

VStar

User Manual



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Introduction

VStar is a free, open source, multi-platform application for visualizing and analyzing time-series data. It is primarily intended for use with variable star data, but any data that can be represented in terms of a Julian Date and magnitude (or other numerical values) is suitable.



The following figure shows that VStar is able to load observations from a variety of sources, including the AAVSO International Database (AID), files, and other sources (such as http streams). The *Observation Sources* section describes what kinds of data can be loaded into VStar and from what sources.



When in “online mode” (loading from AID), VStar also requests information from the Variable Star Index (VSX) about the requested object, such as period, variable type, spectral type and discoverer.

VStar allows data to be viewed and filtered via plots or tables. The details of each observation can be viewed individually or in tabular form. Data and plots can be saved or printed. A mean plot can counteract the effect of noise in data and provide an indication of the extent of signal present.

Phase plots can be created, assuming a certain period, revealing the shape of a star’s light curve in a way that a raw data plot may not.

Period analysis can be applied to reveal one or more periods in the data. Models can be created to represent the essence of a light curve and to subtract one or more periods from the data, the “residual data” being amenable to further analysis. Time-frequency analysis permits changes in period over time to be discovered.

VStar is written in the Java programming language, requiring at least version 1.6 to run on Windows, Mac OS X, Linux, and OpenSolaris. New functionality can be added via plug-ins (from a plug-in library) or can be written by anyone with knowledge of Java or other languages that target the Java Virtual Machine.

The following figure shows the main window’s toolbar and tabs.



The two toolbar buttons below switch between “raw data mode” (or “raw mode” or “light curve mode”) and “phase plot mode” (or “phase mode”).



The tabs labeled Plot, Observations, Means, Model, and Residuals show the same data in different forms or reveal different facets of the underlying data.

In both raw and phase mode the Plot pane contains a multi-series plot of the loaded dataset along with any “synthetic” series, such as means, model, residuals, or filtered observations. Various attributes of the plot can be customized. See the *Plot Pane* section for more detail.

The Observations pane reflects the Plot pane in tabular form. It also provides the ability to search the data and create filtered subsets from search results or a manual selection of observations. See the *Observation List* section for more detail.

The remaining 3 panes – Means, Model, and Residuals – show “synthetic” data in simple tabular form.

The remainder of this document describes different aspects of VStar’s functionality. To learn more about the context in which VStar was developed, see Benn (2012).

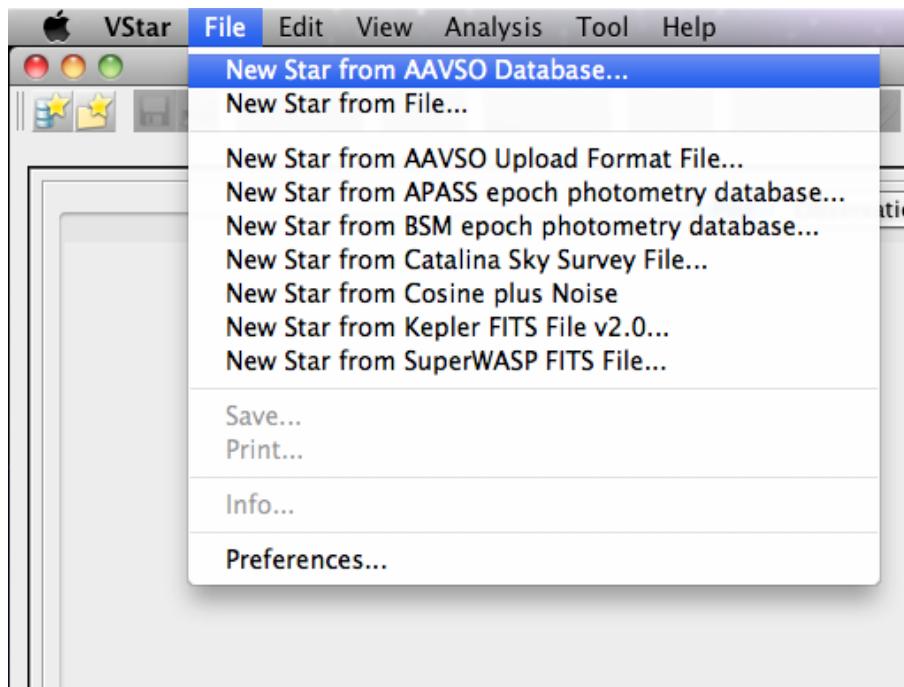
Observation Sources

Observations can be loaded into VStar from a variety of sources. Out of the box, observations can be loaded from the AAVSO International Database (AID) or files conforming to one of two formats (described in the *Load from File* section).

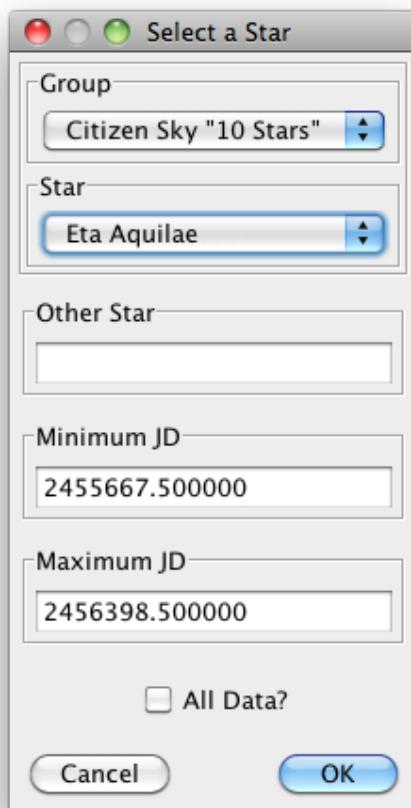
Observations from other sources (e.g. other AAVSO sources, Kepler, SuperWASP, ASAS) can be loaded via plug-ins found in the AAVSO Plug-in Library (<http://www.aavso.org/vstar-plugin-library>).

Load from AID

To load a set of observations from the AID, ensure you are connected to the Internet, then select **New Star from AAVSO Database...** from the File menu or click the button at the left of the toolbar.



The following dialog will appear:



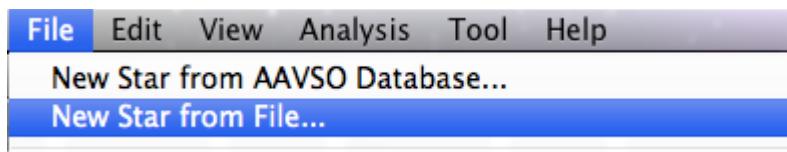
If a name or AAVSO Unique ID (AUID) of a variable star is entered into the Other Star text box, this will be used instead of the currently selected object in the Star drop-down. The Citizen Sky group (see the Group drop-down) contains all stars from the “10 star” tutorial. Other groups can be created via the Preferences dialog.

Either a Julian Date range must be specified or all data requested. This determines the time range over which data should be loaded.

If the requested object does not exist or there is no data for the object in AID or for the specified date range, a message box will appear to that effect, otherwise the available data will be loaded, the light curve plot, observation list, and mean series created. If visual or V band data exists, corresponding observations will be initially visible; this can subsequently be changed via the Plot Control Dialog (see View menu). Otherwise, the band with the most observations will be displayed initially.

Load from File

To load observations from an AAVSO download format file or “simple” file, select the **New Star from File...** menu item from the **File** menu or click the second-to-left-most button in the toolbar.



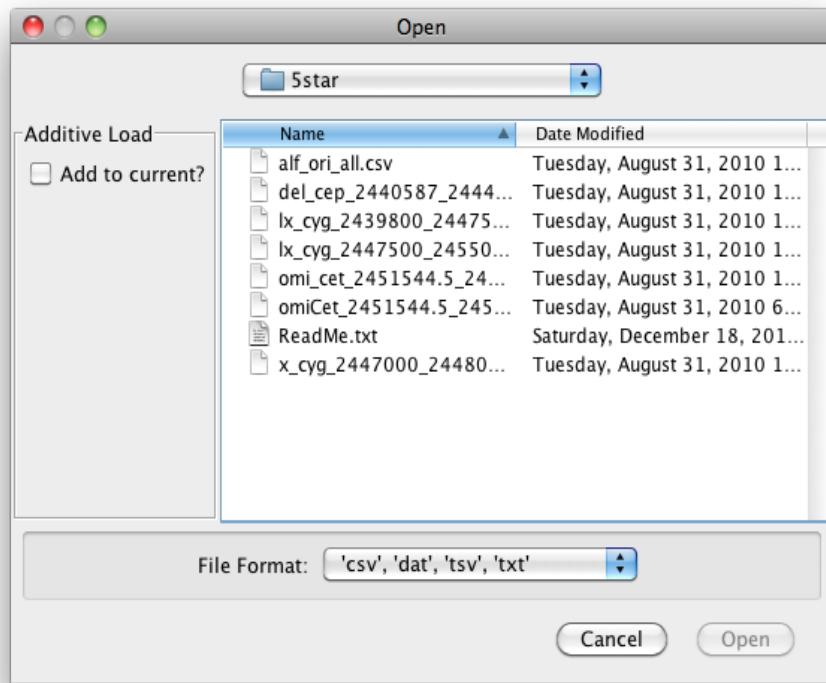
Files must conform to one of two formats:

- AAVSO download format
 - <http://www.aavso.org/data-download>
 - <http://www.aavso.org/format-data-file>
- Simple format, consisting of lines of the form:

Julian Date,Magnitude[,Uncertainty][,Observer Code][,Validation Flag]
or simply:
Julian Date,Magnitude

In both cases, lines can either be comma or tab separated. Observations of other kinds can be loaded via observation source plug-ins. See the AAVSO plug-in library for more.

Selecting this menu item invokes the following file chooser:



Selecting a conformant file and clicking the `Open` button will load the observations, leading to the same result as a Load from AID operation. If the file is not of a suitable format, an error dialog will be displayed.

Additive Loads

If the `Add to current?` checkbox is selected in the file chooser, the loaded observations will be added to any already loaded, irrespective the object. In this way, observations from different sources can be combined in a single VStar session. AID load is currently not additive, but this restriction is likely to be lifted in future.

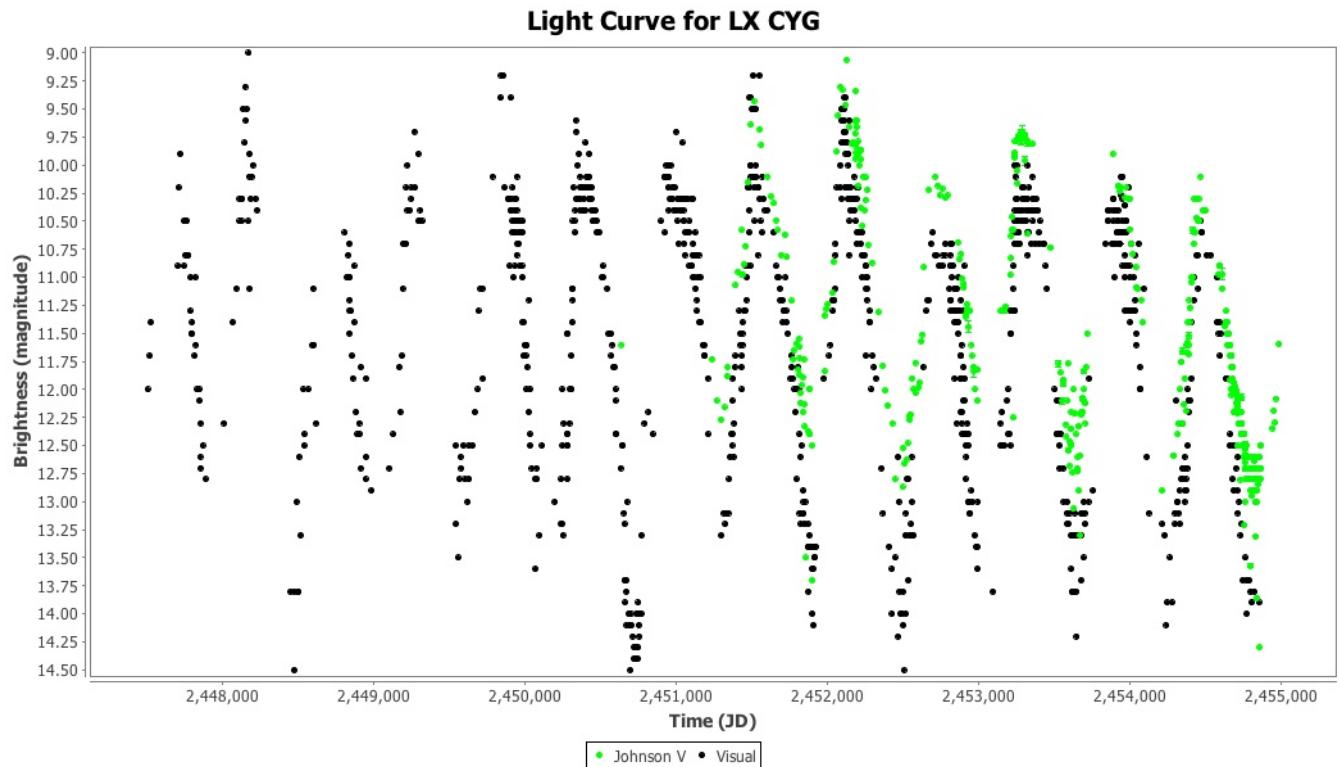
If the newly loaded observations use Heliocentric Julian Dates (HJD), the previously loaded observations will be converted to use HJD or vice versa. Currently, the first data set loaded must be from the AID, since the object's Right Ascension (RA) and Declination (Dec) are necessary for the HJD conversion and for an AID load, VSX is consulted to retrieve the RA and Dec. In future, a dialog may be used to permit the entry of the RA and Dec for non-AID loads.

Further, AAVSO Upload Format files (used for the online WebObs tool; see AAVSO Plug-in Library) and ASAS are currently the only supported observation source that can contain explicit HJD observations, so the data must be loaded from AID and then from one of these file types for HJD conversion to occur. If HJD conversion takes place for a given additive load, a dialog will

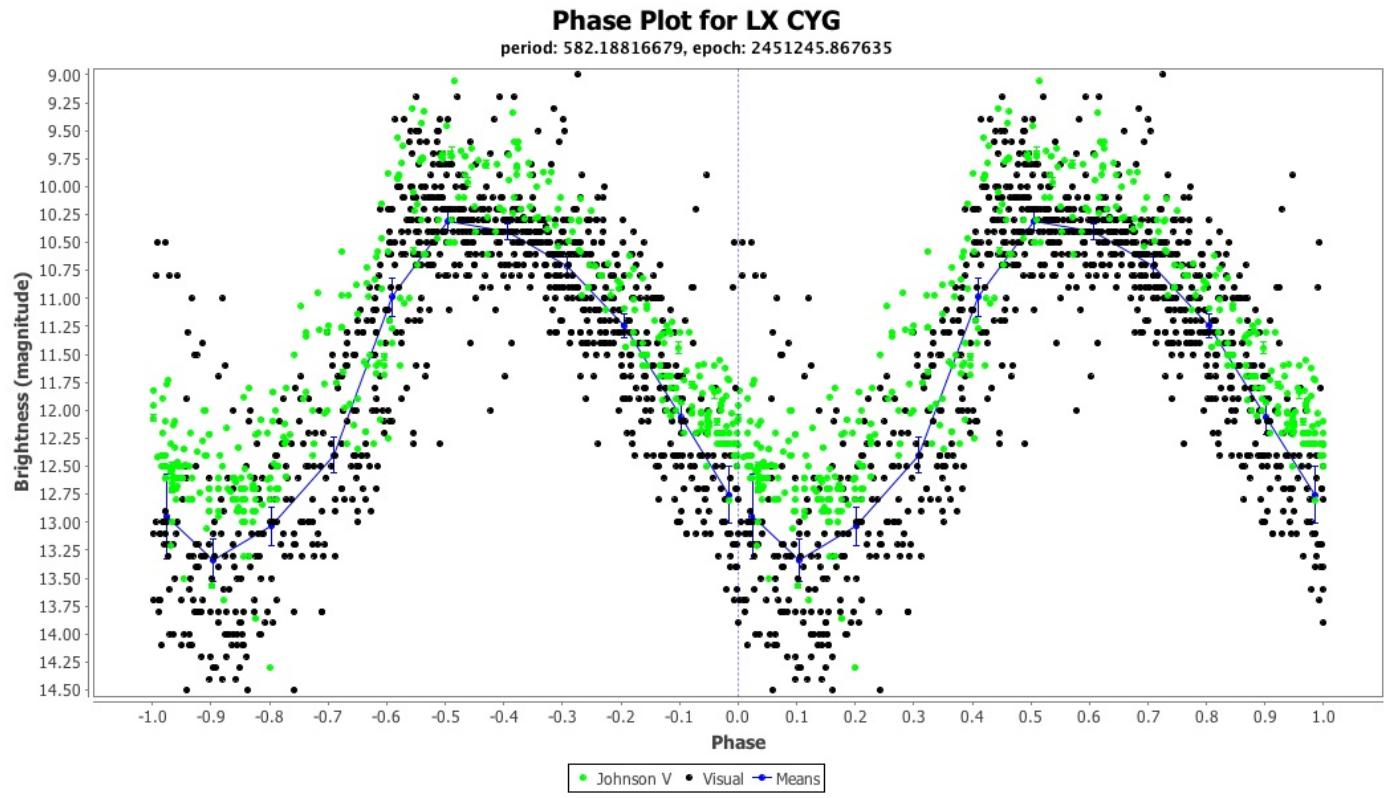
inform the user of this.

The Plot Pane

The Plot Pane provides a view of the currently loaded dataset as a scatterplot of observations, graphing brightness vs time (or phase). This is referred to as a “light curve”. Here is a light curve for LX Cyg:



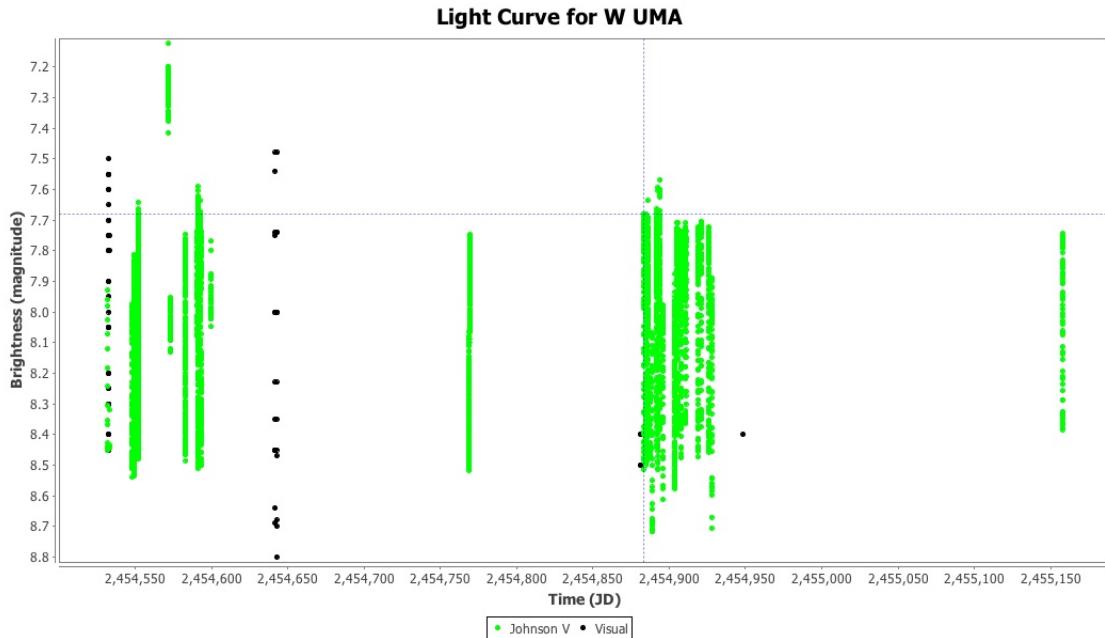
The same dataset in phase plot mode, folded on a period of ~582.188 days, is shown below:



Other than the fact that this shows a phase plot, the other differences from the light curve view are that:

- A mean series has been added. See *Plot Control Dialog* section.
- Period and epoch are shown as a sub-title.
- The X axis is labelled “Phase” instead of “Time (JD)”.

Selecting a data point with the mouse sets the cross-hairs, e.g.

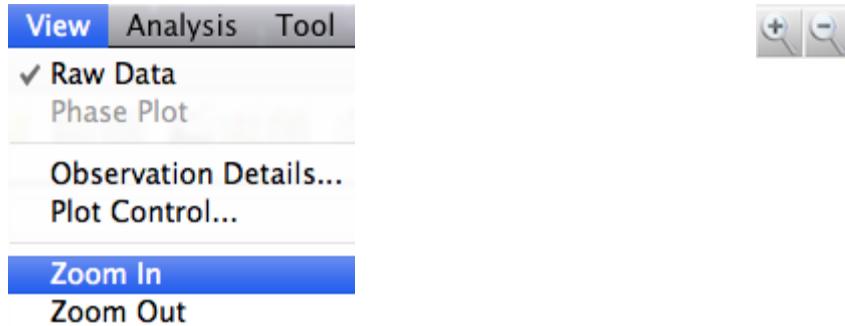


Once selected, a number of operations can be carried out:

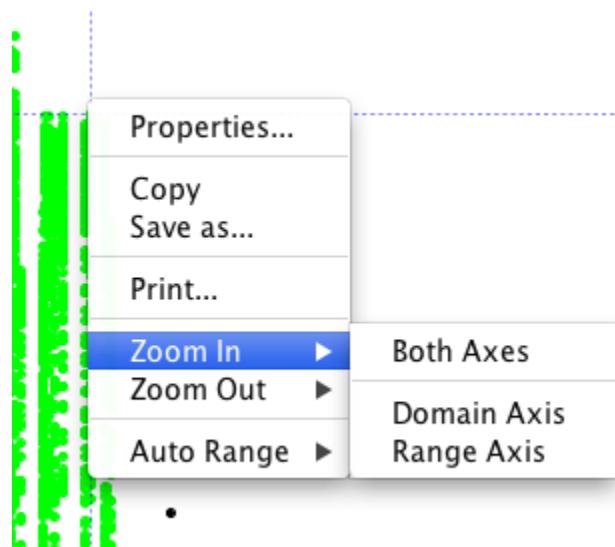
- the observation can be excluded via `Edit → Exclude Selection`;
- the details of the observation can be viewed via the `Observation Details Dialog` or `Observation List` (see the corresponding sections for more detail);
- attributes of the selected observation can be made available for observation filtering (see the *Filtering Observations* section);
- the Plot Pane can be zoomed or panned; these operations are covered next.

On the W UMa plot above, it would be useful to be able to see more detail on a smaller time scale. This is possible in a number of ways:

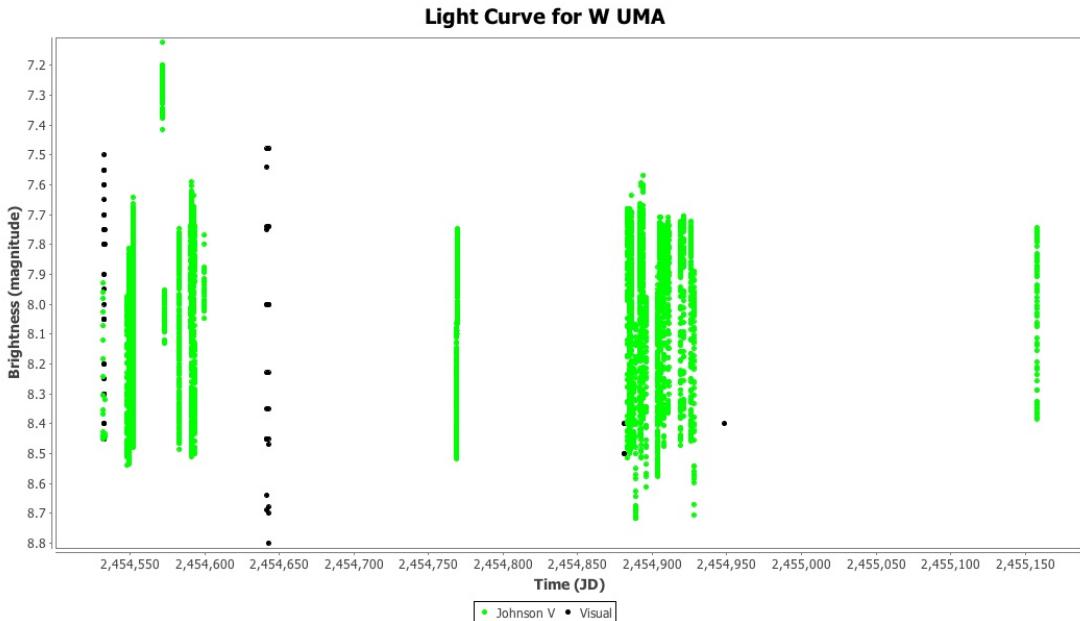
- Drawing a bounding box with the mouse (click and drag) around an area of the plot. To zoom out again, a click-drag-left mouse gesture is used.
- Selecting a point on the plot then using the zoom buttons in the toolbar or the `View` menu's `Zoom In` and `Zoom Out` items:



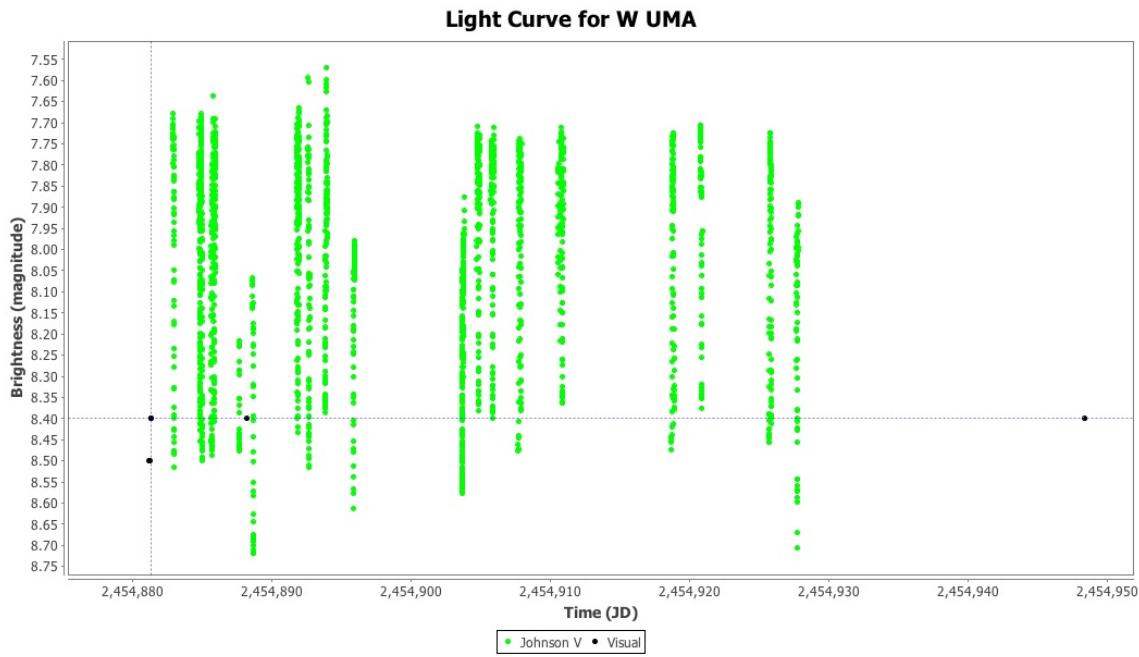
- Use the zoom items of the context menu (e.g. by right-clicking on Windows, ctrl-click on Mac OS X):



Consider the following W UMa light curve.

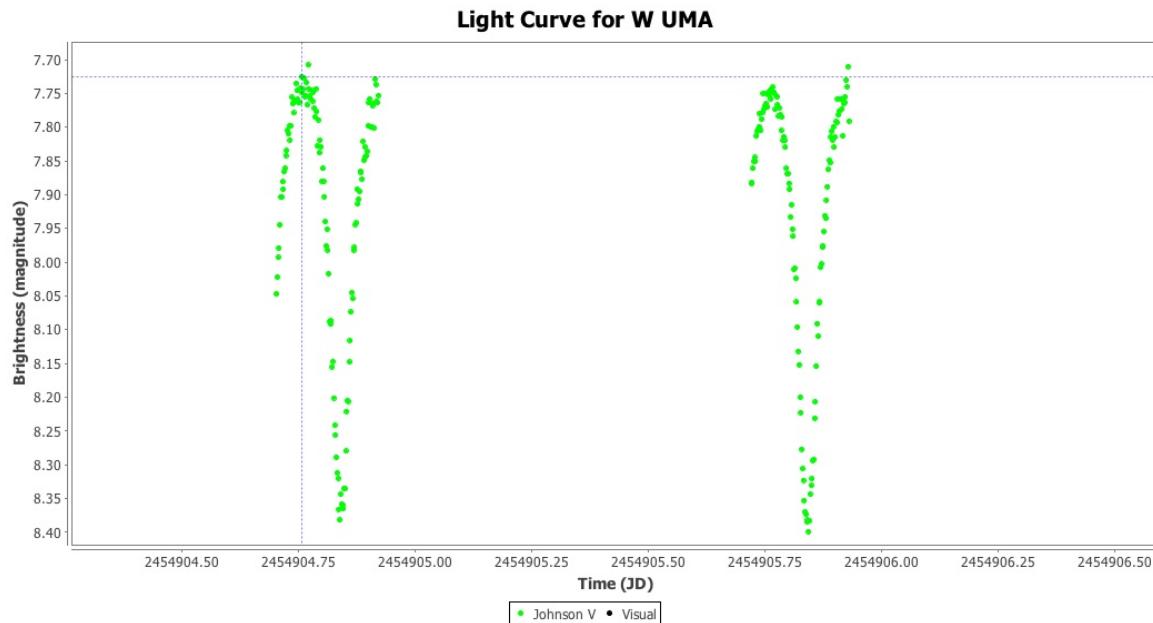


Using the first method above, drawing a bounding box around the section near the cross-hairs to the lower left (JD range: 2454881 to 2454949, magnitude range: 7.5 to 8.75), gives a view like this:

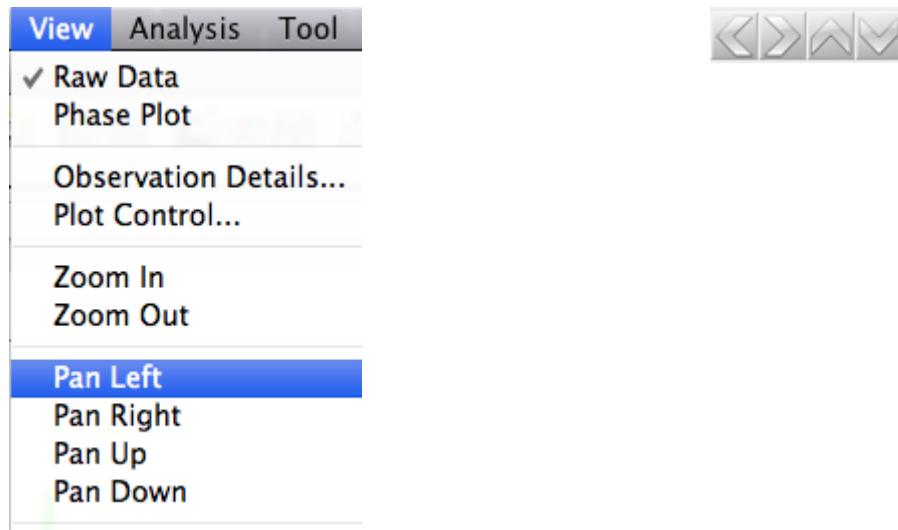


Even at this “magnification” level, one can begin to glimpse that something other than vertical lines exists in this plot. Zooming in further, this time from around JD 2454904 to 2454905

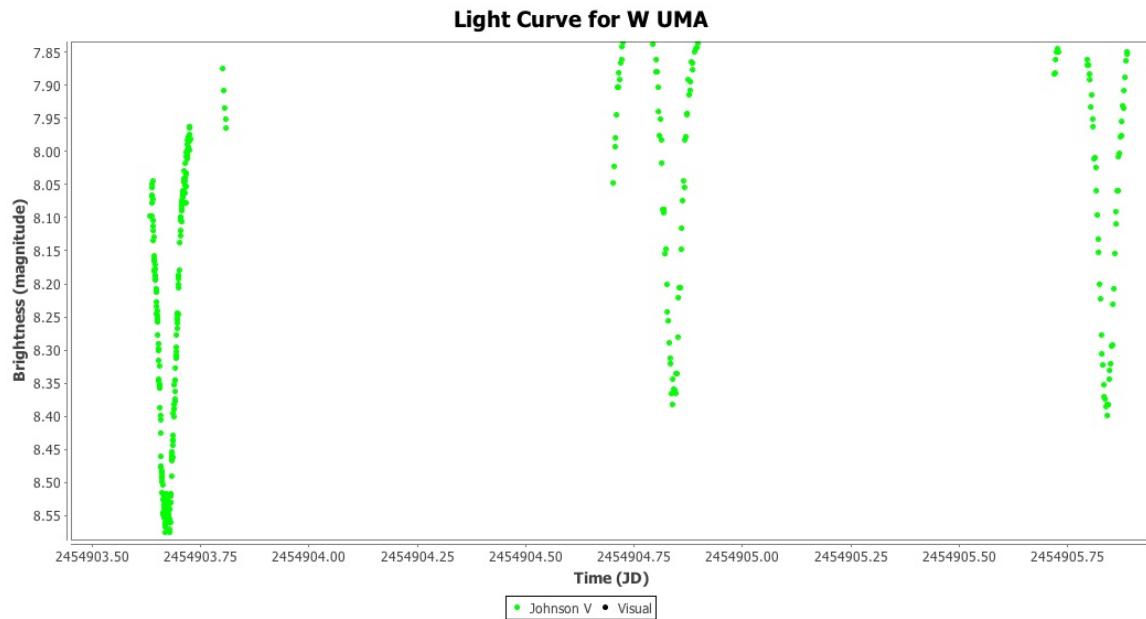
(magnitude 7.7 to 8.4), the following is revealed:



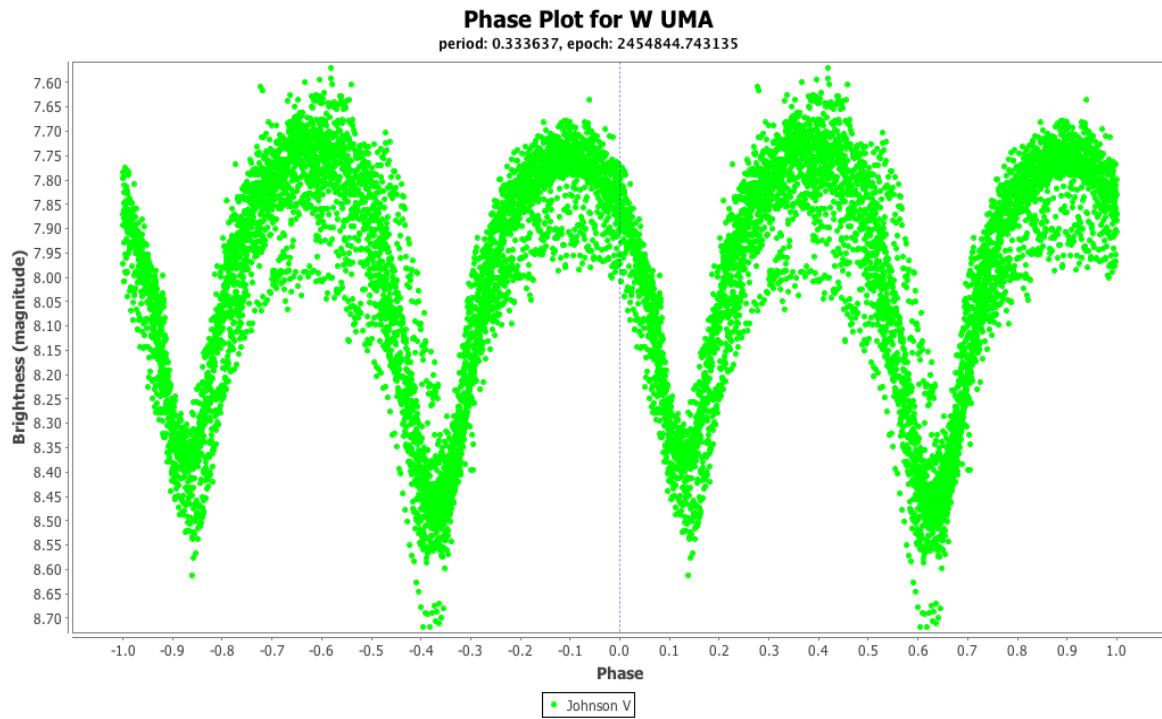
When looking at a magnified view, the pan left, right, up, down View menu items or toolbar buttons can be used to move to regions of the plot not visible.



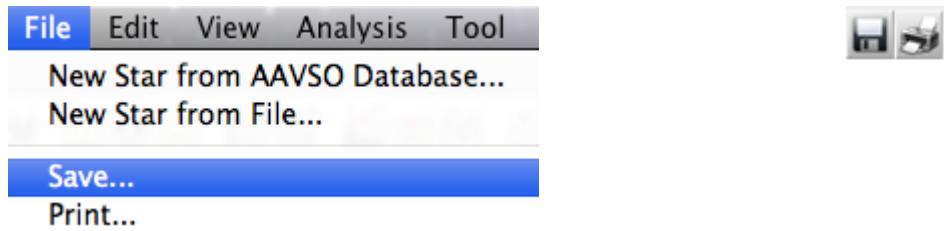
For example, the following is the result of panning the above plot left and down several times:



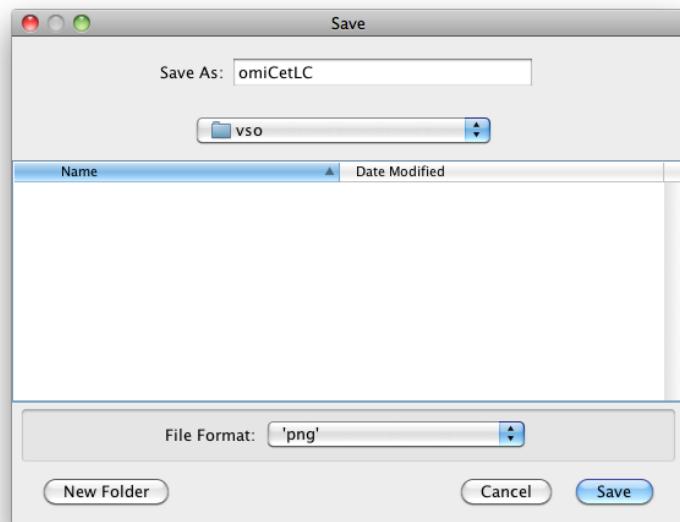
The following shows a phase plot of this dataset at a period (from VSX) of 0.33363749 days:



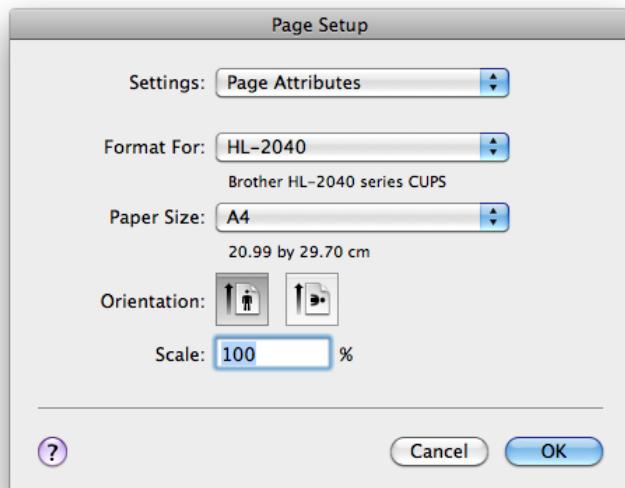
The contents of the Plot Pane can be saved to a file as a PNG image or printed via the **File** menu's **Save...** and **Print...** items or the corresponding toolbar buttons.



