events among all stars in the database. The database consisted of nearly a quarter of a billion of objects, therefore it required a more optimized and revised approach. In order to deal with such vast data set, we relied on Machine Learning (ML) techniques, in particular, on the Random Forest (RF) classifier (Breiman 2001), which finds its own the most successful multi-dimensional selection criteria. Like most of the ML methods, the RF requires all objects to be described in a homogenous way with various properties or features. The advantage of the RF is that it conveniently selects which features are the most meaningful and carry the most information needed to correctly perform the classification. Therefore, the risk that too many features will blur the classification outcome is minimized, which allows for preparing a large number of various features for each object.

3.1. Preparatory steps

Each individual light curve in the database was first pre-processed. This included outlier removal: single data points outlying by more than 3σ from the mean, with the two nearest neighbors lying within 1σ . Then, the photometric error-bars were corrected, following the method described in Wyrzykowski et al. (2009) and Skowron (2009), which returned a correction coefficients as a function of observed magnitude, derived from non-variable stars. Once the error-bars were corrected, we were able to filter out all non-variable light curves because our main goal was to find singular brightening episodes in each light curve. As a variability indicator for a light curve we used ratio of the standard deviation to the mean error, $\sigma_{\text{reletot}} > 1.05$. There were 105 million objects left from the original 150 million. Please note, that not all of them were genuinely variable, the list included also many artifacts, e.g., differential refraction effects or seeing-dependent variability caused

by nearby bright and saturated stars.

In the next step, we were looking for objects exhibiting a positive flux increase over some period of time with respect to the rest of the light curve. For each light curve we applied a running window search for peaks, following the concept used in Sumi et al. (2006) and Wyrzykowski et al. (2009). In this method a window sized half of the time-span of the light curve was run over an entire light curve and for each data point i a value of sigma was computed:

$$\sigma_i = \frac{I_{med,w} - I_i}{\sqrt{\Delta I_i^2 + \sigma_w^2}}, \quad (2)$$

where I_i is the magnitude of the point i and ΔI_i is its error bar, while $I_{med,w}$ and σ_w are the median and the *rms* in the outer window, respectively.

Before computation we masked out all observations taken with airmass > 2 (in most cases they were outlying due to differential refraction effects) and we averaged data points taken within the same night. The OGLE-III sampling was typically one observation per 2-3 nights, however, occasionally there were many observations taken within a single night, when the telescope switched to the follow-up mode to cover some interesting microlensing anomalies more densely.

From the computation of running windows for each variable light curve we obtained a set of parameters (features):

1. **N_{peaks}**, number of detected separate peaks, each defined as a series of points with $\sigma_i > 2.0$
2. **maxsigma**, maximum value of σ_i for the most pronounced peak (the highest σ_i)
3. **N_{seq}**, number of sequential data points in the most pronounced peak, all being above a threshold of $\sigma_{thresh} = 2.0$
4. **peaksum**, sum of σ_i , within the most pronounced peak, i.e., sum of N_{seq} sigmas.
5. **peaksum7**, sum of σ_i , of the maximum and 3 adjacent points from each side of the maximum (7 points)
6. σ_{relB} , variability indicator for outside window (window B) for the most pronounced peak

7. σ_{reltot} , variability indicator for the entire light curve

At this stage, we narrowed the sample only to those objects, which had at least 4 subsequent points in a detected peak and we were left with about 8.5 million objects. Light curves of those were then fitted¹ with the standard Paczyński (1996) microlensing model, (i.e., a point-source – point-lens microlensing event) which is described as:

$$I = I_0 - 2.5 \log [f_S A + (1 - f_S)], \quad (3)$$

where

$$A = \frac{u^2 + 2}{u\sqrt{u^2 + 4}} \quad \text{and} \quad u = \sqrt{u_0^2 + \frac{(t - t_0)^2}{t_E^2}}. \quad (4)$$

The fitted parameters are: t_0 – the time of the maximum of the peak, t_E – the Einstein radius crossing time (event's time-scale), u_0 – the event's impact parameter, I_0 – the baseline magnitude in the *I*-band and f_S – the blending fraction (ratio of lensed source flux to total blends' flux in the *I*-band). The fits were performed in two ways, namely with blending parameter fixed $f_S = 1$, i.e., with no blending, and with f_S being free. For clarity, parameters of the non-blended (4 parameters) model are given the subscript $\mu 4$. Because the standard microlensing model is symmetrical for parameter u_0 , we only used its positive value in further analysis.

Based on the visual inspection of a selection of variable objects from test fields BLG100.1 and BLG206.1 we derived “common-sense cuts” to the set of variable objects. A cut on $N_{seq} \geq 4$ was imposed to remove a vast majority of short duration artifacts, however, it also removed poorly sampled short time-scale events, typically with time-scales shorter than 3 days. A dedicated search for very short duration events will be presented elsewhere. Other cuts required reasonable microlensing fit parameters, i.e. $|u_0| \leq 2$, $2150 < t_0 - 2450000 < 5000$, $1 < |t_E \mu 4| < 400$ days, $1 < |t_E| < 400$ days, $f_S < 1.4$. Limits on the time of the maximum ensured we only

¹fitting was performed using CERN's MINUIT package, <http://wwwasdoc.web.cern.ch/wwwasdoc/minuit/minmain.html>

dealt with microlensing-like episodes of brightening, where the peak time t_0 remained within the available data span. The cut at 400 days in time-scale is induced by the fact of using the running-window method to find brightening episodes over light curves, in which the window size was equal to the half of the total span of the data. Events with time-scale of 400 days may in practice span for more than 1600 days, depending on the u_0 , therefore, our detection efficiency for those will be smaller. Also, there is a large number of false events with alleged time-scale above 400 days, which turn out to be low amplitude slowly varying variable stars. The blending parameter, f_s , limited at 1.4 allowed for some amount of so called “negative” blending (Smith et al. 2007). The border value was computed as a maximum value expected for crowded bulge data assuming background fluctuations due to unresolved stars at about 21 mag.

This pre-filtering narrowed our sample to 194,000 objects.

3.2. Ghost events filtering

It is very common in the crowded stellar fields that the variability of one of the stars can affect the photometry of nearby objects, often called “ghosts” or “children”. The radius of influence depends on the seeing conditions as well as on the amplitude of the variability, therefore some of strongly amplified microlensing events can cause an effect that many nearby stars undergo a very similar brightening episode. Figure 3 shows an example of such effect in the OGLE-III data in case of a microlensing event with an amplitude of many magnitudes.

Production of “ghost” events can significantly disturb the statistics and the quality of detected microlensing events. Therefore, it is essential to recognize which star is actually linked to the genuine microlensing event and filter out all the neighboring ghosts. In the case shown in Fig. 3 it is obvious which one is the main event, given much better quality of the light curve. However, not in every situation the conclusion is that simple, as in dense bulge areas there could be stars very close to each other with similar brightness, thus the wings of PSF of the additional flux may spread equally to all nearby objects and produce light curves of similar quality.

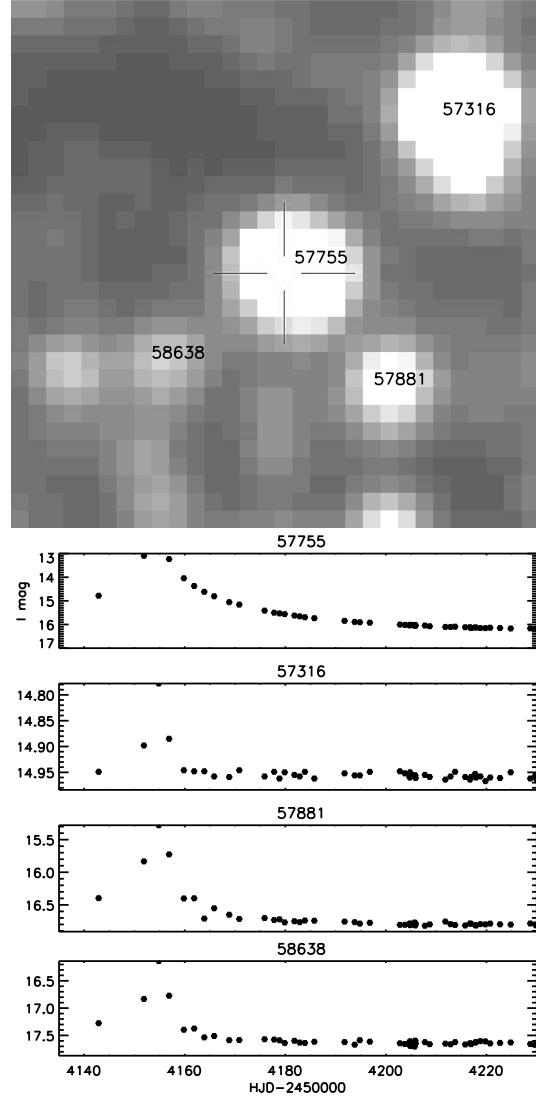


Fig. 3.— Example of “ghost” microlensing events present in the database. The microlensing event occurred on star number 57755, marked with the cross on the chart sized 30×30 pixels ($\sim 8 \times 8$ arc-seconds). Its light curve is shown in the top panel of the lower figure. At least three other database objects had their photometry affected by the extra flux from the brightening event. Their light curves resemble that of the “real” event, however the quality deteriorates with the distance from the source event. Difference imaging technique used for data reductions in the OGLE-III allows for a very accurate determination of the centroid of the microlensed flux, hence we are able to recognize which of the objects should be catalogued as a microlensing event.

Here we employed one of the very useful features of the difference imaging data reductions. Within the DIA pipeline there is a stage at which the subtracted images are obtained and all variable sources can be identified on them with their surplus flux and exact position measured. Then, in order to construct a temporal light curve, the difference flux is added to the objects from the reference (template) image. For each of 194,000 variable objects we found, we searched for the nearest differential object and read its centroid position for the frame taken at the maximum brightness of the object. Then if the offset from the template position was smaller than 2 pix (0.52 arc sec) we assumed that given reference object was the one with the brightening event. About 25% of variable objects survived that selection and we were left only with 48,112 light curves for further classification.

3.3. Features for classification

For all the light curves remaining after the steps described above, we derived a set of features, used later for Random Forest classification. First seven features were listed above and came from the running-window analysis of a light curve.

Next set of features came from the microlensing fits to each light curve:

8. $t_{E\mu 4}$, time scale of the event fitted with blending fixed to 1, describing an overall longevity of the event,
9. $\log u_{0\mu 4}$, logarithm of the impact parameter, sensitive to the actual height of the event,
10. $\chi_{\mu 4}^2$, goodness of fit of the 4-parameter model (no blending),
11. $\chi_{\mu 4}^2/\text{dof}$, reduced goodness of fit of the 4-parameter model,
12. t_E , time scale of the event with free blending,
13. $\log u_0$, logarithm of the impact parameter in 5-parameter model,
14. $\log f_s$, logarithm of the blending parameter,
15. χ^2 , goodness of fit of the 5-parameter model,
16. χ^2/dof , reduced goodness of fit of the 5-parameter model.

For each light curve we also fitted a constant line to the out-of-event data, defined as data outside of $t_{0\mu 4} \pm 5t_{E\mu 4}$ region. Such fit returned a mean magnitude, *rms* scatter (sigma) and its goodness of fit. To the set of features we added the following:

17. σ_{out}
18. χ_{out}^2
19. $\chi_{\text{out}}^2/\text{dof}$

The following features were then derived, using combinations of microlensing model parameters. They were reflecting any dramatic changes between the 4- and 5-parameter models, which often signals a non-microlensing origin of the event.

20. $\Delta t_0 = \log |(t_{0\mu 4} - t_0)|$,
21. $\Delta u_0 = \log |(u_{0\mu 4} - u_0)|$,
22. $\Delta I_0 = \log |(I_{0\mu 4} - I_0)|$,
23. $\Delta \chi^2 = \chi_{\mu 4}^2 - \chi^2$.

Microlensing modeling which gave the above mentioned parameters for each event was performed over the entire light curve. On top of that, we also determined the χ^2 of the 5-parameter microlensing model solely for the peak and the baseline data. Such features carried information on details of how good the microlensing model recovers the data and were sensitive to small anomalies, to which a global model was much less sensitive. Here the peak was defined as $t_0 \pm 3t_E$ and contained N_{peak} data points. The rest of the data (baseline) contained N_{base} data points.

24. $\chi_{\text{peak}}^2/N_{\text{peak}}$,
25. $\chi_{\text{base}}^2/N_{\text{base}}$,
26. $\chi_{\text{peak}}^2/N_{\text{peak}}/\chi_{\text{base}}^2/N_{\text{base}}$, ratio of both χ^2 values.

The final feature to include was the $V - I$ color of the blended source, obtained from the OGLE-III photometric maps (Szymański et al. 2011):

27. $V - I$, color of the blended source.

This set of features was derived for all light curves, which passed the basic filters described above.

3.4. Random Forest classifier

For the classification of events described with 27 features derived above we decided to use the Random Forest classifier from its Java implementation in Weka package (version 3.6.5)² developed at the University of Waikato in New Zealand. For training the classifier we manually selected candidate microlensing events from the pre-filtered sample of objects from the fields with various sampling properties: BLG100 (about 2400 observations), BLG180 (about 1400 observations) and BLG206 (about 1300 observations). During the visual classification we divided the events into three classes: *ULENS* for standard microlensing events, *EXOTIC* for non-standard microlensing events (e.g., binary lens events, events with parallax effect) and *OTHER* for all the remaining light curves. Among the latter there were various types of outbursting variables like Dwarf Novae or Be-type stars, but also numerous artificial events or other types of long-term variable stars (e.g., Soszyński et al. 2013). The visual selection of the events for the training set was performed independently by a few authors (LW, AER and MP) in order to cross-validate the results. The training set was also supplemented by 30 manually selected standard events reported by the EWS.

Our final training set had the following composition: *ULENS*: 977, *EXOTIC*: 53, *OTHER*: 2135. The number of exotic examples was significantly smaller than the remaining classes, however, in this study we did not intend to pick the exotic events with very high effectiveness. The main reason for adding that middle class was to differentiate those type of events (typically with high signal-to-noise and large values of χ^2) from standard events and other types of variables and artifacts. The 10-fold cross-validation test of the RF classifier, setup to use 15 best features and 15 trees, resulted in 96.7% correctly classified instances.

Because our primary goal was to provide a sample of standard microlensing events with the lowest contamination, we added a second stage of classification to our search pipeline. It was designed to analyze only those events which were marked as *ULENS* during the first stage of the classification. The training set comprised of only two

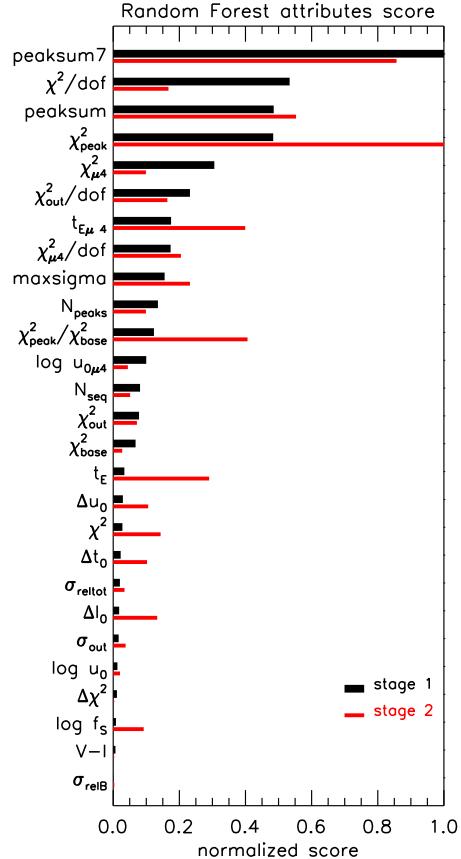


Fig. 4.— Score of each feature returned by the Random Forest for stage 1 and stage 2 classification. The scores for stage 1 are normalized to the score of **peaksum7** feature, whereas scores for stage 2 are normalized to the score of χ^2/dof .

classes: *GOOD* (706 instances) and *BAD* (271 instances), and was composed after the visual inspection of test runs of the stage 1 classifier on a fraction of the input list. For the verification we used field BLG104, in which we manually identified 60 genuine events, after removal of the ghost events. Running both stages of the RF classifier on the entire data of field BLG104 returned 56 events, with the remaining 4 being of lowest quality. This shows a very high efficiency of the Random Forest classifier with simultaneous high purity of the output.

Fig. 4 shows the normalized score for each of the features used in the RF classifier. The features are sorted from the most useful during the

²<http://www.cs.waikato.ac.nz/ml/weka/>

first stage of classification at the top to the least useful at the bottom. The most powerful is the `peaksum7` feature, which describes how strong is the event in the light curve. Not surprisingly the second most useful feature is χ^2/dof , the goodness of fit of the microlensing model. Also shown in this figure are the scores of the same features during the second stage. Because the classification here is more detailed than at stage 1, slightly different features play an important role in the classification. For example, the most important one at stage 2 is the goodness of fit at the peak as well as its ratio to χ^2_{base} . It shows that hierarchical division of the classification was justified as the two steps functioned on different grounds.

4. Catalog

The classification procedure employing the Random Forest classifier returned exactly 3700 events. We then cross-matched the catalog with itself checking within 3 arc seconds radius and requiring t_0 to be within 10 days, and found 95 events as duplicates due to overlapping OGLE-III fields. We did not find any pair of events located within 3 arc seconds and which occurred at significantly different moments of time, which could have been caused by the same lensing object or a wide binary lens (e.g., Skowron et al. 2009). From each pair of duplicate events we selected the event with better quality of the light curve i.e., the one with higher `maxsigma`, but the information on the multiplicity is stored in the table of events. Therefore, the catalog contains 3560 unique standard microlensing events, dubbed the “class A” sample. The density of events distributed over the sky in the Galactic coordinates in each OGLE-III field is shown in Figure 5. Table 1 lists the class A events, coordinates, OGLE-III field and star number identification, cross-match with the EWS and indication on any duplicated entries. The events will be labeled OGLE3-ULENS-*nnnn*, where *nnnn* is the number of the event, with the leading zeroes.

Each event from the catalog was fitted with the microlensing model using the Markov Chain Monte Carlo (MCMC) method³. The priors on all model parameters were assumed uniform around

³we used Python module `pymc` from <https://pypi.python.org/pypi/pymc/>

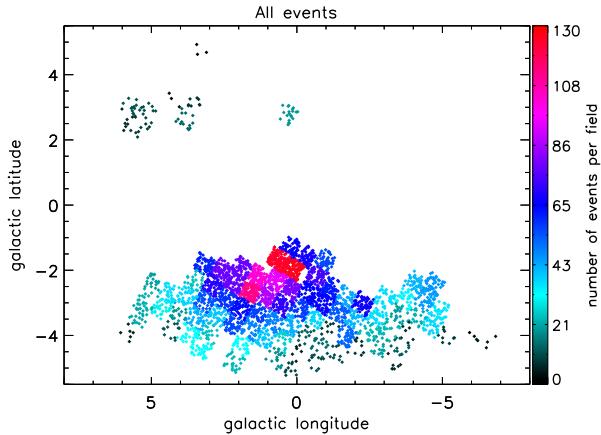


Fig. 5.— Location of all 3560 class A standard events found in the OGLE-III Bulge data from years 2001-2009. The color of each dot indicates the number of events found in the OGLE-III field containing that event. Each field had 8 independent CCD chips and covered 0.34 sq.deg.

the initial values, taken from the standard χ^2 modeling with MINUIT code.

Therefore, for each event we obtained the posteriori distributions for each parameter and the median and one sigma asymmetric errors (84.1345 and 15.8655 percentiles). Example result of the MCMC fitting to an event along with the distributions and relations for t_E , f_S and u_0 are shown in Figure 6. Such figures are available for all 3560 class A events in the on-line version of the paper, allowing for considering the degenerations of the model parameters in each of the events. Distributions of all fitted microlensing parameters for the class A sample are shown in Fig. 7. Table 2 shows all the standard Paczyński microlensing model parameters and their errors found in the MCMC models of the events from class A.

The full machine-readable catalog, the light curves of all events and their MCMC models are available in the on-line version of the paper and from the OGLE internet archive:

<http://ogle.astroww.edu.pl>

The light curves published there were reprocessed taking the exact position of the event on the DIA image into account. That assures their best possible photometric accuracy. The photometric data were not averaged nightly, however, some cleaning was applied, similarly to the proce-

TABLE 1
OGLE-III STANDARD EVENTS OF CLASS A.