An example of the Save Dialog is shown below. All plot image files are saved in PNG format. If the filename entered does not have a “.png” or “.PNG” suffix, VStar will add this. In the example below, when the Save button is clicked, the current plot will be saved to a file called “omiCetLC.png”.



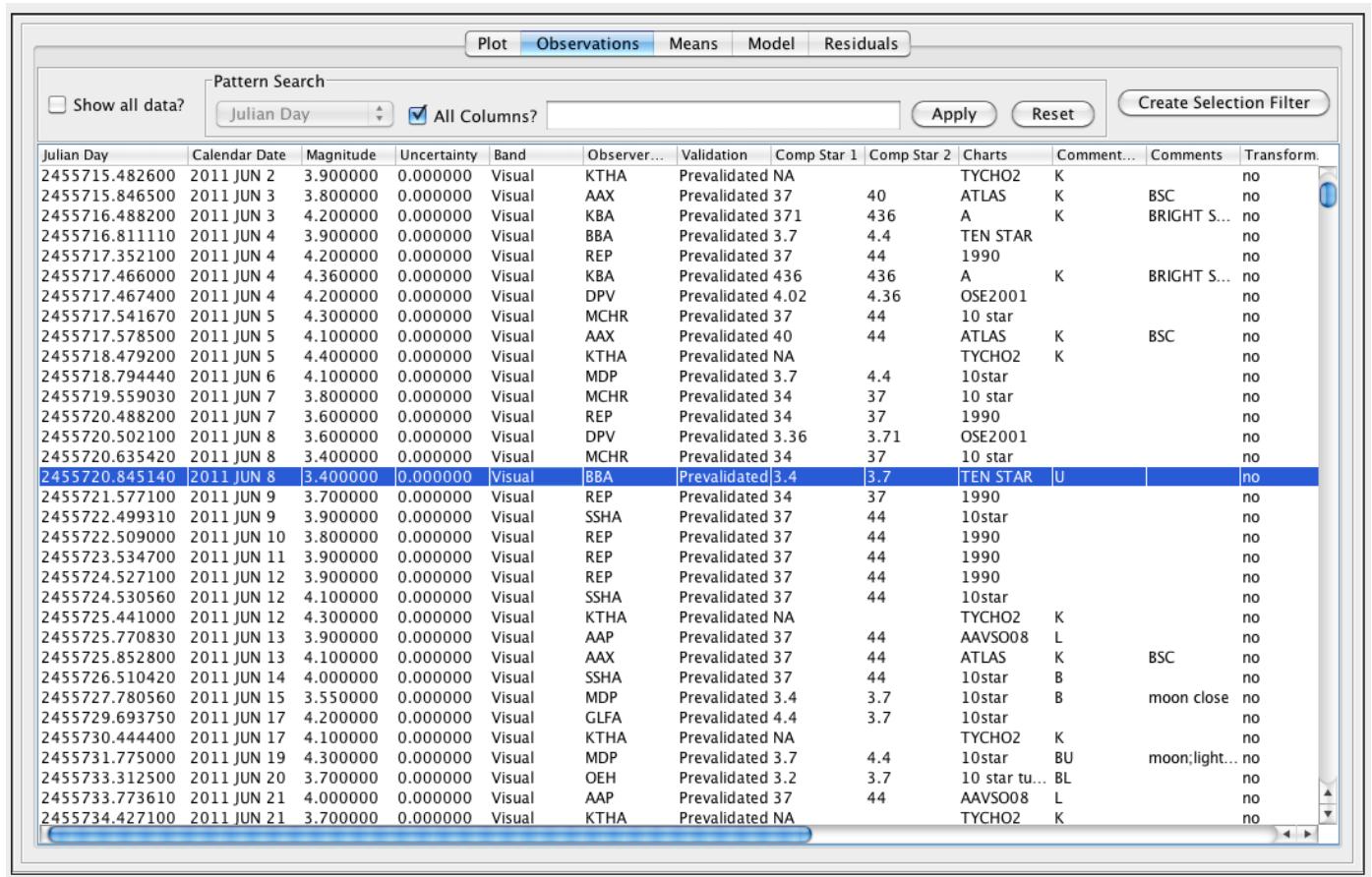
The Print Dialog looks like this on a Mac OS X machine:



Observation List

VStar's observation list is selected via the **Observations** tab in the main window and consists of a tabular view of the loaded observations and a Search Creation pane.

The following shows the observations list with eta Aql observations loaded from AID.



The screenshot shows the VStar software interface with the 'Observations' tab selected. At the top, there are tabs for Plot, Observations, Means, Model, and Residuals. Below the tabs is a search bar labeled 'Pattern Search' with options to 'Show all data?' and 'All Columns?'. The main area is a table of observations with the following columns:

Julian Day	Calendar Date	Magnitude	Uncertainty	Band	Observer...	Validation	Comp Star 1	Comp Star 2	Charts	Comment...	Comments	Transform.
2455715.482600	2011 JUN 2	3.900000	0.000000	Visual	KTHA	Prevalidated NA			TYCHO2	K	BSC	no
2455716.488200	2011 JUN 3	3.800000	0.000000	Visual	AAX	Prevalidated 37	40	436	ATLAS	K	BRIGHT S...	no
2455717.352100	2011 JUN 4	4.200000	0.000000	Visual	KBA	Prevalidated 371			A	K		
2455717.466000	2011 JUN 4	4.360000	0.000000	Visual	KBA	Prevalidated 436	436	44	TEN STAR			
2455717.467400	2011 JUN 4	4.200000	0.000000	Visual	DPV	Prevalidated 4.02	4.36		OSE2001			
2455717.541670	2011 JUN 5	4.300000	0.000000	Visual	MCHR	Prevalidated 37	44		10 star			
2455717.578500	2011 JUN 5	4.100000	0.000000	Visual	AAX	Prevalidated 40	44		ATLAS	K	BSC	no
2455718.479200	2011 JUN 5	4.400000	0.000000	Visual	KTHA	Prevalidated NA			TYCHO2	K		
2455718.794440	2011 JUN 6	4.100000	0.000000	Visual	MDP	Prevalidated 3.7	4.4		10star			
2455719.559030	2011 JUN 7	3.800000	0.000000	Visual	MCHR	Prevalidated 34	37		10 star			
2455720.488200	2011 JUN 7	3.600000	0.000000	Visual	REP	Prevalidated 34	37		1990			
2455720.502100	2011 JUN 8	3.600000	0.000000	Visual	DPV	Prevalidated 3.36	3.71		OSE2001			
2455720.635420	2011 JUN 8	3.400000	0.000000	Visual	MCHR	Prevalidated 34	37		10 star			
2455720.845140	2011 JUN 8	3.400000	0.000000	Visual	BBA	Prevalidated 3.4	3.7		TEN STAR	U		no
2455721.577100	2011 JUN 9	3.700000	0.000000	Visual	REP	Prevalidated 34	37		1990			
2455722.499310	2011 JUN 9	3.900000	0.000000	Visual	SSHA	Prevalidated 37	44		10star			
2455722.509000	2011 JUN 10	3.800000	0.000000	Visual	REP	Prevalidated 37	44		1990			
2455723.534700	2011 JUN 11	3.900000	0.000000	Visual	REP	Prevalidated 37	44		1990			
2455724.527100	2011 JUN 12	3.900000	0.000000	Visual	REP	Prevalidated 37	44		1990			
2455724.530560	2011 JUN 12	4.100000	0.000000	Visual	SSHA	Prevalidated 37	44		10star			
2455725.441000	2011 JUN 12	4.300000	0.000000	Visual	KTHA	Prevalidated NA			TYCHO2	K		
2455725.770830	2011 JUN 13	3.900000	0.000000	Visual	AAP	Prevalidated 37	44		AAVS008	L		
2455725.852800	2011 JUN 13	4.100000	0.000000	Visual	AAX	Prevalidated 37	44		ATLAS	K	BSC	no
2455726.510420	2011 JUN 14	4.000000	0.000000	Visual	SSHA	Prevalidated 37	44		10star	B		
2455727.780560	2011 JUN 15	3.550000	0.000000	Visual	MDP	Prevalidated 3.4	3.7		10star	B	moon close	no
2455729.693750	2011 JUN 17	4.200000	0.000000	Visual	GLFA	Prevalidated 4.4	3.7		10star			
2455730.444400	2011 JUN 17	4.100000	0.000000	Visual	KTHA	Prevalidated NA			TYCHO2	K		
2455731.775000	2011 JUN 19	4.300000	0.000000	Visual	MDP	Prevalidated 3.7	4.4		10star	BU	moon;light...	no
2455733.312500	2011 JUN 20	3.700000	0.000000	Visual	OEH	Prevalidated 3.2	3.7		10 star tu...	BL		
2455733.773610	2011 JUN 21	4.000000	0.000000	Visual	AAP	Prevalidated 37	44		AAVSO08	L		
2455734.427100	2011 JUN 21	3.700000	0.000000	Visual	KTHA	Prevalidated NA			TYCHO2	K		

The most obvious feature is the tabular list of observations consisting of columns relating to each observation, e.g. Magnitude, Band, Comp Star 1. The figure shows a row having been selected by the user. Switching to the Plot tab would reveal that the cross-hair was pointing to the same observation in the light curve or phase plot.

Columns can be rearranged by dragging the column headers, and observations can be sorted by column by single-clicking on the column header. In any mode (raw or phase plot), the default is for observations to be sorted by Julian Date.

The set of columns that appear in the observation list varies with the observation source (e.g. AID, files of particular type).

Something not obvious at first glance is that the observation list mirrors the plot. In other words, if a particular band/series is visible on the plot, it will also be visible in the observation list. It is possible to override this by selecting the `Show all data?` checkbox. Note this will not affect the plot. The same is true for `Pattern Search` feature. Typing a *regular expression* into the text box to the left of the `Apply` button, followed by clicking that button will:

- Look for a match across values in all columns.
- Filter the list based upon the match.

If no match is found, no observations will be displayed. Clicking the `Reset` button will restore the normal observation list. Note that when `Show all data?` is selected, the `Search` pane is disabled.

The following shows a pattern search for the word “moon”, resulting in a listing that consists only of observations in which the word “moon” appears somewhere, specifically: in the `Comments` column. In this example, the comments column needs to be expanded to see some occurrences of the pattern, e.g. those near the middle beginning with “Clear. Wa...” when expanded, show “Clear. Waxing moon.”

Plot Observations Means Model Residuals

Show all data? All Columns? moon Create Selection Filter

Julian Day	Calendar Date	Magnitude	Uncertainty	Band	Observer ...	Validation	Comp Star 1	Comp Star 2	Charts	Comment...	Comments	Transform.
2455727.780560	2011 JUN 15	3.550000	0.000000	Visual	MDP	Prevalidated 3.4	3.7	10star	B	moon close no		
2455731.775000	2011 JUN 19	4.300000	0.000000	Visual	MDP	Prevalidated 3.7	4.4	10star	BU	moon;light... no		
2455755.718750	2011 JUL 13	3.600000	0.000000	Visual	MDP	Prevalidated 3.4	3.7	10star	B	moon close no		
2455756.729170	2011 JUL 14	3.600000	0.000000	Visual	MDP	Prevalidated 3.4	3.7	10star	B	moon very... no		
2455757.697220	2011 JUL 15	4.100000	0.000000	Visual	MDP	Prevalidated 3.7	4.4	10star	B	full moon... no		
2455763.690970	2011 JUL 21	3.650000	0.000000	Visual	MDP	Prevalidated 3.7	3.4	10star	BU	moon;smog no		
2455765.766670	2011 JUL 23	3.650000	0.000000	Visual	MDP	Prevalidated 3.7	3.4	10star	BU	moon;smog no		
2455790.621530	2011 AUG 17	4.200000	0.000000	Visual	MDP	Prevalidated 3.7	4.4	10star	BU	moon;haze no		
2455813.583330	2011 SEP 9	3.550000	0.000000	Visual	MDP	Prevalidated 34	37	10startuto... B		close to m... no		
2455814.600000	2011 SEP 10	3.600000	0.000000	Visual	MDP	Prevalidated 3.4	3.7	10star	B	moon near no		
2455816.637500	2011 SEP 12	4.200000	0.000000	Visual	MDP	Prevalidated 3.7	4.4	10star	B	full moon no		
2455818.651390	2011 SEP 14	4.600000	0.000000	Visual	MDP	Prevalidated 4.4	4.7	not sure	B	moon no		
2455835.544440	2011 OCT 1	3.400000	0.000000	Visual	KJMA	Prevalidated 32	37	Star Train... B		Clear. Wa... no		
2455836.513890	2011 OCT 2	3.650000	0.000000	Visual	MDP	Prevalidated 3.7	3.4	10star	B	moon near no		
2455837.135420	2011 OCT 2	3.300000	0.000000	Visual	KJMA	Prevalidated 32	34	Star Train... B		Clear. Wa... no		
2455838.121530	2011 OCT 3	4.200000	0.000000	Visual	KJMA	Prevalidated 37	44	Star Train... B		Clear. Wa... no		
2455839.648610	2011 OCT 5	4.450000	0.000000	Visual	MDP	Prevalidated 44	47	10star;oth... B		moon clos... no		
2455840.604170	2011 OCT 6	4.550000	0.000000	Visual	MDP	Prevalidated 4.4	4.7	10star+fo... B		moon no		
2455842.654170	2011 OCT 8	3.600000	0.000000	Visual	MDP	Prevalidated 34	37	10startuto... B		moon close no		
2455850.690970	2011 OCT 16	3.600000	0.000000	Visual	EJDA	Prevalidated 37	34	10 star		moon no		
2455867.517360	2011 NOV 2	4.000000	0.000000	Visual	MDP	Prevalidated 37	44	10startuto... B		moon no		
2455871.510420	2011 NOV 6	3.650000	0.000000	Visual	MDP	Prevalidated 34	37	10startuto... B		moon no		
2455899.467360	2011 DEC 3	4.300000	0.000000	Visual	MDP	Prevalidated 37	44	10startuto... BUL		moon; city... no		
2456110.718750	2012 JUL 2	4.000000	0.000000	Visual	MDP	Prevalidated 37	44	10 star	B	moon close no		
2456113.716670	2012 JUL 5	4.200000	0.000000	Visual	MDP	Prevalidated 37	44	10 star	BU	moon close no		
2456114.722920	2012 JUL 6	3.650000	0.000000	Visual	MDP	Prevalidated 37	44	10 star	BU	moon close no		
2456144.493700	2012 AUG 4	3.600000	0.000000	Visual	BPEB	Prevalidated 34	37	10star	B	moon no		
2456144.510900	2012 AUG 5	3.500000	0.000000	Visual	GCDA	Prevalidated 34	44	10 star		Observati... no		
2456168.644440	2012 AUG 29	4.200000	0.000000	Visual	MDP	Prevalidated 44	37	10 star	B	moon close no		
2456200.392360	2012 SEP 29	3.800000	0.000000	Visual	TSP	Prevalidated 37	44	8914BUN	B	Full moon no		
2456249.459720	2012 NOV 17	4.150000	0.000000	Visual	MDP	Prevalidated 37	44	10 star tu... B		moon close no		
2456385.879860	2013 APR 3	4.450000	0.000000	Visual	MDP	Prevalidated 44	47	10 star tu... B		moon no		
2456409.881940	2013 APR 27	3.500000	0.000000	Visual	MDP	Prevalidated 34	37	10 star tu... B		moon, dawn no		

Rather than asking VStar to match across all columns, a particular column can be targeted, as in the following example that looks for observations whose Calendar Date column matches “2011 OCT 31” or “2013 APR 23” via the pattern **2011 OCT 31|2013 APR 23** where the vertical bar character (“|”) means “or” or “alternatively”.

Pattern Search													
<input type="checkbox"/> Show all data?		Julian Day	<input checked="" type="checkbox"/> All Columns?	2011 OCT 31 2013 APR 23		<input type="button" value="Apply"/>	<input type="button" value="Reset"/>	Create Selection Filter					
Julian Day	Calendar Date	Magnitude	Uncertainty	Band	Observer...	Validation	Comp Star 1	Comp Star 2	Charts	Comment...	Comments	Transform...	
2455865.507260	2011 OCT 31	4.529000	0.015000	Johnson B	SAH	Prevalidated ENSEMBLE		6054AK				no	
2455865.510410	2011 OCT 31	3.742000	0.012000	Johnson V	SAH	Prevalidated ENSEMBLE		6054AK				no	
2455865.513530	2011 OCT 31	3.124000	0.010000	Cousins R	SAH	Prevalidated ENSEMBLE		6054AK				no	
2455865.516640	2011 OCT 31	2.915000	0.016000	Cousins I	SAH	Prevalidated ENSEMBLE		6054AK				no	
2455866.281940	2011 OCT 31	4.000000	0.000000	Visual	RJG	Prevalidated 34	41	SA1986				no	
2455866.283330	2011 OCT 31	4.000000	0.000000	Visual	CAI	Prevalidated 34	41	SA1986				no	
2455866.290280	2011 OCT 31	3.780000	0.000000	Visual	VUG	Prevalidated 321	436	SPA/VSS/ BUK				no	
2455866.368800	2011 OCT 31	3.700000	0.000000	Visual	REP	Prevalidated NA		NA				no	
2456405.800000	2013 APR 23	3.800000	0.000000	Visual	AAX	Prevalidated 37	40	ATLAS K	BSC			no	
2456405.812500	2013 APR 23	3.900000	0.000000	Visual	AAX	Prevalidated 37	40	ATLAS K	BSC			no	

The following shows the result of applying the pattern `^\d{2}$` to the Comp Star 2 column so that only those observations whose second comparison star consists of 2 digits will be shown:

Julian Day	Calendar Date	Magnitude	Uncertainty	Band	Observer...	Validation	Comp Star 1	Comp Star 2	Charts	Comment...	Comments	Transform.
2455695.843100	2011 MAY 14	4.100000	0.000000	Visual	AAX	Prevalidated 40	44	ATLAS	K	BSC	no	
2455698.868100	2011 MAY 17	3.600000	0.000000	Visual	AAX	Prevalidated 33	37	ATLAS	BK	MOON BSC	no	
2455704.861800	2011 MAY 23	3.700000	0.000000	Visual	AAX	Prevalidated 33	37	ATLAS	K	BSC	no	
2455705.675700	2011 MAY 24	3.600000	0.000000	Visual	REP	Prevalidated 34	37	1990			no	
2455709.396500	2011 MAY 27	4.100000	0.000000	Visual	REP	Prevalidated 37	44	1990			no	
2455709.866000	2011 MAY 28	3.900000	0.000000	Visual	AAX	Prevalidated 37	40	ATLAS	K	BSC	no	
2455710.427080	2011 MAY 28	4.300000	0.000000	Visual	KSQ	Prevalidated 37	44	AAVSO			no	
2455710.489600	2011 MAY 28	4.400000	0.000000	Visual	REP	Prevalidated 37	44	1990			no	
2455711.427100	2011 MAY 29	4.500000	0.000000	Visual	REP	Prevalidated 44	47	1990			no	
2455711.838200	2011 MAY 30	3.900000	0.000000	Visual	AAX	Prevalidated 37	40	ATLAS	K	BSC	no	
2455712.378500	2011 MAY 30	4.300000	0.000000	Visual	REP	Prevalidated 37	44	1990			no	
2455712.857600	2011 MAY 31	3.500000	0.000000	Visual	AAX	Prevalidated 33	37	ATLAS	K	BSC	no	
2455713.848600	2011 JUN 1	3.400000	0.000000	Visual	AAX	Prevalidated 33	37	ATLAS	K	BSC	no	
2455714.848600	2011 JUN 2	3.500000	0.000000	Visual	AAX	Prevalidated 33	37	ATLAS	K	BSC	no	
2455715.846500	2011 JUN 3	3.800000	0.000000	Visual	AAX	Prevalidated 37	40	ATLAS	K	BSC	no	
2455717.352100	2011 JUN 4	4.200000	0.000000	Visual	REP	Prevalidated 37	44	1990			no	
2455717.541670	2011 JUN 5	4.300000	0.000000	Visual	MCHR	Prevalidated 37	44	10 star			no	
2455717.578500	2011 JUN 5	4.100000	0.000000	Visual	AAX	Prevalidated 40	44	ATLAS	K	BSC	no	
2455719.559030	2011 JUN 7	3.800000	0.000000	Visual	MCHR	Prevalidated 34	37	10 star			no	
2455720.488200	2011 JUN 7	3.600000	0.000000	Visual	REP	Prevalidated 34	37	1990			no	
2455720.635420	2011 JUN 8	3.400000	0.000000	Visual	MCHR	Prevalidated 34	37	10 star			no	
2455721.577100	2011 JUN 9	3.700000	0.000000	Visual	REP	Prevalidated 34	37	1990			no	
2455722.499310	2011 JUN 9	3.900000	0.000000	Visual	SSHA	Prevalidated 37	44	10star			no	
2455722.509000	2011 JUN 10	3.800000	0.000000	Visual	REP	Prevalidated 37	44	1990			no	
2455723.534700	2011 JUN 11	3.900000	0.000000	Visual	REP	Prevalidated 37	44	1990			no	
2455724.527100	2011 JUN 12	3.900000	0.000000	Visual	REP	Prevalidated 37	44	1990			no	
2455724.530560	2011 JUN 12	4.100000	0.000000	Visual	SSHA	Prevalidated 37	44	10star			no	
2455725.770830	2011 JUN 13	3.900000	0.000000	Visual	AAP	Prevalidated 37	44	AAVSO08	L		no	
2455725.852800	2011 JUN 13	4.100000	0.000000	Visual	AAX	Prevalidated 37	44	ATLAS	K	BSC	no	
2455726.510420	2011 JUN 14	4.000000	0.000000	Visual	SSHA	Prevalidated 37	44	10star	B		no	
2455733.773610	2011 JUN 21	4.000000	0.000000	Visual	AAP	Prevalidated 37	44	AAVSO08	L		no	
2455734.704860	2011 JUN 22	3.600000	0.000000	Visual	MCHR	Prevalidated 34	37	10 star			no	
2455734.769440	2011 JUN 22	3.600000	0.000000	Visual	AAP	Prevalidated 34	37	AAVSO08	L		no	

The caret (“^”) means “start of string”, dollar (“\$”) means “end of string”, “\d” means decimal digit characters (0 to 9), “{2}” means two of those (digit characters). This reduces to: “comp star 2 must only consist of 2 decimal digits to be included in the observation list.”

Regular expressions can be simple strings or extremely complex patterns, but a VStar user need not be troubled by using anything more complex than he or she requires.

A complete treatment of regular expressions is beyond the scope of this document. See *Regular Expressions* subsection in *References and Further Reading*.

A search will disrupt the synchronization between plot and list, as will clicking the Reset button.

If one or more observations are selected in the list by single-clicking, shift-clicking, or ctrl-clicking, those observations may:

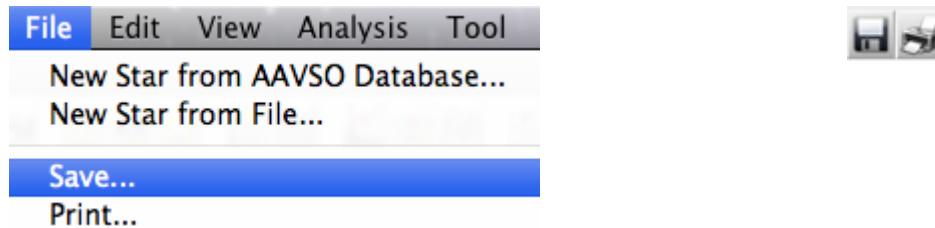
- be excluded via *Edit* → *Exclude Selection*;
- form the basis of an observation filter via the *Create Selection Filter* button.

See the *Filtering Observations* section for more information.

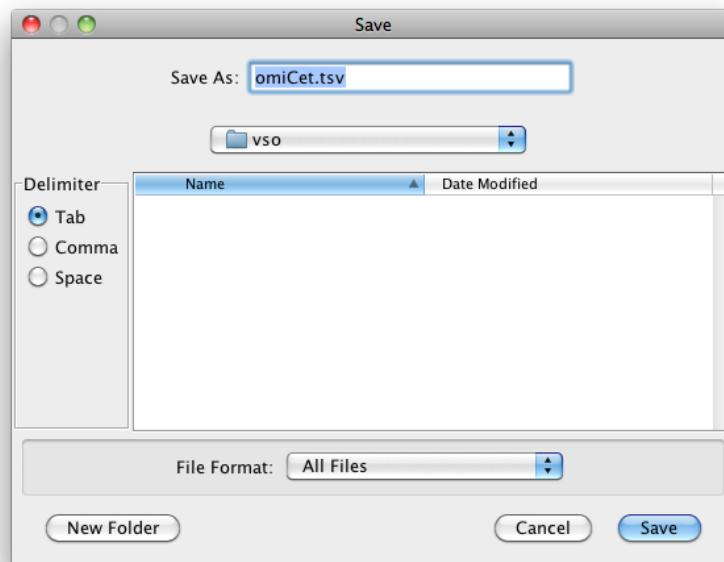
Each Observation List row also contains a **Discrepant?** checkbox indicating whether an observation is “discrepant” and permitting an observation to be reported as discrepant. See the *Observation Details Dialog* section for details.

Comp Star 1	Comp Star 2	Charts	Comment...	Comments	Transformed	Airmass	CMag	KMag	HJD	HQ Uncert...	MType	Discrepant?	Name
37	44	10 star	BL	In city lig...	no						Standard	<input type="checkbox"/>	eta Aql
4.0	4.4	TEN STAR	UB		no						Standard	<input type="checkbox"/>	ETA AQL
3.0	3.5	TEN STAR	B		no						Standard	<input checked="" type="checkbox"/>	ETA AQL
37	44	10star	L	Elevation...	no						Standard	<input type="checkbox"/>	ETA AQL
37	44	10star	U		no						Standard	<input type="checkbox"/>	ETA AQL

The contents of the Observation List can be saved to a text file or printed via the **File** menu’s **Save...** and **Print...** items, or the corresponding toolbar buttons.



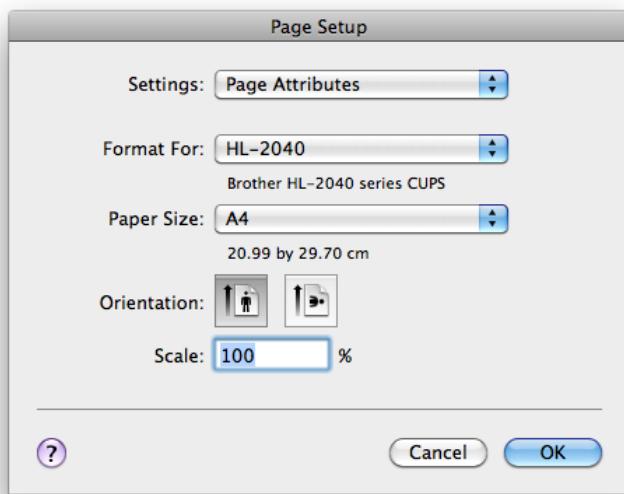
An example of the Save Dialog is shown below.



In the example above, when the Save button is clicked, the current Observation List will be saved to a file called “omiCetLC.tsv”.

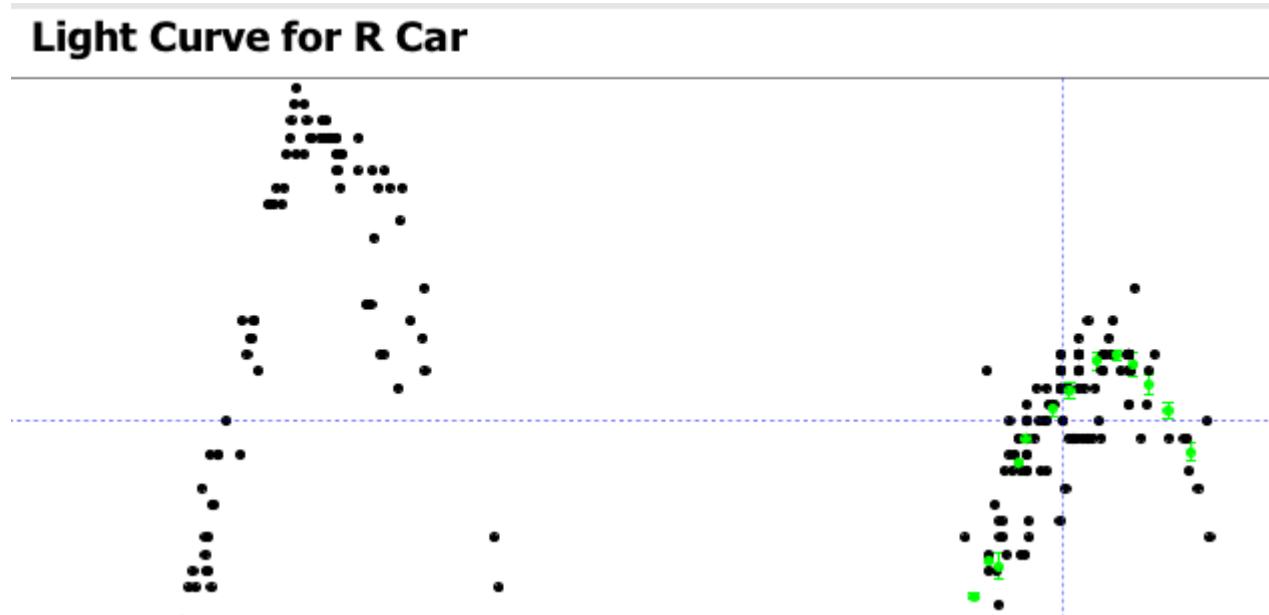
Files can be saved as rows of tab, comma, or space separated values. The first two are preferable over the latter since spaces in some fields lead to ambiguity if loading back into VStar or other software is required. From experience, tabs are less frequent, compared to commas, in fields such as comments (as found in AAVSO download files), so tab separation is a good choice.

The Print Dialog looks like this on a Mac OS X machine:



Observation Details Dialog

The details of an individual observation can be viewed by selecting an observation in the Plot pane such that the cross hair is upon it, e.g.

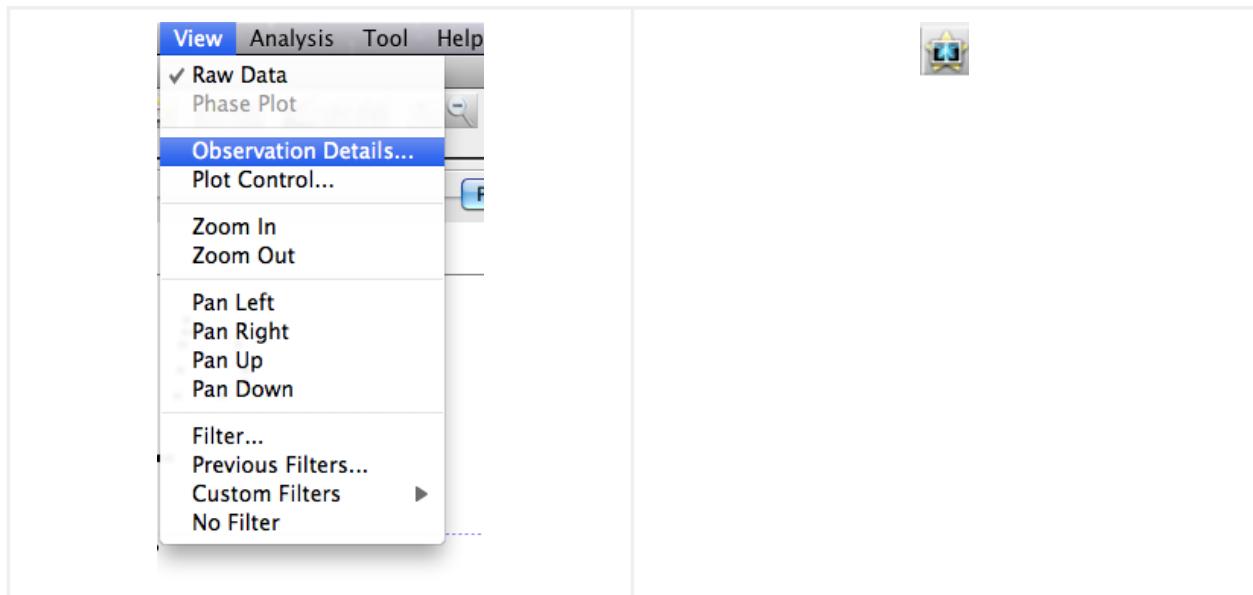


Selecting a row in the observation list is equivalent to this, e.g.

A screenshot of a software interface showing an "Observations" table. The table has columns for Julian Day, Calendar Date, Magnitude, Uncertainty, Band, and Observer. A blue row is selected, corresponding to the data point highlighted in the light curve plot above.

Julian Day	Calendar Date	Magnitude	Uncertainty	Band	Observer...
2456422.291670	2013 MAY 9	5.700000	0.000000	Visual	GEZA
2456422.291670	2013 MAY 9	5.900000	0.000000	Visual	CJMB
2456422.291670	2013 MAY 9	5.800000	0.000000	Visual	TNAA
2456422.333330	2013 MAY 9	5.800000	0.000000	Visual	IPA
2456423.060420	2013 MAY 10	5.900000	0.000000	Visual	PAW
2456423.062860	2013 MAY 10	6.100000	0.000000	Visual	BDJB
2456423.434720	2013 MAY 10	6.100000	0.000000	Visual	RFP

The observation details dialog is invoked from the View menu or via the corresponding toolbar button:



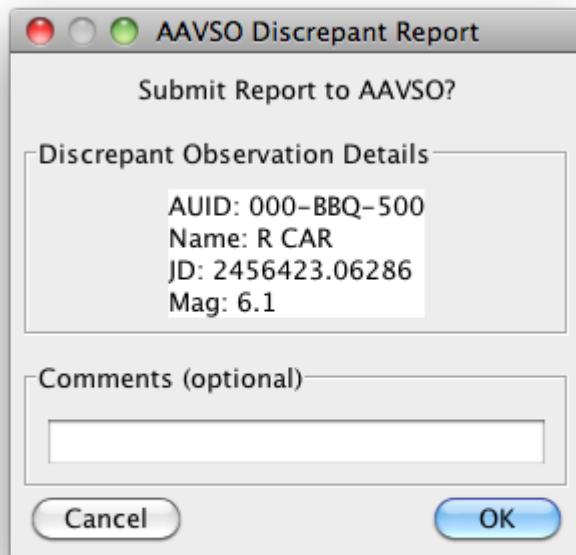
The observation in this example was a visual estimate of R Car made by the author:



If a phase plot has been created, the phase will also be shown. If the time is a Heliocentric Julian Date rather a Julian Date, this will be indicated. Other details will differ depending upon the kind of observation (e.g. visual vs photometric), the observer, comment codes, band, and so on.

Note that the same information is available in the Observation List; however it can be useful to have multiple observation details dialogs open simultaneously, especially when observations are separated widely in time.

Selecting the `Discrepant?` checkbox relegates the observation to the discrepant series. This can be made visible on the plot via the Plot Control Dialog. If the dataset was loaded from the AID, another dialog will be opened, asking whether to submit a discrepant report to AAVSO HQ, e.g.

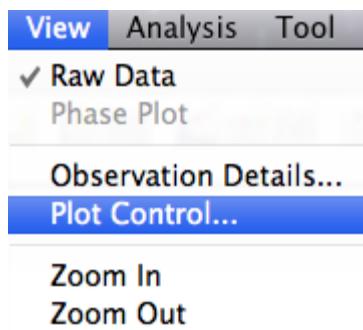


Of course, this particular observation is fine and no discrepant report was actually submitted. Before the observation report is sent, you will be asked to authenticate with your AAVSO web credentials.

This same Discrepant checkbox and reporting capability also exists in each Observation List row.

Plot Control Dialog

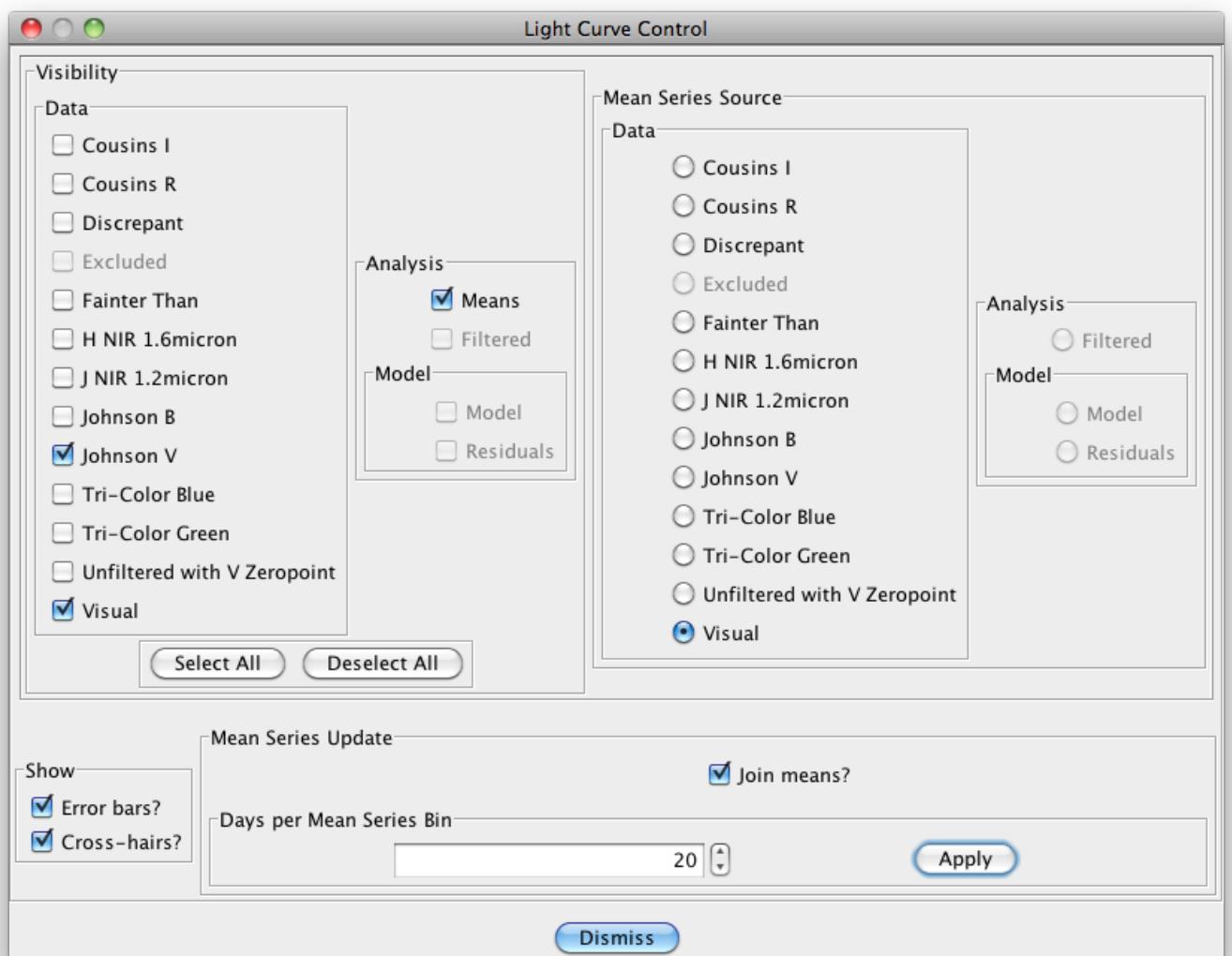
Selecting the View → Plot Control... item:



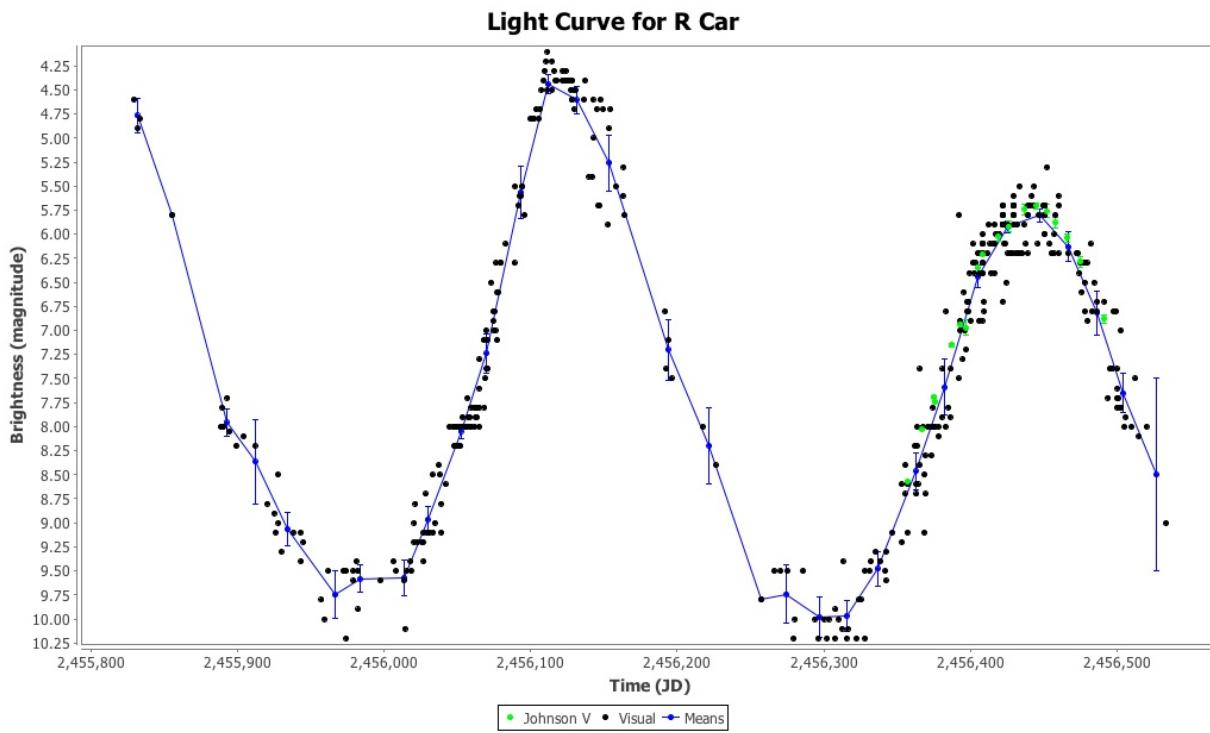
opens the plot control dialog. The example below shows the Plot Control Dialog resulting from an AID load of R Car. The dialog has a few distinct features:

- Visibility checkboxes. Toggling these immediately changes the series that are visible on the plot and in the observation list. Select All and Deselect All buttons make it easier to work with many series.
- Mean series source radio buttons. Selecting a radio button changes the series that is used to create a mean series. If the mean series is not already visible, changing this selection makes it so. A mean series source change is acted upon immediately.
- Checkboxes for showing or hiding error bars or cross-hairs.
- Binned mean series update:
 - Checkbox to join mean data points.
 - Spinner and text-box for number of days or phase steps per bin. The Apply button must be clicked for a change to this value to be acted upon.

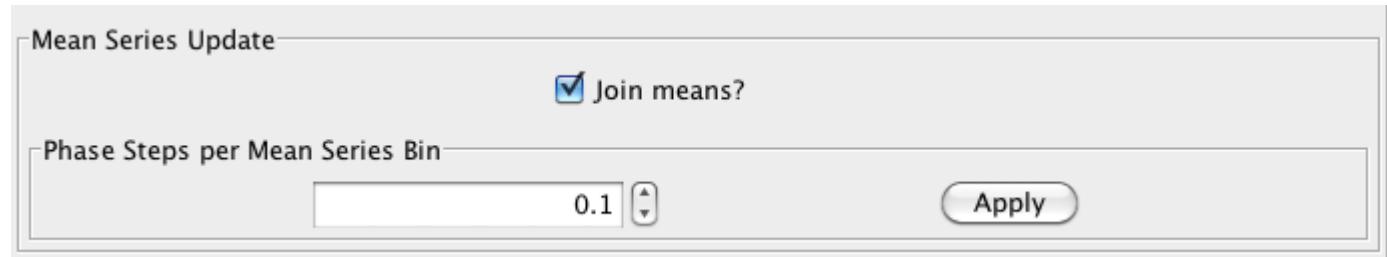
The first two sections vary according to the number of series in the currently loaded dataset.



The resulting plot (after clicking **Apply**) would look somewhat like this, depending upon JD range.

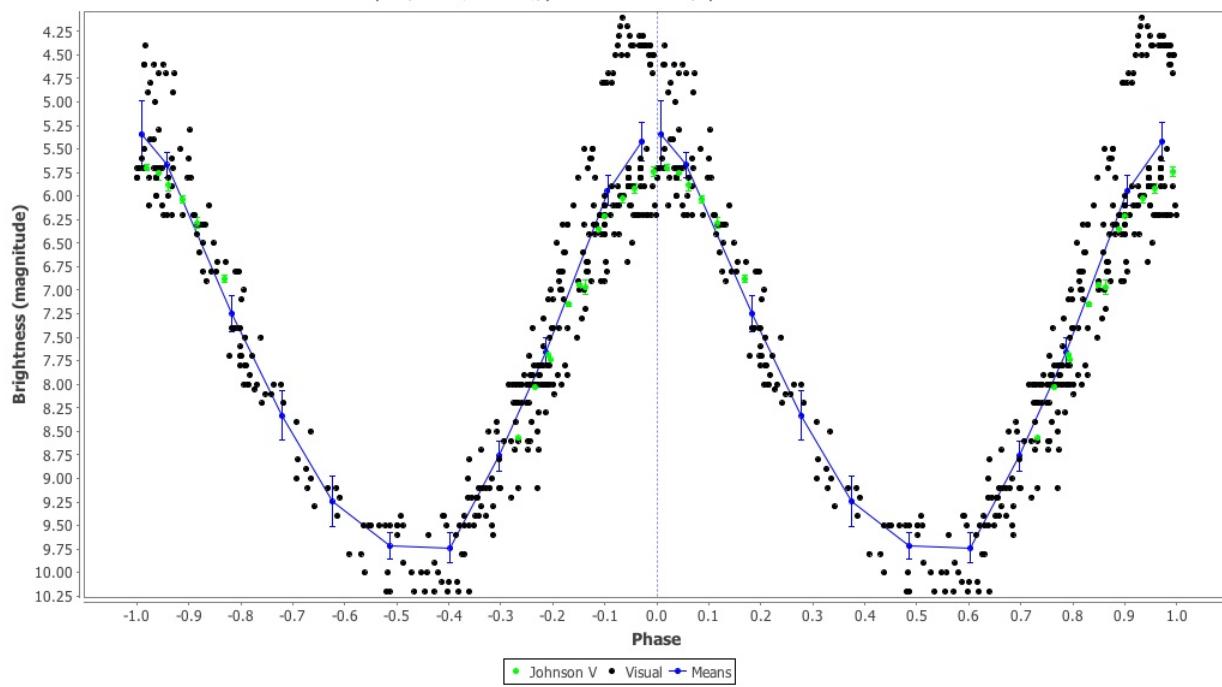


When in Phase Plot view mode, the Mean Series Update section changes to refer to Phase steps per Mean Series bin instead of Days per Mean Series bin as shown below.



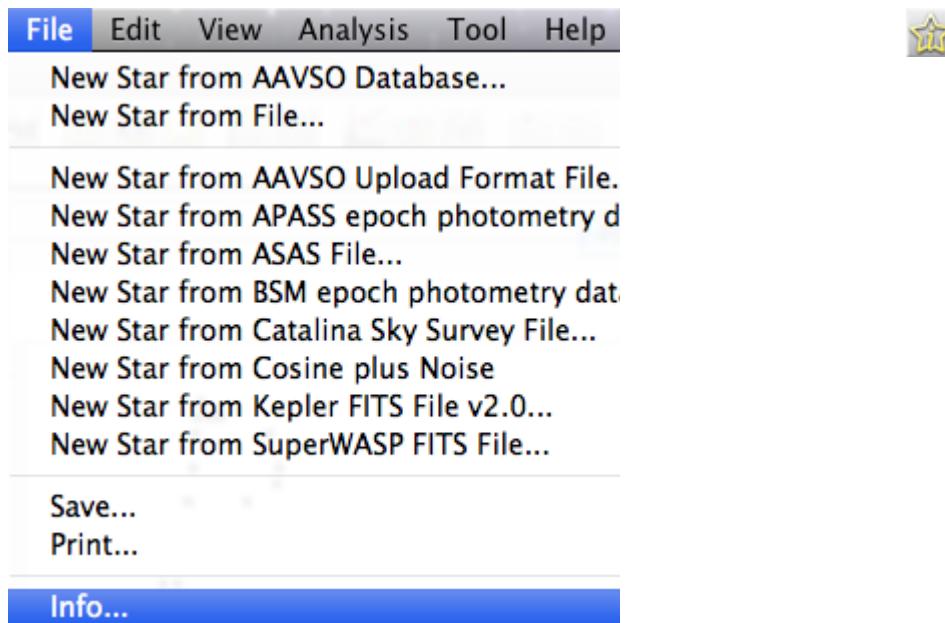
The resulting phase plot (after clicking Apply) would look somewhat like this.

Phase Plot for R Car
Sep 21, 2013 (database), period: 307.000000, epoch: 2454597.000000

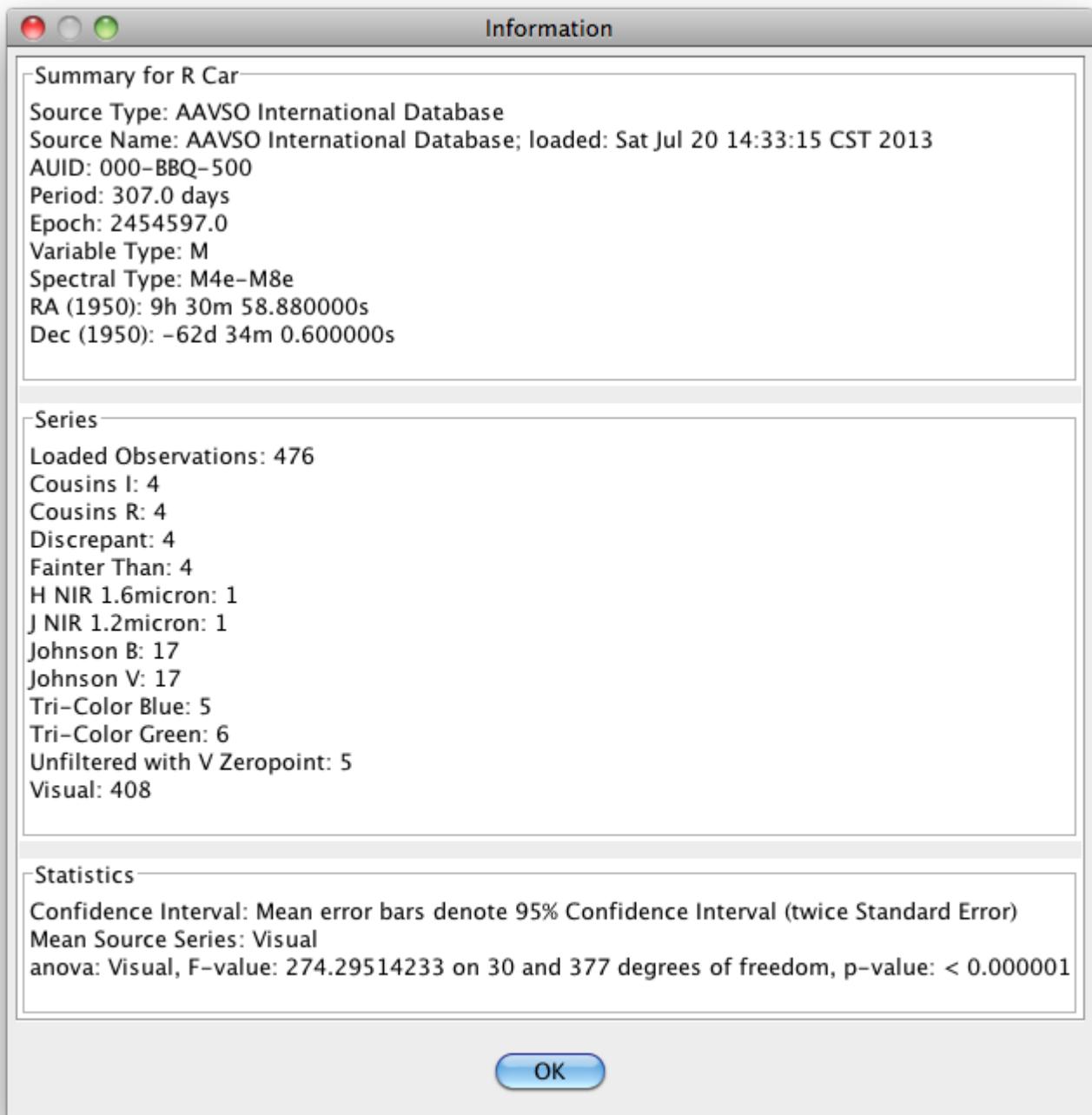


Information Dialog

Selecting **File** → **Info...** invokes the Information Dialog for the currently loaded dataset or from the corresponding toolbar button.



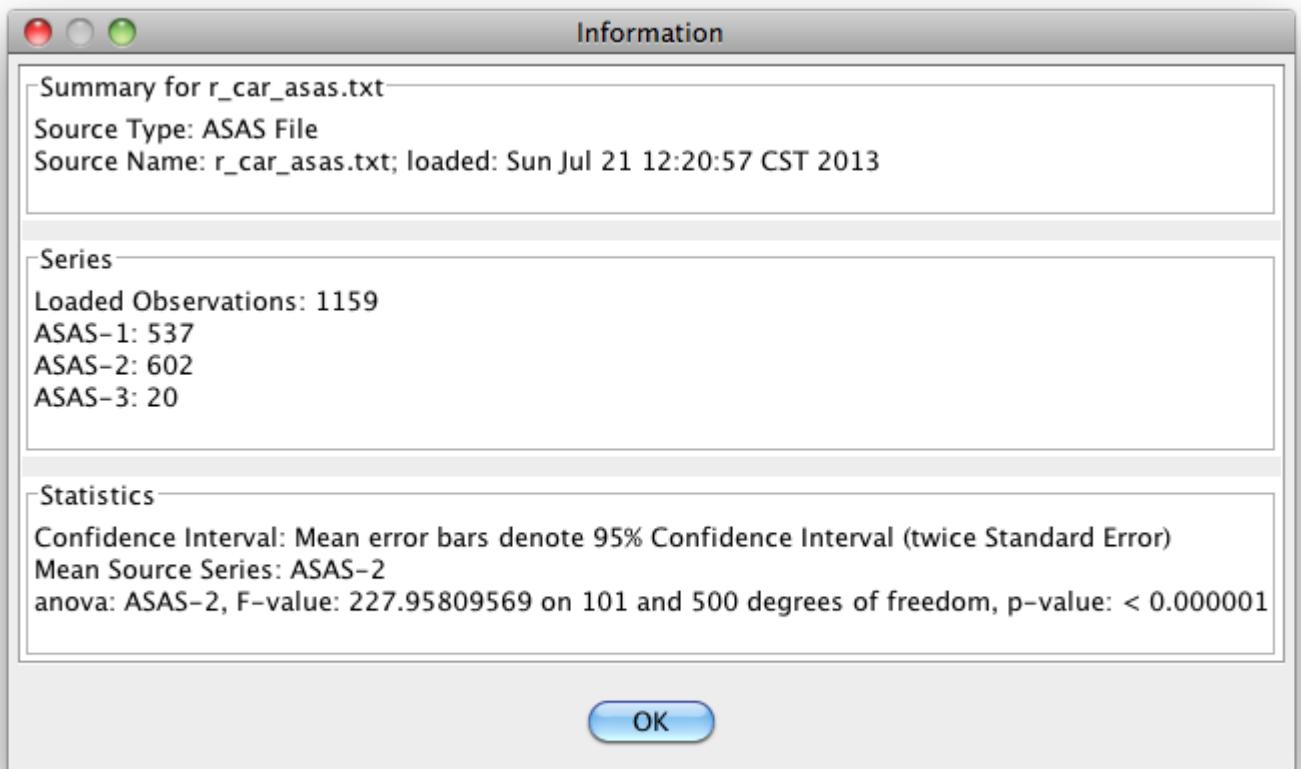
Here is an example Information Dialog for 2 years of R Car data:



The dialog is divided into Summary, Series, and Statistics sections.

The Summary section includes information about the source of the dataset, AAVSO Unique ID (AUID), and if known, period, epoch, variable type, spectral type, and coordinates. For an AID loaded dataset like this VSX is consulted for such information.

For an observation source other than AID (e.g. AAVSO upload or download files, ASAS, Kepler), VSX is not consulted so the summary information will be different. Here is an example of an Information Dialog for ASAS R Car data:



The Series section lists all series present in the dataset (e.g. the filter band or visual observations) along with the number of observations in each series.

The Statistics section provides statistical information including the standard confidence interval of mean error bars, the series that is currently selected to yield the binned mean series, and the ANOVA (analysis of variance) statistic based upon this mean series.

The ANOVA p-value indicates whether there is sufficient evidence to reject the *null hypothesis* that there is no non-trivial signal (of some kind). A low p-value indicates that the probability of there being no non-trivial signal is low. A high value indicates that the probability of there being no non-trivial signal is high. A low p-value does *not mean* that there is a *high probability of a signal being present*. Also, the question of what kind of signal there may be is a separate issue, e.g. it could be, for example, the result of actual variability in the star's brightness, observer bias,

or an artifact observation frequency. See the *References and Further Reading* section (especially Foster (2010)) for more information.

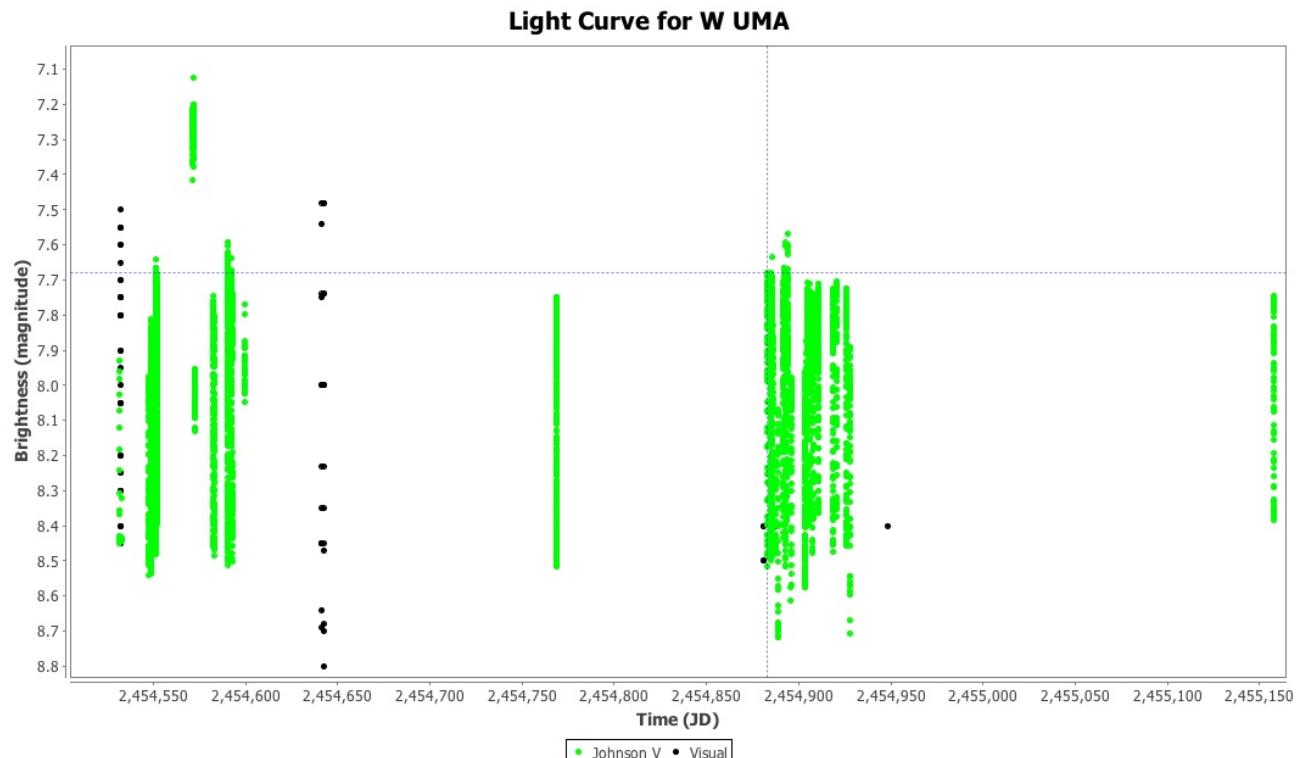
Filtering Observations

It is often useful to work with a subset of the data based upon criteria other than the series in which the data resides, e.g. magnitude or time range, observer code, or some combination.

VStar has three mechanisms by which to create a filtered subset:

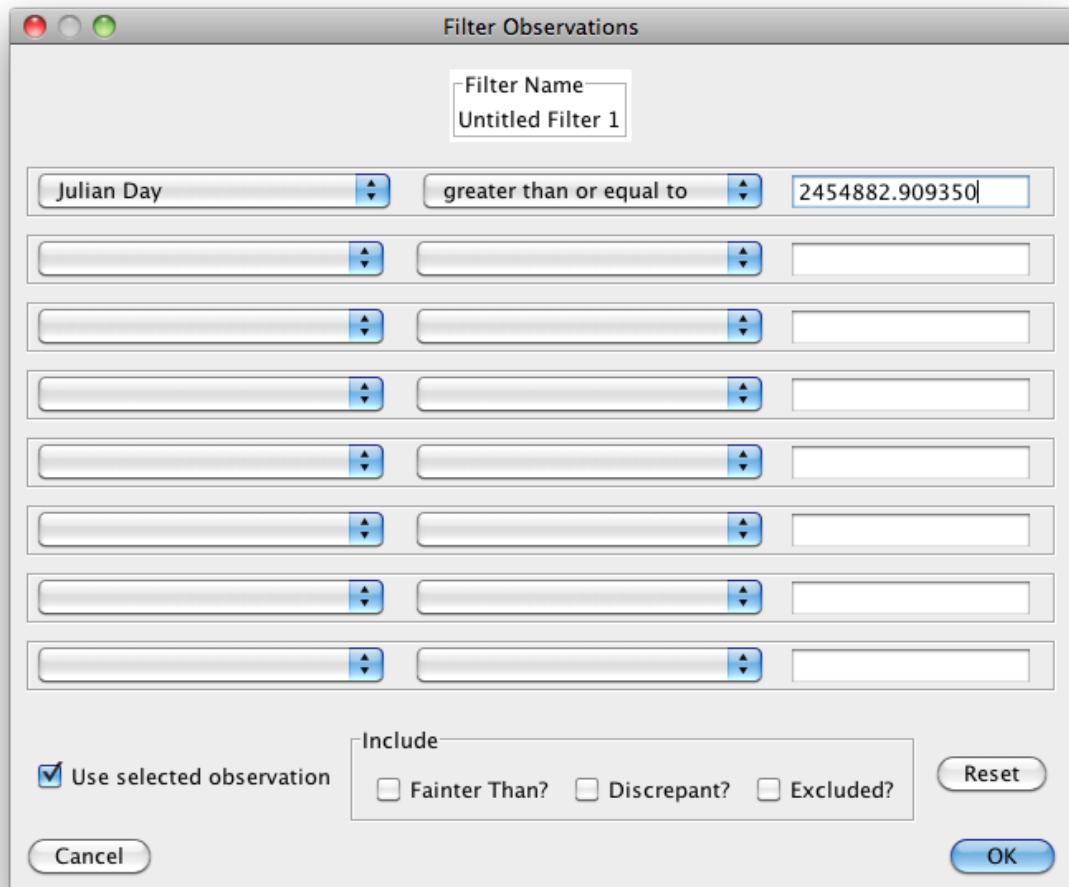
- The Observation Filter Dialog (`View → Filter...`).
- An Observation List selection filter (`Create Selection Filter` button).
- A custom observation filter plug-in.

Consider the following W UMa observations:

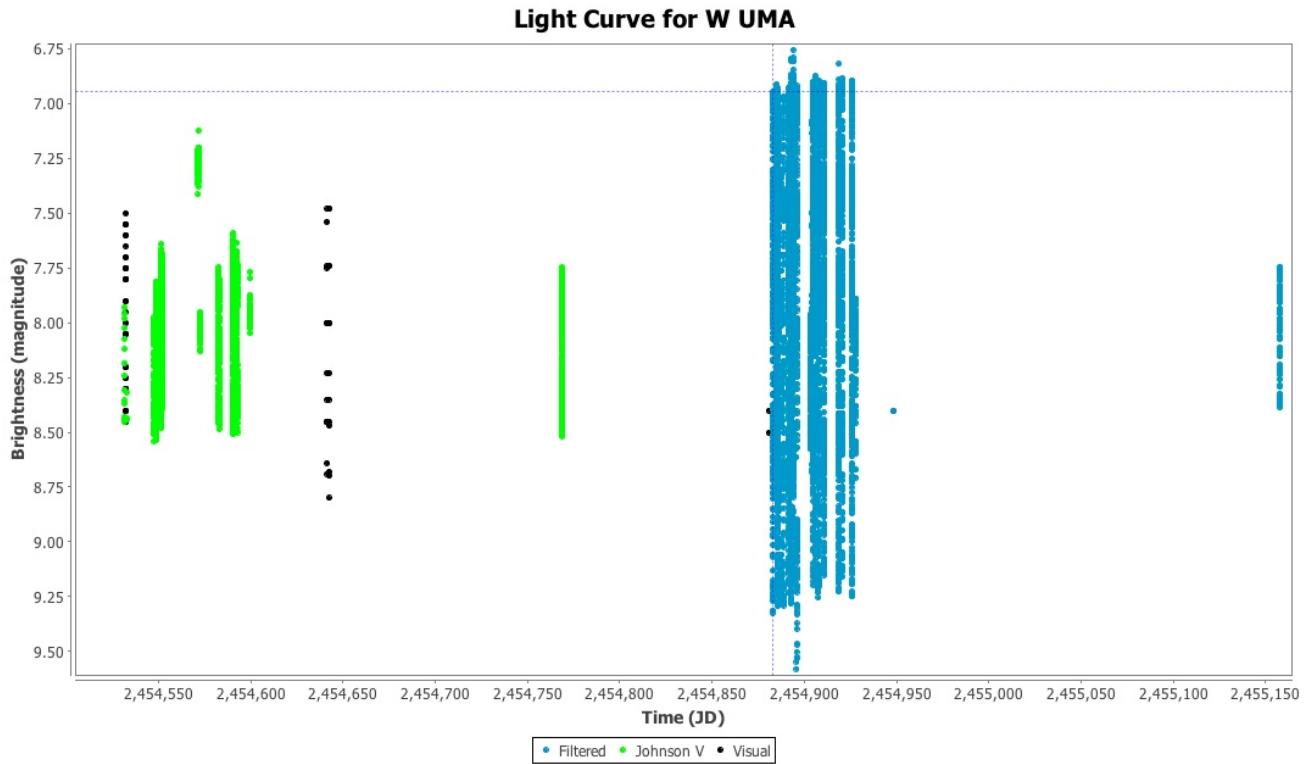


Notice that the cross-hairs are on an observation at the leftmost time extent of a particular block of Johnson V observations.

Selecting `View → Filter...` or the corresponding toolbar button opens the Observation Filter Dialog.



Here, “Use selected observation” has been checked such that when “Julian Day” is selected from a drop-down field menu, the text box is populated with the JD of the selected observation (under the cross-hairs). Selecting “greater than or equal” from the middle drop-down menu then dismissing the dialog with the OK button yields the following:



Notice that all observations with JDs greater than or equal to the selected observation are shown as being in the Filtered series (see legend at bottom of plot), regardless of series.

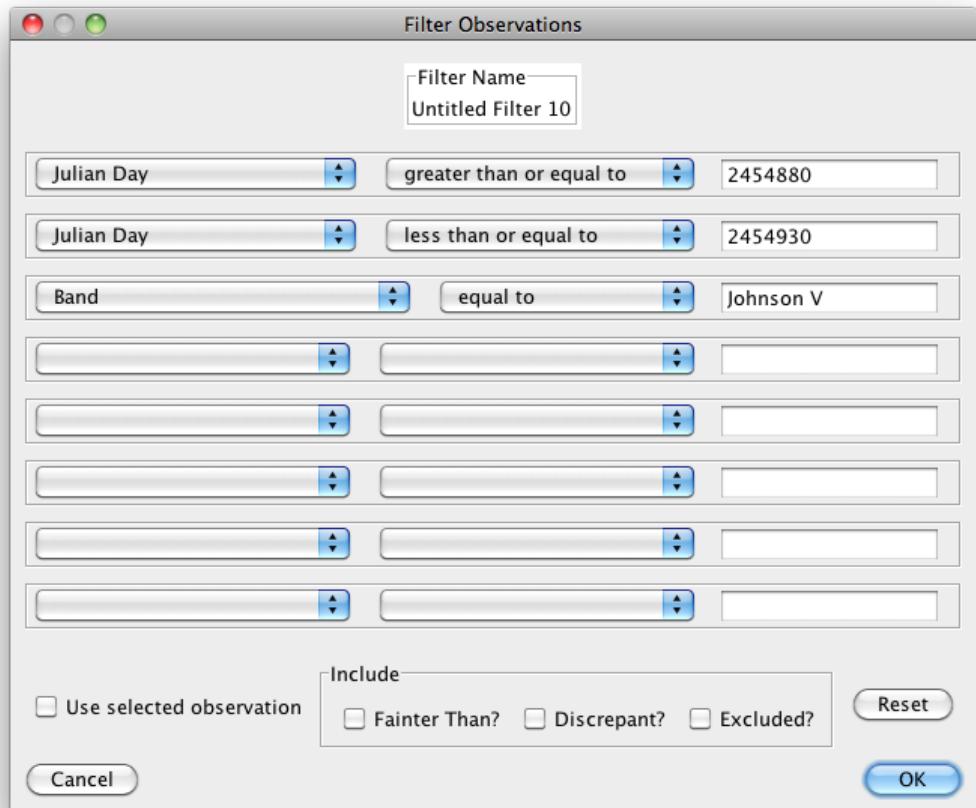
This filter could be refined to restrict the JD range and series included. The following example creates a filtered subset of observations where the JD is greater than or equal to 2454880 and the JD is less than or equal to 2454930, and the band is Johnson V. Operators differ with field type:

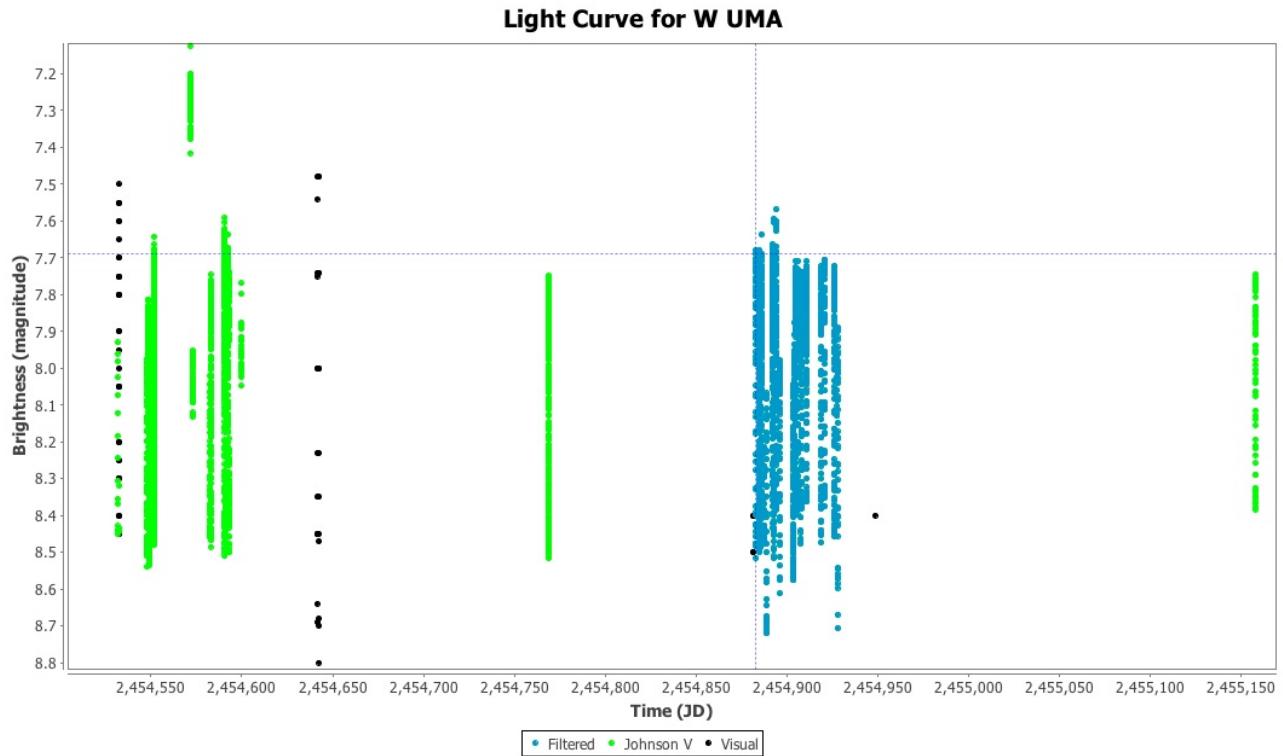
- For numeric fields such as JD, magnitude, and error, the usual relational operators are provided (equal, not equal, less than, greater than, less than or equal, greater than or equal).
- For string fields such as observer code or object name, equal and not equal are permitted.

The constrained filter mentioned above and the resulting observation plot are shown below.

Notice the implicit “and’s represented by the filter dialog selections. Currently, 8 such

“conjunctive” terms are permitted in a single filter.

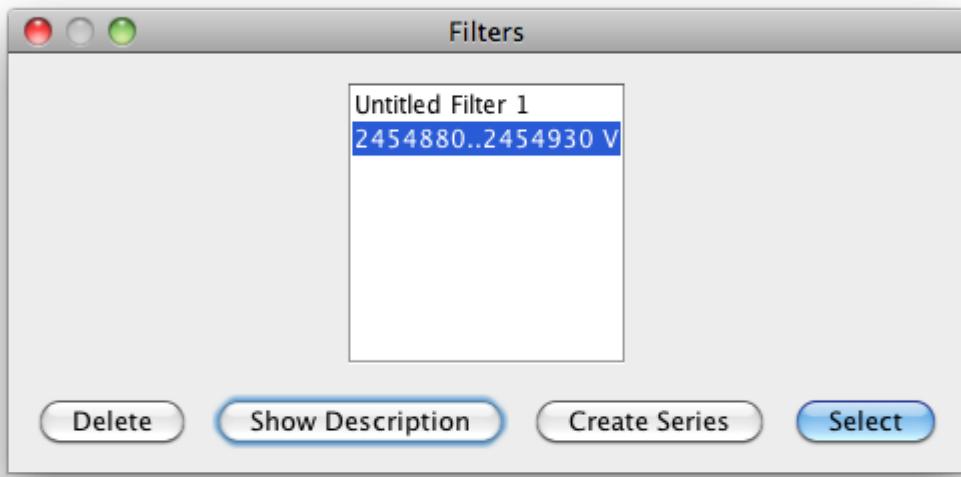




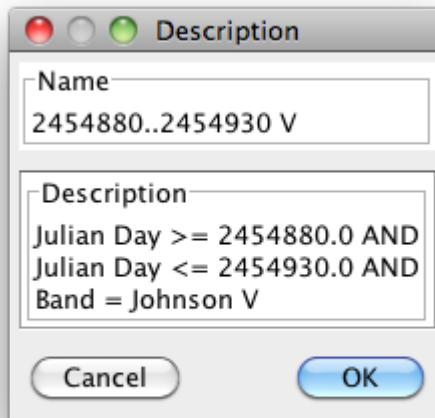
Note that the “Filter Name” text box at the top of the observation filter dialog can be modified to give a filter a meaningful name, e.g. “2454880...2454930 V.”

A filter series can be hidden from view just like any other series, via the Plot Control Dialog. The current filter can be completely removed by selecting View → No Filter.

Previously created filter series for the currently loaded dataset can be retrieved and viewed by selecting View → Previous Filters...



Filters can also be deleted and new series can be created from them, to be treated just like any other series in the Plot Control Dialog, preferences and so on. Each filter has a description that can be viewed via the Show Description button, e.g.



The second method of creating an observation filter is via an Observation List selection filter. The following observation list screenshot shows a pattern search (see *Observation List* section) being used to narrow down observations to just Johnson V (this could also have been done via the Plot Control dialog) and a selection (via shift + mouse-click) of multiple observations.