ID OGLE3-ULENS-	RA _{J2000} [h:m:s]	Dec _{J2000} [deg:m:s]	field	starno	EWS id	Duplicate
0001	17:51:44.00	-30:17:20.7	BLG100.1	27898	2007-BLG-258	-
0002	17:52:21.54	-30:16:19.0	BLG100.1	47983	-	-
0003	17:51:01.18	-30:15:28.9	BLG100.1	55988	2007-BLG-459	-
0004	17:51:11.06	-30:14:20.6	BLG100.1	57414	-	-
0005	17:51:43.85	-30:14:43.3	BLG100.1	89325	-	-
0006	17:52:13.46	-30:13:22.9	BLG100.1	101269	2005-BLG-404	-
0007	17:52:08.57	-30:14:56.3	BLG100.1	102672	-	-
0008	17:52:19.37	-30:13:13.9	BLG100.1	106173	2005-BLG-294	-
0009	17:51:16.18	-30:12:31.9	BLG100.1	119596	2002-BLG-001	-
0010	17:51:16.71	-30:11:10.7	BLG100.1	124027	-	-
0011	17:52:13.40	-30:12:32.8	BLG100.1	160378	2004-BLG-435	-
0012	17:51:17.94	-30:09:12.7	BLG100.1	180707	2007-BLG-602	-
0013	17:51:29.14	-30:09:13.4	BLG100.1	186215	2007-BLG-532	-
0014	17:51:34.36	-30:10:07.3	BLG100.1	192835	2004-BLG-484	-
0015	17:51:47.89	-30:09:17.7	BLG100.1	198644	2005-BLG-251	-
0016	17:52:03.54	-30:10:42.4	BLG100.1	206957	-	-
0017	17:52:08.26	-30:09:44.1	BLG100.1	214894	-	-
0018	17:51:12.37	-30:05:04.1	BLG100.2	70077	2004-BLG-153	-
0019	17:51:43.83	-30:04:59.5	BLG100.2	91698	2004-BLG-229	-
0020	17:52:03.51	-30:05:32.6	BLG100.2	97813	-	-
0021	17:51:01.97	-30:02:00.2	BLG100.2	119749	2004-BLG-115	-
0022	17:52:09.43	-30:03:01.3	BLG100.2	164216	2008-BLG-303	-
0023	17:52:13.07	-30:03:30.3	BLG100.2	164721	-	-
0024	17:51:28.62	-30:01:00.2	BLG100.2	188384	-	-
0025	17:51:48.79	-30:00:15.9	BLG100.2	201309	2004-BLG-192	-
0026	17:51:05.19	-29:58:58.9	BLG100.3	1858	2006-BLG-097	-
0027	17:51:35.90	-29:59:08.8	BLG100.3	21115	-	-
0028	17:51:20.69	-29:55:52.1	BLG100.3	63514	2004-BLG-167	-
0029	17:51:34.32	-29:55:19.7	BLG100.3	74159	2002-BLG-193	-
0030	17:52:06.70	-29:55:16.2	BLG100.3	98639	2004-BLG-028	-
0031	17:51:29.95	-29:54:19.7	BLG100.3	123110	2007-BLG-250	-
0032	17:51:25.78	-29:53:50.6	BLG100.3	123957	2004-BLG-034	-
0033	17:52:01.31	-29:53:56.2	BLG100.3	144687	2004-BLG-281	-
0034	17:52:00.25	-29:54:39.7	BLG100.3	147281	2007-BLG-166	-
0035	17:51:05.26	-29:52:11.6	BLG100.3	167671	2006-BLG-018	-
0036	17:52:14.55	-29:50:51.9	BLG100.3	210689	2007-BLG-219	-
0037	17:51:32.62	-29:48:38.2	BLG100.4	17039	2004-BLG-135	-
0038	17:51:38.28	-29:47:16.3	BLG100.4	73670	2005-BLG-172	-
0039	17:52:16.62	-29:44:25.7	BLG100.4	159146	2006-BLG-098	-
0040	17:51:42.83	-29:42:44.0	BLG100.4	187869	2008-BLG-454	-
0041	17:52:15.81	-29:42:16.4	BLG100.4	206107	2006-BLG-019	-
0042	17:50:48.09	-29:49:30.3	BLG100.5	29000	-	-
0043	17:49:38.48	-29:55:46.5	BLG100.6	37372	-	-
0044	17:49:47.24	-29:56:42.2	BLG100.6	38608	2004-BLG-131	-
0045	17:50:33.56	-29:56:15.4	BLG100.6	57621	2008-BLG-144	-
0046	17:49:49.54	-29:54:05.4	BLG100.6	79092	2009-BLG-093	-
0047	17:50:13.79	-29:52:42.1	BLG100.6	87913	2004-BLG-301	-
0048	17:50:09.90	-29:52:59.0	BLG100.6	91567	2007-BLG-440	-
0049	17:50:25.92	-29:52:48.4	BLG100.6	93188	2008-BLG-265	-
....
3560	17:35:05.50	-23:24:39.9	BLG354.4	131	2002-BLG-192	-

NOTE.—Full table in the machine-readable format is available in the on-line version of the paper.

TABLE 2
MCMC MICROLENSING PARAMETERS FOR 3560 STANDARD OGLE-III MICROLENSING EVENTS OF CLASS A.

ID	field,star	t_0	HJD-2450000 [days]	t_E	a_0	f_S	I_0	χ^2_{mag}	N_dof	V
0001	BLG100.1.27898	4247.209697 \pm 0.003005	2.584243 \pm 0.120834	0.079670 \pm 0.010063	0.945561 \pm 0.063004	17.81770 \pm 0.000728	5321.163243	2384	20.932 \pm 0.101	
0002	BLG100.1.47983	3276.397031 \pm 0.004192	53.835570 \pm 0.005755	0.071218 \pm 0.005854	0.027828 \pm 0.000234	16.76628 \pm 0.000216	3566.665557	2473	18.257 \pm 0.018	
0003	BLG100.1.55088	4327.089713 \pm 0.004050	0.0227387	14.708127 \pm 1.713289	0.664801 \pm 0.161500	0.761287 \pm 0.002213	18.817654 \pm 0.002285	1934.388674	1515	99.999 \pm 9.999
0004	BLG100.1.57414	4230.125995 \pm 0.006646	6.285773 \pm 1.779811	0.148585 \pm 0.042069	0.464988 \pm 0.140634	18.801048 \pm 0.001737	3311.499669	2379	21.242 \pm 0.260	
0005	BLG100.1.89325	3577.831703 \pm 0.0015767	2.244371 \pm 0.416897	0.049491 \pm 0.010467	0.7255830 \pm 0.053156	19.992170 \pm 0.004028	2689.08779	2382	99.999 \pm 9.999	
0006	BLG100.1.101269	10.286849 \pm 0.002515	0.03852630	0.01770072 \pm 0.000352	0.557021 \pm 0.053156	18.879324 \pm 0.001428	4225.67130	2388	99.999 \pm 9.999	
0007	BLG100.1.102672	3466.212591 \pm 0.001977	12.216754 \pm 0.417556	0.406672 \pm 0.047053	0.557021 \pm 0.053156	19.377784 \pm 0.001780	3653.140716	2387	99.999 \pm 9.999	
0008	BLG100.1.106173	8.023514 \pm 0.000113	0.0235072 \pm 0.000113	0.517530 \pm 0.0476509	0.4071634 \pm 0.043456	17.19198 \pm 0.000447	2888.139368	2374	20.001 \pm 0.005	
0009	BLG100.1.119596	66.802986 \pm 0.152678	66.802986 \pm 0.149083	0.365573 \pm 0.011041	0.410459 \pm 0.012626	15.495008 \pm 0.000145	5805.342886	2388	19.131 \pm 0.031	
0010	BLG100.1.1240277	3120.436704 \pm 0.0021928	12.800806 \pm 0.249928	0.320332 \pm 0.02526	1.037836 \pm 0.019154	19.804619 \pm 0.001470	2921.226040	2382	99.999 \pm 9.999	
0011	BLG100.1.160378	3213.992068 \pm 0.000311	24.7781573 \pm 0.253433	0.094060 \pm 0.027038	0.488810 \pm 0.019159	18.978050 \pm 0.0011919	3041.733473	2388	99.999 \pm 9.999	
0012	BLG100.1.180707	4414.837567 \pm 0.078817	43.750745 \pm 3.670396	0.528626 \pm 0.124207	0.875749 \pm 0.087044	18.869671 \pm 0.001020	2197.703384	1982	99.999 \pm 9.999	
0013	BLG100.1.186215	4350.821796 \pm 0.07292	9.141864 \pm 0.175292	0.166307 \pm 0.077126	0.590168 \pm 0.259562	19.062267 \pm 0.001881	2858.225140	2085	99.999 \pm 9.999	
0014	BLG100.1.192835	3230.394075 \pm 0.00133	25.956916 \pm 0.2920630	0.110588 \pm 0.065377	0.7191638 \pm 0.281864	18.532044 \pm 0.002288	2818.919150	2386	99.999 \pm 9.999	
0015	BLG100.1.198644	3516.324122 \pm 0.002132	4.475618 \pm 0.363962	0.000101 \pm 0.011236	0.887918 \pm 0.113917	17.000783 \pm 0.000118	3164.73294	2340	20.685 \pm 0.059	
0016	BLG100.1.200957	4318.982668 \pm 0.005559	5.918266 \pm 0.126559	0.324206 \pm 0.107114	0.576050 \pm 0.124575	18.371787 \pm 0.001029	5094.61848	2387	99.999 \pm 9.999	
0017	BLG100.1.214894	4549.661247 \pm 0.658040	16.358670 \pm 7.332213	0.575791 \pm 0.219489	0.550016 \pm 0.205528	19.125495 \pm 0.002447	2921.606741	2323	99.999 \pm 9.999	
0018	BLG100.2.70077	3110.076035 \pm 0.075639	17.278477 \pm 3.583333	0.202231 \pm 0.065377	0.721755 \pm 0.25718	19.062267 \pm 0.002365	2666.073479	2383	99.999 \pm 9.999	
0019	BLG100.2.91698	3137.960796 \pm 0.0017360	11.498126 \pm 0.263241	0.105307 \pm 0.011236	0.211990 \pm 0.092153	18.670055 \pm 0.001416	3065.777318	2388	99.999 \pm 9.999	
0020	BLG100.2.97813	4922.779174 \pm 0.077951	41.157209 \pm 1.230317	0.721773 \pm 0.199362	0.119636 \pm 0.124575	17.004479 \pm 0.001044	3815.125221	2388	18.787 \pm 0.016	
0021	BLG100.2.119749	3094.515382 \pm 0.095944	9.532813 \pm 0.126559	0.458382 \pm 0.098794	0.512758 \pm 0.235056	17.078059 \pm 0.001040	19945.974495	2085	21.588 \pm 0.0223	
0022	BLG100.2.164216	4635.153622 \pm 0.099609	37.971852 \pm 0.728795	0.015255 \pm 0.010225	0.7358435 \pm 0.309435	19.231689 \pm 0.0002365	3269.302151	2388	99.999 \pm 9.999	
0023	BLG100.2.164721	3703.421201 \pm 0.076735	6.025499 \pm 0.184497	0.288698 \pm 0.039455	0.841835 \pm 0.313945	19.853586 \pm 0.0002385	2686.073479	2383	99.999 \pm 9.999	
0024	BLG100.2.188384	3912.877724 \pm 0.082390	5.840109 \pm 0.348654	0.020276 \pm 0.012574	0.9036243 \pm 0.086254	18.670055 \pm 0.001413	6201.018084	2388	99.999 \pm 9.999	
0025	BLG100.2.201309	3135.171722 \pm 0.117585	12.810128 \pm 1.910004	0.282024 \pm 0.067166	0.563405 \pm 0.186266	18.349130 \pm 0.000329	3720.610376	2344	99.999 \pm 9.999	
0026	BLG100.3.1.1858	3819.777782 \pm 0.123129	21.901028 \pm 3.132355	0.051051 \pm 0.052400	0.700065 \pm 0.119455	19.053450 \pm 0.001073	4440.646941	2388	99.999 \pm 9.999	
0027	BLG100.3.1.21115	3194.345685 \pm 0.095940	40.149985 \pm 1.453622	0.015255 \pm 0.01255	0.7358435 \pm 0.309435	19.231689 \pm 0.0002365	3269.302151	2388	99.999 \pm 9.999	
0028	BLG100.3.6.33514	3115.795536 \pm 0.123112	5.653413 \pm 0.086512	0.581420 \pm 0.044552	0.902276 \pm 0.215232	18.670055 \pm 0.001413	2853.568006	2388	21.569 \pm 0.184	
0029	BLG100.3.7.41519	2443.200937 \pm 0.060732	38.457010 \pm 0.262165	0.002225 \pm 0.016117	0.010822 \pm 0.002508	18.645523 \pm 0.001413	4337.863876	2388	20.088 \pm 0.076	
0030	BLG100.3.9.98339	3095.940688 \pm 0.095944	22.558979 \pm 0.662425	0.448675 \pm 0.072395	0.862350 \pm 0.209792	19.014171 \pm 0.001051	4440.646941	2388	99.999 \pm 9.999	
0031	BLG100.3.123110	4243.516713 \pm 0.095940	6.6311141 \pm 0.143535	0.118756 \pm 0.061619	0.7358435 \pm 0.309435	19.140657 \pm 0.0002365	3269.302151	2388	99.999 \pm 9.999	
0032	BLG100.3.123957	3088.500011 \pm 0.095940	25.932759 \pm 3.45352	0.020276 \pm 0.01255	0.902276 \pm 0.215232	18.771149 \pm 0.0002365	2402.898762	2385	99.999 \pm 9.999	
0033	BLG100.3.144687	3165.835691 \pm 0.302196	281.118346 \pm 72.550404	0.018223 \pm 0.004056	0.001763 \pm 0.000737	18.771149 \pm 0.0002365	4693.397757	2379	19.976 \pm 0.077	
0034	BLG100.3.147281	4209.446231 \pm 0.065142	9.944647 \pm 0.865160	0.080389 \pm 0.025282	0.564791 \pm 0.148907	19.182254 \pm 0.0002351	3195.601152	2388	99.999 \pm 9.999	
0035	BLG100.3.167671	3802.264525 \pm 0.061360	19.227266 \pm 0.846307	0.051051 \pm 0.052400	0.862480 \pm 0.209792	17.607753 \pm 0.0002365	3261.911190	2379	21.338 \pm 0.144	
...
3560	BLG354.4.131	2445.245 \pm 0.083	12.10 \pm 1.01	0.496992 \pm 0.098749	1.1141 \pm 0.3650	16.693 \pm 0.0001	401.274383	302	19.137 \pm 0.024	