Plot Observations Means Model Residuals

Show all data? Band All Columns? Johnson V

Julian Day	Calendar Date	Magnitude	Uncertainty	Band	Observer...	Validation	Comp Star 1	Comp Star 2	Charts	Comment...	Comments	Transform
2454768.986440	2008 OCT 29	7.796000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454768.986700	2008 OCT 29	7.781000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454768.986960	2008 OCT 29	7.770000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454768.987220	2008 OCT 29	7.784000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454768.987480	2008 OCT 29	7.778000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454768.987740	2008 OCT 29	7.776000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454768.988010	2008 OCT 29	7.796000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454768.988260	2008 OCT 29	7.765000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454768.988530	2008 OCT 29	7.781000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454768.988800	2008 OCT 29	7.790000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454768.989050	2008 OCT 29	7.792000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.892250	2009 FEB 20	7.763000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.893660	2009 FEB 20	7.735000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.895090	2009 FEB 20	7.725000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.896520	2009 FEB 20	7.715000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.897940	2009 FEB 20	7.703000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.899360	2009 FEB 20	7.690000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.900790	2009 FEB 20	7.714000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.902220	2009 FEB 20	7.796000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.903650	2009 FEB 20	7.756000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.905070	2009 FEB 20	7.738000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.906490	2009 FEB 20	7.733000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.907920	2009 FEB 20	7.728000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.909350	2009 FEB 20	7.679000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.910780	2009 FEB 20	7.706000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.912200	2009 FEB 20	7.739000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.913620	2009 FEB 20	7.738000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.915050	2009 FEB 20	7.764000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.916480	2009 FEB 20	7.734000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.917910	2009 FEB 20	7.737000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.919330	2009 FEB 20	7.732000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.920750	2009 FEB 20	7.762000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	
2454882.922180	2009 FEB 20	7.804000	0.000000	Johnson V	TRE	Good	000BBQ747	000BBQ651	1010APY		no	

Clicking the Create Selection Filter button would yield the same filtered subset as above (assuming the end date was appropriately selected).

See the *Plug-ins* section for details of custom filters, also available in the View menu.

Phase Plots

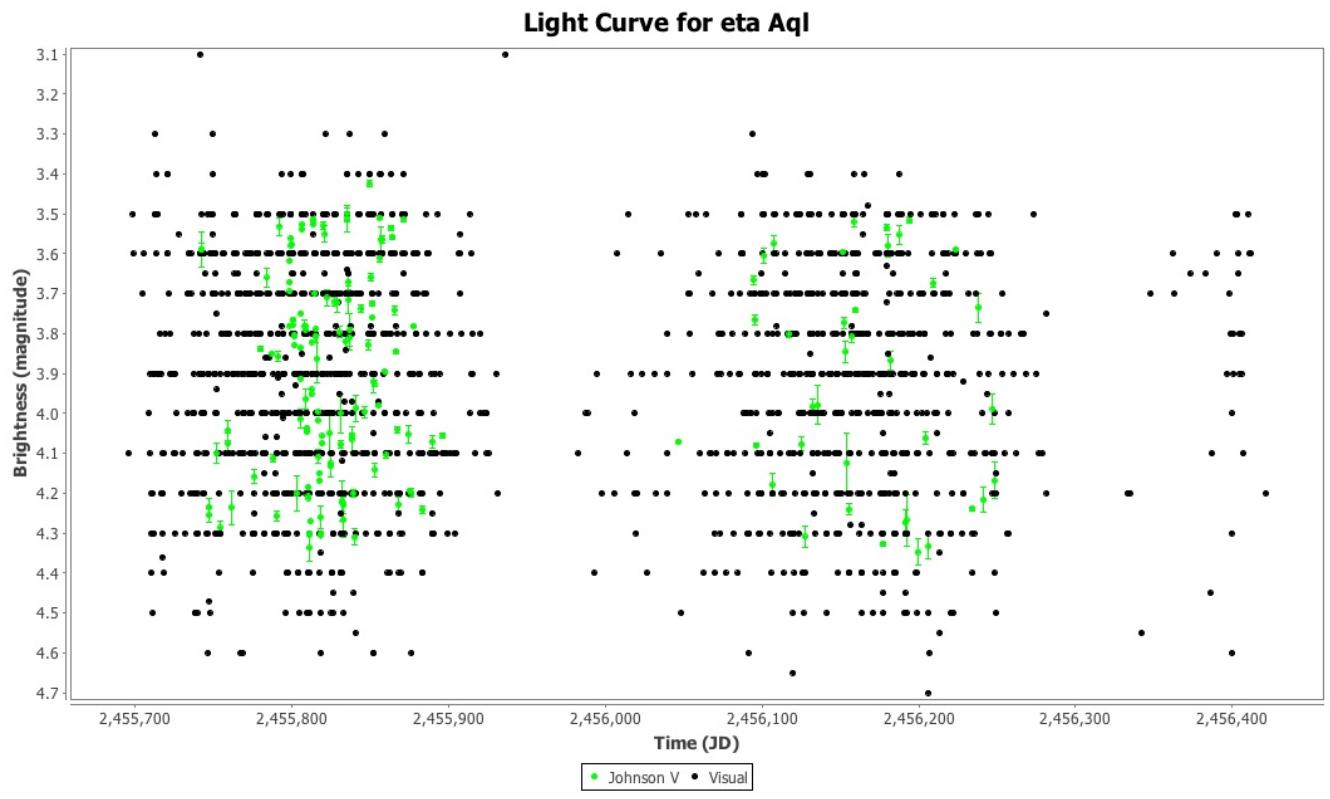
The *Period Analysis* section details how VStar can be used to analyze a dataset for candidate periods (the time taken for a star’s variability to complete one full cycle). If a star is assumed to have a particular period, VStar can be asked to create a so-called phase plot or folded light curve, in which all observations are plotted into a “window” of a particular size: the period in days.

To accomplish this, each observation’s time is converted to a phase according to the following equation:

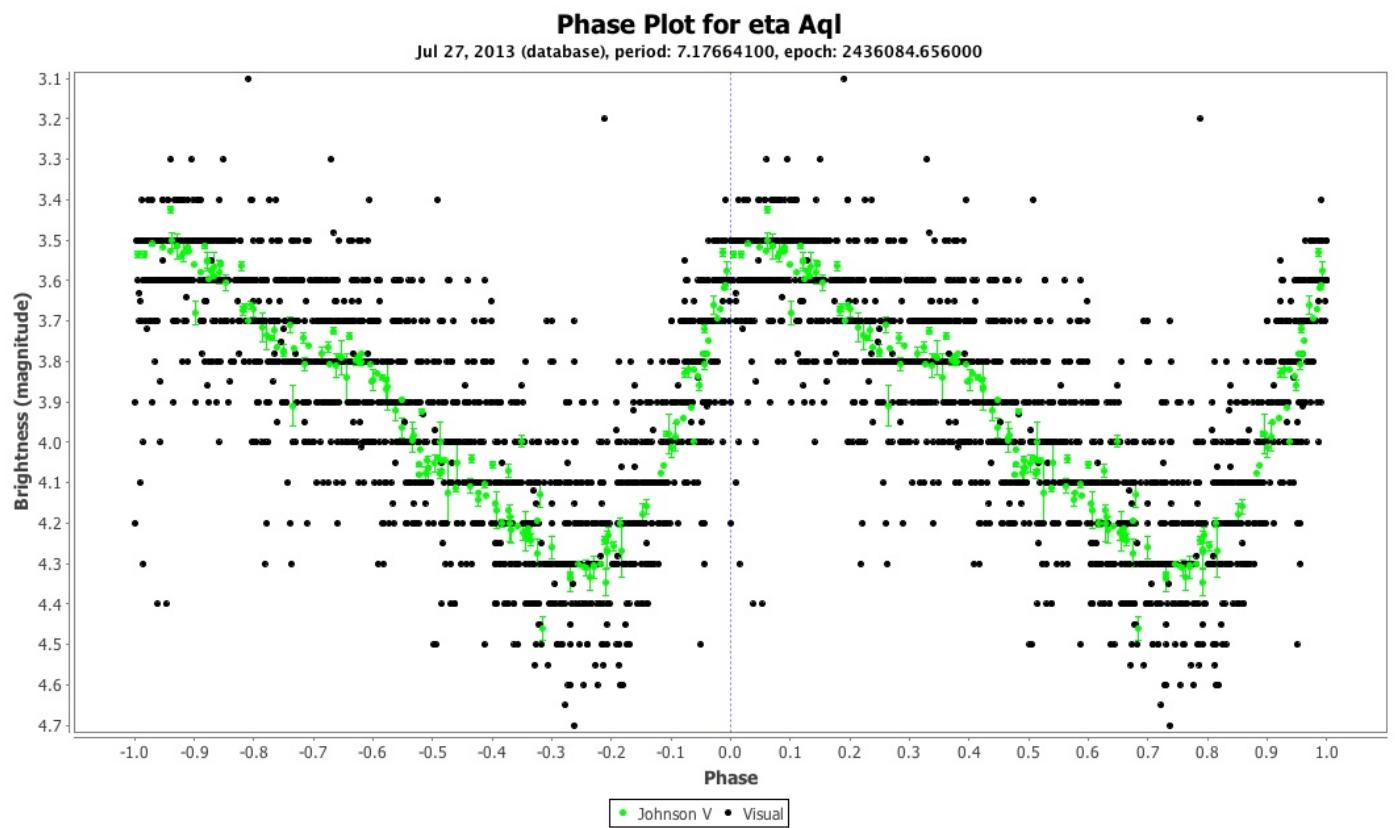
$$\phi = \frac{t - epoch}{P}$$

where ϕ is the phase, t is the observation time, *epoch* is some initial time, and P is the assumed period. The *epoch* can be one of a number of values from the current dataset, such as the time of the first observation, the time closest to the mid-point, the mean of the first and last time, a time of minimum or maximum. The most appropriate *epoch* to use may depend upon the binary type. If one is not available from VSX for the star (and only for an AID loaded dataset), the default epoch computed by VStar is the mean of the first and last time in the current dataset.

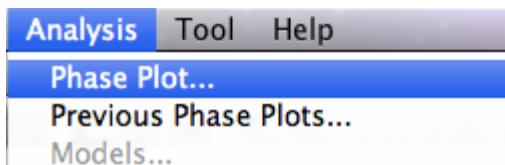
Assume the following eta Aquilae dataset is loaded:



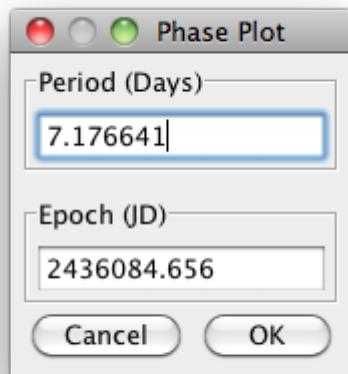
Using the period and epoch for eta Aquilae shown on VSX (<http://www.aavso.org/vsx/index.php?view=detail.top&oid=2802>), the following phase plot can be created.



To create a phase plot like this one, select Analysis → Phase Plot...

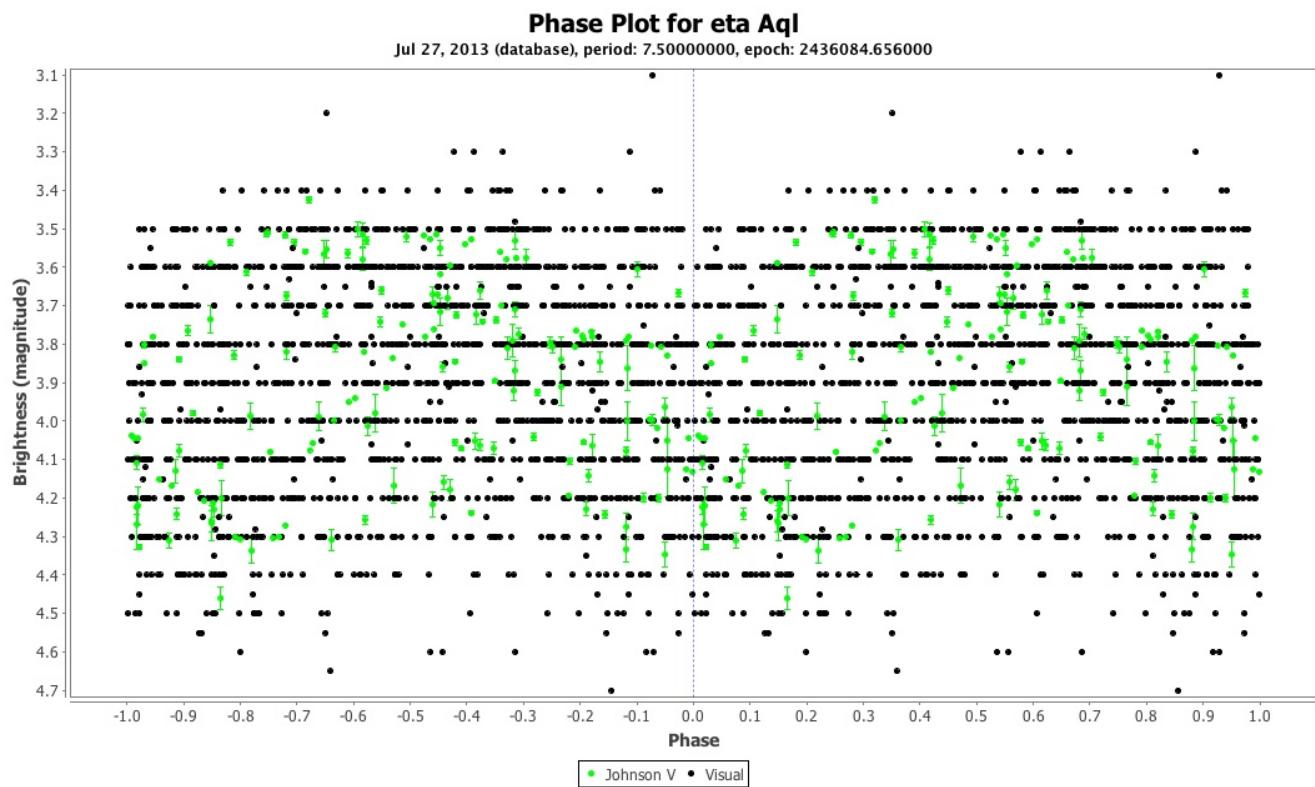


yielding the phase plot parameter dialog and click OK to create the phase plot:

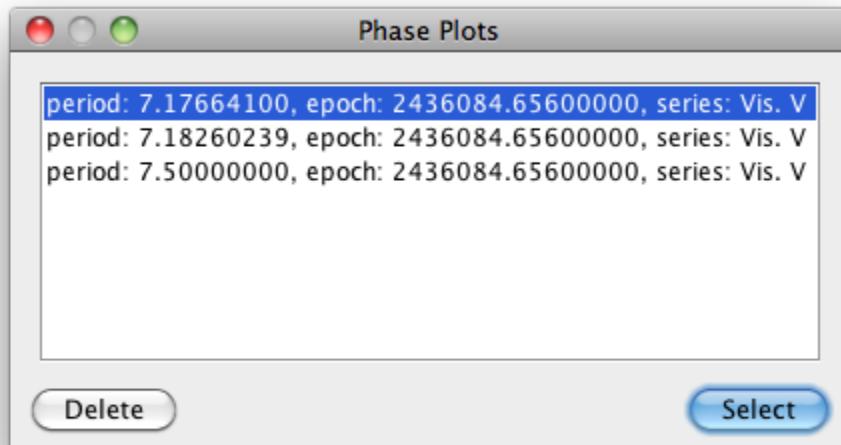


If a dataset is loaded from AID, VStar will retrieve this information from VSX and populate the phase plot parameter dialog text fields.

A phase plot for which the specified period is incorrect compared to the actual period, results in a “messy” (or out of phase) plot. For example, here is eta Aquilae with a phase plot using a period of 7.5 days instead of 7.176641 days:



The Analysis menu's Previous Phase Plots... item yields a dialog that permits previously created phase plots to be selected or deleted, e.g.



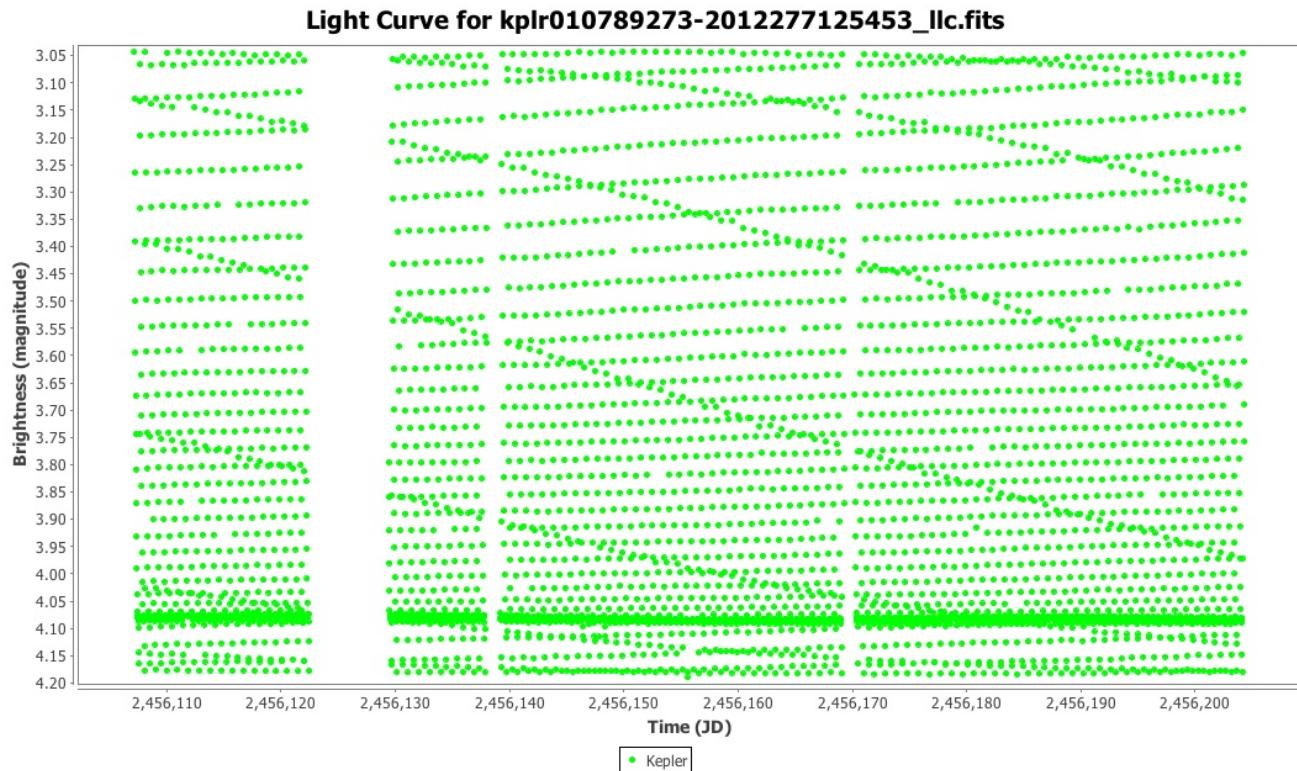
The first two items in the View menu allow you to switch between phase plot and raw view

modes. There are corresponding toolbar buttons.

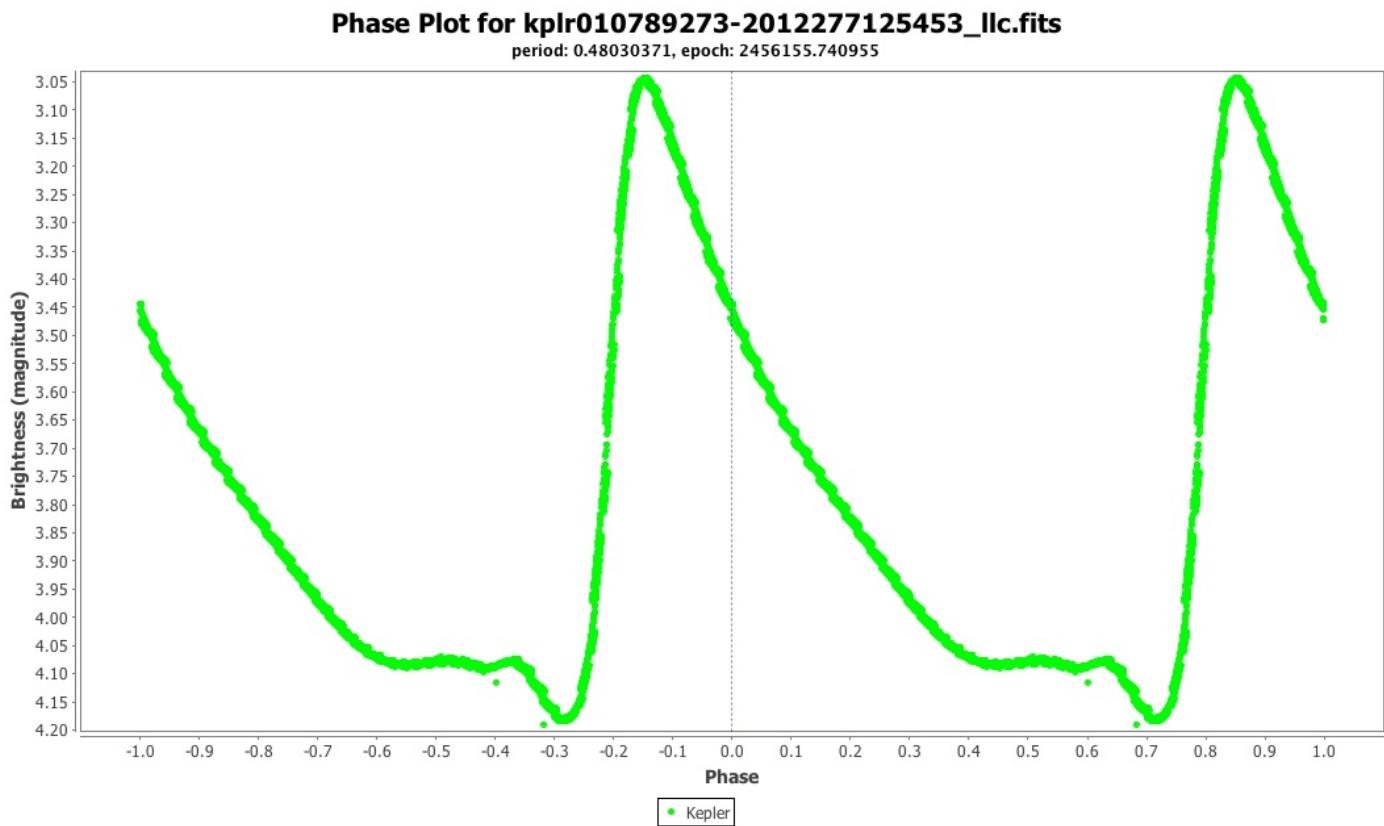


If no phase plot exists the first time Phase Plot view mode is selected, the phase plot parameter dialog will be invoked.

Here is another example that demonstrates the visualization power of a phase plot given the right period. The following shows a V0838 Cyg dataset loaded via the Kepler observation source plugin.



As described in detail elsewhere (<http://dbenn.wordpress.com/2013/03/29/obtaining-and-analysing-kepler-data-with-vstar>), the following phase plot reveals the nature of the star as an RR Lyr star variable (RRAB) with a period of around 0.48 days.



The observation list is also affected when phase plot mode is enabled: a phase column is added, e.g.

Pattern Search							
<input type="checkbox"/> Show all data?	Phase	<input checked="" type="checkbox"/> All Columns?	<input type="button" value="Apply"/>				
Phase	Julian Day	Calendar Date	Magnitude	Uncertainty	Band	Record	
0.853557	2456107.159943	2012 JUN 28	3.043647	0.000119	Kepler		
0.896101	2456107.180377	2012 JUN 28	3.129627	0.000129	Kepler		
0.938645	2456107.200811	2012 JUN 28	3.264647	0.000147	Kepler		
0.981189	2456107.221245	2012 JUN 28	3.389898	0.000166	Kepler		

By default, observation rows are still sorted by JD, but clicking the Phase column will order by phase.

The *Period Analysis* section shows how to create a phase plot from the result dialog of a period search. See also the *Further Reading* section.

Period Analysis

Introduction

The purpose of period analysis is, as the term suggests, to identify candidate periods for a variable star from the available data. There are a number of ways to do this. Sometimes, one can simply inspect the light curve to find the time taken for the brightness change to go through one full cycle. For less obvious cases, an algorithm such as AoV, PDM, or Fourier analysis can be used to search for candidate periods.

VStar implements the Date Compensated Discrete Fourier Transform (DCDFT) algorithm (Ferraz-Mello 1981), yielding a power spectrum and a table of “top-hits” given a specified series, frequency (or period) range, and resolution. The “date compensated” part of the name indicates that gaps in the data, common for variable star observations, are compensated for by the algorithm.

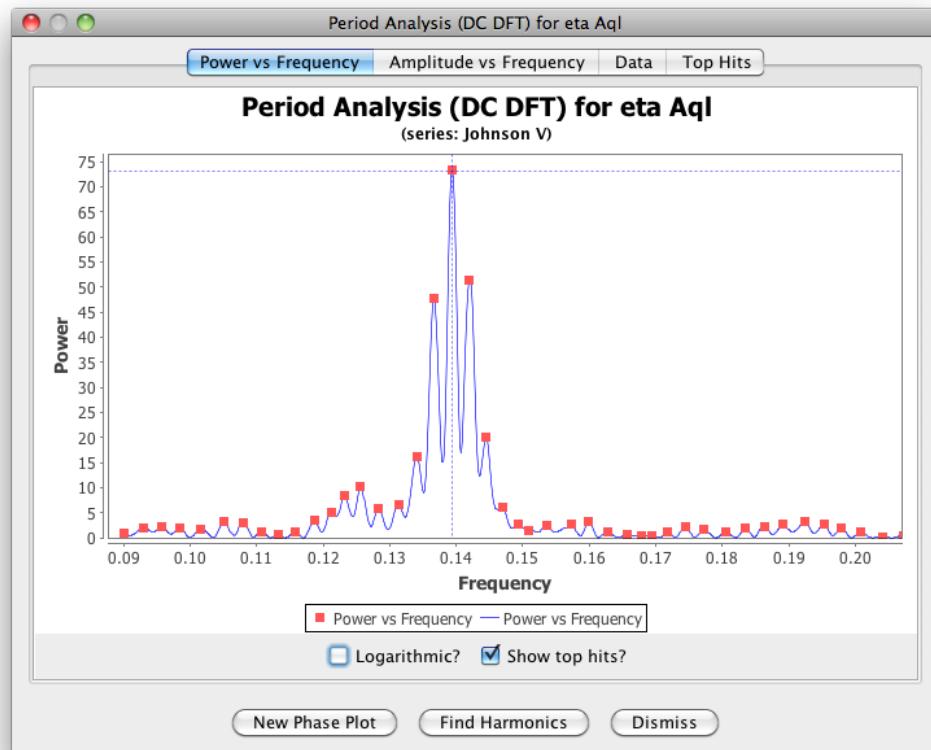
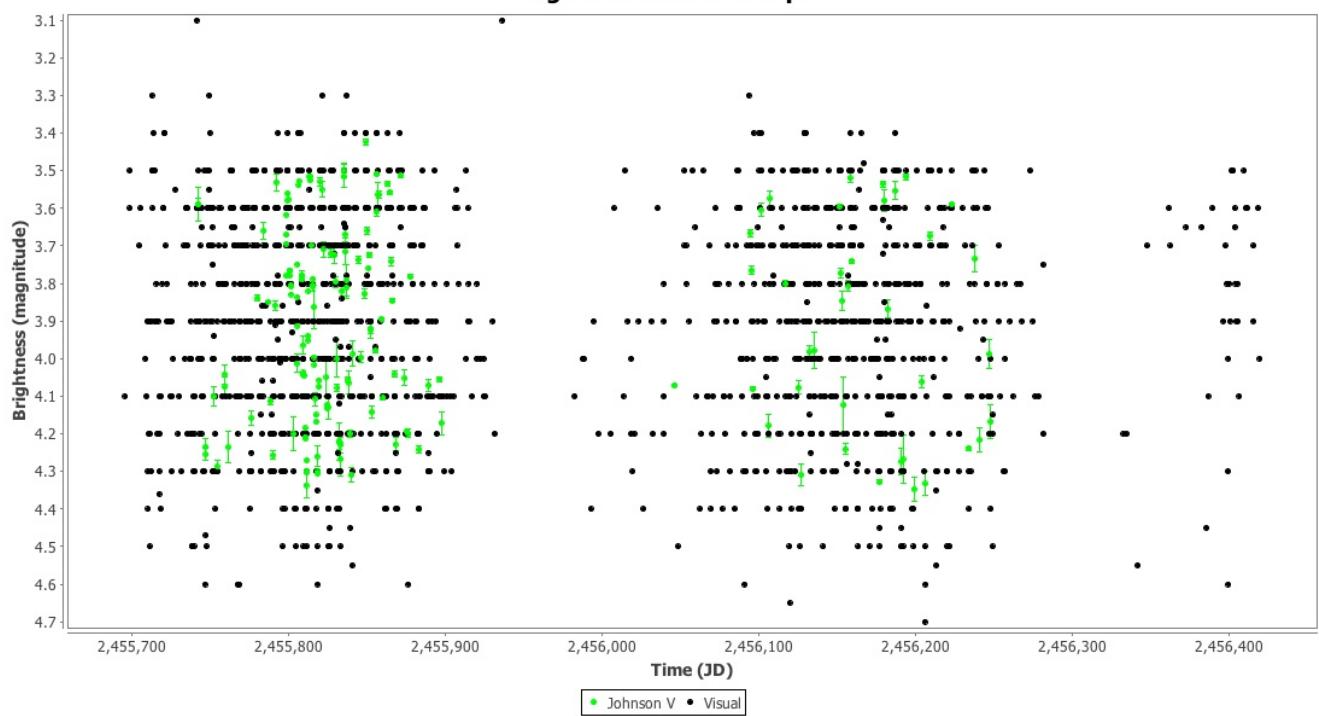
From within the DCDFT result window, a phase plot can be created (see *Phase Plots* section for more information). In addition, one or more periods each with one or more harmonics can be selected to create a model. A model’s Fourier coefficients can be viewed along with relative amplitudes and phases.

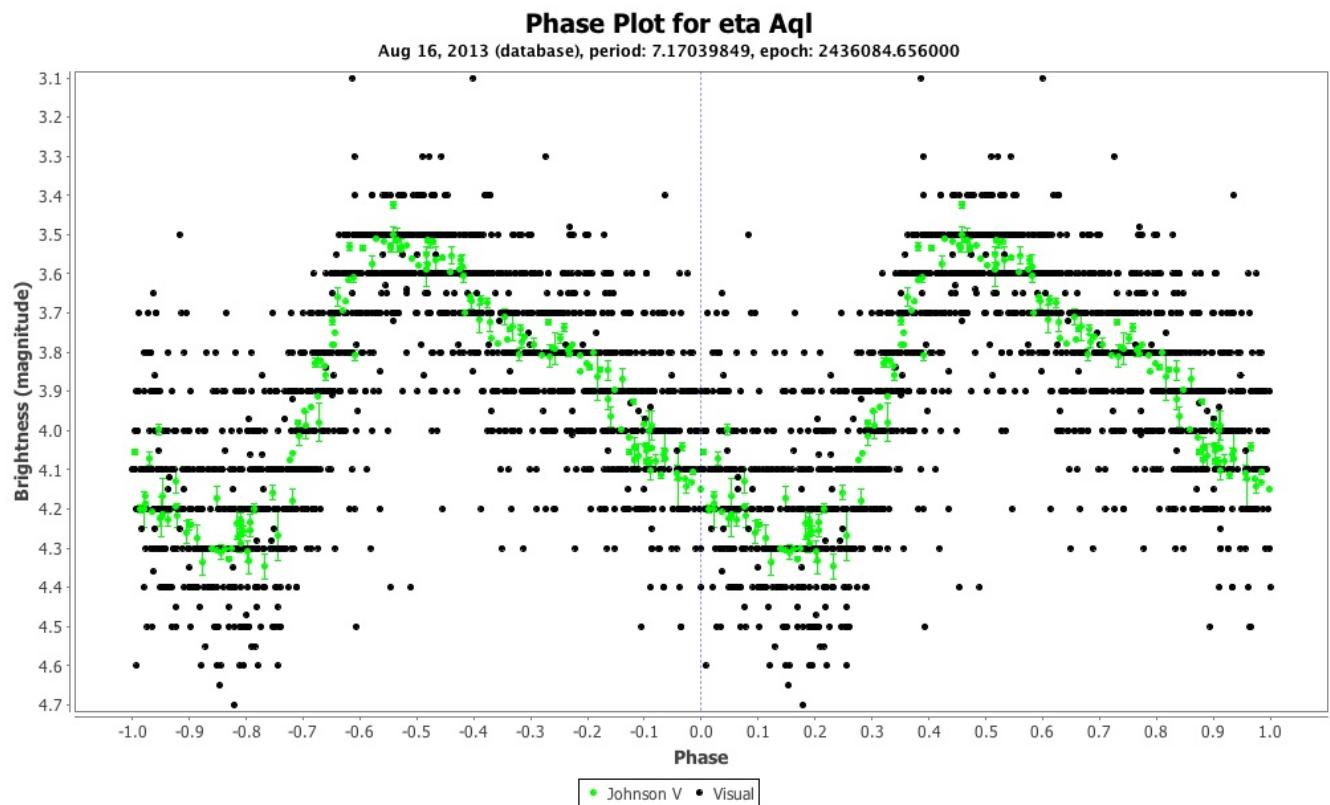
Optionally, multiple periods found by DCDFT can subsequently be refined via the CLEANest (Foster 1995) algorithm.

When a model is created, it is also subtracted from observations in the series on which the DCDFT was performed to yield a second, “residuals,” series. DCDFT can then be applied to these residuals to look for further signals (periods), a process often called “pre-whitening.”

The following shows eta Aquilae data from AAVSO International Database (AID) in the JD range 2455695.8431 to 2456419.8604, a DCDFT “power spectrum”, and a phase plot resulting from a top-hit selection.

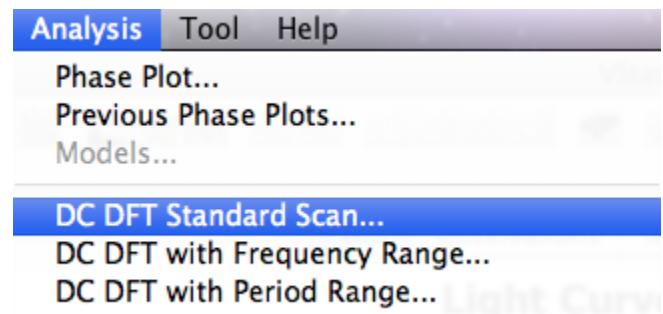
Light Curve for eta Aql



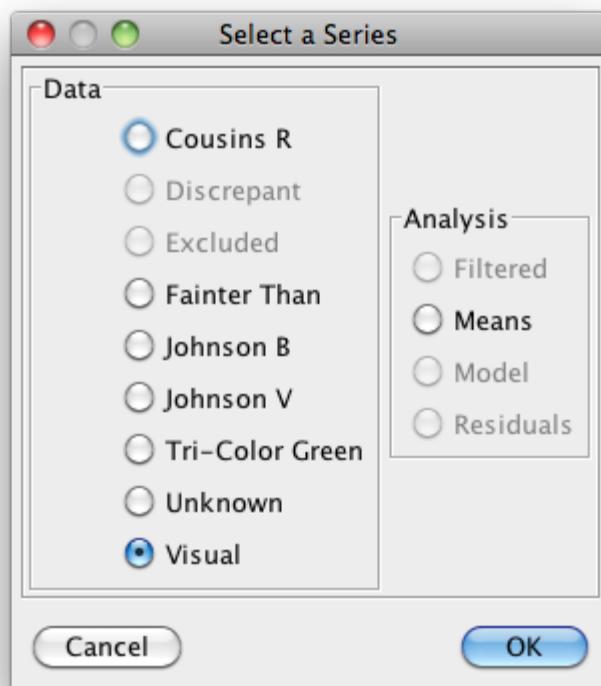


DCDFT in Detail

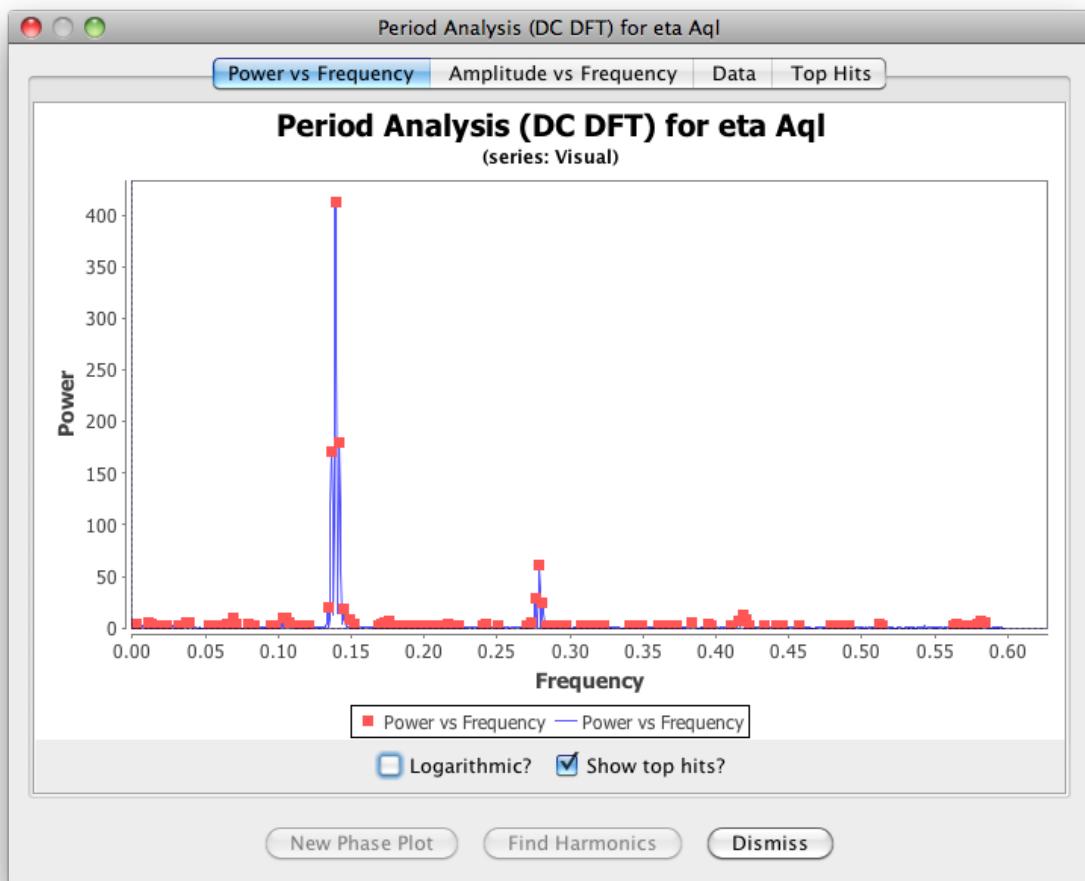
DCDFT menu items are available in the `Analysis` menu. As shown below, the options are DCDFT standard scan, DCDFT with frequency range, or with period range.



Looking at the eta Aql example above in more detail, selecting `DCDFT Standard Scan...` yields the series selection dialog, since period search is only applied to a single series:



Here, the `Visual` series is selected. Clicking the `OK` button leads, after a short time, to the result dialog:



By default, the `Power vs Frequency` plot (sometimes called the “power spectrum”) is shown. Notice that `Show top hits?` is checked here, again, by default. The small brightly colored squares correspond to “top hits,” those frequencies/periods at the peaks of the plot (those with highest power value). These are the best candidates for the actual frequency (period).

As with all plots in VStar, zooming in to a DCDFT plot can be done by click-dragging a rectangular area of the plot or using the context menu (e.g. via right-click in Windows or ctrl-click on Mac OS X) `Zoom In/Out` items, as was done for the power spectrum shown in the *Introduction* section.