NOTE.—Full table in the machine-readable format is available in the on-line version of the paper.

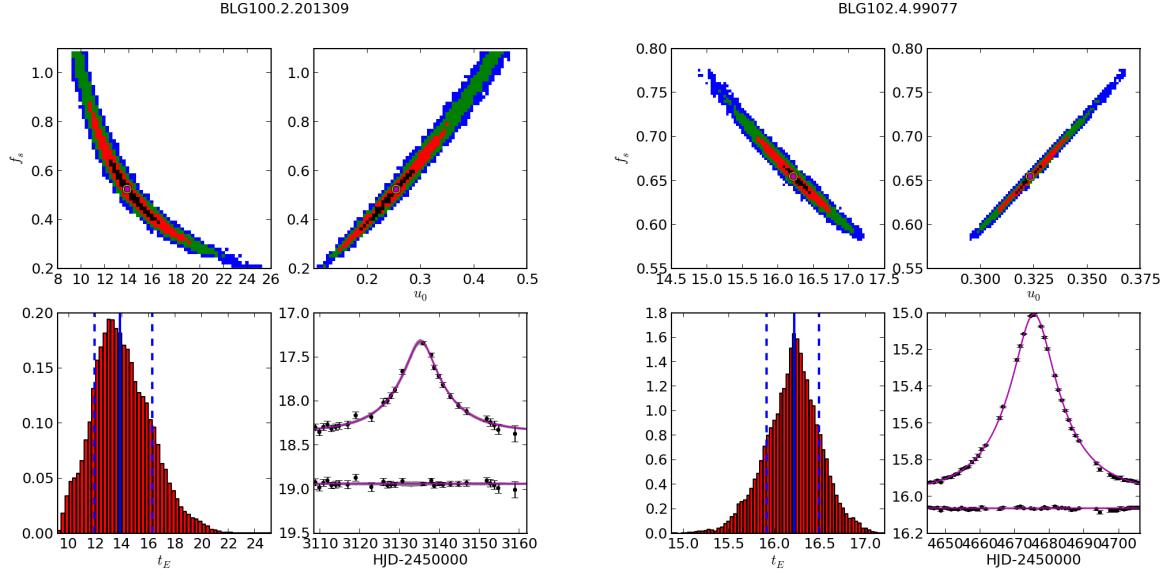


Fig. 6.— MCMC fitting results (density plots) for example events from the OGLE-III Bulge microlensing events catalog. The top panels show relations between distributions of blending parameter (f_s), time-scale (t_E) and impact parameter (u_0). Colors denote sigma contours: black, red and green for 1, 2 and 3 sigmas. A magenta-yellow circle shows where the median of each distribution is located. Lower left panel in each of the plots is the distribution of the time-scales with the solid and two dashed lines marking the 50%, 15% and 85% percentiles. Lower right panel shows the light curve of the event with the median model and its residuals (below the light curve with the offset) shown in magenta and 1-sigma models and their residuals in gray. Full catalog and the MCMC models for all events are available in the on-line version of the paper and on the OGLE web site.

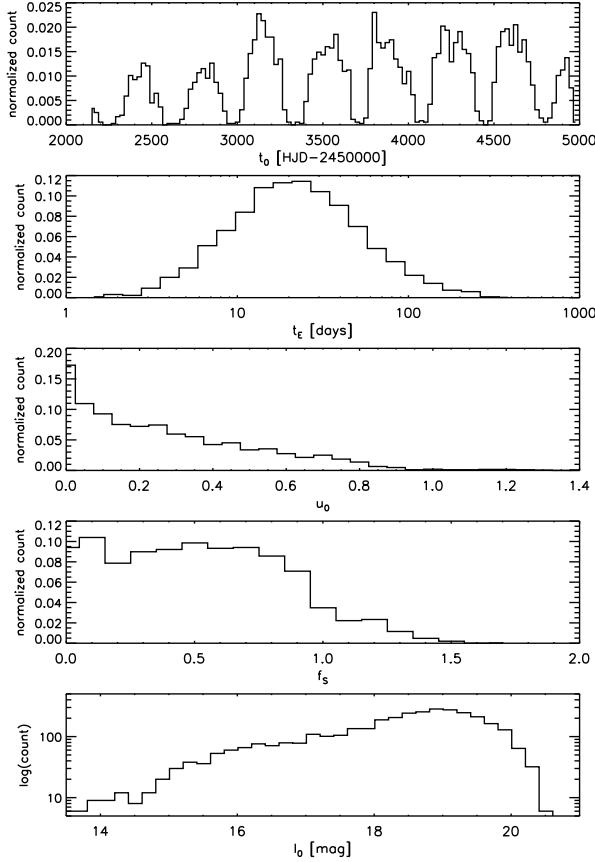


Fig. 7.— Distributions of microlensing parameters for the class A sample of 3560 microlensing events found in the OGLE-III bulge data.

dures described above, i.e. bad observations were masked out and the single outlying points were removed. The error-bars were also corrected.

4.1. Class B events

Apart from the fully automatized search procedure to find microlensing events, we also applied a separate search for standard events for which the full blended fit was not converging. From a sample of light curves rejected at the pre-filtering stage we selected those for which the $\chi^2_{\mu 4}/dof$ for the non-blended model was smaller than 2.5 if the baseline was fainter than 16 mag, and $\chi^2_{\mu 4}/dof < 5$ for $I_{0\mu 4} > 16$. To ensure that only high quality events passed through that filter we also required that the peak was covered with at least 10 data points and selected only events with significant amplification,

i.e., with $u_{0\mu 4} < 1$ in the non-blended model.

There were about 300 candidates, which were then inspected visually and an additional sample of 158 “class B” standard events was selected. Four additional events passed the criteria, but were duplicates due to overlapping OGLE-III fields. The main reason for non-converging microlensing models for those events was in many cases some small, often gradual, long-term variability on top of the microlensing bump, likely caused by the intrinsic variability of the source or high proper motion of either the source or nearby stars. We provide the list of those candidate events in Table 3 and their light curves are also made available on-line for further studies. The events from class B are dubbed following the schema: OGLE3-ULENS-9nnnn, where the individual number encoded in nnnn is preceded by number 9, in order to distinguish from class A events.

Because the microlensing models of class B events were not full and did not include blending, the derived time-scales were not reliable in vast majority of the cases. Moreover, those events would not be detected by the automated pipeline, hence the detection efficiency did not include them and they were not used in the statistical studies below. However, because class B events are most likely the genuine standard microlensing events, they might be a useful reference for future searches of events in their vicinity.

4.2. Comparison with the EWS

Comparison with the OGLE’s Early Warning System (EWS⁴) (Udalski 2003) was tricky, as the reference images used for the real-time analysis and production of the final photometry we relied on had a slightly better quality and depth. Thus, many stars got resolved in the new data, or the blending has changed. Nevertheless, we compared the catalog with the events from all years of the EWS (2002-2009).

Within 91 fields analyzed here, EWS has detected 3796 events, among which were 162 duplicates due to overlapping fields, therefore there were 3634 unique events reported. Using a 3 arc sec matching radius and the time of maximum,

⁴<http://ogle.astrowu.edu.pl/ogle3/ews/NNNN/ews.html>, where N=(2002,2003,2004,2005,2006,2007,2008,2009)

TABLE 3
OGLE-III CANDIDATE STANDARD EVENTS (CLASS B).