Before leaving this tab, notice the `Logarithmic?` checkbox. If selected, the power (or amplitude in the case of the `Amplitude vs Frequency` plot) axis will become a base-10 log scale. This can sometimes help to reveal frequencies of interest.

Selecting the `Top Hits` tab shows the top hits, referred to above, in tabular form.

Period Analysis (DC DFT) for eta Aql

Frequency	Period	Power	Amplitude
0.13946226	7.17039849	413.05793301	0.29895566
0.14176109	7.05412176	179.53117465	0.19552786
0.13678030	7.31099454	171.11829754	0.19041868
0.27892453	3.58519925	60.30037260	0.11155846
0.27624256	3.62000701	29.29911506	0.07826145
0.28122335	3.55589244	24.43931680	0.07189908
0.13409833	7.45721443	20.68094814	0.06493511
0.14444306	6.92314337	17.99564602	0.06098598
0.41838679	2.39013283	12.32965424	0.05127045
0.10536297	9.49100019	9.81780191	0.04593121
0.06858172	14.58114554	9.59572999	0.04426406
0.10306414	9.70269536	9.42094550	0.04461421
0.42068562	2.37707200	8.48601531	0.04242292
0.14942385	6.69237193	7.59515641	0.04047890
0.58198675	1.71825217	7.38488388	0.03921081
0.17586038	5.68632909	7.05278187	0.03813793
0.41570482	2.40555304	6.34992068	0.03621076
0.17317841	5.77439171	5.66317097	0.03427980
0.03946322	25.34004905	5.47302860	0.03343941
0.03678126	27.18776096	5.21778103	0.03358764
0.01072787	93.21518043	4.98850751	0.03265935
0.27356059	3.65549727	4.78728188	0.03193053
0.10766180	9.28834538	4.76483948	0.03174396
0.58428558	1.71149184	4.65395022	0.03087283
0.38390436	2.60481542	4.60280399	0.03078519

[Create Model](#) [CLEANest](#)

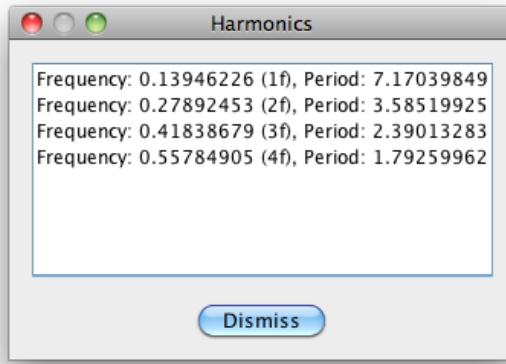
[New Phase Plot](#) [Find Harmonics](#) [Dismiss](#)

Clicking the New Phase Plot button with the shown top-hit selected will result in the phase seen earlier in this section.

Harmonics

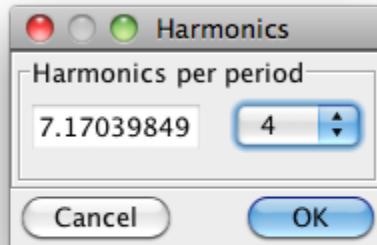
The Find Harmonics button will show a list of harmonics of the main frequency (or period

$\frac{1}{frequency}$) in the data. VStar adopts the convention (from Foster 2010) that the first harmonic is the fundamental frequency. This can be used as a decision support aid for model creation.

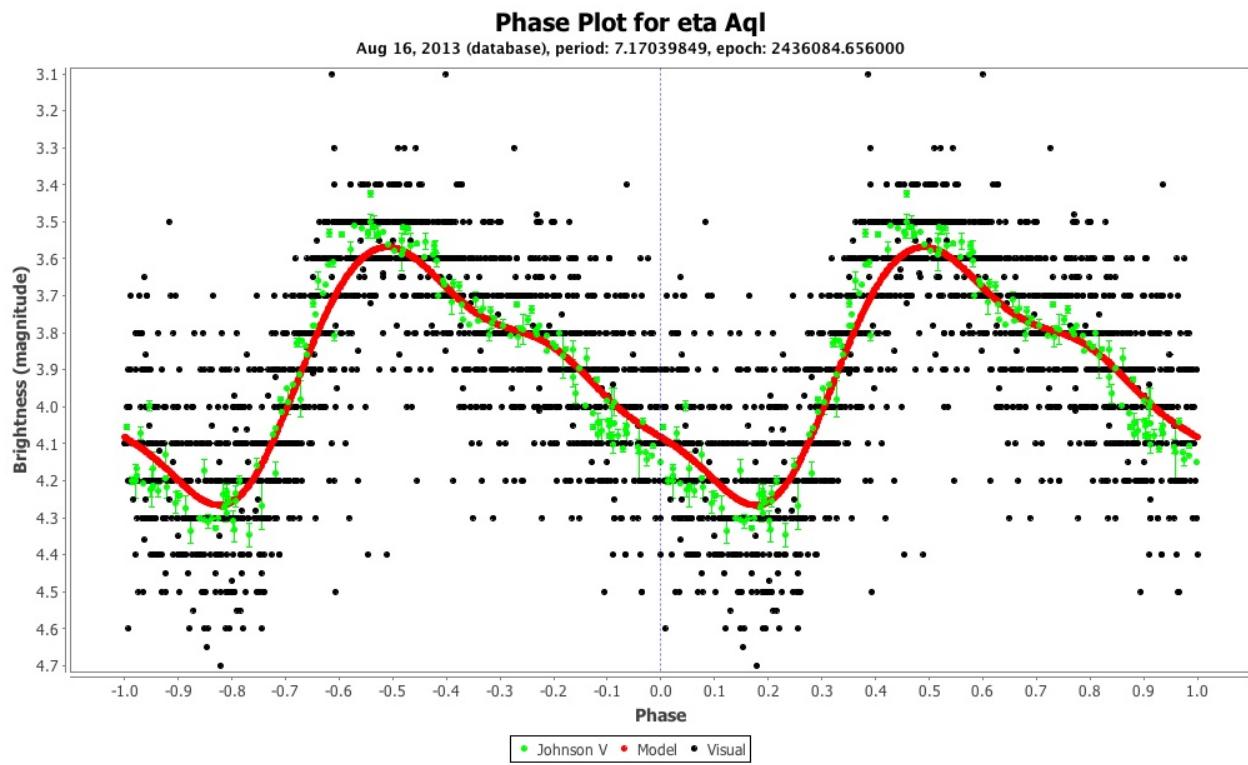


Creating Fourier Series Models

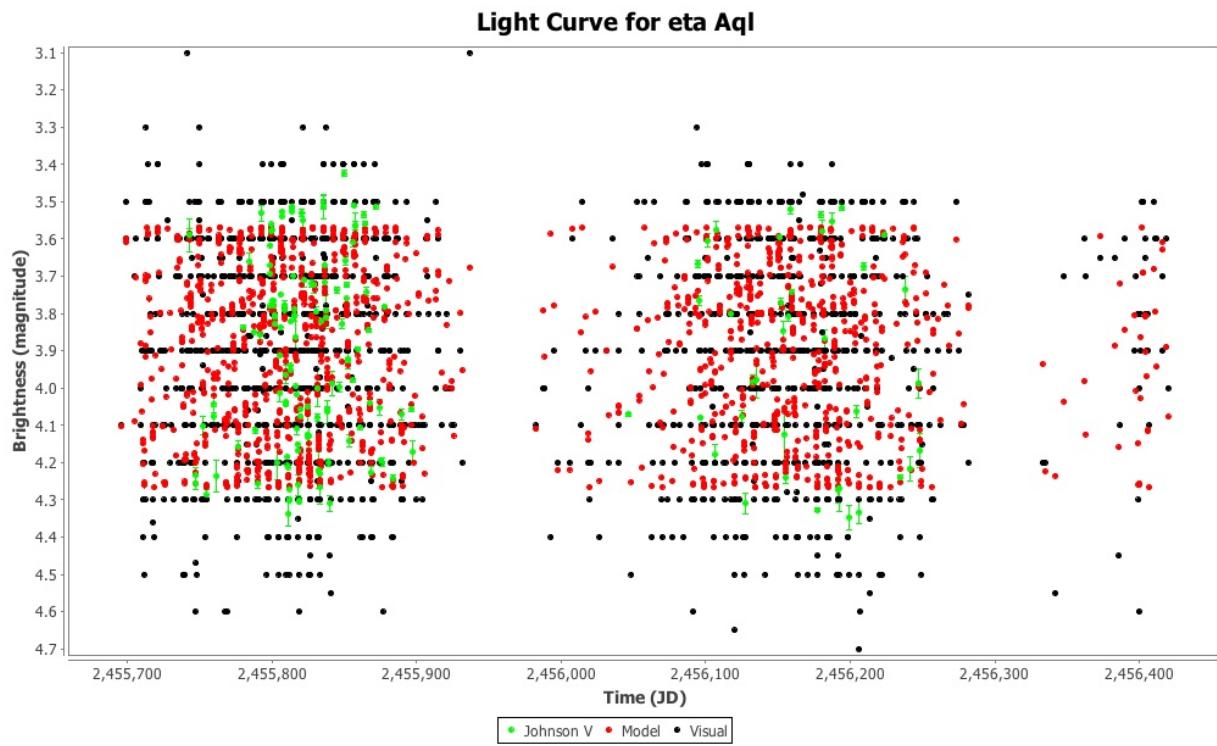
The Create Model button creates a Fourier series from the selected period and one or more harmonics selected from the following dialog. In this case, 4 harmonics are selected, which is to say, the fundamental and the next 3 harmonics.



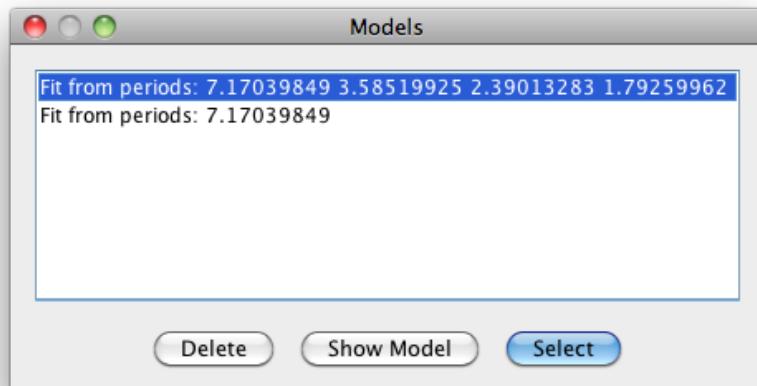
Clicking OK gives the following if phase plot mode is selected:



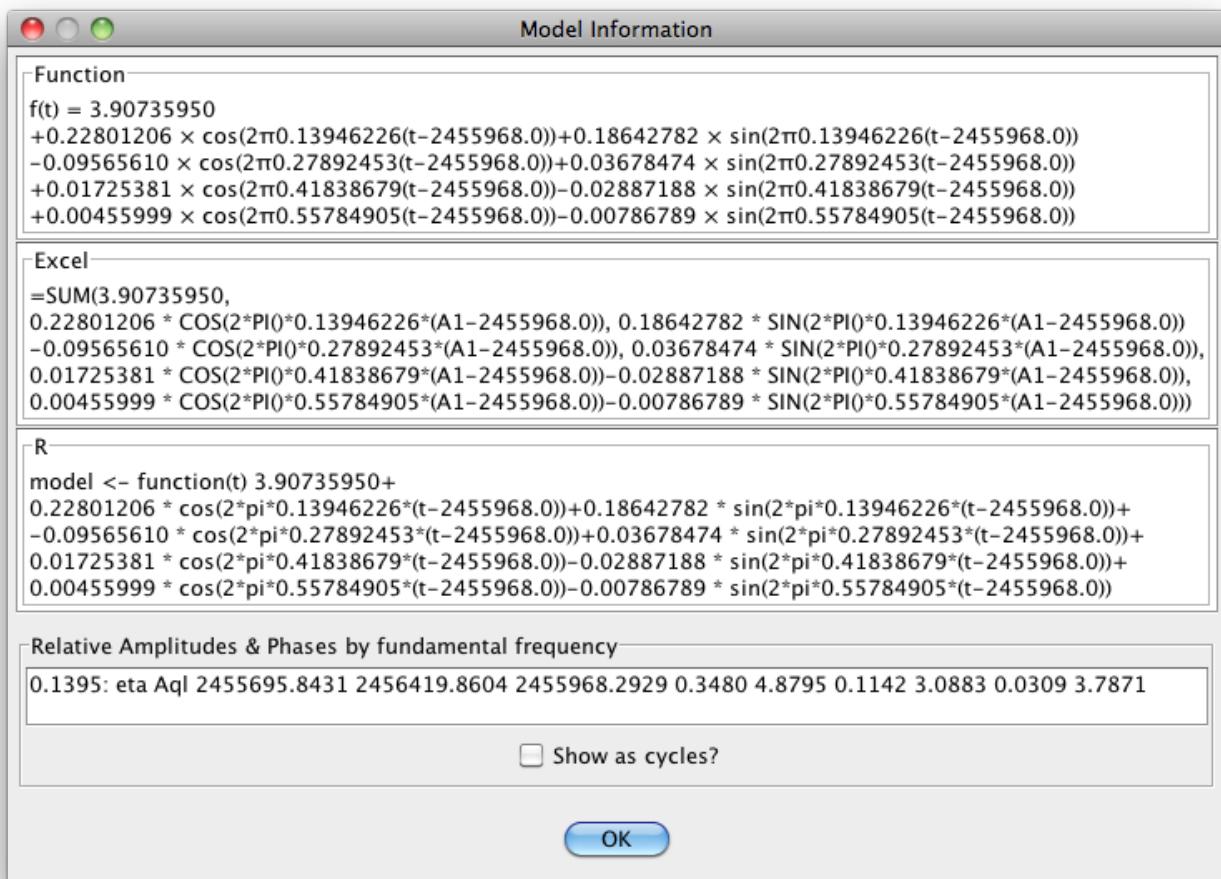
or the following if VStar is in raw data mode, by means of contrast immediately revealing the allure of the combination of period analysis, phase plots and models:



Previous models can be revisited via the models dialog, invoked by Analysis → Models...



Models can be deleted or selected for plotting or inspection. Clicking the Show Model button opens a dialog that shows the selected model function in different forms.



The model equations can be selected and copied with the mouse (via manual selection or select-all/copy keyboard shortcuts) and used as part of a script. For example, from the dialog above, the R Fourier series equation:

```
model <- function(t) 3.90735950+
0.22801206 * cos(2*pi*0.13946226*(t-2455968.0))+0.18642782 *
sin(2*pi*0.13946226*(t-2455968.0))+
-0.09565610 * cos(2*pi*0.27892453*(t-2455968.0))+0.03678474 *
sin(2*pi*0.27892453*(t-2455968.0))+
0.01725381 * cos(2*pi*0.41838679*(t-2455968.0))-0.02887188 *
sin(2*pi*0.41838679*(t-2455968.0))+
0.00455999 * cos(2*pi*0.55784905*(t-2455968.0))-0.00786789 *
sin(2*pi*0.55784905*(t-2455968.0))
```

can be used in conjunction with observations, model data and the script available at the following

location to plot the model equation:

https://sourceforge.net/p/vstar/code/HEAD/tree/trunk/script/plot_model.R

The Excel formula should be copied and pasted into a text editor and linefeeds removed before pasting into Excel.

In the presence of a model with harmonics, the model information dialog will include a pane titled **Relative Amplitudes & Phases by fundamental frequency**. This is a standard way of describing the shape of a waveform, especially for pulsating stars with close-to-perfect periodicity. For more information about this, see Foster (2010).

Having created a model, looking at the residuals series can provide information about any remaining signal (via binned means, ANOVA). The residuals series can, like any other series, have DCDFT applied to it to search for additional periods.

Other DCDFT options: Period and Frequency Range

Up to this point, only the DCDFT standard scan has been mentioned. As can be seen from the Analysis menu, DCDFT with frequency range and period range are also available.

VStar's DCDFT Java implementation is a direct translation of AAVSO's TS program. The following is adapted from the documentation accompanying the Fortran TS implementation created by Matthew Templeton, which in turn was translated from a BASIC implementation by Grant Foster (see *References* section):

The simplest choice is DCDFT Standard Scan. The smallest frequency (longest period) tested by the standard scan is $\frac{1}{4}T$, where T is the total time span of the data. The standard scan will test frequencies from $\frac{1}{4}T$ to $\frac{N}{4}T$, where N is the number of data, in steps of $\frac{1}{4}T$.

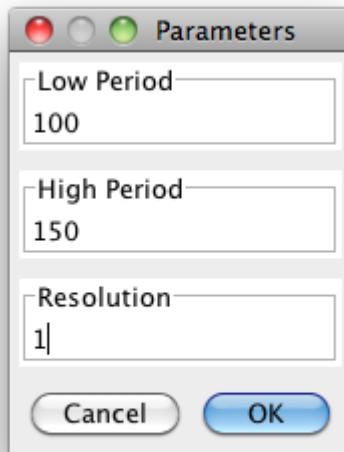
Because VStar utilizes a discrete form of the Fourier transform, the computation time will increase as the square of the number of data points. For 100,000 data points or more, consider averaging or filtering the data points prior to processing.

As it tests each frequency, VStar records the frequency, period, and power. Frequencies with the highest power levels are the most likely possibilities for the actual frequency of a periodic fluctuation in the data (if it has one; after all, not all data are periodic). VStar keeps a record of

the top (highest power) 100 frequencies/periods (see Top Hits pane in the DCDFT result dialog).

To specify a range of frequencies or periods to test, choose option the DCDFT with Frequency Range or DCDFT with Period Range options. VStar requests the low frequency or period to test, the high frequency or period, and for the resolution. The resolution is the spacing between test frequencies or periods. For instance, to test a range of periods, from as low as 100 days to as high as 150 days, if a resolution of 1 is specified, VStar will test all periods from 100 to 150 in 1-day steps.

Note that non-zero values must be entered in the parameter dialog. The following figure shows an example of the parameter dialog for the scenario described above:



The DCDFT with Frequency Range option populates the dialog with an initial frequency range and resolution as described for Standard Scan above.

CLEANest

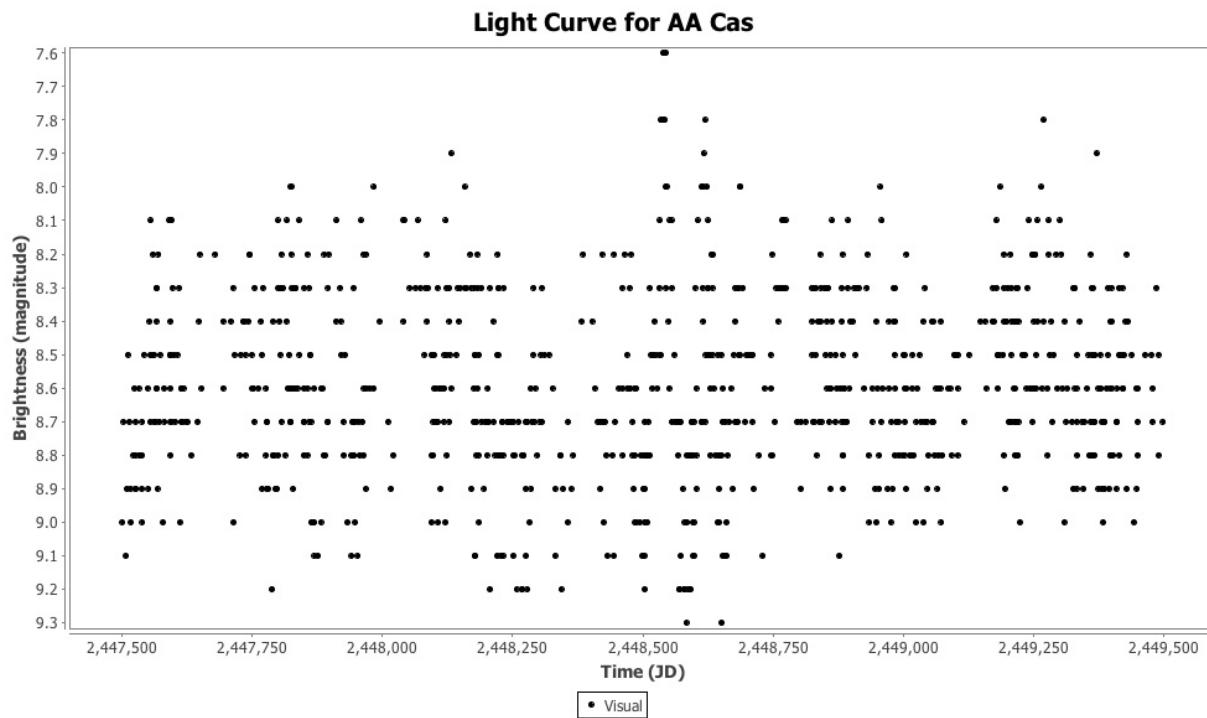
The TS documentation cited above goes on to make the following comment:

When you run a Fourier analysis of the data, it is possible that the "peak" signal or signals you detect may not be the precise frequency actually detected in the data set, because the sampled frequencies tested might be offset slightly from the true signals.

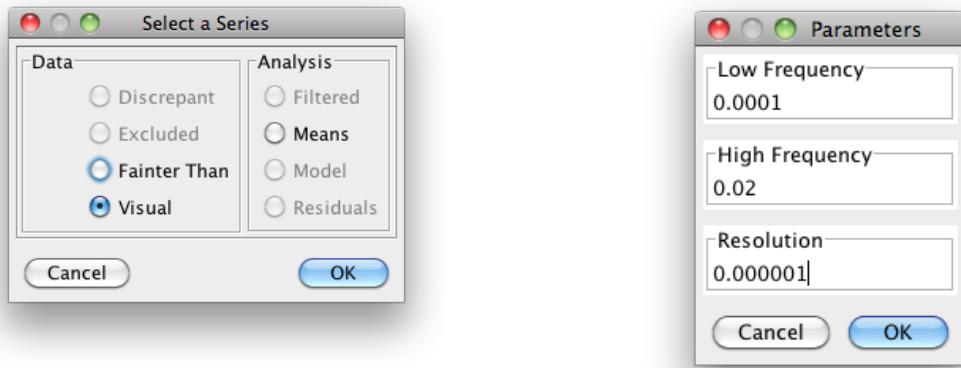
Like TS, VStar is capable of refining one or more periods simultaneously via the CLEANest (Foster 1995) algorithm.

Grant Foster's 1995 CLEANest Fourier Spectrum paper (Foster 1995) gives a number of examples of applying the CLEANest algorithm to datasets, artificial and real. Two of these use AAVSO visual magnitude estimates: S Ori and AA Cas.

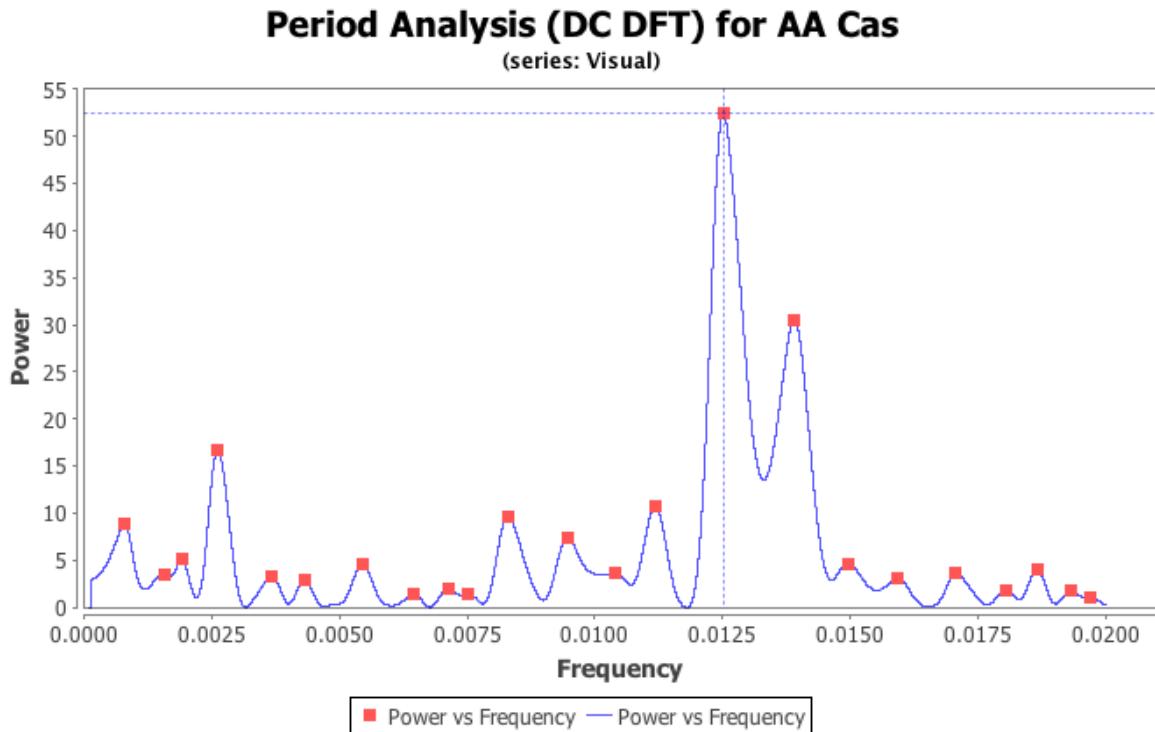
What follows will show VStar's CLEANest implementation applied to AA Cas. Foster (1995) uses an AA Cas dataset in the JD range 2447500 to 2449500. Here is that dataset loaded from the AAVSO International Database (AID).



A DCDFT with frequency range can be initiated from VStar's Analysis menu, selecting the Visual band and specifying minimum and maximum frequencies, the range over which to scan (0.0001 to 0.02), and frequency resolution (0.000001) over the range.



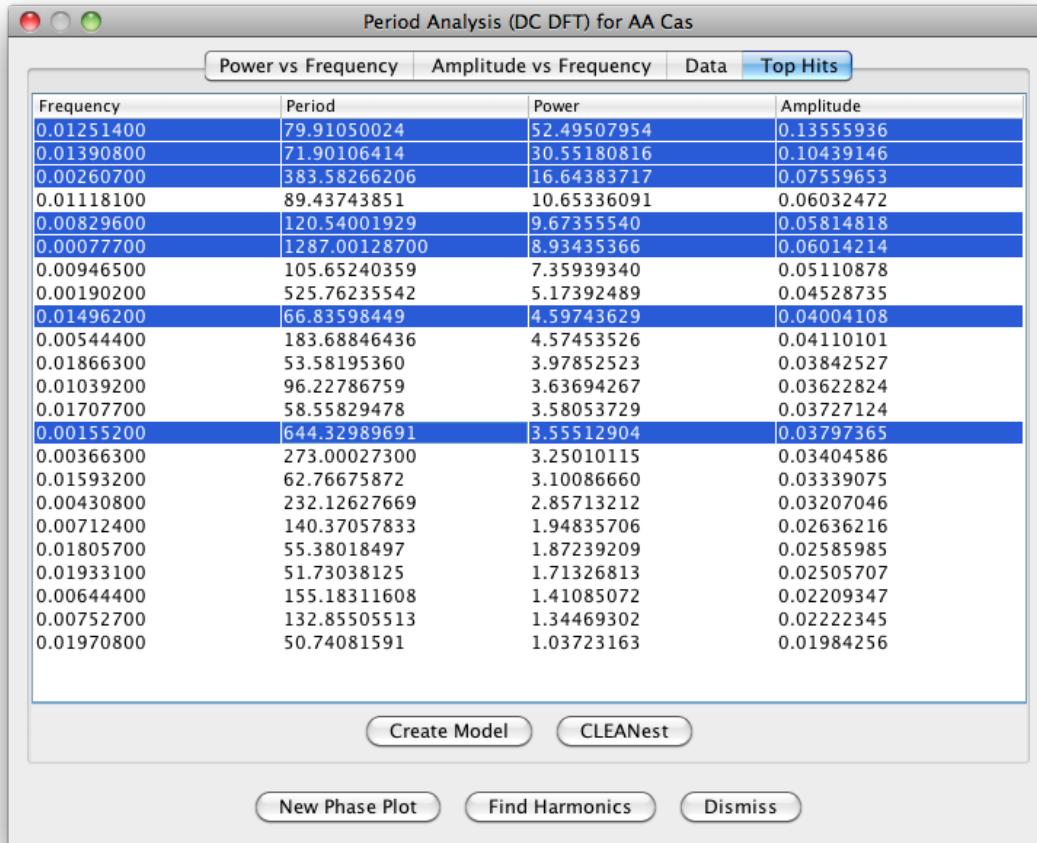
This results in the following power spectrum (in the Power vs Frequency pane) with the orange squares showing peaks or “top hits”.



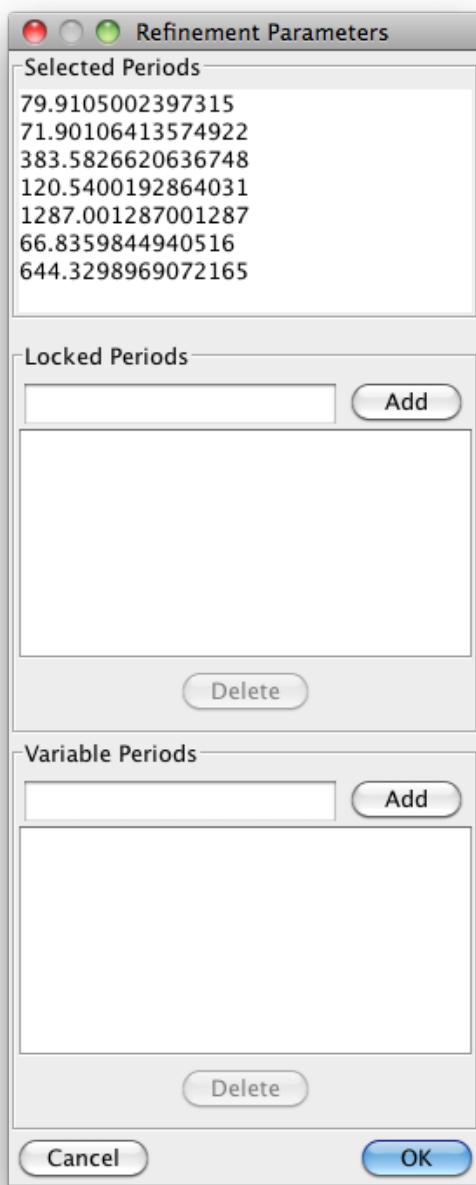
These top hits are shown in the next diagram in tabular form.

Notice that seven top hits are shown as being selected using combinations of shift-click and control-click (Windows) or command-click (Mac). The initial input values to CLEANest are not stated in Foster (1995; section 5, page 1900), but the rows selected above fairly closely

correspond to what the paper presents.



Clicking the CLEANest button from the Top Hits pane opens this dialog.



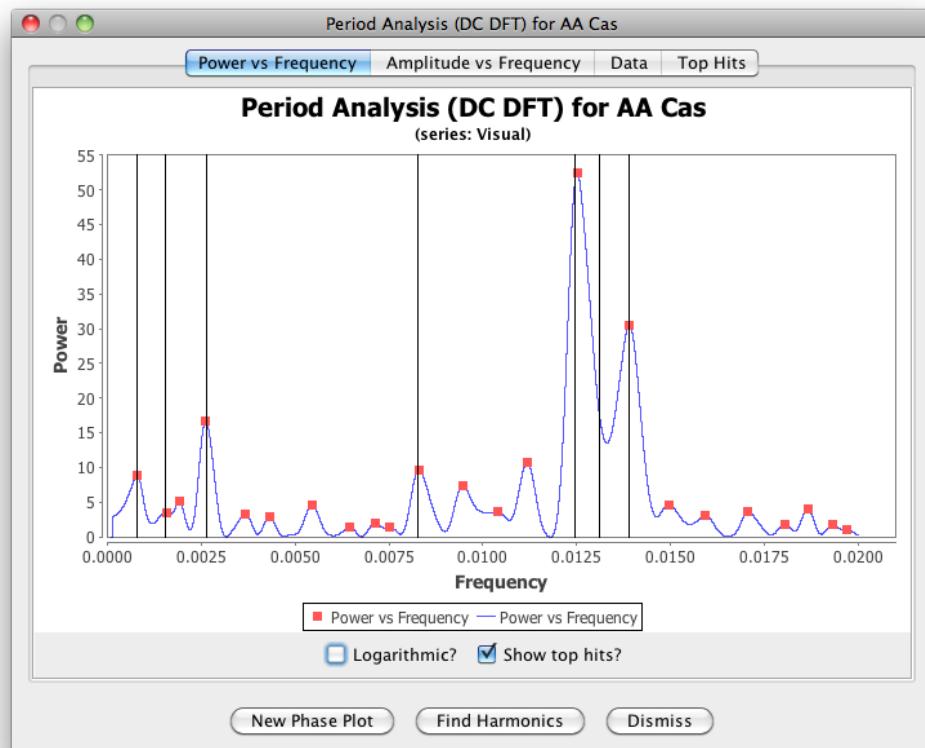
Clicking OK here adds seven new top hits with the same power value, shown multiply-selected in the top hits list and annotated on the power spectrum.

Period Analysis (DC DFT) for AA Cas

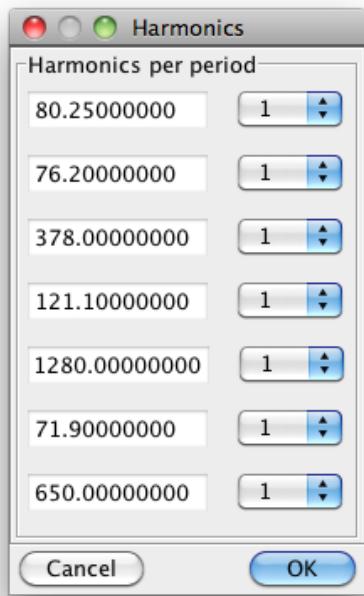
Power vs Frequency | Amplitude vs Frequency | Data | **Top Hits**

Frequency	Period	Power	Amplitude
0.01246106	80.25000000	155.03700075	0.00902804
0.01312336	76.20000000	155.03700075	0.00902804
0.00264550	378.00000000	155.03700075	0.00902804
0.00825764	121.10000000	155.03700075	0.00902804
0.00078125	1280.00000000	155.03700075	0.00902804
0.01390821	71.90000000	155.03700075	0.00902804
0.00153846	650.00000000	155.03700075	0.00902804
0.01251400	79.91050024	52.49507954	0.13555936
0.01390800	71.90106414	30.55180816	0.10439146
0.00260700	383.58266206	16.64383717	0.07559653
0.01118100	89.43743851	10.65336091	0.06032472
0.00829600	120.54001929	9.67355540	0.05814818
0.00077700	1287.00128700	8.93435366	0.06014214
0.00946500	105.65240359	7.35939340	0.05110878
0.00190200	525.76235542	5.17392489	0.04528735
0.01496200	66.83598449	4.59743629	0.04004108
0.00544400	183.68846436	4.57453526	0.04110101
0.01866300	53.58195360	3.97852523	0.03842527
0.01039200	96.22786759	3.63694267	0.03622824
0.01707700	58.55829478	3.58053729	0.03727124
0.00155200	644.32989691	3.55512904	0.03797365
0.00366300	273.00027300	3.25010115	0.03404586
0.01593200	62.76675872	3.10086660	0.03339075
0.00430800	232.12627669	2.85713212	0.03207046
0.00712400	140.37057833	1.94835706	0.02636216

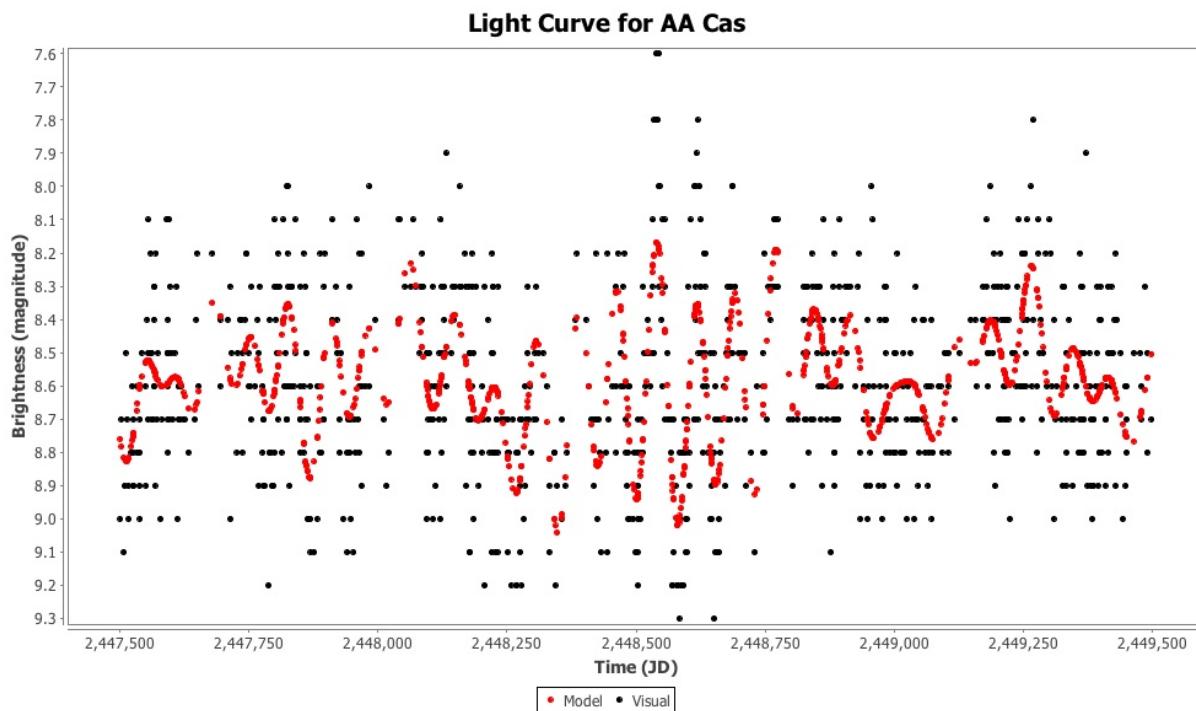
Create Model | CLEANest | New Phase Plot | Find Harmonics | Dismiss



Now click Create Model in the Top Hits pane and the following dialog will open.