ID OGLE3-ULENS-	RA _{J2000} [h:m:s]	Dec _{J2000} [deg:m:s]	field	starno	EWS id
90001	17:52:12.90	-30:13:38.5	BLG100.1	101182	2005-BLG-359
90002	17:51:19.10	-30:02:47.0	BLG100.2	132033	-
90003	17:50:32.77	-29:52:13.7	BLG100.6	137109	2007-BLG-194
90004	17:50:06.77	-30:08:46.5	BLG100.7	9419	-
90005	17:53:52.53	-29:59:55.0	BLG101.1	238410	2007-BLG-090
90006	17:52:43.65	-29:39:20.2	BLG101.5	20164	-
90007	17:52:45.56	-29:49:46.9	BLG101.6	22645	-
90008	17:53:29.42	-29:56:57.0	BLG101.7	59322	-
90009	17:53:24.17	-29:56:32.7	BLG101.7	124729	-
90010	17:55:55.53	-29:21:14.5	BLG102.5	46864	2008-BLG-552
90011	17:55:13.27	-29:16:53.1	BLG102.5	135824	-
90012	17:55:04.70	-29:25:14.0	BLG102.6	139751	-
90013	17:55:30.55	-29:24:39.4	BLG102.6	161973	2007-BLG-295
90014	17:55:50.73	-29:45:40.6	BLG102.8	108332	-
90015	17:56:47.26	-30:01:28.4	BLG103.3	151443	2006-BLG-053
90016	17:56:24.60	-29:50:19.7	BLG103.4	124317	-
90017	17:55:14.38	-29:53:14.0	BLG103.5	79851	-
90018	17:56:05.22	-30:11:04.9	BLG103.7	113081	-
90019	17:55:11.11	-30:22:04.3	BLG103.8	54768	2004-BLG-486, 2004-BLG-490
90020	17:59:23.08	-29:45:01.5	BLG104.1	17566	-
90021	17:59:47.01	-29:18:13.3	BLG104.4	35755	-
90022	17:59:16.67	-29:15:14.1	BLG104.4	76137	-
90023	17:59:43.35	-29:14:37.8	BLG104.4	105201	-
90024	17:57:53.42	-29:12:03.9	BLG104.5	194584	-
90025	17:58:40.44	-29:11:52.7	BLG104.5	225289	2006-BLG-202
90026	18:02:28.53	-29:43:26.2	BLG105.1	105229	-
90027	18:01:13.71	-29:27:30.1	BLG105.6	40335	2002-BLG-357
90028	18:01:38.68	-29:20:28.5	BLG105.6	240417	-
90029	17:46:33.26	-36:42:32.2	BLG108.2	3390	2004-BLG-517
90030	17:46:57.18	-34:27:57.6	BLG130.1	56220	-
90031	17:45:09.46	-33:56:48.3	BLG130.5	98747	2007-BLG-275
90032	17:46:19.48	-34:13:45.2	BLG130.7	200260	2006-BLG-262
90033	17:49:19.54	-34:23:00.9	BLG131.1	141016	-
90034	17:48:15.25	-34:06:53.7	BLG131.6	104924	2004-BLG-099
90035	17:52:40.90	-34:26:59.4	BLG132.1	72963	-
90036	17:53:09.51	-34:08:26.3	BLG132.3	89578	-
90037	17:56:59.79	-33:55:15.6	BLG134.5	193547	2005-BLG-061
90038	17:45:31.29	-33:46:15.7	BLG138.1	163835	-
90039	17:46:14.66	-33:47:27.9	BLG138.1	185675	-
90040	17:44:55.31	-33:25:03.2	BLG138.5	70885	2008-BLG-238
90041	17:45:03.92	-33:52:29.2	BLG138.8	81987	2004-BLG-499
90042	17:47:05.93	-33:40:46.9	BLG139.7	75690	2006-BLG-054
90043	17:50:42.23	-33:30:57.8	BLG140.6	138482	2007-BLG-476
90044	17:57:00.80	-33:53:58.7	BLG142.1	24051	2007-BLG-311
90045	17:51:05.78	-32:45:47.4	BLG147.4	157489	-
90046	17:52:07.67	-33:11:44.1	BLG148.8	193131	2006-BLG-363
90047	17:53:51.79	-33:14:12.4	BLG149.8	72373	2004-BLG-567
90048	17:57:23.66	-33:06:11.0	BLG150.7	75044	2004-BLG-016
90049	17:54:51.71	-32:21:46.7	BLG156.6	116230	2008-BLG-496
90050	17:54:04.49	-32:40:38.2	BLG156.8	68730	-
....
90158	17:41:15.09	-23:59:52.3	BLG346.5	139563	-

NOTE.—Full table in the machine-readable format is available in the on-line version of the paper.

we identified 2309 events in our catalog which were previously found. Therefore, 1409 (1333 class A + 76 class B) events are newly discovered standard microlensing events.

The main reason for the significant majority of EWS events not recovered in our analysis was lack of convergence in the microlensing model with free blending. The converged modeling was among the preliminary pre-filtering requirements, i.e., we requested that both blending parameter and time-scale were within sensible ranges. EWS was marking events exhibiting some problems with data modeling with free blending by forcing f_S to be constant at 1. In many cases, the models did not converge because the events were not due to single-lens-single-source, or exhibited additional effects like parallax or binary lens caustics.

The remaining EWS events which were rejected after pre-filtering at the later stage (RF classification), were inspected visually and most of them were found to be of too low signal-to-noise to be detected (e.g., 2007-BLG-044). There were also clear non-microlensing outbursts (e.g., 2003-BLG-266), most likely caused by Be stars or cataclysmic variables (e.g., Mroz et al. (2014)), and a significant number of exotic events with strong parallax, binary lens or finite source effects (e.g., 2003-BLG-067, 2008-BLG-199).

4.3. Detection efficiency

The events selection procedure described above was fully automatized, therefore we were able to derive the detection efficiency for finding the standard microlensing events in the OGLE-III data. Following the method of Wyrzykowski et al. (2009), we simulated 300,000 light curves of events for three regions of the Bulge covered by the OGLE-III survey, with different stellar density and different sampling. In order to account for blending in the Bulge fields in our simulations, we used archival Hubble Space Telescope (HST) *I*-band images of the field BLG206 to obtain the distribution of blending parameter f_S . The distribution was derived by matching the OGLE objects to individual stars present on HST *I*-band image and by finding the relative brightness of all HST components to the OGLE brightness. The 91 OGLE-III fields used in this analysis were distributed over the densest region of the Bulge and, despite the changes in the stellar counts, we as-

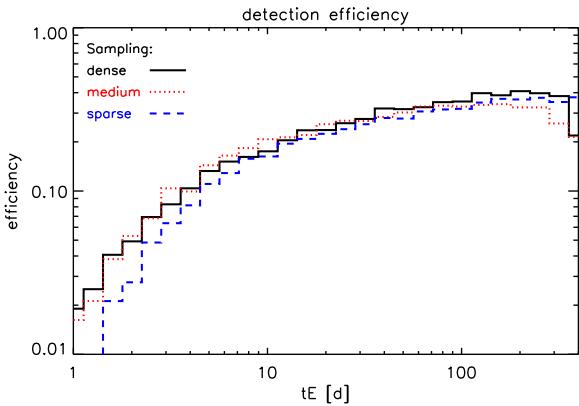


Fig. 8.— Microlensing events detection efficiency curves as a function of simulated time-scale. Shown are the efficiencies for three fields with different sampling density.

sumed that the overall distribution of blending parameter remained similar in all the fields. However, a more detailed analysis of the blending over a range of OGLE-III fields is currently ongoing as it is required to compute the real number of monitored stars, which is a necessary component of the microlensing event rate and the optical depth.

For each simulated light curve we applied the pre-filtering procedure as described above, derived 27 features, and checked if the event passed through the Random Forest classifier. Figure 8 shows the derived efficiency curves for three regions of the OGLE-III fields observed with different sampling: dense, medium and sparse. The detection efficiency for the bulk of the events is at the level of 30-40% and drops below 10% for events with time-scales shorter than 5 days. At the long end the efficiency starts dropping from about 200 days to 400 days, which was the truncation in the simulations. The search procedure was not optimized for events with $t_E > 400$ days and was susceptible to numerous low amplitude contaminants. A dedicated search for very long events will be presented separately. Because the three detection efficiency curves are similar, for further analysis we use the mean for all the fields.

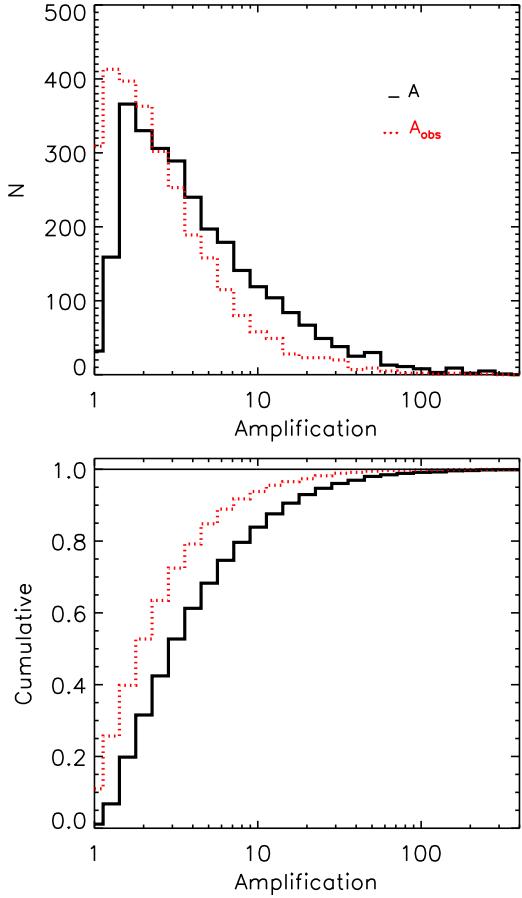


Fig. 9.— Distributions of the maximum microlensing amplification (A) (solid line) and the observed amplification which takes blending into account (red dotted line). There are 24 events with $A > 100$ and 5 with $A > 300$.

5. Discussion

The large number of standard microlensing events allows for deriving various statistical properties of the events toward the Galactic Bulge. For each event we derived the microlensing model parameters using the MCMC modeling and each event in class A sample is accompanied with the plot with relations between the main parameters: f_S , u_0 and t_E , allowing for detailed investigation of degeneracies in the models and the statistical properties of the events.

5.1. Properties of the events

Figure 9 shows distributions of maximum amplifications for 2812 events selected from the entire sample to have well constrained impact parameter, i.e., $\Delta u_0/u_0 < 1$. The drop between 1 and 2 is caused by the lower detection efficiency for events which were not magnified significantly above the noise. There are 32 events with amplification greater than 100 and 6 with amplification greater than 300. The cumulative histogram on Fig. 9 shows that half of the events had $A > 3$. Also shown (red dotted line) is the distribution of the observed amplification, $A_{\text{obs}} = (A - 1)f_S + 1$, i.e., the actual observed rise in brightness given the blending. There were 8 and 1 events which got 100 and 300 times brighter, respectively, typically with $f_S \sim 1$. Half of the events had their observed magnification greater than 2.0.

Figure 10 shows how the MCMC errors (approximated here as symmetric) of microlensing parameters are spread. The events with poorly constrained amplification are forming a separate group on most of the panels, however, it can be seen that their time-scales are still well recovered, with relative error on t_E below 1. Only about 2% of events have the relative errors in t_E greater than 100% and about 17% have it larger than 50%. This allows to select a subsample of events with well constrained Einstein radius crossing time, the events which can then be used for statistical studies, computing the optical depth or constraining the shape of the IMF.

We notice that larger relative error on f_S and u_0 does not necessarily lead to a bad derivation of the t_E , however, it is clear that majority of badly constrained time-scales are related to $f_S < 0.5$, i.e., where the blending is severe. Strong blending causes the light curve to show only the tip of the actual event when the observed amplification becomes high enough to overcome the noise of the baseline.

We also find that vast majority of poorly constrained time-scales is related to the sparse sampling and is not to the time-scale itself. The bottom panels of Fig. 10 show the dependence of the error on t_E and f_S as a function of the average number of data points taken per 1 day unit of the time-scale of the event. It can be seen that all events with more than $1.5 \times t_E$ points collected had

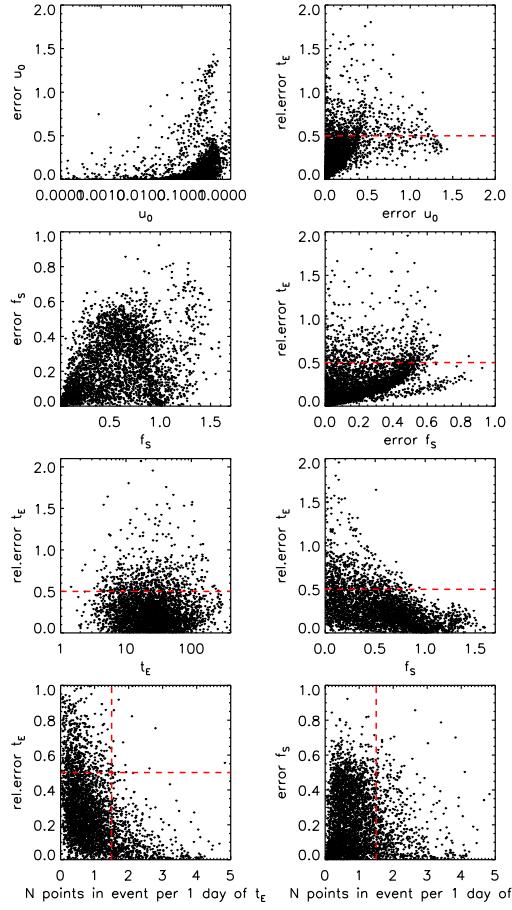


Fig. 10.— Relative errors on the main microlensing parameters: u_0 , f_s and t_E . Despite the error in determination of u_0 and f_s spans a very wide range, well above 100%, the error on t_E stays mostly within 50%. Only 2% of events have relative error in t_E larger than 1.

their time-scales derived with small relative error. This property can be useful when designing future microlensing experiments, however, we emphasize that is only valid for the time-scales longer than about $t_E > 5$, as there were not enough events below that range in our sample. Determination of the blending parameter is somewhat less sensitive to the average number of points taken during the event, however, it depends strongly on the coverage of the wings of the event, as shown already in Woźniak & Paczyński (1997).

Figure 11 shows blending parameter distribution with relation to u_0 , I_0 and t_E . We clearly see

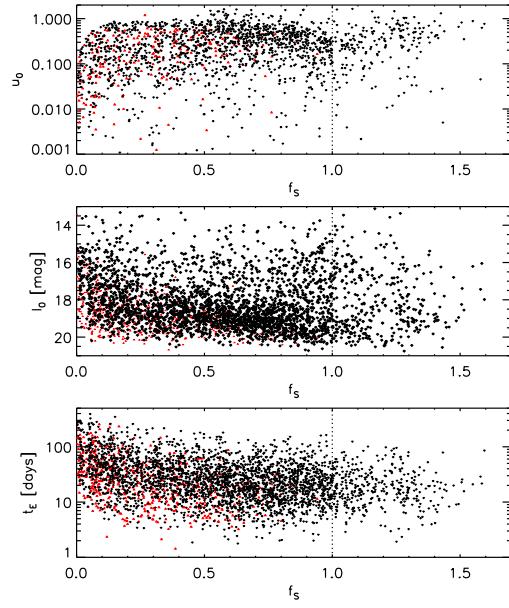


Fig. 11.— Dependence of u_0 , I_0 and t_E on the blending parameter f_s , as obtained in the MCMC modeling of the microlensing events. Red dots point out events for which the relative error on t_E was larger than 50% and which were excluded from time-scale distribution analysis.

the observational bias that highly blended events ($f_s \ll 1$) can only be detected if highly magnified, i.e., with very small u_0 . Then, the observed (blended) baseline magnitude follows roughly the results of the blending simulations from Smith et al. (2007). Agreement with simulations indicates the blending parameter obtained for in our microlensing models follow the expected distribution, which, in turns, means we derive the time-scales of events correctly. The relation shows that for bright events blending tends to be around 1, however can also take much smaller values. As already pointed by Smith et al. (2007), events with bright baseline can not be assumed to suffer less blending than fainter events. For fainter events, below 18 mag, the blending parameter can be of any value. Events with errors on time-scale larger than 50% tend to have smaller f_s , indicating the blending uncertainty is mainly responsible for larger errors in t_E .

There is a signature visible on the t_E - f_s plot that a fraction of long events tends to have small f_s . This could be an artifact of modeling and can

lead to a bias in derived distributions of t_E toward longer time-scales. However, as we show later, we found a way to avoid such biases. The same plot also shows that there is hardly any event with t_E shorter than 15 days as very small f_S , which, in turn, is expected, as severely blended short time-scale events have small A_{obs} and hence small detection efficiency.

5.2. Time-scale distribution

One of the main applications for a large ensemble of microlensing events is studying the structure of the inner parts of the Galaxy. This is usually done with the microlensing events rate (Γ) and the optical depth (τ) (e.g., Evans & Belokurov 2002, Mao 2012). Both of those quantities rely on the time-scales, however, they also require the knowledge of the number of monitored stars. As noted in e.g., Sumi et al. (2006) and Wyrzykowski et al. (2009), obtaining that value unambiguously is not straightforward for the dense stellar fields. Microlensing can occur on stars which are very faint and are blended with brighter foreground stars, therefore in order to know the exact number of monitored stars it is necessary to understand the blending in observed fields, typically by comparing the observed luminosity function (LF) to the one from the images of much higher resolution (e.g., from the HST) in which the individual stars are resolved. In Wyrzykowski et al. (2011a) and Wyrzykowski et al. (2011b) such analysis was performed for the OGLE-III data of the LMC and the SMC, assuming no major variations in the LF between fields. However, repeating it for the Bulge data is much more complicated. OGLE-III covered a wide area of the sky toward the Galactic Center, where the extinction and the mix of stellar populations vary significantly from field to field. A detailed analysis of the number of stars and blending distribution over a range of OGLE-III fields will be performed in near future, here however, we still can use the observed time-scales, corrected for the detection efficiency, to investigate the effects of the structure of the Galaxy on the distribution of t_E .

5.2.1. Comparison to Besançon model

The most recent model of the Galaxy and its microlensing yield in the Bulge (Kerins et al. 2009) derived the mean time-scale for events

which occurred on sources with resolved magnitude brighter than $I_S < 19$ mag. From sample A of our catalog we selected 1184 events for which the source magnitude (computed using I_0 and f_S) was brighter than 19 mag and the relative error on t_E was less than 50%. We binned the events into $1 \times 1 \text{ deg}^2$ bins, requiring at least 5 events in a bin – most bins contained actually about 30 events, with the maximum of about 100 in the densest area. Table 4 contains the values of mean time-scale and the number of events in each bin.

Within each bin we created a distribution of t_E in log-space, convolved it with the detection efficiency, and computed the arithmetic mean. Additionally, in order to minimize the effects of single long- or short-time-scale events affecting the computation of the mean t_E , we also fitted the distribution in each bin with a log-normal model, and then computed the mean for that model. Figure 12 shows maps of the mean time-scale obtained in those two ways. The maps are also overlaid with the expectations for the Besançon model from Kerins et al. (2009).

Computation of the mean time-scale based on the log-Gaussian model clearly is less prone to outliers in the t_E distribution, as can be seen in the bins at high galactic longitudes, where there were very few events in each bin. The averages computed from a log-normal fit are somewhat smaller in most of the bins, also proving that the regular mean is affected by long events.

As can be seen in Fig. 12, OGLE-III events are located almost solely between two isolines of the Kerins et al. (2009) model, between 20 and 25 days. However, in the central parts, at $l \sim 0$, our values tend to be significantly below 20 days, on both maps. The mean time-scale then increases with increasing $|l|$, but with clear asymmetry and larger values on the negative galactic longitudes. This is most likely a signature of the bar geometry, also present as an asymmetry on the synthetic map of Kerins et al. (2009) – on the negative l (far side of the bar) our line-of-sight crosses larger section of the bar than it does on the near side, resulting in an increase of bar-bar events (Paczyński et al. 1994), which tend to have longer time-scales due to small velocity dispersion in the bulge.

The effect of asymmetry is also visible if the events are binned into three broad regions: for positive, central and negative galactic longitudes,

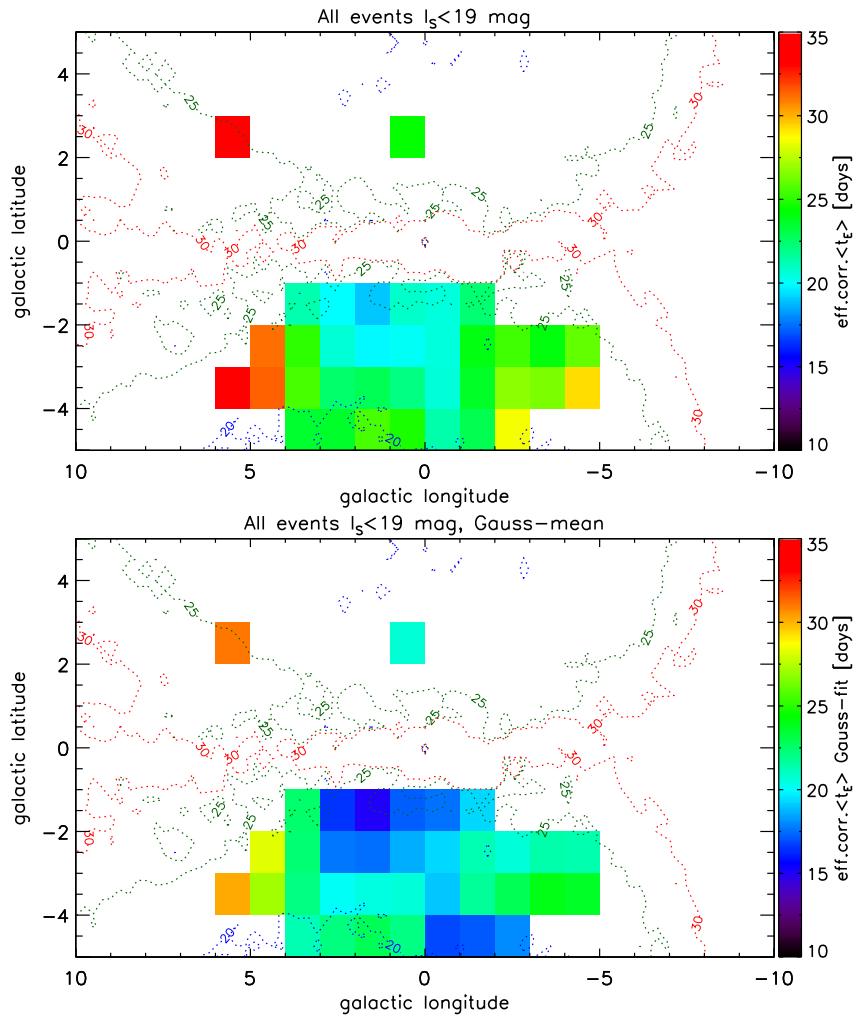


Fig. 12.— Efficiency corrected mean t_E map for 1184 events with source(resolved) magnitude brighter than 19 mag and relative error of t_E better than 50%. Dashed contours show the expected mean time-scale (at 20, 25 and 30 days) as computed in Kerins et al. (2009). Within each bin (size 1×1 degree) the mean time-scale was computed from the actual values of t_E (upper map) and from the log-normal fit to the distribution (lower map). Both maps were smoothed with a Gaussian with FWHM=1.75 deg.