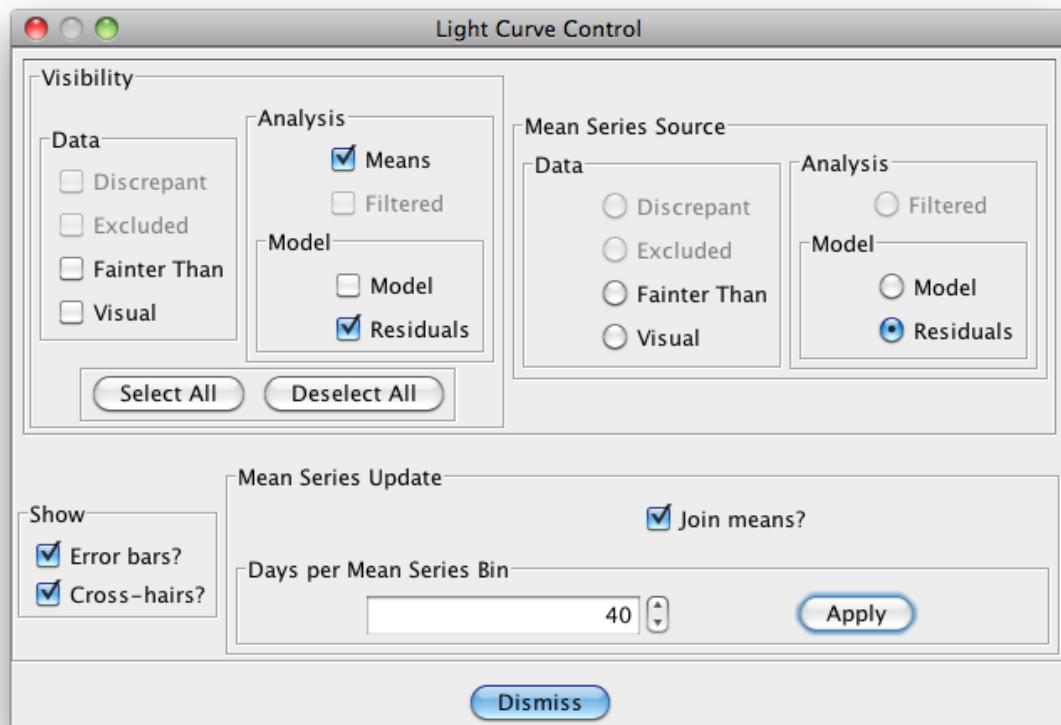


Click OK and the main plot will have an additional “model” series added. Dismiss the main DCDFT dialog to return to the main VStar window.

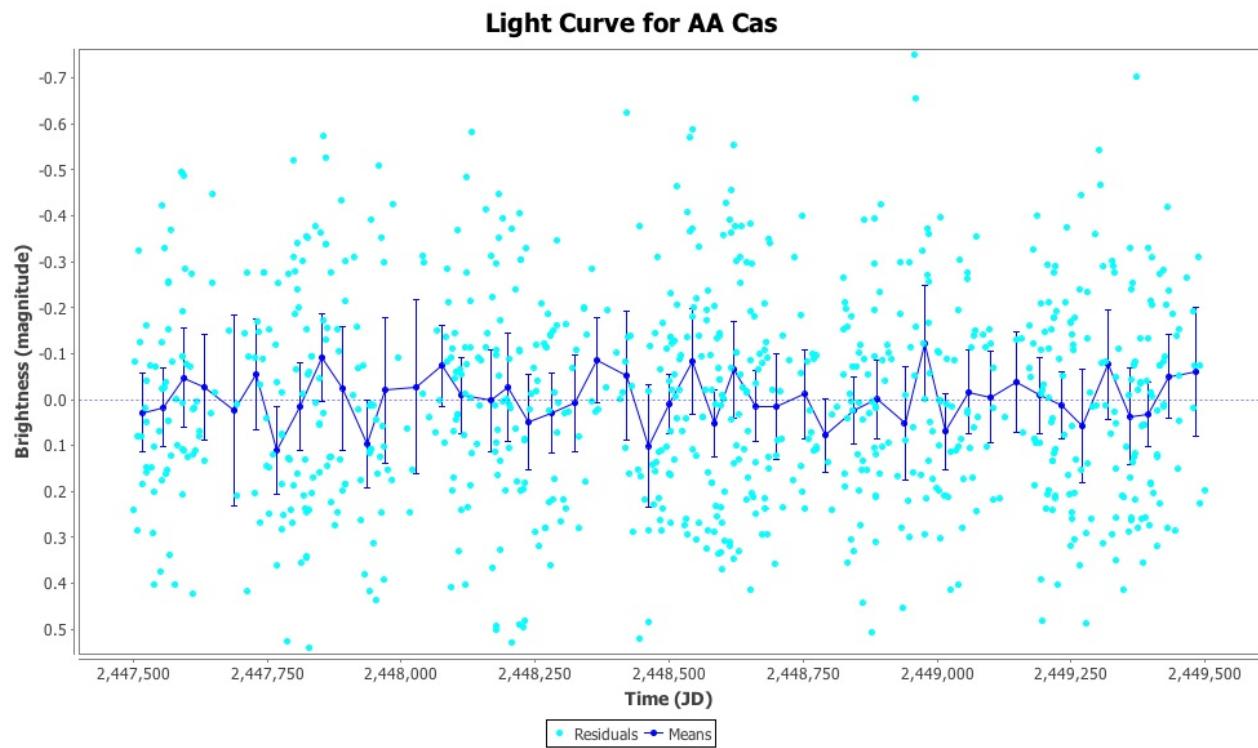


The residuals for this model can be viewed by opening the Plot Control dialog from the View

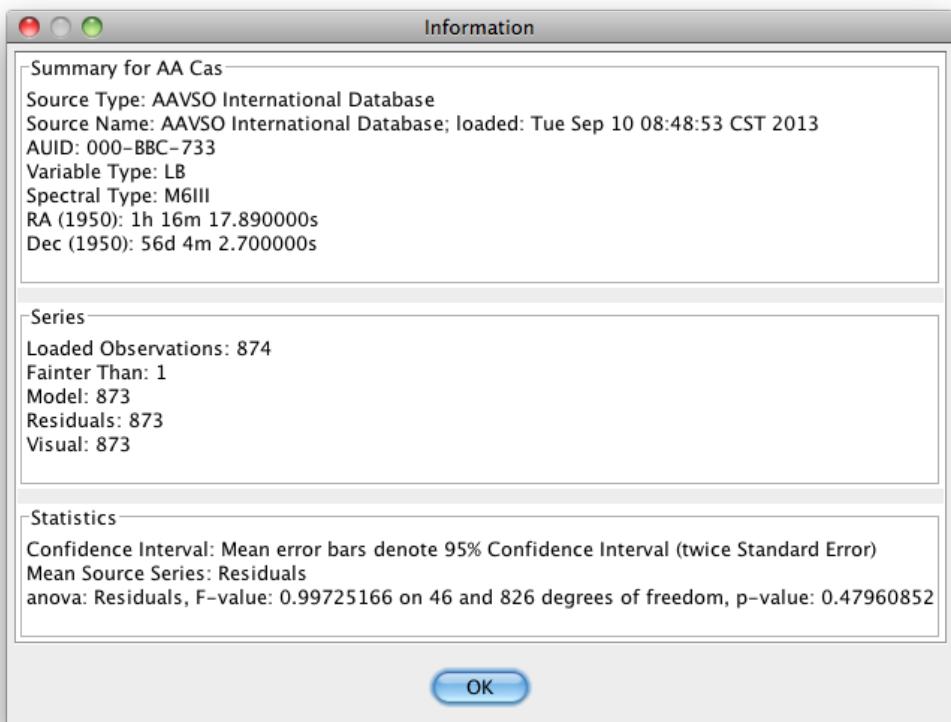
menu and setting it as shown, including changing the Days per Mean Series Bin (and clicking Apply).



Dismissing the dialog changes the plot to look like this:



The ANOVA value in the Info dialog (File menu) also suggests that there is unlikely to be any signal remaining to be found in the residuals.



We can say this since the p-value of ~ 0.48 suggests that what is observed in the residuals is likely to be due to noise rather than the presence of a non-trivial signal.

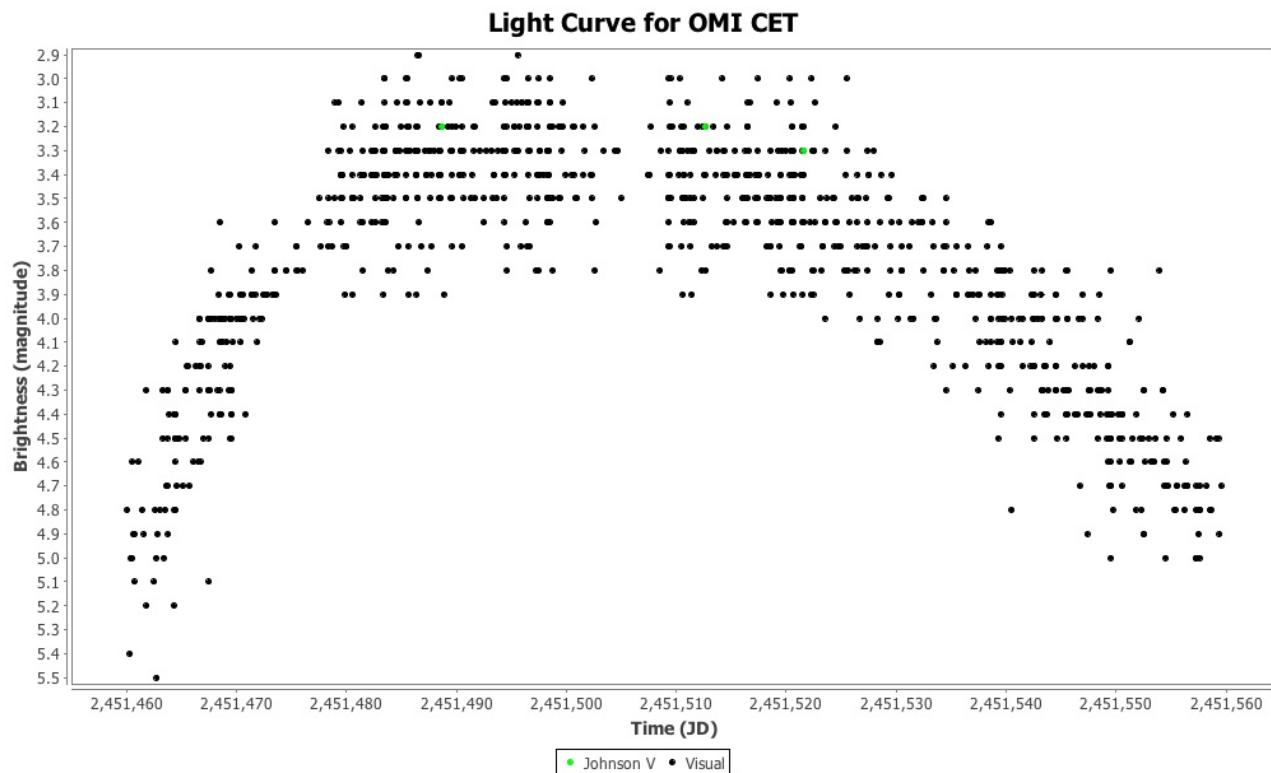
For another example of using CLEANest in VStar, see the following:

<http://dbenn.wordpress.com/2011/07/13/bz-uma-model-and-cleanest/>

Polynomial Fit

Just as a mean curve can be thought of as a model of a dataset that has a smoothing effect, a *polynomial model* can also be used as a smoothing mechanism to capture key aspects of a dataset without all the “rough edges.”

Suppose you want to determine the time of maximum for the following Mira maximum light curve segment.



One way to do this is by *fitting a polynomial* to the data and using that to determine what the time of maximum is. The relationship between a sequence of x (time) values and corresponding dependent y (magnitude) values can be modeled as an *nth order* polynomial.

A polynomial model of *degree n* is defined as a function of time with the following form:

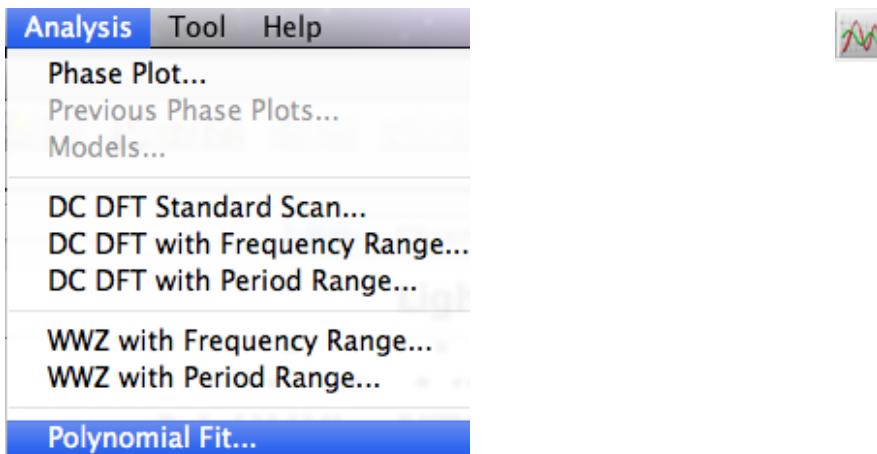
$$y = f(t) = \beta_0 + \beta_1 t + \beta_2 t^2 + \beta_3 t^3 + \dots + \beta_n t^n$$

where t is time (e.g. Julian Date) and β_n is a *coefficient*. The sum of these coefficients multiplied

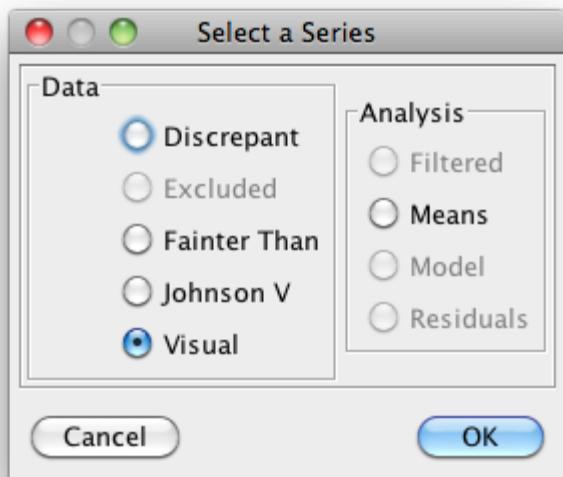
by corresponding time values, each raised to a power (the number or index of the coefficient: $0 \dots n$), creates a curve that approximates the data. Notice that β_0 is just $\beta_0 t^0$. The process through which the coefficients are determined is beyond the scope of this document. See *References and Further Reading* for more background information about polynomial models and data fitting.

The *degree* of the polynomial will have a bearing on the values of the minima or maxima (*critical points* or *extrema*) found since (roughly speaking) the higher the degree, the better the polynomial approximates the features of the data.

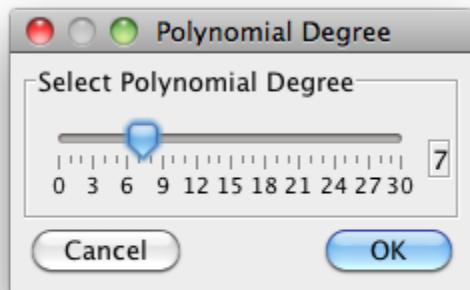
In VStar, a polynomial fit can be created via the Analysis → Polynomial Fit... menu item or the corresponding toolbar button.



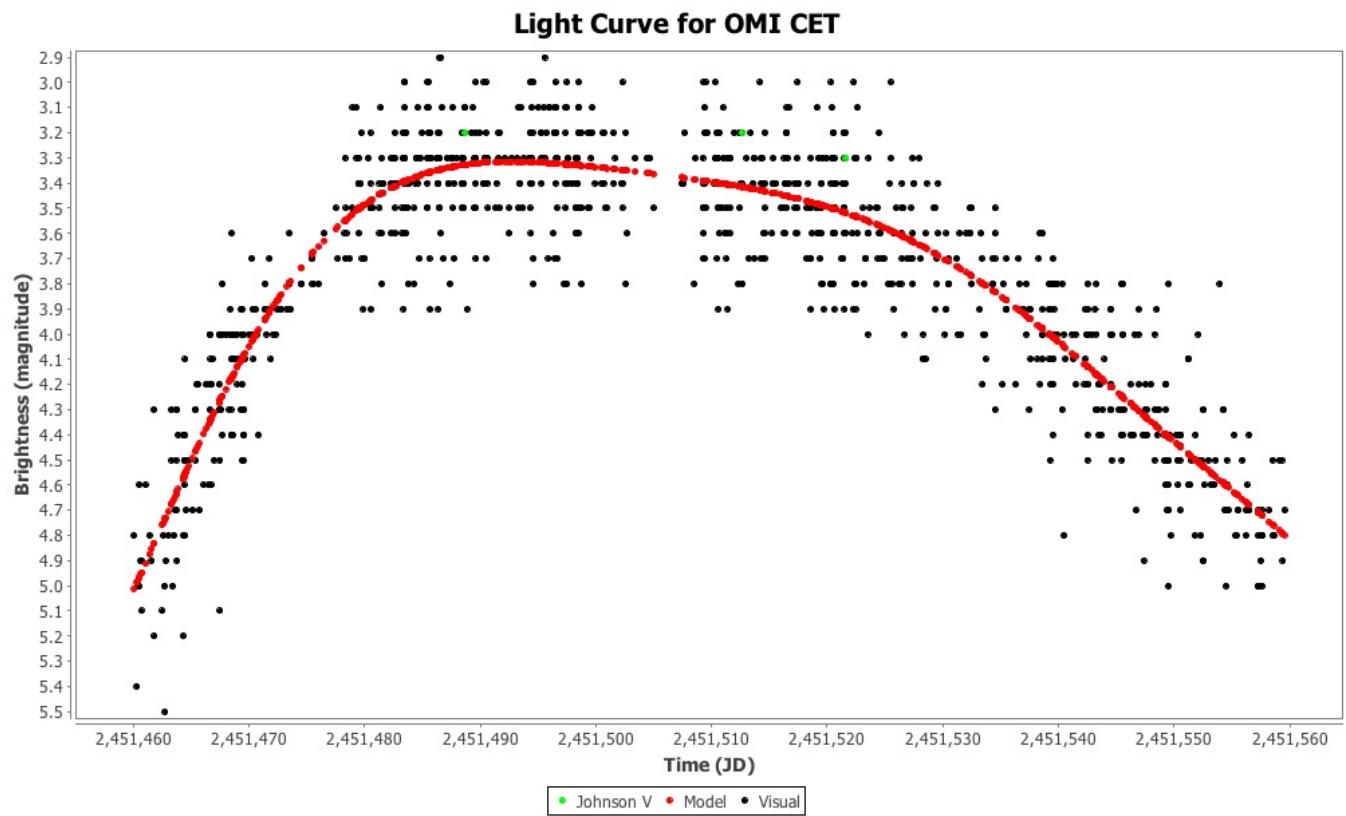
Selecting one of these will result in a Series Selection Dialog being invoked



In this case, the Visual series is selected and clicking the OK button yields the polynomial degree dialog.



This is used to specify the number of degrees for the polynomial to be created, 7 in this case. Clicking the OK button creates the polynomial fit of the Visual series and adds a Model series as follows:



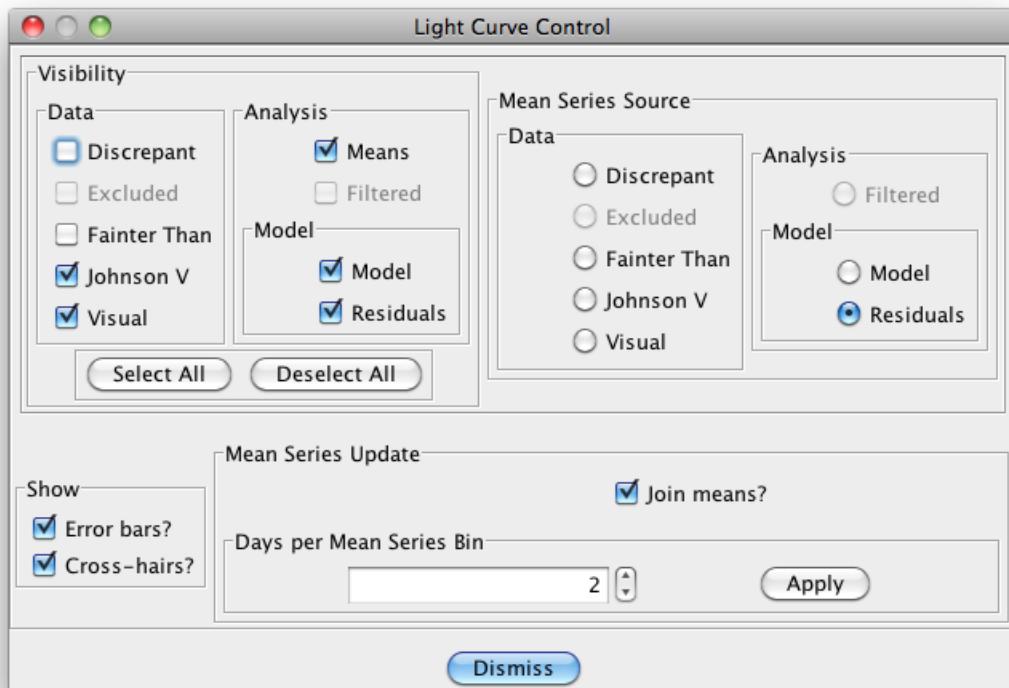
Foster (2010) discusses essentially the same example except that R (a statistical programming language) is used rather than VStar. The **Model** tab contains a list of the polynomial fit data points visible on the plot.

Plot Observations Means Model Residuals			
Julian Day	Calendar Date	Magnitude	
2451460.076400	1999 OCT 8	5.011615	
2451460.310000	1999 OCT 8	4.987632	
2451460.400000	1999 OCT 8	4.978363	
2451460.500000	1999 OCT 9	4.968048	
2451460.533300	1999 OCT 9	4.964610	
2451460.638900	1999 OCT 9	4.953693	
2451460.700000	1999 OCT 9	4.947370	
2451460.700700	1999 OCT 9	4.947297	
2451461.048600	1999 OCT 9	4.911200	

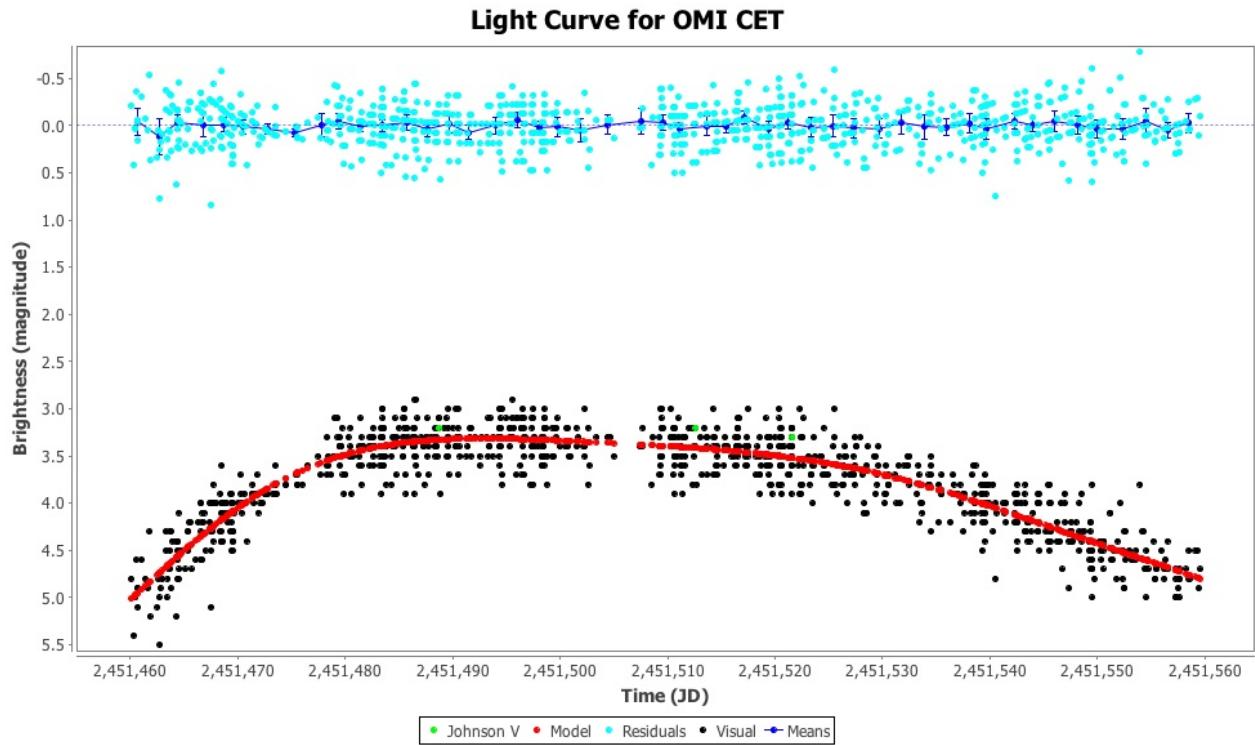
The **Residuals** tab contains a list of the residuals (observation minus model for each such pair of values) from the polynomial fit.

			Plot	Observations	Means	Model	Residuals
Julian Day	Calendar Date	Magnitude					
2451460.076400	1999 OCT 8	-0.211615					
2451460.310000	1999 OCT 8	0.412368					
2451460.400000	1999 OCT 8	0.021637					
2451460.500000	1999 OCT 9	0.031952					
2451460.533300	1999 OCT 9	-0.364610					
2451460.638900	1999 OCT 9	-0.053693					
2451460.700000	1999 OCT 9	-0.047370					
2451460.700700	1999 OCT 9	0.152703					

This series can also be viewed on the Plot pane by selecting the Residuals series in the Plot Control Dialog series visibility checkboxes.



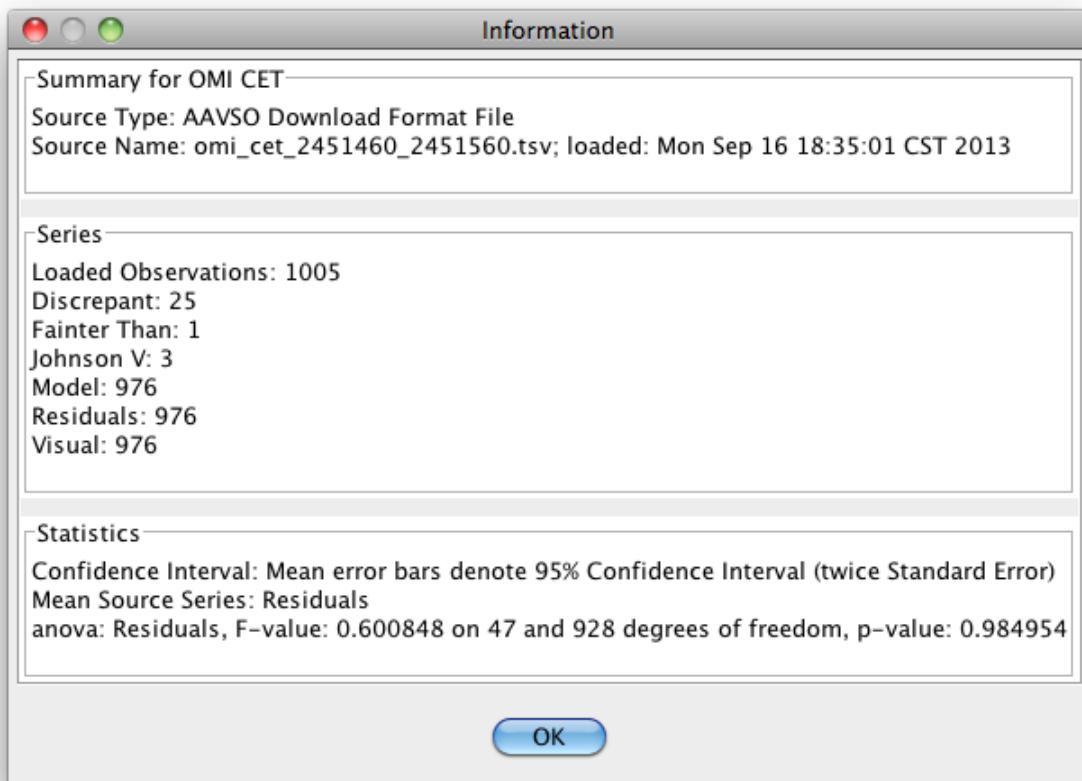
In addition, as shown above, if the Residuals mean series source radio button is selected and Days per Mean Series bin set to 2, as shown, the Plot Pane will look like this:



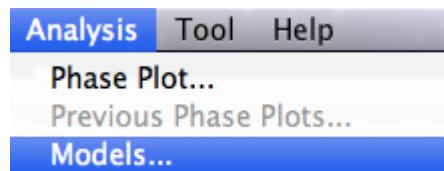
As for a model or residuals plot, mean plot data points can be viewed in tabular format in the Means list by selecting the corresponding tab, for example:

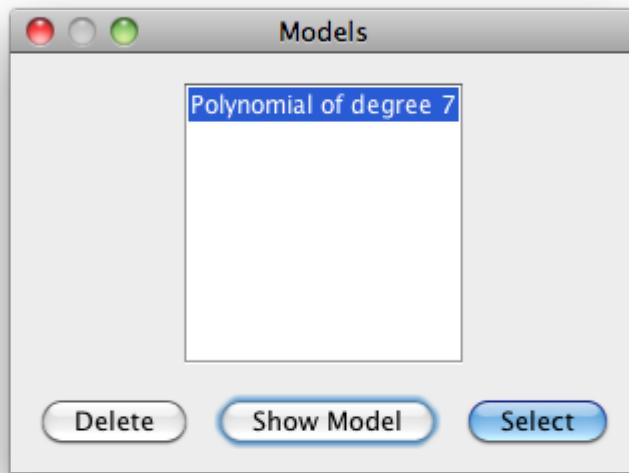
Julian Day	Calendar Date	Mean Magnitude	Standard Error of the Average
2451460.630790	1999 OCT 9	-0.044445	0.071520
2451462.703323	1999 OCT 11	0.114983	0.093679
2451464.458864	1999 OCT 12	-0.033523	0.041960
2451466.812405	1999 OCT 15	-0.009216	0.062839
2451468.737686	1999 OCT 17	-0.008915	0.039249
2451470.514332	1999 OCT 19	0.012571	0.037948
2451472.748873	1999 OCT 21	0.033468	0.028781
2451475.285429	1999 OCT 23	0.079523	0.019781
2451477.762800	1999 OCT 26	-0.005127	0.060773
2451479.431043	1999 OCT 27	-0.045181	0.042450

Looking at the ANOVA statistic in the Information Dialog (File → Info...), it would appear that there is no significant signal remaining, so the fit is apparently a good one. We can say this since the p-value of ~ 0.98 suggests that what is observed in the residuals is likely to be due to noise rather than the presence of a non-trivial signal.



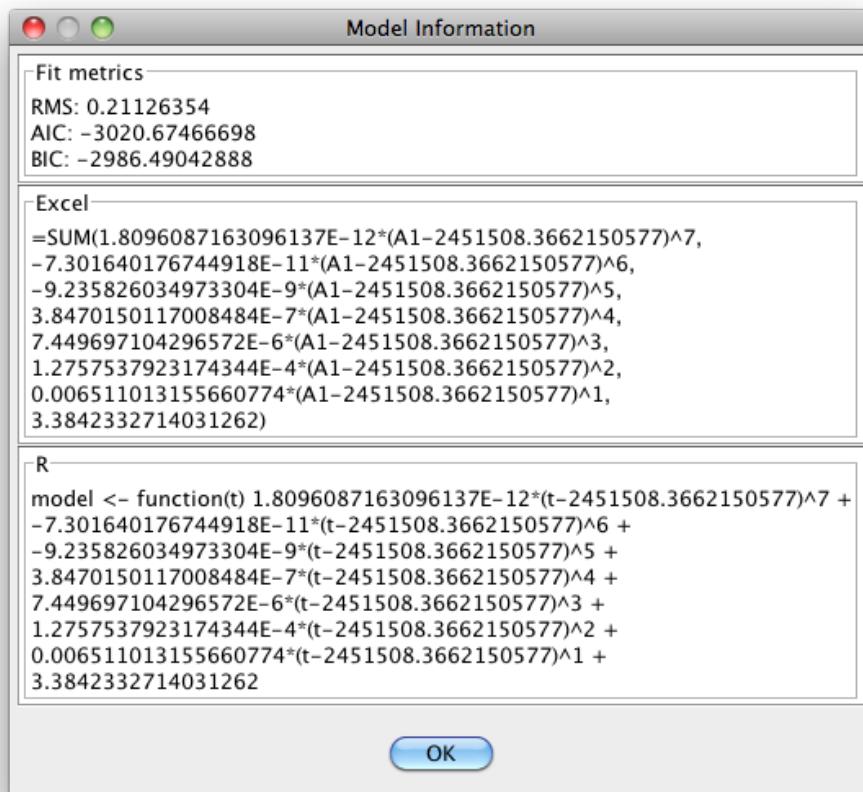
Additional information about created models can be obtained via the Model Information dialog. This can be invoked from the Analysis menu's Models... item:





A model can be deleted via the `Delete` button. The `Select` button is used to select the highlighted model in the list for viewing on the Plot Pane and the Model and Residuals lists.

Selecting the `Show Model` button invokes the following dialog:



The Model Information dialog gives:

- Excel and R functions for the polynomial.
- The Root Mean Square (RMS) value for the fit.
- Akaike and Bayesian Information Criteria, goodness of fit measures (AIC, BIC).

Notice that the Excel and R functions reveal the polynomial degree in the number of coefficients as explained earlier in this section. Notice also that each time term has a Julian Date *zero-point* subtracted from it, reflecting the way in which the polynomial fit was created, in order to reduce the magnitude of values (and possible loss of precision) when higher powers are involved.

The R function can be used in conjunction with observations, model data and the script available at the following location to plot the model equation:

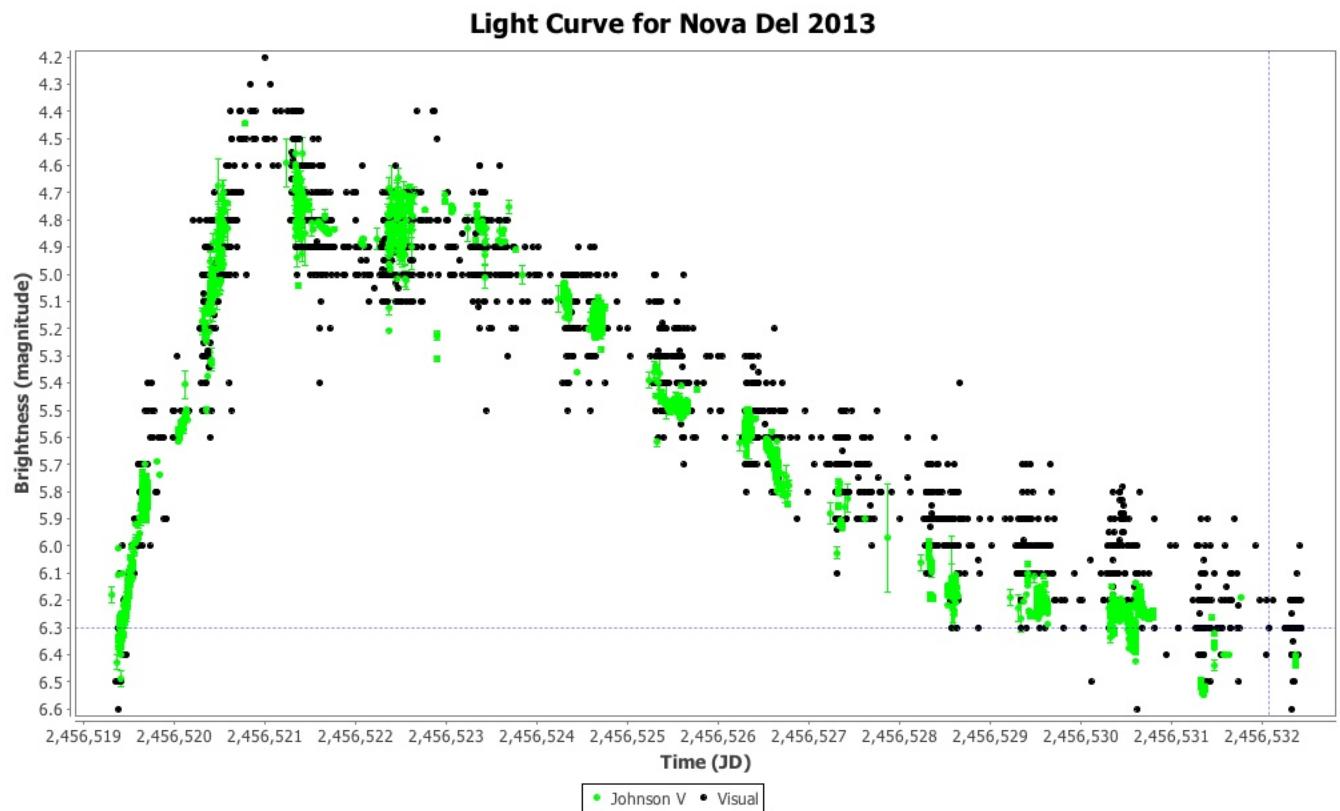
https://sourceforge.net/p/vstar/code/HEAD/tree/trunk/script/plot_model.R

The Excel formula should be copied and pasted into a text editor and newlines removed before

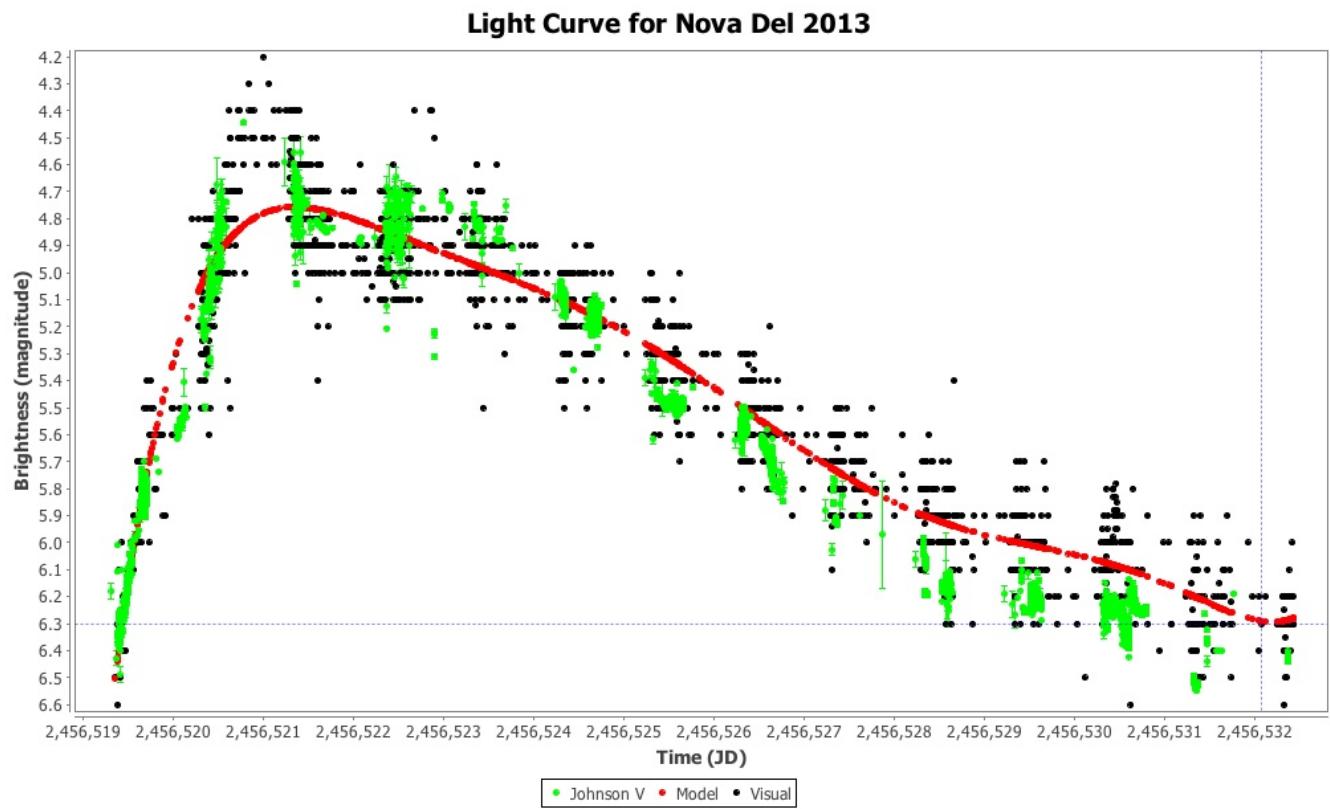
pasting it into Excel.