TABLE 4
MEAN TIME-SCALE BINS FOR OGLE-III EVENTS.

l_{central} [deg]	b_{central} [deg]	$\langle t_E \rangle_{\text{uncorr}}$ [days]	$\langle t_E \rangle_{\text{eff,corr}}$ [days]	$\langle t_E \rangle_{\text{Gauss}}$ [days]	N_{events}
5.5	-3.5	34.4	32.8	31.9	9
5.5	2.5	46.8	37.5	30.9	10
4.5	-3.5	47.1	42.5	23.4	17
4.5	-2.5	42.8	32.2	29.4	18
3.5	-4.5	29.1	24.1	19.7	18
3.5	-3.5	33.5	24.0	19.7	39
3.5	-2.5	34.9	27.2	18.6	33
3.5	-1.5	30.8	15.6	24.5	11
2.5	-4.5	28.5	22.4	23.1	19
2.5	-3.5	26.7	20.3	17.5	40
2.5	-2.5	23.7	16.5	12.4	75
2.5	-1.5	29.4	24.4	15.5	18
1.5	-4.5	29.5	26.4	22.1	23
1.5	-3.5	30.1	24.6	21.2	38
1.5	-2.5	26.9	20.8	19.9	98
1.5	-1.5	15.6	11.4	12.7	20
0.5	-4.5	27.8	17.4	28.0	17
0.5	-3.5	28.0	24.3	23.8	53
0.5	-2.5	26.0	19.3	16.6	92
0.5	-1.5	31.8	25.5	19.5	97
0.5	2.5	27.9	22.4	15.1	6
-0.5	-4.5	22.4	20.8	11.4	12
-0.5	-3.5	21.6	17.2	16.0	35
-0.5	-2.5	26.4	18.3	22.4	56
-0.5	-1.5	24.4	19.3	17.4	47
-1.5	-4.5	28.5	20.8	17.5	16
-1.5	-3.5	33.5	23.3	27.9	38
-1.5	-2.5	34.3	27.4	25.0	37
-1.5	-1.5	26.0	22.1	18.0	15
-2.5	-4.5	46.8	35.0	16.3	10
-2.5	-3.5	38.7	29.1	28.9	31
-2.5	-2.5	32.0	24.4	16.1	17
-3.5	-3.5	27.0	22.2	23.5	16
-3.5	-2.5	28.5	18.6	36.5	12
-4.5	-3.5	39.2	33.8	23.9	28
-4.5	-2.5	35.1	28.6	19.7	32

NOTE.—Machine-readable table is available in the on-line version of the paper.

within the same galactic latitude band (from -4 to -2 degrees). The distributions for those regions are shown in figure 13 along with the best fit Gaussian models. The standard mean t_E for positive and central regions is 23.0 ± 0.7 days and 21.4 ± 0.4 days, respectively. For the negative longitudes bin the mean time-scale is 28.0 ± 1.1 days, which clearly stands out with respect to the other bins. The Gaussian-fit mean t_E values for positive and central bins is 20.5 and 20.7 days, respectively. For the negative bins it is 24.4 days, somewhat smaller than the standard mean, however, still clearly higher than the mean time-scale in positive and central parts of the map. Moreover, the mean time-scales for negative galactic longitudes are higher than expected for the Galactic model of Kerins et al. (2009).

OGLE-III fields do not cover the regions beyond $|l| > 5$, where the increase in the mean t_E should be even more pronounced, however, the fact that the mean time-scale increases somewhat quicker with galactic longitude may indicate either that the bar is tilted somewhat more toward the line-of-sight. Another reason of such rise could be that the boxy bar is wider than assumed in the Besançon model, producing more bar-bar events at lower l than in the center. A future comprehensive analysis of the OGLE-IV survey data, which is monitoring the region up to $|l| = 10$ since 2010, will give even more clues to verify the shape of the galactic bar.

When comparing the mean time-scale structure to the theoretical predictions from Evans & Belokurov (2002) for the Freudenreich's bar model (their Fig. 5), it resembles the model without bar streaming included. On the other hand, the values of the time-scales here are from 50 to 75% larger than predicted in Evans & Belokurov (2002), but still significantly lower than those expected for bar streaming.

5.2.2. All events sample

Figure 14 shows maps of mean time-scales computed in a standard way and from the Gauss-fit for all class A events with relative error in t_E better than 50%, with no restriction regarding their source or baseline magnitude. There were 3019 such events and the maps were binned with small bins of 0.4×0.25 degrees, a similar size as shown on map of mean time-scales obtained from MOA-

II data (Sumi et al. 2013). The main difference between standard and Gauss-fit mean time-scales is clearly visible again on the edges of the observed region, where there were very few events (even just 2 per bin) to compute the $\langle t_E \rangle$ and the arithmetic average was very sensitive to long events. There is some broad structure visible again on both maps, with the lower mean time-scale values on the positive galactic longitudes and higher on the negative longitudes, however, the finer elements of the map are of very low statistical significance due to low number of events in each bin. The all-events sample is a mixture of events from various combinations of populations: bar-bar, disk-bar and disk-disk, therefore the observed distributions are generally more blurred, as different pairs cause different duration events, with bar-disk being typically the shortest, whereas bar-bar and disk-disk the longest.

The overall Gauss-fit mean time-scale still stays within predicted 25 days in the central parts, but for large $|l|$ reaches close to 30 days. This map could be compared to Fig.4 of Evans & Belokurov (2002) for Freudenreich's model (Freudenreich 1998), which shows the mean time-scale for sources and lenses from either disk or the bar, however due to different resolutions it is difficult to conclude on the comparison. Nevertheless, the values on our map tend to agree with the expectations for the model with the contribution of the spiral structure for the Freudenreich's model.

We can also compare our map of standard mean time-scales for all class A events to the results of MOA-II (Sumi et al. 2013) (their Fig. 3, uppermost panel), which were based on 474 events. First, we see that our longest mean time-scale is about 35 days, the value which actually appears only at the edges of our observing region and should be ignored as induced by very low number statistics in those areas (those regions completely disappear when using Gauss-fit mean time-scales, see bottom map in Fig.14). Therefore, our longest value to compare with MOA-II map is actually closer to 25 days, whereas the map in Sumi et al. (2013) shows areas with $\langle t_E \rangle$ as high as 40 days. Moreover, on our map we do not see any of the regions of high $\langle t_E \rangle$ values, especially at $l \sim +3$, claimed by MOA-II and previous results (Alcock et al. (1997), Popowski et al. (2005)), indicating it was likely a statistical fluctuation driven by small

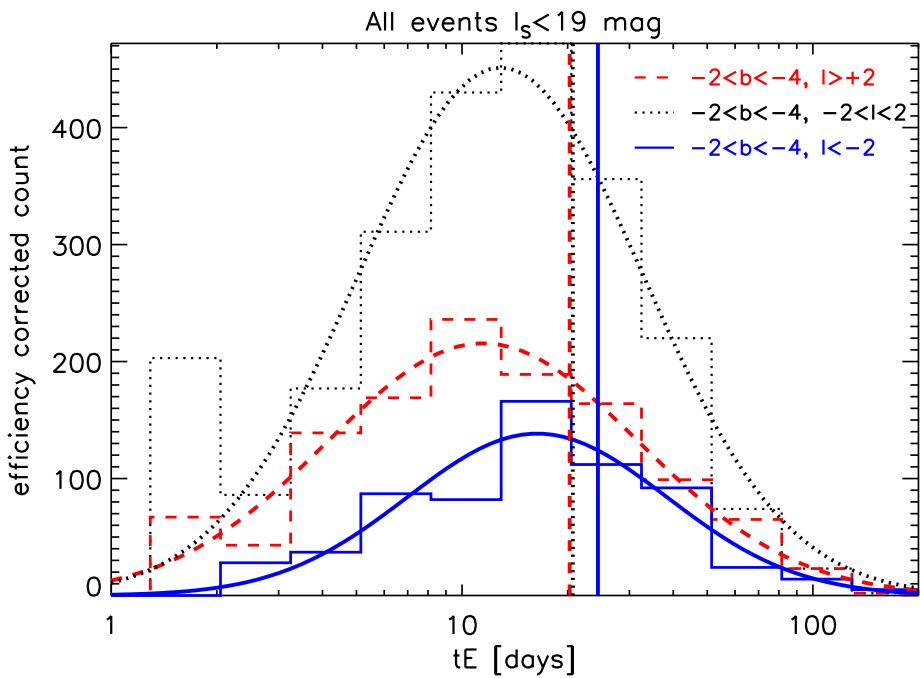


Fig. 13.— Distribution of the time-scales of events for three regions between -2 and -4 galactic latitudes: positive galactic longitudes (dashed red), central (dotted black) and negative (solid blue). Only events with $I_0 < 19$ mag and relative error on $t_E < 50\%$ are used for the histograms. Also shown are the best fit Gaussians. The vertical lines show Gaussian-fit mean time-scales for the corresponding bins. The highest mean is measured as 24.4 days for the negative bins (solid blue).

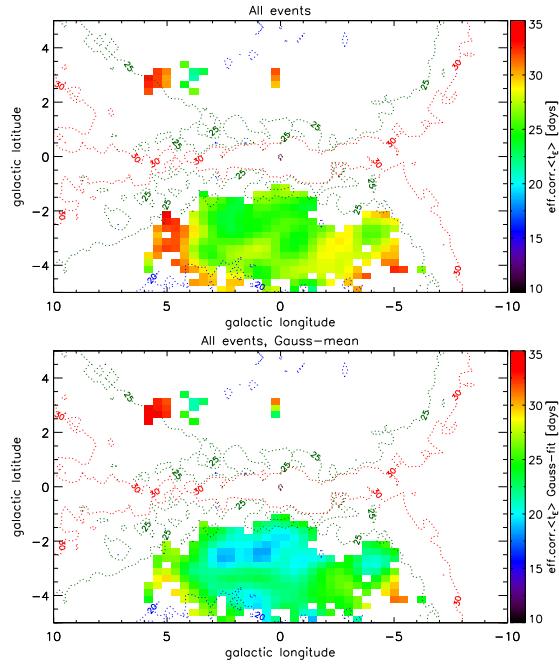


Fig. 14.— Maps of efficiency corrected mean t_E from computed in standard way (top) and based on Gaussian fit to the t_E distribution in each bin (bottom). Maps were build for 3019 events with no restriction to the source magnitude and with the relative error of t_E better than 50%. The bins are 0.4×0.25 degrees and the maps are smoothed with a Gaussian with FWHM= $3 \times$ bin size. Dashed contours show the expected mean time-scale for events with $I_S < 19$ mag (Kerins et al. 2009), shown here for reference only.

number of events available for those studies. The only part where maps from OGLE-III and MOA-II agree is the decrease of mean t_E at $l \approx +1$, $b \approx -2$, which is even more clearly visible on the Gauss-fit mean t_E map for OGLE-III. However, as mentioned above, the All-sample maps present a blurred view on mean time-scales due to mixing populations and distances of sources and lenses.