A treatment of RMS and Information Criteria is beyond the scope of this document. See Foster (2010) and the relevant “Polynomial Fit” sub-section of *References and Further Reading*.

A polynomial fit can also be useful to characterize the overall shape of some light curves. Take for example the following light curve for Nova Delphini 2013:



A polynomial of degree 8 gives this:

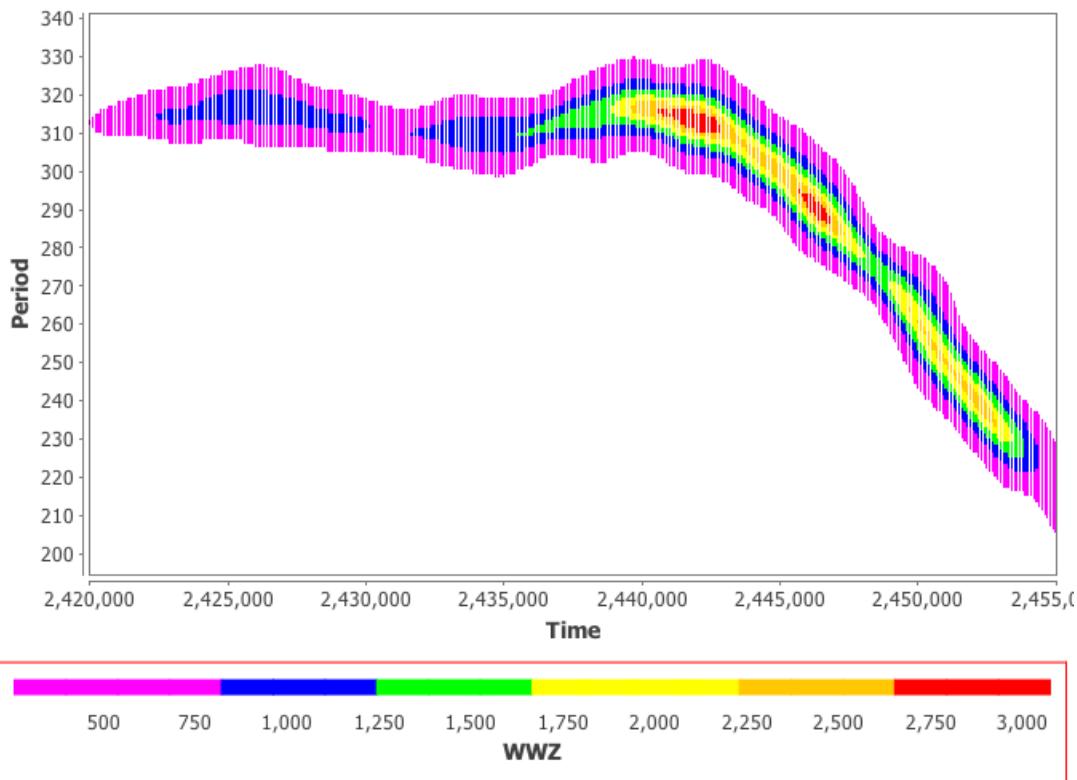


Note that this could also be useful in determining the rate of decline after maximum magnitude is reached.

Time-Frequency Analysis

VStar provides time-frequency analysis functionality in the form of Weighted Wavelet Z-Transform (WWZ) (Foster 1996). The user specifies a series (e.g. Visual, Johnson V), a frequency or period range, and a resolution; the result is a visualization of how the period changes over time. This can be viewed as a 2D graph, a contour plot, a rotatable 3D graph, or in tabular form. Periods at particular points in time can be selected for phase plot creation. A detailed explanation of the statistical analysis techniques being utilized here is beyond the scope of this document. The interested reader is referred to Foster (2010) and *References and Further Reading*.

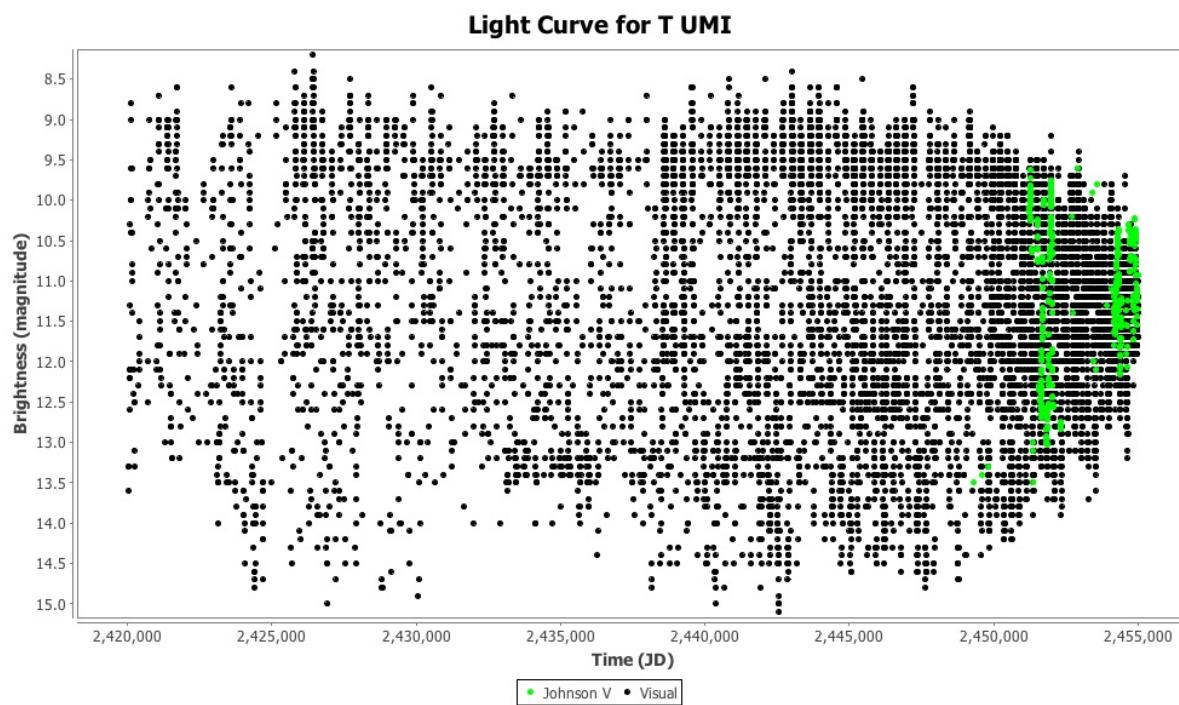
The following shows the period change for T UMi in the year range 1913 to 2009. Here the color represents the WWZ statistic, the strength of a periodicity, at a particular time. This example is discussed in Foster (2010).



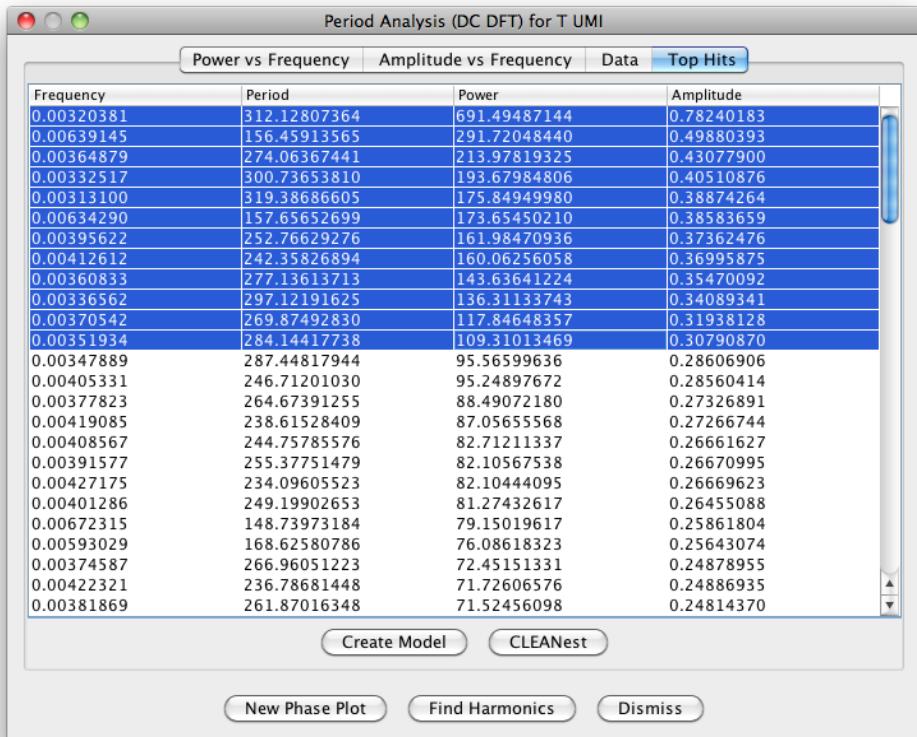
This section explores WWZ via T UMi.

The T UMi dataset below, taken from the AAVSO International Database (AID), spans the JD range 2,420,000 to 2,455,000. Even a visual inspection suggests amplitude and possible period

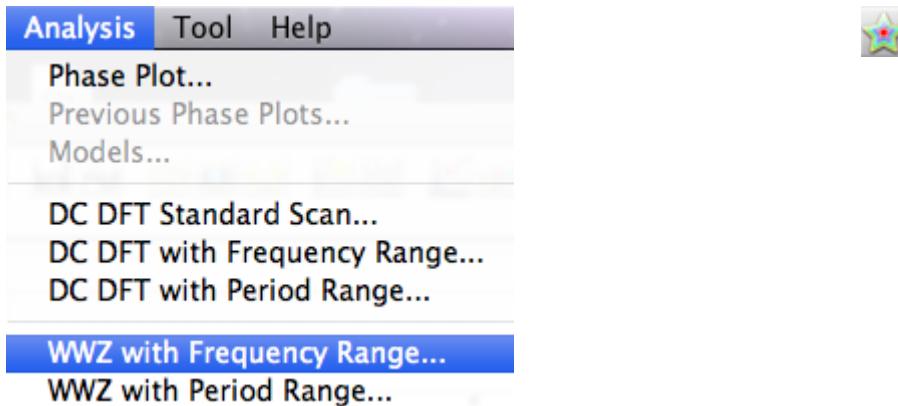
change.



Creating a phase plot with periods taken from DCDFT (Date Compensated Discrete Fourier Transform) in VStar does not result in an obviously "clean" fit over the time range, as evidenced by DCDFT high-power top-hits (via Analysis → DCDFT Standard Scan...):



WWZ helps to explain why. It is available via two **Analysis** menu items and a toolbar button.

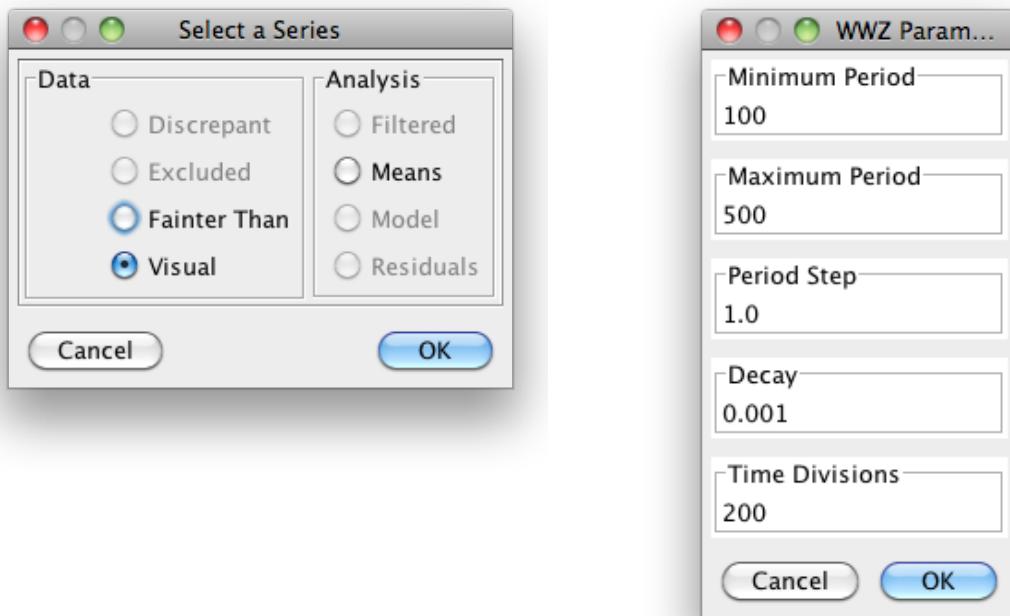


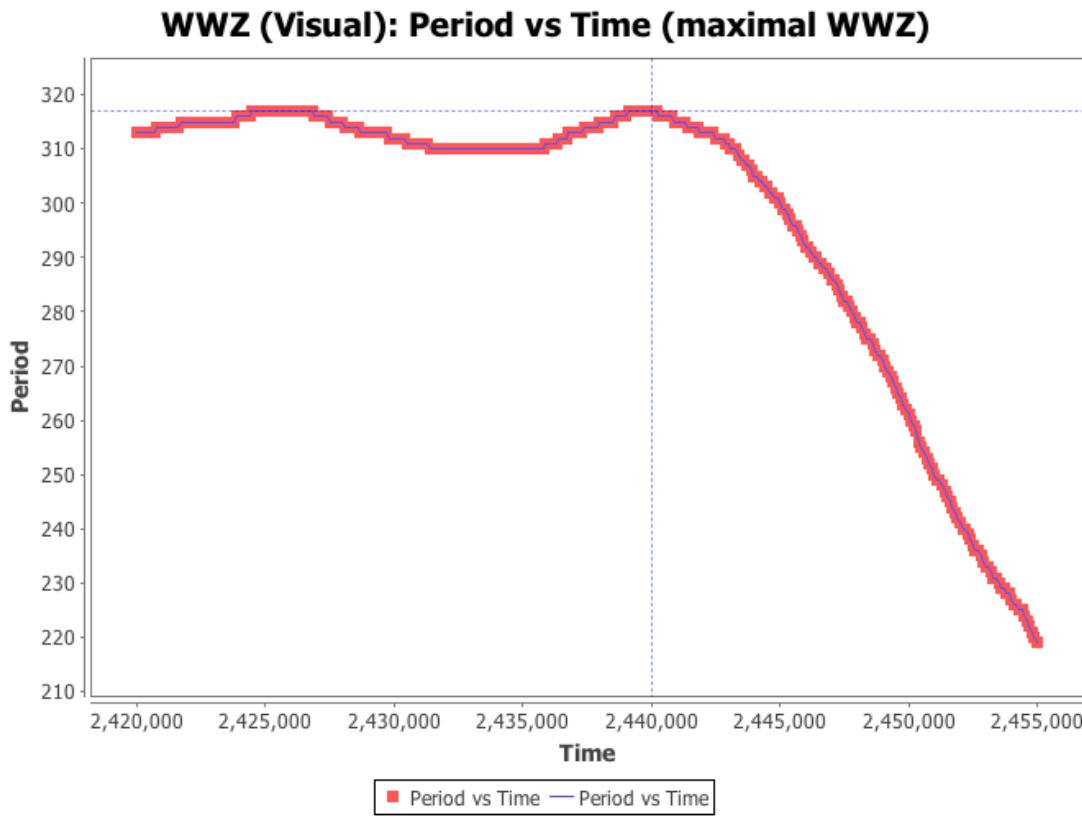
The toolbar button corresponds to the **WWZ with Period Range** item of the **Analysis** menu.

Applying WWZ to the dataset requires selecting the series to be analyzed, then the following parameters:

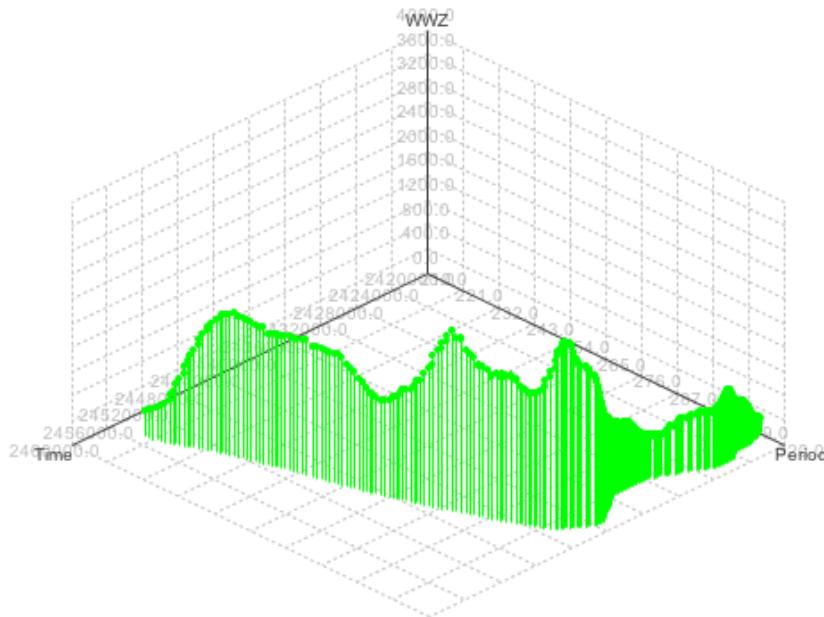
- Minimum frequency or period.
- Maximum frequency or period.
- Frequency or period step: this is the resolution in the frequency domain.
- Decay: this is the wavelet window; smaller values yield better resolution of variation.
- Time divisions: this gives the time resolution, defined as $\frac{JD_{max} - JD_{min}}{\text{time divisions}}$; this is revealed by the extent of “vertical banding” in the contour plot (as we see in the first figure of this section).

For T UMi, selecting Analysis → WWZ with Period Range, selecting the Visual series and entering values of 100 days (min period), 500 days (max period), 1 day (period resolution), 0.001 (decay) and 200 (time divisions) for the above parameters, gives a plot of period vs time:





The contour plot at the start of this section adds the WWZ statistic to the Period vs Time plot above. The Result dialog provides a number of other plots and tables. One of these is a 3D plot that provides the same information as the contour plot but allows rotation in all axes.



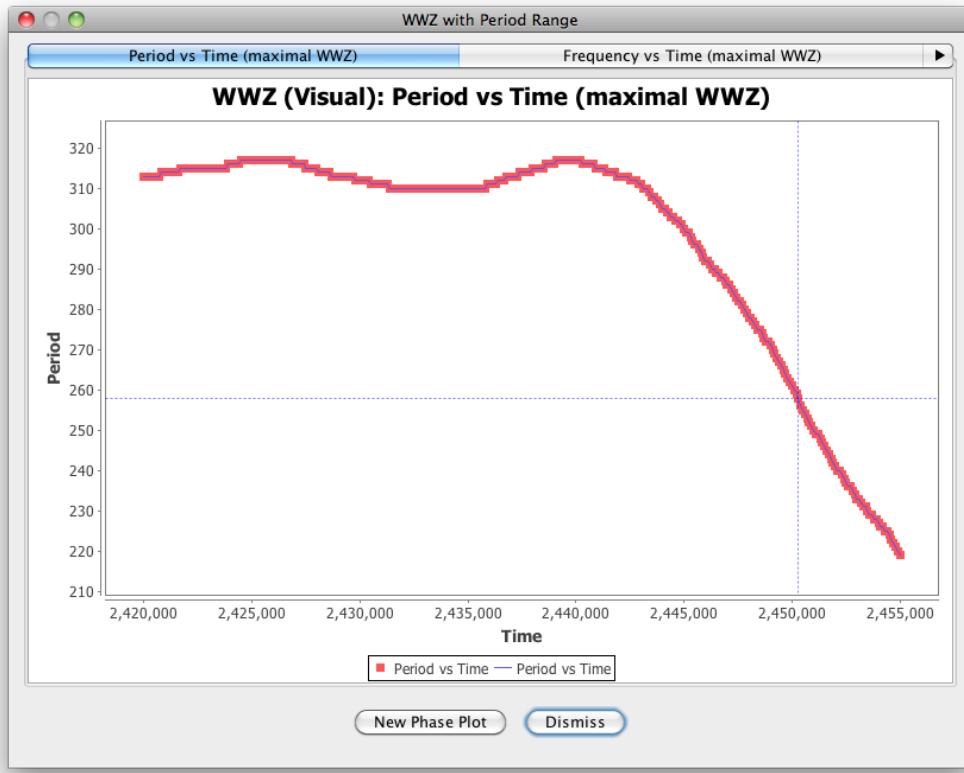
The WWZ algorithm generates the following statistics for the specified series and frequency or period range and time resolution which VStar presents in plot or tabular form (adapted from the AAVSO Fortran WWZ implementation documentation):

- Time: The time being examined, in JD.
- Frequency: The frequency being tested, in cycles per time unit.
- Period: The corresponding period under test, in JD.
- WWZ: Value of the WWZ; this is approximately an F-statistic with Effective number of data and 2 degrees of freedom, and expected value 1. In short, it indicates whether or not there is a periodic fluctuation of the given frequency at the given time.
- Semi-amplitude: Weighted wavelet amplitude; if the signal is periodic at the frequency being tested, this gives the (real semi-) amplitude of the corresponding best-fit sinusoid.
- Mean magnitude: Mean apparent magnitude of the object at the corresponding time.
- Effective number of data: The effective number of data for the given time and frequency being tested.

The maximal WWZ table shows the points at which the WWZ statistic was maximal:

Period vs Time vs WWZ 3D (maximal WWZ)				WWZ Results		Maximal WWZ Results	
Time	Frequency	Period	WWZ	Semi-amplitude	Mean Magnitude	Effective number of data	
2420000.000000	0.003195	313.000000	254.772811	2.643359	12.259818	140.769974	
2420100.000000	0.003195	313.000000	266.853592	2.630088	12.242746	147.206801	
2420200.000000	0.003195	313.000000	278.767516	2.615324	12.225057	153.607063	
2420300.000000	0.003195	313.000000	290.451820	2.599119	12.206843	159.921423	
2420400.000000	0.003195	313.000000	301.872704	2.581556	12.188219	166.112167	
2420500.000000	0.003195	313.000000	313.033314	2.562755	12.169330	172.157607	
2420600.000000	0.003195	313.000000	323.979413	2.542866	12.150344	178.055437	
2420700.000000	0.003195	313.000000	334.802121	2.522074	12.131454	183.824585	
2420800.000000	0.003185	314.000000	346.875131	2.526676	12.140931	189.921968	
2420900.000000	0.003185	314.000000	360.068145	2.503954	12.121856	195.598931	
2421000.000000	0.003185	314.000000	374.041085	2.481008	12.103491	201.328956	
2421100.000000	0.003185	314.000000	389.046472	2.458116	12.086105	207.197608	
2421200.000000	0.003185	314.000000	405.348642	2.435548	12.069956	213.294638	
2421300.000000	0.003185	314.000000	423.209725	2.413556	12.055284	219.706159	
2421400.000000	0.003185	314.000000	442.871871	2.392362	12.042299	226.505352	
2421500.000000	0.003185	314.000000	464.534555	2.372150	12.031166	233.741655	
2421600.000000	0.003185	314.000000	488.325850	2.353062	12.021999	241.428536	
2421700.000000	0.003175	315.000000	515.761819	2.331226	12.017918	250.227332	
2421800.000000	0.003175	315.000000	545.949042	2.314019	12.011903	258.657815	
2421900.000000	0.003175	315.000000	577.679847	2.298206	12.008035	267.230949	
2422000.000000	0.003175	315.000000	610.330411	2.283790	12.006238	275.709610	
2422100.000000	0.003175	315.000000	643.068868	2.270762	12.006388	283.797238	
2422200.000000	0.003175	315.000000	674.904519	2.259106	12.008320	291.163513	
2422300.000000	0.003175	315.000000	704.783115	2.248810	12.011840	297.484676	
2422400.000000	0.003175	315.000000	731.004533	2.235383	12.009335	298.036027	
2422500.000000	0.003175	315.000000	748.646541	2.230028	12.016003	299.183124	
2422600.000000	0.003175	315.000000	762.919087	2.225684	12.025351	300.882208	

A phase plot can be created from a selected result table row or plot datapoint in the WWZ results dialog via the New Phase Plot button.



The AAVSO Fortran WWZ implementation's documentation has a section on suggested parameter values, minimally adapted below:

The chosen frequency range should only cover the range of frequencies of astrophysical interest to reduce computation time. For Mira and Semiregular variables, choosing a frequency range between 0.0001 ($P = 10000$ days) and 0.02 ($P = 50$ days) with Δf of 0.00001 is reasonable, and should not oversample the frequency spectrum too severely. Be sure to choose frequency values that are physically relevant to the system you are studying, and to the data available.

The decay constant, c , defines the width of the wavelet "window". It defines the number of cycles of a given frequency f expected within the window. Smaller values of c will produce wider windows. Reasonable values of c are between 0.001 and 0.0125. Note that using small values of c will result in improved frequency resolution of variations, but will smear out temporal variations. Conversely, large values of c will improve the temporal resolution, but will generate larger uncertainties in peak frequency.

WWZ scans the data set starting from the earliest data and progressing to the latest. If you notice that the program returns zero values of the WWZ statistic, then you probably have a large data

gap just prior to the point where the zero values begin. Consider truncating the data set to include only data before or after the gap, or split the data and analyze both sets separately.

For another example of using WWZ in VStar, see the following which includes the R Dor example from Foster (2010) mentioned at the start of this section.

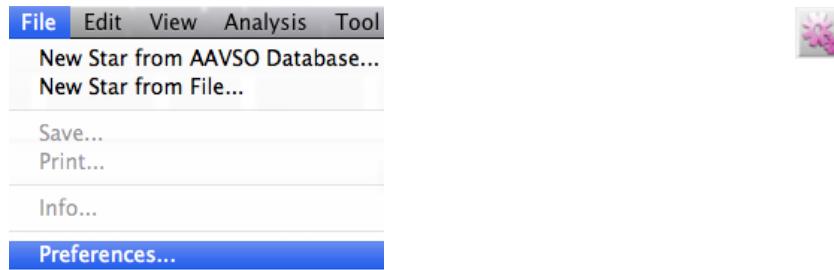
<http://dbenn.wordpress.com/2011/08/30/weighted-wavelet-z-transform-wwz-in-vstar>

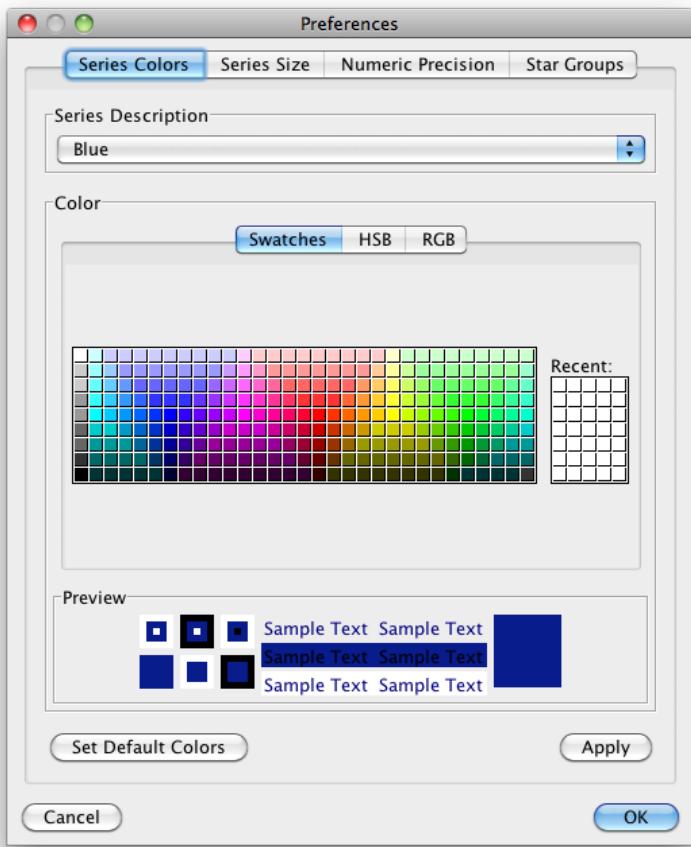
Preferences

VStar allows some aspects of its functionality to be customized; currently these are:

- The color and size of observations in a series as they appear on a plot.
- The precision of numbers, i.e. the number of decimal places.
- The objects that appear in the Load from AAVSO International Database (AID) dialog.

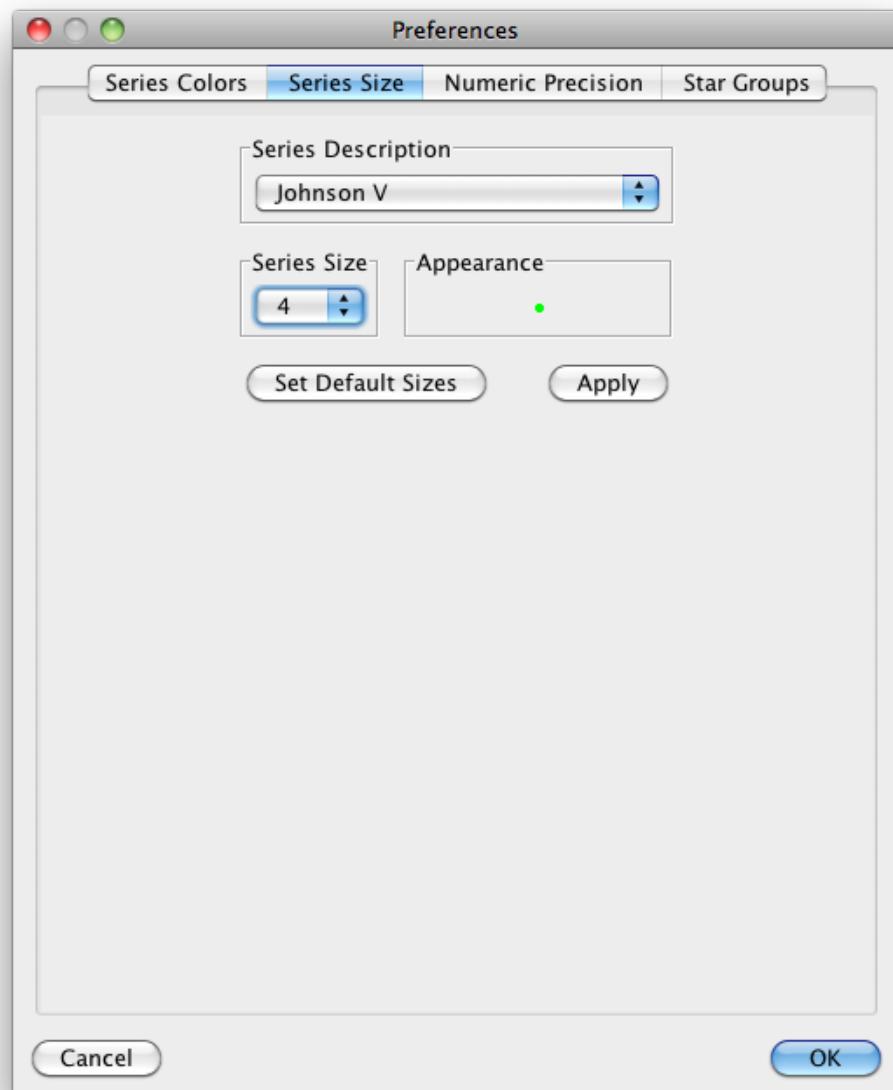
To open the Preferences dialog, select the `File → Preferences...` menu item or the corresponding toolbar button.



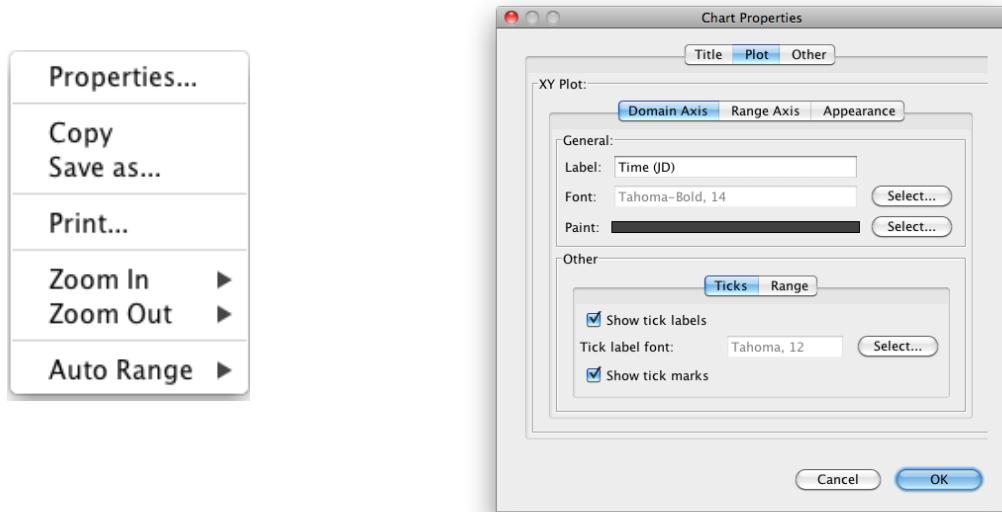


The first tabbed pane shown above controls the color of a series observation as it appears on a plot. The series is selected from a list and the color selected. Clicking the Apply or OK buttons confirms the color change. Like all preference settings, this change persists across VStar invocations. A change can be cancelled via the Cancel button. The Set Default Colors button resets all series colors back to their “factory settings.” This is particularly important for series colors since AAVSO defines a set of standard colors for many series. This does not include synthetic series such as Filtered, Model, Residuals, Means, Excluded.

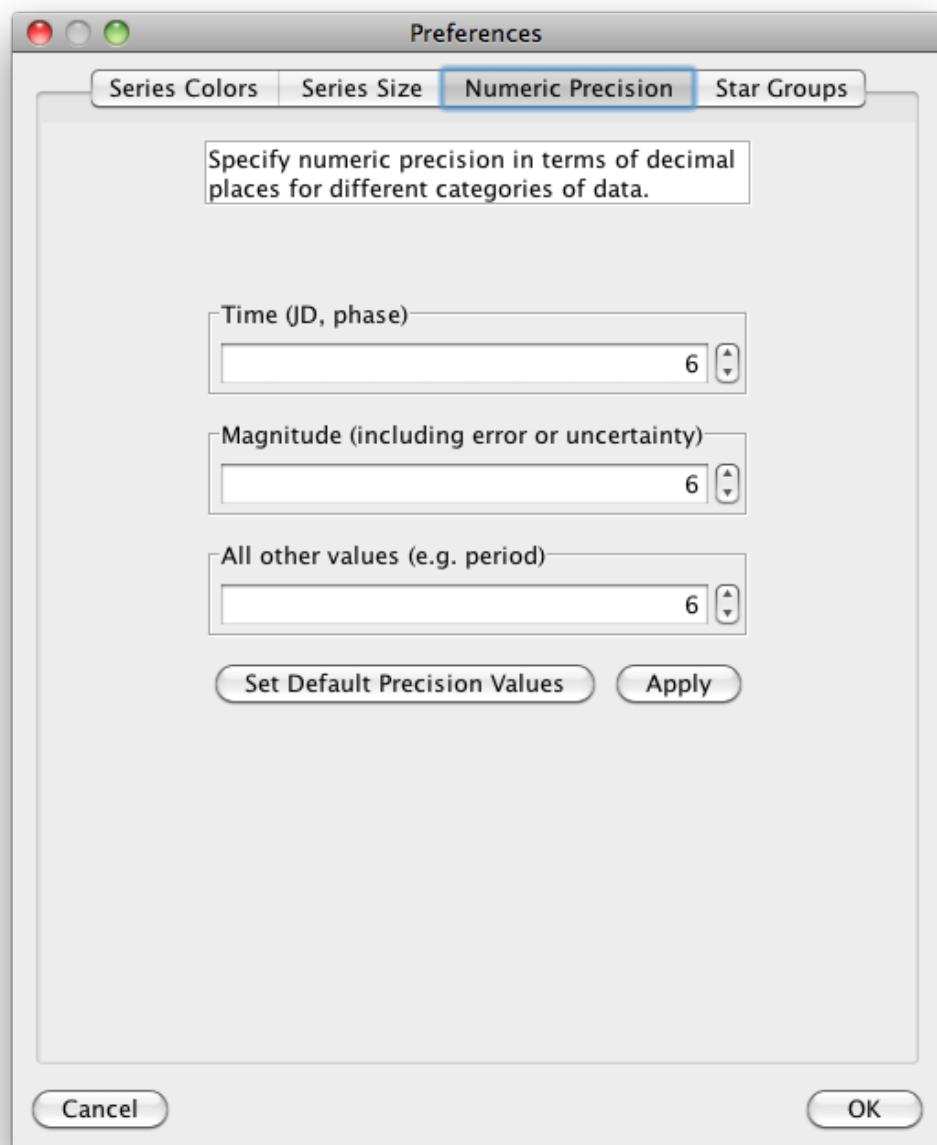
The Series Size tab controls the size of a series observation as it appears on a plot. The series is selected from a list and the size selected from another list. Clicking the Apply or OK buttons confirms the size change. The Set Default Sizes button resets all series sizes back to their “factory settings.”