5.2.3. Bulge-source events sample

In order to minimize the impact on population mixing, we restricted our events to those with Red Clump (RC) stars from the bulge as the most likely sources. Figure 15 shows the color-magnitude diagram of sources of all events, marking the region of Red Clump sources. The I -band magnitude

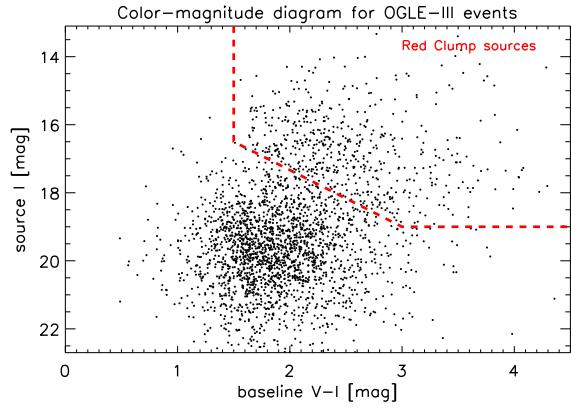


Fig. 15.— Color-Magnitude Diagram of sources in All Sample microlensing events, composed of source I -band magnitude computed from the blended microlensing fit and approximate $V - I$ from the baseline. Region with sources belonging to the bulge Red Clump population is marked.

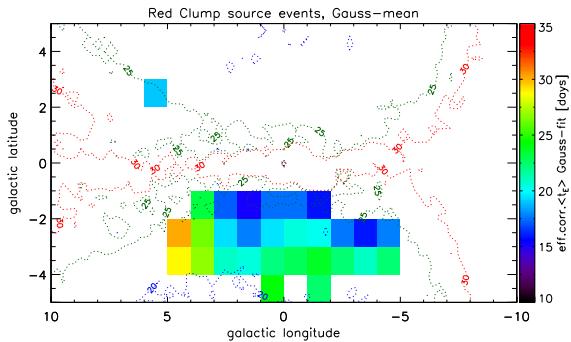


Fig. 16.— Map of efficiency corrected mean t_E from the Gaussian fit to the t_E distribution in each bin. Map was build for 556 events with Red Clump stars from the bulge as most likely sources and with the relative error of t_E better than 50%. Dashed contours show the expected mean time-scale for events with $I_S < 19$ mag (Kerins et al. 2009), shown here for reference only. Bin size is 1×1 degrees.

was computed using I_0 and f_S , however, because we did not have enough observations taken in the V -band, the color of the source is assumed to be similar to the color of the baseline (i.e., source and blends). This approximation is close to reality because Red Clump giants are typically much brighter than other (bluer) stars in the direction of the Bulge, hence the color of the red giant dominates in the baseline. Note that the opposite assumption, that the non-RC sources are bluer in the baseline, however, would not be true for the same reason.

Selecting RC sources allowed us to construct another map of mean time-scales, shown in Fig. 16. The values of $\langle t_E \rangle$ are clearly much smaller than in the previous maps, most likely because in this case most of events have sources in the bar and lenses are from the disk. Such combinations of lenses and sources typically have higher relative proper motion, hence tend to produce shorter t_E (e.g., Di Stefano 2012. However, Fig. 16, has a clearly distinctive two regions with different values of mean time-scale, with the split at about $b \sim -3$ deg. The events closer to the Galactic Center tend to have shorter t_E (well below 20 days), whereas those at higher Galactic latitudes have t_E above 20 days. The duality, or a gradient of $\langle t_E \rangle$ with galactic latitudes, can be explained that at low b the dominating lensing configuration contains a source from the Galactic bar and a lens from the disk, as the volume of the disk we cross with the line-of-sight toward the bar increases with lower b . On the other hand, the further we look away from the Galactic plane, the less likely it is to have events caused by disk lenses, hence the bar-bar events dominate and their mean time-scale is typically longer than bar-disk events.

Our values of $\langle t_E \rangle$ for the Red Clump sample do not match either of two scenarios presented in Fig.5 of Evans & Belokurov (2002), showing mean time-scales for bar sources in Freudreich's model including and excluding the contribution of bar streaming. It might mean that there is very little streaming in the bar, however, our sample of RC events is small and is distributed over a small area of the sky, which might be influencing the conclusions. More microlensing data is needed for larger galactic longitudes to verify this issue.

5.2.4. Mass Function

Finally, in Fig. 17 we show the distributions of efficiency corrected time-scales of events in three subgroups: All, Red Clump and within the 2 degrees around the Baade's Window (BW, $l \sim 1$, $b \sim -2$). Thick dashed lines show the expected power-law slopes for short and long tail of the distribution of the time-scales, valid for a range of mass functions from Mao & Paczyński (1996). The slope of the distribution for long end agrees with the prediction, at least for the Red Clump sources. For the BW and All events, the long event tail seems to have more shallow slope, which could be caused by either flatter mass function of the disk, or small relative velocity dispersion for disk-disk events, as the BW and All samples contain more disk-disk events, which have typically longer t_E . This could also just be the modeling bias, mentioned earlier (see sec.5.1). For the short time-scales the slope in the data is more flat, which could be a sign of some additional population of either light or fast lenses (Sumi et al. 2011). Gauss-fit average values for t_E were computed as 27.0, 24.6 and 24.6 days for All, RC and BW samples, respectively.

5.2.5. Very long and very short events

We note that our sample contains a significant number of long and short events: there are 9 events with t_E longer than 300d, 205 (7%) events longer than 100d and 722 (24%) longer than 50d. Because our search pipeline was designed to select purely standard events, the events with strong parallax signal were excluded, however, still there remains a couple of events exhibiting some weak systematic deviation from the standard microlensing model likely due to the parallax effect. The difference in t_E with and without parallax model included in those cases was within the reported error-bars for t_E , however. Therefore, those events can be safely included in our sample of standard events. The entire OGLE-III data will be searched separately for long events with strong parallax.

For short events, we see that 561 (19%) events have $t_E < 10$ d, 118 (4%) have $t_E < 5$ d and there are even 6 events with time-scales between 1 and 2 days. Given a relatively sparse sampling of the OGLE-III survey and low detection efficiency for very short events, this might represent a high

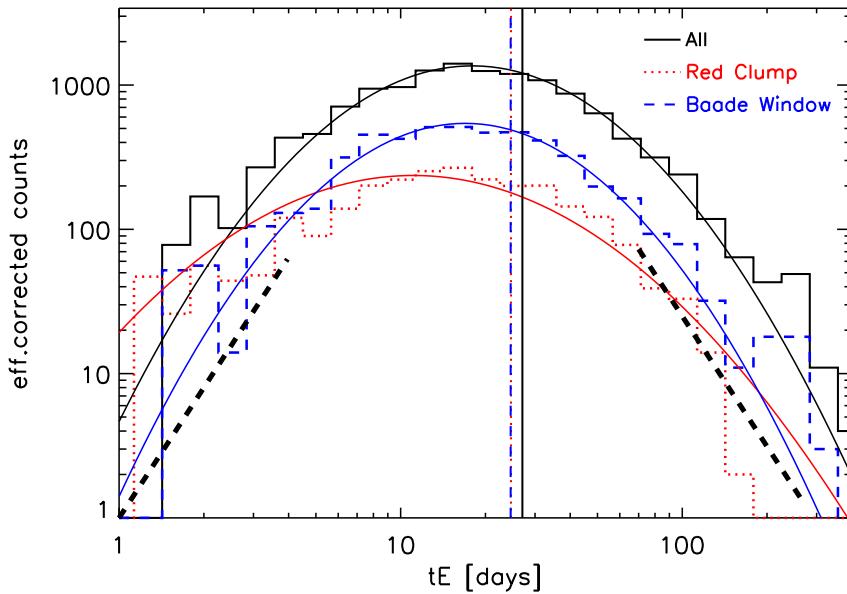


Fig. 17.— Distribution of the logarithm of the Einstein Radius crossing times (t_E) for All events (black solid), Red Clump source events (red dotted) and around the Baade's Window (blue dashed). We only show events for which the relative error on the time-scale is smaller than 50%. There are 3019, 556 and 1144 events in All, RC and BW samples, respectively. The Gauss-fit mean values of t_E are marked with corresponding vertical lines (27.0, 24.6, 24.6 days, respectively). Thick dashed lines indicate power-law slopes of +3 and -3 for short and long time-scale tails, respectively.

abundance of low-mass lensing objects like brown dwarfs or unbound or wide-orbit planets, or just a surplus of very fast lenses. A systematic optimized search for very short OGLE-III and OGLE-IV events will be conducted in near future, along with the study of potential instrumental or numerical systematics causing the time-scales becoming shorter than in reality.

6. Summary

We have prepared and presented the largest homogenous catalog of 3718 good quality standard microlensing events toward the Galactic Bulge using 8 years of the OGLE-III survey (2001-2009). 1409 of the events are new and were not previously reported by the OGLE EWS. The full catalog is available for the astronomical community, with all the light curves, probability distributions plots for the microlensing parameters and values of all parameters with their uncertainties in a machine-readable form.

Applications of such catalog are numerous, from studying the individual events, to large sample analyzes. Because of their number and distribution over a wide range of galactic longitudes and latitudes, it was possible to measure the mean t_E per region of the sky robustly enough to provide a direct comparison to the expectations. Our data are in good agreement with recent models of the inner parts of the Milky Way, except that the asymmetry in the mean time-scale due to the bar angle seems to be somewhat more pronounced. This might be an evidence of a slightly different geometry or shape of the Galactic bar and may call for new models of the very center of our Galaxy. In near future, the catalog will be the basis for the optical depth and microlensing rate computation for the OGLE-III data, which will further allow for the verification of the models of the Milky Way.

OGLE-IV survey during its first three years of operation has already detected about 5000 events, reported by the EWS. An analysis similar to the one presented here on the archival OGLE-IV data will supplement the catalog of events from the OGLE-III and form a colossal sample of standard microlensing events, spread over a wide range of galactic coordinates, allowing for very detailed studies of the inner parts of the Galaxy. OGLE-IV is hoped to continue its functioning for many years

ahead hence there is room for synergy with the ESA's corner stone astrometric mission Gaia. This will create a first time ever opportunity to combine astrometry and photometry and derive masses and distances to individual lenses in microlensing events (Belokurov & Evans 2002, Wyrzykowski & Hodgkin 2012), bringing a completely new perspective to studies of the Galactic structure.

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