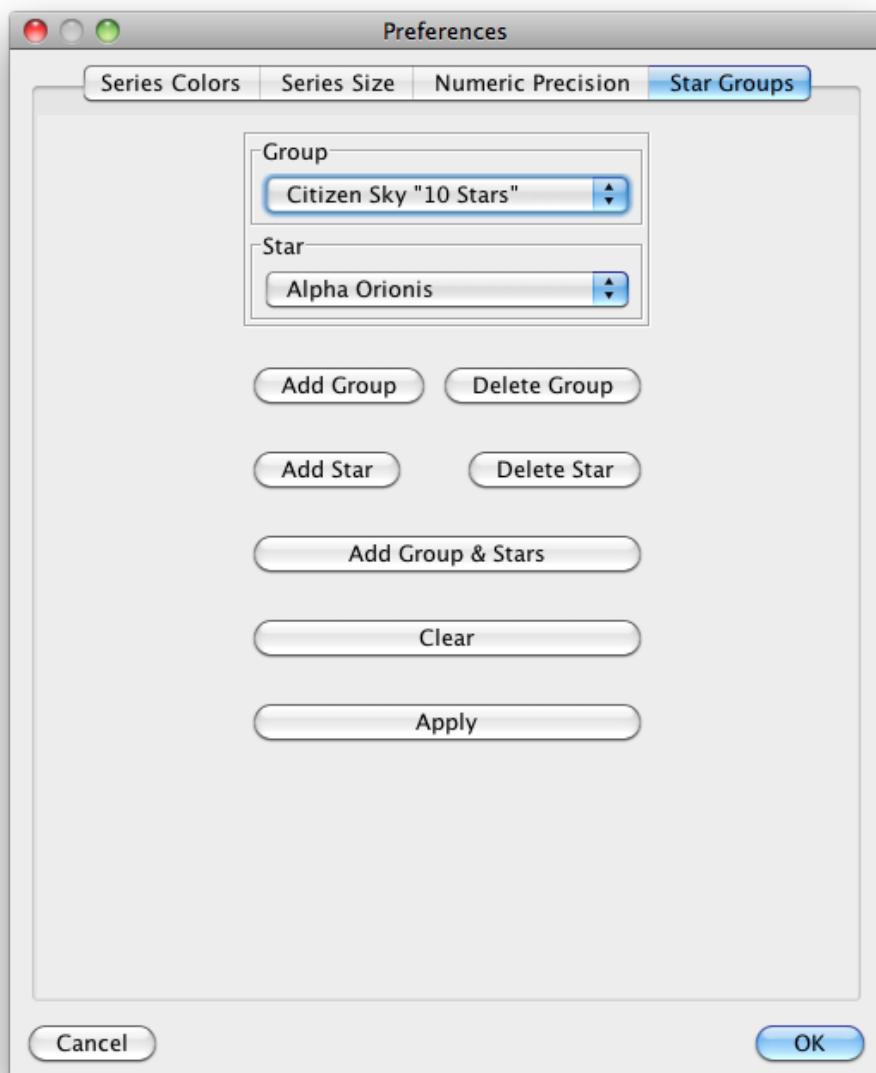
On a related note, while not part of VStar Preferences, some other aspects of a plot can be configured (for the current VStar session only) by selecting `Properties...` from the context menu (e.g. by right-clicking on Windows, ctrl-click on Mac OS X):



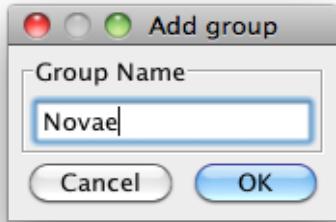
The Numeric Precision tab permits the number of decimal places to be customized for various categories of numeric values that appear as in VStar's user interface, e.g. parameter dialog fields. The Set Default Precision Values button resets all precision values back to their “factory settings.”



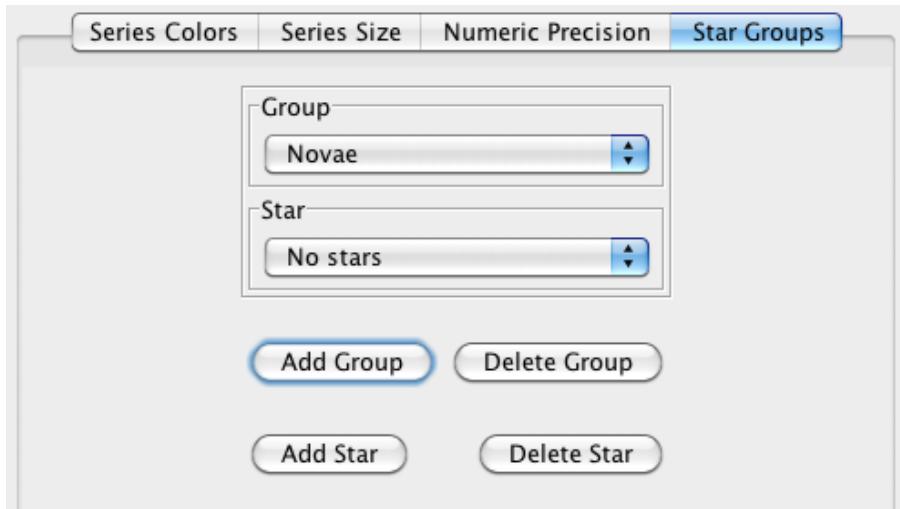
The Star Groups tab permits new groups of stars to be added to the dialog that opens when File → New Star from AAVSO Database... is selected.



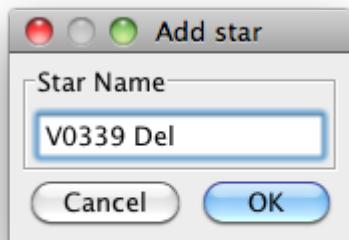
Suppose we want to add a new group for novae. To do so, click the Add Group button and enter the text “Novae” into the dialog, as shown.



The Preferences dialog now shows an empty Novae group.



The next step is to click the Add Star button and enter the text “V0339 Del” into the dialog, as shown.



VStar will show the busy mouse cursor as it goes away to the Variable Star Index (VSX) to ask if this object exists. If so, it will be added to the group. The Group and Stars sections of the pane

should now look like this:

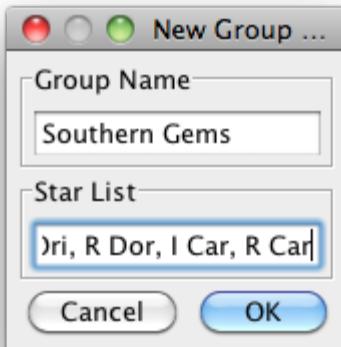


Click the **Apply** or **OK** button to save the change. Doing so will make the change available immediately and across VStar runs.

The **Delete Star** button deletes a star from the currently selected group.

The **Delete Group** button deletes the currently selected group.

Now suppose we want to add a new group with multiple stars all at once. To do so, click the **Add Group & Stars** button and populate the dialog that opens as shown:



See the *References and Further Reading* section for details of “Southern Gems”. The **Star List** text to be entered is:

X Sgr, W Sgr, eta Aql, kap Pav, zet Phe, bet Dor, V Pup, alf Ori, R Dor, l Car, R Car

After each object has been checked (the busy mouse cursor will stop), click the **Apply** button to save the new group and its stars. Doing so will make the change available immediately and across VStar runs.

The New Star from AAVSO Database dialog will now have Novae and Southern Gems groups from which objects can be selected for loading.

The Clear button will set the groups back to “factory settings” (the Citizen Sky “10 Stars” group remains). As with all the other operations above, clicking the Apply or OK button will make the change available beyond the current run of VStar.

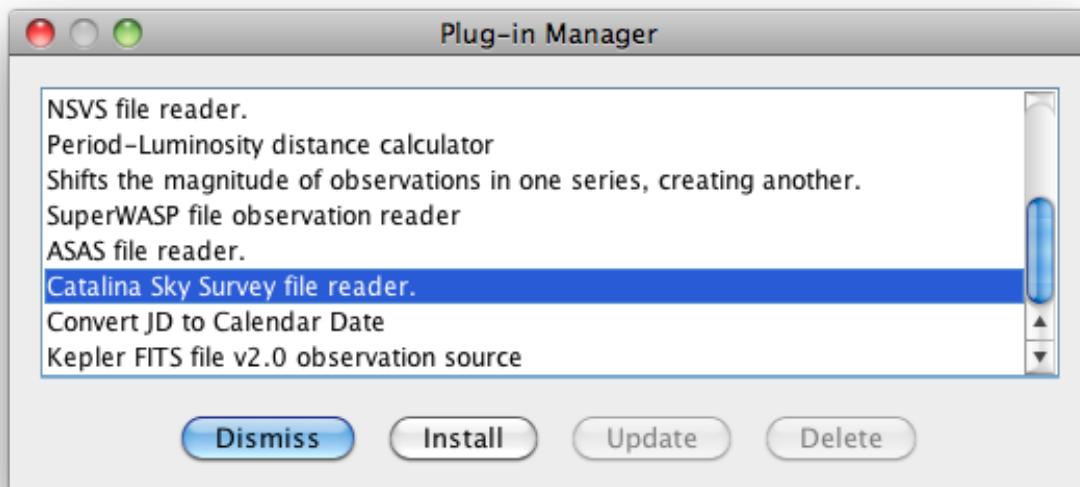
Plug-ins

VStar's capabilities can be extended in various ways by installing plug-ins, e.g.

- Loading observations from sources such as Kepler and ASAS.
- Filtering the loaded dataset in a custom way.
- Creating models, e.g. Loess fit.
- Performing period analysis using an algorithm other than the in-built ones.

See <http://www.aavso.org/vstar-plugin-library> for details of how to install plug-ins manually.

Alternatively, use VStar's plug-in manager, which is accessible from the Tool menu. After selecting the `Plugin-Manager...` menu item, VStar will ask for an AAVSO web user name and password. The plug-in manager dialog will then open:



Scrolling through the list of plug-in descriptions and selecting each one will indicate whether a plug-in can be installed or updated. If neither the `Install` nor the `Update` button is enabled, then the plug-in is up-to-date on your local machine with respect to the AAVSO plug-in archive. A plug-in can be deleted from your system via the `Delete` button.

Some plug-ins require AAVSO membership status to run (e.g. APASS, BSM), so these will only be listed if the user is a member.

After installing, updating, or deleting a plug-in, dismiss the dialog, then close and restart VStar to see the changes.

The menu in which a plug-in will appear in VStar depends upon its type. Observation source plug-ins appear in the File menu. Custom filter plug-ins appear in the View menu's Custom Filters sub-menu. Model creation and period analysis plug-ins appear in the Analysis menu. Observation and general tools appear in the Tool menu.

Follow these links to learn about how to develop plug-ins for VStar:

- https://sourceforge.net/p/vstar/code/HEAD/tree/trunk/plugin/doc/vstar_plugin_dev.pdf
- https://sourceforge.net/p/vstar/code/HEAD/tree/trunk/doc/vstar_docs.zip

The first is a document titled “Developing VStar Plug-ins.” The second is a zip file containing the latest VStar Java API in standard “javadoc” form.

Scripting VStar

VStar permits some of its operations to be automated by writing a script:

- Loading observations from a file.
- Loading observations from the AAVSO International Database (AID).
- Switching between raw and phase plot mode.
- Creating a phase plot given a period and epoch.
- Saving a phase plot or light curve as a PNG image file.
- Saving a light curve to a data file.

The VStar API documentation should be consulted for full function details:

http://sourceforge.net/p/vstar/code/HEAD/tree/trunk/doc/vstar_docs.zip

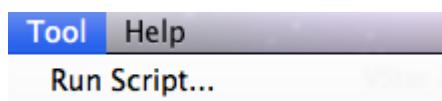
It is intended that more functionality will be made available via the scripting API (Application Programming Interface) in future. The following gives a summary of the available underlying API functions.

Method Summary

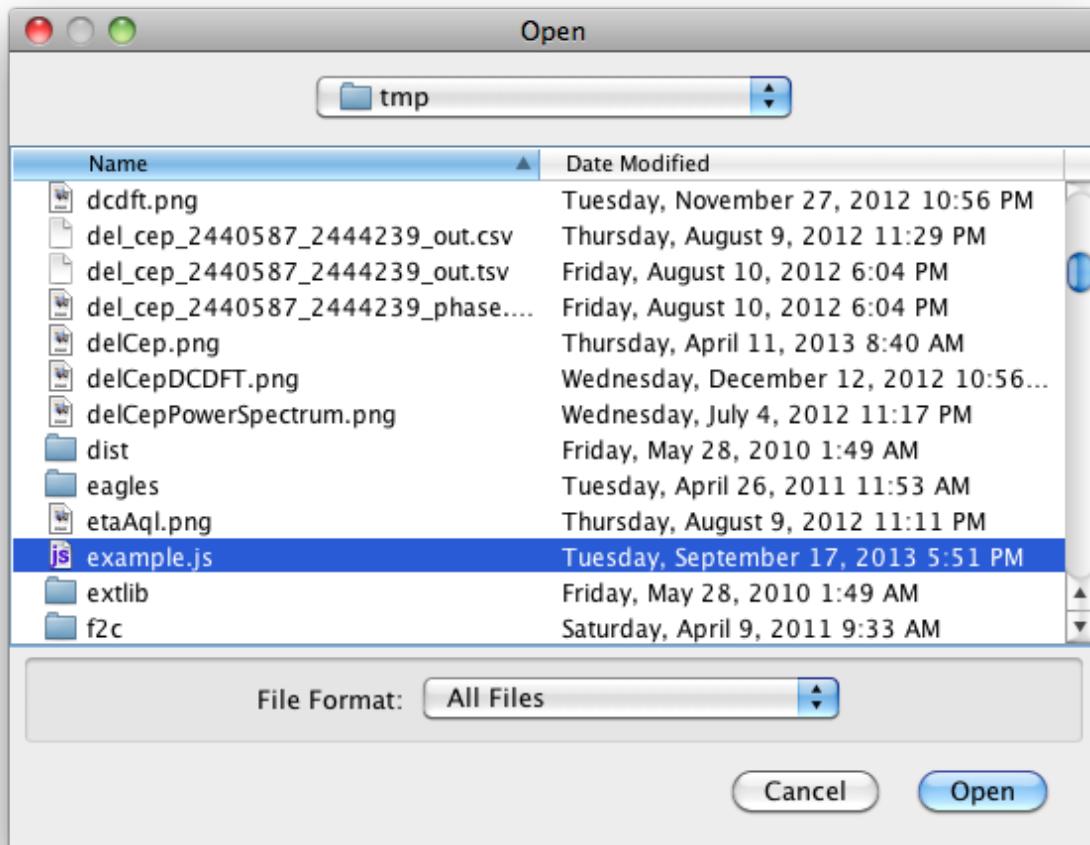
	<code>void additiveLoadFromFile(java.lang.String path)</code> Load a dataset from the specified path, adding it to the existing dataset.
	<code>void exit()</code> Exit VStar.
<code>static VStarScriptingAPI</code>	<code>getInstance()</code> Return Singleton.
	<code>void lightCurveMode()</code> Switch to phase raw (light curve) mode.
	<code>void loadFromAID(java.lang.String name, double minJD, double maxJD)</code> Load a dataset from the AAVSO international database.
	<code>void loadFromFile(java.lang.String path)</code> Load a dataset from the specified path.
	<code>void pause(long millis)</code> Pause for the specified number of milliseconds.
	<code>void phasePlot(double period, double epoch)</code> Create a phase plot given period and epoch.
	<code>void phasePlotMode()</code> Switch to phase plot mode.
	<code>void saveLightCurve(java.lang.String path, int width, int height)</code> Save the light curve to a PNG image file.
	<code>void savePhasePlot(java.lang.String path, int width, int height)</code> Save the phase plot to a PNG image file.
	<code>void saveRawData(java.lang.String path, java.lang.String delimiter)</code> Save the raw dataset (light curve observation list) to a file of rows of values separated by the specified delimiter.
	<code>void showBands()</code> Shows the band names in the current dataset.

Currently, only the JavaScript language is supported but others may be permitted in future, such as Python or Ruby.

To run a script in VStar, select Tool → Run Script...



Doing so opens a file chooser requesting the script file to be executed.



The file chooser shows `example.js` being selected. Clicking the Open button will cause the script to be executed. The example script will load data for eta Aquilae in the JD range 2455821.5...2456552.5, create a phase plot, and save it as a PNG file. It will repeat this for beta Lyrae.

To try this, create a file called `example.js` (or whatever name you choose) with a text editor (e.g. Notepad, vi), paste in the following code, and save the file.

```
with (vstar) {
    root = "/Users/david/tmp/"

    objs = ["eta Aql", "bet Lyr"]
    periods = [7.176641, 12.94061713]
    epochs = [2436084.656, 2455434.8702]
```

```

for (i=0;i<objs.length;i++) {
    obj = objs[i]

    loadFromAID(obj, 2455821.5, 2456552.5)

    phasePlot(periods[i], epochs[i])

    path = root + obj.replace(" ", "_") + "phase_plot.png"
    savePhasePlot(path, 600, 400)
}
}

```

The root path should be replaced with a path of your choosing:

```
root = "/Users/david/tmp/"
```

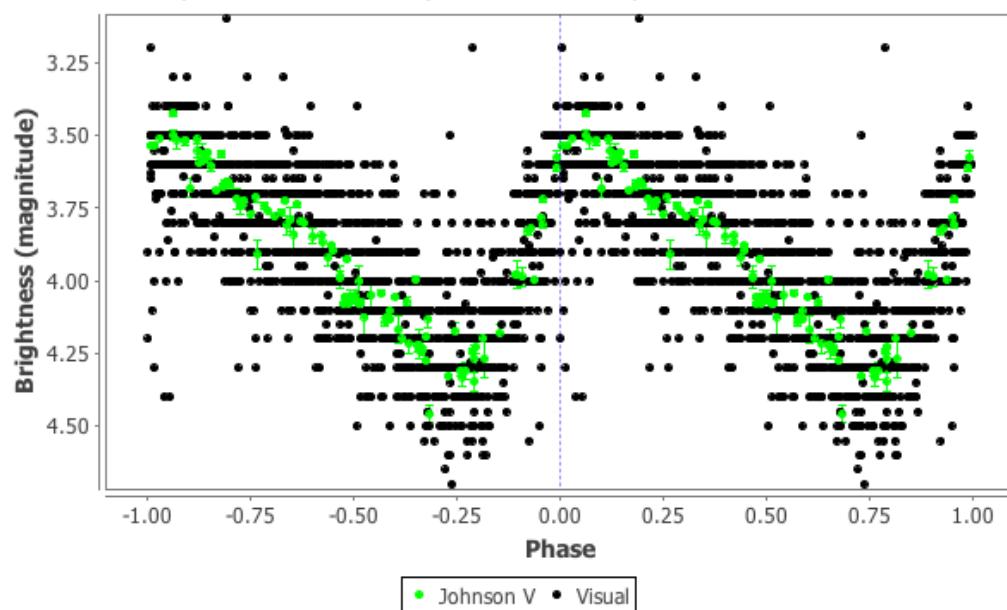
Once the script has completed executing, you will see a phase plot for beta Lyrae appear in VStar and the following files will be in the location corresponding to the `root` directory:

```
bet_Lyr_phase_plot.png
eta_Aql_phase_plot.png
```

Opening the second of these will reveal a phase plot for eta Aquilae, as follows:

Phase Plot for eta Aql

Sep 17, 2013 (database), period: 7.176641, epoch: 2436084.656000



Variable star names used in VSX

(contributed by Mike Simonsen)

Introduction

This section details the names, prefixes and designations that are used specifically for naming variable stars. The one major exception is the Greek letter names used by Bayer. There are many variables whose names are derived from the sky survey they were cataloged in, such as Guide Star Catalog (GSC), USNO CCD Astrograph Catalog (UCAC), Bonner Durchmusterung (BD), etc. These are not included here because many of these stars will eventually receive a VSX designation or are also known by other names.

This list is not exhaustive, but it includes most of the prefixes and naming conventions typically encountered in searching VSX or using VSTAR.

Greek Letter Names (Bayer Designations)

A Bayer designation is a stellar designation in which a specific star is identified by a Greek letter, followed by the genitive form of its parent constellation's Latin name. Most of the brighter stars were assigned their first systematic names by the German astronomer Johann Bayer in 1603, in his star atlas *Uranometria*. Bayer assigned a lower-case Greek letter, such as alpha (α), beta (β), gamma (γ), etc., to each star he catalogued, combined with the Latin name of the star's parent constellation in genitive (possessive) form. For example, Aldebaran is designated α Tauri (alpha Tauri), which means "alpha of the Bull".

A single constellation may contain fifty or more stars, but the Greek alphabet has only twenty-four letters. When these ran out, Bayer began using lower-case Latin letters: hence s Carinae and d Centauri. Within constellations having an extremely large number of stars, Bayer eventually advanced to upper case Latin letters, as in G Scorpii and N Velorum. The last upper-case letter used by Bayer was Q .

As luck would have it, many of these bright stars turned out to be variable stars, so they are often known by their Bayer designations. Examples are alpha Orionis, beta Lyrae, delta Cephei and eta Aquilae. AAVSO and VSX use a specific set of abbreviations for these names. To learn more see- <http://www.aavso.org/greek-letters>

Argelander's System for Naming Variable Stars

Frederick W. A. Argelander invented the traditional system used to name variable stars today. Under this system, variable stars are named using a variation on the Bayer designation format, combining an identifying label with the Latin genitive (possessive) of the name of the constellation in which the star resides. He started with the letter R to avoid confusion with the Bayer designated stars. As each new variable in a constellation was discovered it received the next letter combination in line.

The rules for naming stars in order of their discovery is this:

- Stars with existing Greek letter Bayer designations are not given new designations.

- Otherwise, start with the letter R and go through Z.
- Continue with RR through RZ, and then use SS through SZ, TT through TZ and so on until ZZ.
- Then it reverts to the beginning of the alphabet, AA through AZ, BB through BZ, CC through CZ and so on until reaching QZ, omitting J in both the first and second positions.
- Note that the first letter is never further up the alphabet than the second, that is to say no star can be BA, CA, CB, DA or so on.
- The Latin script is abandoned after 334 combinations of letters and numbers are used, starting with V335, V336, and so on.

Other Naming Conventions

As if that weren't confusing enough, there are now a host of other prefixes and numbers assigned to variable stars and objects. The following is an alphanumerical guide to help you understand what these names mean and where they came from.

1SWASP Jxxxxxx.xx+xxxxxx.x – Variables discovered serendipitously by SuperWASP, a UK extra-solar planet detection program. SuperWASP consists of two robotic observatories (LaPalma and South Africa) that operate continuously all year round. The observatories each consist of eight wide-angle cameras that simultaneously monitor the sky for planetary transit events. <http://www.superwasp.org/>

2MASS Jxxxxxx+xxxxxx – Variables discovered by The Two Micron All-Sky Survey (2MASS), a survey of the whole sky in three infrared wavebands around 2 micrometers.

<http://irsa.ipac.caltech.edu/Missions/2mass.html>

2QZ Jhhmmss.s-ddmmss – Objects discovered by the 2dF QSO Redshift Survey. The aim is to obtain spectra of QSOs out to redshifts so high the visible light emitted by these objects has shifted into the far infrared. The observations are actually of the ultra-violet part of the spectrum that has been redshifted into the visible. As with most QSO surveys, a serendipitous byproduct is the discovery of CVs and other blue stars. A description and awesome pictures of the equipment can be found here: http://www.2dfquasar.org/Spec_Cat/basic.html

Home site: <http://www.2dfquasar.org/index.html>

3C xxx – These are objects from the Third Cambridge (3C) catalog (Edge et al. 1959), based on radio-wavelength observations at 158 MHz. There are 471 3C sources, numbered sequentially by right ascension. All 3C sources are north of -22 declination. The 3C objects of interest to variable star observers are all active galaxies (quasars, BL Lacs, etc.).

Antipin xx – Variable stars discovered by Sergej V. Antipin, a junior researcher working for the General Catalogue of Variable Stars Group.

ASAS hhmmss+ddmm.m – This is the acronym for All Sky Automated Survey, which is an ongoing survey monitoring millions of stars down to magnitude 14. The survey cameras are located at the Las Campanas Observatory in Chile, so it covers the southern sky from the pole to

about +28 degrees declination. <http://www.astrouw.edu.pl/asas/>

ASASSN-XXaa – Supernovae and cataclysmic variables discovered by the All Sky Automated Survey for SuperNovae, a survey running instruments in Hawaii and Chile. Some galactic novae have been discovered and are given the prefix ASASN.

CSS yymmdd: hhmmss+ddmmss – The prefix stands for Catalina Sky Survey followed by the date in yymmdd form, then the coordinates in an unscientific, unsatisfactory format. These are assigned to transient events or cataclysmic variables, mostly. <http://www.lpl.arizona.edu/css/>

CSS_Jxxxxxx.x+xxxxxx – Variables from the Catalina Surveys Periodic Variable Star Catalog, an on-line catalog containing periods, amplitudes, and classifications for ~61,000 periodic variables. <http://arxiv.org/pdf/1405.4290.pdf>

EUVE Jhhmm+ddmm – These are objects detected by NASA's Extreme Ultraviolet Explorer, a satellite dedicated to studying objects in far ultraviolet wavelengths. The first part of the mission was dedicated to an all-sky survey using the imaging instruments that cataloged 801 objects. Phase two involved pointed observations, mainly with the spectroscopic instruments. One of the highlights of the mission was the detection of Quasi Periodic Oscillations (QPOs) in SS Cyg.

FBS hhmm+dd.d – Stands for First Byurakan Survey and the coordinates of the object. The First Byurakan Survey (FBS), also known as the Markarian survey, covers about 17,000 square degrees.

FSVS Jhhmm+ddmm – Discoveries from the Faint Sky Variability Survey, the first deep wide field, multi-color, time-sampled CCD photometry survey. It was specifically aimed at detecting point sources as faint as 25th magnitude in V and I and 24.2 in B. Targets were faint CVs, other interacting binaries, brown dwarfs and low mass stars and Kuiper Belt Objects.

HAT-xxx-xxxxxxx – Stars with exoplanet transits discovered by the Hungarian-made Automated Telescope Network (HATNet) Exoplanet Survey. <http://hatnet.org/>

HadVxxx – This represents variables discovered by Katsumi Haseda. Haseda's most recent discovery was Nova 2002 in Ophiuchus, V2540 Oph.

He-3 xxxx – Variables from Henize, K. G. 1976, "Observations of Southern Emission-Line Stars", Ap.J. Suppl. 30, 491.

HS hhmm+ddmm – The Hamburg Quasar Survey is a wide-angle objective prism survey searching for quasars in the northern sky, avoiding the Milky Way. The limiting magnitude is approximately 17.5B. The taking of the plates was completed in 1997. <http://www.hs.uni-hamburg.de/EN/For/Exg/Sur/hqs/index.html>

HVxxxx – Preliminary designations of variables discovered at Harvard Observatory.

Lanning xx – Discoveries of UV-bright stellar objects by H. H. Lanning from Schmidt plates centered primarily on the galactic plane. In all, seven papers entitled "A finding list of faint UV-bright stars in the galactic plane" were published.

LD xxx – Variables discovered by Lennart Dahlmark, a Swedish retiree living in southern France are given this prefix. Dahlmark has been conducting a photographic search for new variable stars; discovering several hundred to date.

Markarian xxxx – The widely used abbreviation for Markarian objects is Mrk. These are active galaxies from lists published by the Soviet Armenian astrophysicist B.E. Markarian. Markarian looked for galaxies that emit unusually strong UV radiation, which comes from either pervasive star-formation HII regions or from active nuclei. In 1966, Markarian published 'Galaxies With UV Continua'. Around that time, he started the First Byurakan Spectral Sky Survey (FBS), which is now completed. In 1975, Markarian initiated a Second Byurakan Survey (SBS). His collaborators continued the SBS after his death.

MASTER OT Jxxxxxx.xx+xxxxxx.x – Supernovae, cataclysmic variables and other variable objects discovered by the MASTER (Mobile Astronomical System of TElescope Robots) project. <http://observ.pereplet.ru/>

MisVxxxx – The stars are named MisV after MISAO Project Variable stars. The MISAO Project makes use of images taken from all over the world, searching for and tracking astronomically remarkable objects. Few of these stars have light curves, and the type and range of many are still undetermined. The project website url is: <http://www.aerith.net/misao/>

NSV xxxx – These are stars in the Catalog of New and Suspected Variable Stars, produced as a companion to the Moscow General Catalog of Variable Stars (GCVS) by B.V. Kukarkin et al. All stars in the NSV have reported but unconfirmed variability, in particular, lacking complete light curves. Some NSV stars will eventually prove truly variable; others will be spurious. Information about this and the General Catalog of Variable Stars can be found at:

<http://www.sai.msu.su/groups/cluster/gcvs/gcvs/intro.htm>

NSVS Jxxxxxx+xxxxxx – Variable stars discovered as a result of the Northern Sky Variability Survey (NSVS). <http://skydot.lanl.gov/nsvs/nsvs.php>

OGLE Jxxxxxx.xx-xxxxxx.x – The Optical Gravitational Lensing Experiment is a Polish astronomical project that is chiefly concerned with discovering dark matter using the microlensing technique.

As a byproduct of the constant monitoring of hundreds of millions of stars, one of the largest catalogs of variable stars was constructed, and the first exoplanets discovered using the microlensing technique were detected. <http://ogle.astrow.u.edu.pl/>

OGLE-LMC-RCBxx – R CrB stars discovered by OGLE.

OGLE-RRLYR-xxxxx – RR Lyrae stars discovered by OGLE.

PG hhmm+DDd – Palomar Green Survey conducted to search for blue objects covering 10714 square degrees from 266 fields taken on the Palomar 18-inch Schmidt telescope. Limiting magnitudes vary from field to field, ranging from 15.49 to 16.67. The blue objects detected tend to be quasars and cataclysmic variables. The CVs were documented in Green, R. F., et al. 1986, "Cataclysmic Variable Candidates from the Palomar Green Survey", Ap. J. Suppl. 61, 305.

PKS hhmm+ddd – This was an extensive radio survey (Ekers 1969) of the southern sky undertaken at Parkes (PKS), Australia, originally at 408 MHz and later at 1410 MHz and 2650 MHz. These sources are designated by their truncated 1950 position. For example 3C 273 = PKS 1226+023. This is still the most common, and useful, system of naming quasars.

ROSAT is an acronym for the ROentgen SATellite. ROSAT was an X-ray observatory developed

through a cooperative program between Germany, the United States, and the United Kingdom. The satellite was designed and operated by Germany, and was launched by the United States on June 1, 1990. It was turned off on February 12, 1999.

Prefixes for x-ray sources detected by ROSAT include, 1RXS, RXS and RX. The J2000 coordinates for the source are then stated according to the accuracy of the X-ray position and the density of stars in the field.

arcsecond accuracy → RX J012345.6-765432

tenth-arcmin accuracy → RX J012345-7654.6

arcmin accuracy → RX J0123.7-7654

Distressingly, these can all refer to a single object!

Rosino xxx or N xx – Variables discovered by Italian astronomer L. Rosino, primarily in clusters and galaxies through photographic surveys.

ROTSE1 thru 3 Jhhmmss.ss+ddmmss.s - The Robotic Optical Transient Search Experiment (ROTSE) is dedicated to the observation and detection of optical transients on time scales of seconds to days. The emphasis is on gamma-ray bursts (GRBs). Objects detected by this survey are designated with positions to 0".1 precision. <http://www.rotse.net/>

S xxxxx – These are preliminary designations of variables discovered at Sonneberg Observatory.

SBS hhmm+dd.d – Indicates objects discovered by the Second Byurakan Sky Survey, plus the coordinates of the object.

SDSSp Jhhmmss.ss+ddmmss.s – These are discoveries from the Sloan Digital Sky Survey. The positions of the objects are given in the names. SDSS- (Sloan Digital Sky Survey), p- (preliminary astrometry), Jhhmmss.ss+ddmmss.s (the equinox J2000 coordinates). In subsequent papers on CVs detected by SDSS (Szkody et al) the p was dropped and the names became simply SDSS Jhhmmss.ss+ddmmss.s.

SN YYYYaa – IAU named supernovae.

SVS xxx – Soviet Variable Stars, indicates preliminary designations of Soviet-discovered variables.

TAV hhmm+dd – The Astronomer Magazine, in England, has a program that monitors variable stars and suspected variable stars. TAV stands for The Astronomer Variable, plus the 1950 coordinates.

TASV hhmm+dd – TASV stands for The Astronomer Suspected Variable, plus the 1950 coordinates. The Astronomer Variable star page can be found at this url:
<http://www.theastronomer.org/variables.html>

TKx – TK stands for T.V. Kryachko. The TK numbers of new variables continue a numbering system first introduced in Kryachko and Solovyov (1996). The authors invented this acronym.

VSX Jhhmmss.s+ddmmss – This is the system used by the International Variable Star Index (VSX), maintained by the AAVSO. These J2000 coordinate designations are applied to newly

cataloged objects in VSX. <http://www.aavso.org/vsx/>

XTE Jhhmm+dd – These are objects detected by the Rossi X-Ray Timing Explorer Mission. The primary objective of the mission is the study of stellar and galactic systems containing compact objects. These systems include white dwarfs, neutron stars, and possibly black holes.

<http://www.nasa.gov/centers/goddard/missions/rxte.html>

Glossary

Term or Acronym	Meaning
AAVSO International Database	A collection of tens of millions of variable star observations spanning over a century. Maintained by the AAVSO, it is the largest digital database of variable stars in the world [http://www.aavso.org/aavso-international-database].
AAVSO Unique ID	A unique reference “name” for a star using the format 000-XXX-000, where the 0's are numbers (0-9) and the X's are letters.
AIC	See Akaike Information Criteria.
AID	See AAVSO International Database.
Akaike Information Criteria	A measure of how well a statistical model fits a given set of data, balancing how complicated the model is. Alternate models can be compared in this way. See also BIC.
Analysis of Variance	A statistical test that checks whether or not variations in the data are simply noise or are statistically significant.
ANOVA	See Analysis of Variance.
APASS	The AAVSO Photometric All-Sky Survey; an ongoing project to use automated telescopes to create a photometric catalog of all stars between approximately 10th and 17th magnitude using five filters: Johnson B and V, and Sloan g', r', i'. [http://www.aavso.org/apass].
ASAS	All Sky Automated Survey; an ongoing project to use photometric automated telescopes to discover and observe all stars with variability (including exoplanet systems and variable stars) brighter than 14th magnitude [http://www.astrouw.edu.pl/asas/?page=main].
AUID	See AAVSO Unique ID.
Band	A narrow region of the electromagnetic spectrum in which the magnitude is measured. A number of standard bands are frequently used, including the visual (V) band (centered on 551 Angstroms). See http://en.wikipedia.org/wiki/Photometric_system and http://spiff.rit.edu/classes/phys440/lectures/filters/filters.html .

Bayesian Information Criteria	Like the AIC, a statistical method for selecting which of multiple models best fits the data; it weighs the relative complexity of each model more harshly than AIC.
BIC	See Bayesian Information Criteria.
Binned	Data that is grouped by a certain range of values, for example grouping variable star observations in increments of 5 Julian days. See http://pic.dhe.ibm.com/infocenter/spssstat/v20r0m0/index.jsp?topic=%2Fcom.ibm.spss.statistics.help%2Fidh_webhelp_scatter_options_palette.htm .
BSM Epoch Photometry Data	Data from the Bright Star Monitor AAVSO database; currently there are 14 million photometric observations of approximately one million stars. See http://www.aavso.org/bright-star-monitor-epoch-photometry-database .
Catalina Sky Survey	An automated telescope survey of Near Earth Objects [http://www.lpl.arizona.edu/css/]. Transient phenomenon of many types (such as variable star outbursts) are also surveyed [http://crts.caltech.edu/].
CLEANest	A method of refining multiple periods in variable star data simultaneously. In VStar, this can be applied to a period analysis (DCDFT) result.
Comparison Star	A star of known, constant magnitude that a nearly variable star is compared to in order to estimate the variable's apparent brightness.
Confidence Level	The range of values for a mathematical variable for which there is high confidence (usually 95%) that the actual value lies within that range.
Cousins	A standard set of photometric filters. See http://www.company7.com/library/optec/filter_monograph.pdf .
Discrepant Data	A data point that appears to vary significantly (well outside of normal error bars) from the mean and/or other observations taken at the same time. The AAVSO has a tool called Zapper which can be used to flag discrepant data for further investigation [http://www.aavso.org/zapper].
Epoch	In a phase plot, the initial time selected to begin the cycle (e.g. the time of the first observation or the time of minimum).
Error Bar	On a graph, lines extending above and below and/or to the right and left of a data point that signify uncertainty in the

	measurement (usually to 95% confidence level). In VStar only error bars in magnitude are included, since uncertainty in times. not recorded in AAVSO data.
Fainter Than	A null observation of a variable star that was too faint for the instrument/observer. Since the star was not actually seen/imaged, it can only be said that at that time it was “fainter than” some limiting magnitude.
FITS	Flexible Image Transport System is a digital file format commonly used in astronomy and other sciences for the storage, transmission and processing of images.
Fourier Analysis	A power spectrum (see below) that is based on sinusoids.
Harmonics	The fundamental period divided by an integer (i.e. half the period, one-third the period).
Heliocentric Julian Day	A recalculation of the Julian date to take into account the earth’s distance from the sun; i.e. the Julian Date if the star were measured from the center of the sun. See http://www.physics.sfasu.edu/astro/javascript/hjd.html .
HJD	See Heliocentric Julian Day.
JD	See Julian Day.
Johnson Band	A standard set of photometric filters. See http://www.company7.com/library/optec/filter_monograph.pdf .
Julian Day	A continuous account of the days (and fractions of a day) from noon Universal Time on January 1, 4713 BCE. See http://aa.usno.navy.mil/data/docs/JulianDate.php .
Kepler	An exoplanet-seeking space telescope that records variations in brightness for approximately 150,000 stars in a single field of view near the Cygnus/Lyra border. In addition to discovering exoplanets (via the transit method) it also captures observations of numerous variable stars.
Light Curve	A plot of variations in magnitude over time.
Mean Series	Calculating the mean of observations for equal sized bins of Julian Dates or phase “steps” to construct a mean light curve.
Observer Code	A unique 3 or 4 letter code assigned to each observer by the AAVSO. Individual observations are tagged by the observer code in the database.

Period	The time between two repetitions of a cycle (often peak-to-peak).
Phase Plot	A plot in which the periodic data is “folded” such that, if the correct period is selected, each individual instance of the cycle aligns with the others.
Polynomial model (polynomial fit)	Representing data as a sum of terms of the form $\beta_i t^i$ where $i=0,1,2,\dots$. See Foster (2010).
Power Spectrum	Also known as a periodogram; A plot of frequency versus power that is used when a trial frequency is presumed in order to find periodicity in the data. Here power is a measure of the statistical significance of the fit of the trial frequency to the actual data.
Pre-whitening	In order to examine residuals in the data set, the strongest signal (highest peak) is removed. This permits additional periods to be searched for.
Regular Expressions	In computer programming a regular expression is a special text string used to describe a search pattern. See http://www.regular-expressions.info/ .
Residuals	The difference between the raw data and the model of the data. Residuals can sometimes contain valuable information, such as the existence of a second periodicity.
RMS	See root mean square.
Root Mean Square	In statistics this is the square root of the mean of a sum of squared values (i.e. the square root of the mean of $x_1^2 + x_2^2 + x_3^2 \dots$).
Standard Error	A measure of how precise an average value of data is expected to be.
SuperWASP	A UK-based exoplanet detection consortium. See http://www.superwasp.org/ .
Variable Star Index	An online “clearinghouse” for timely information on variable stars, including suspected suspected. See http://www.aavso.org/vsx/ .
VSX	See Variable Star Index.
Weighted Wavelet Z-Transform	An algorithm designed for analyzing variable star data, especially cases where there are period changes and/or transient phenomena (e.g. mode switching). See Foster (1996)

	and Templeton (2004).
WWZ	See Weighted Wavelet Z-Transform.

References and Useful Links

General Information on VStar and Variable Star Observing

AAVSO VStar Overview and Download page: The starting point for using VStar; includes instructions for downloading the program (including helpful hints for dealing with firewalls and other vagaries) as well as links for further instructions and tutorials.

[<http://www.aavso.org/vstar-overview>]

VStar SourceForge Project page: A source for downloading VStar for those who have difficulty downloading directly from the AAVSO website (for example, when firewalls are in use). Also, the downloadable distribution archive contains additional documents and resources that are not easy to make available via the WebStart™ download method.

[<https://sourceforge.net/projects/vstar>]

AAVSO Plug-in Library page: Instructions for installing various plug-ins for VStar as well as links to those currently available.

[<http://www.aavso.org/vstar-plugin-library>]

AAVSO VStar Forum: The home page for all ongoing and past discussion of VStar, its usage, and problems users have encountered with it.

[<http://www.aavso.org/forums/about-aavso/vstar>]

AAVSO Download Format: A detailed list and explanation of the format and terms used in AAVSO data sets.

[<http://www.aavso.org/format-data-file>]

WebObs: A portal on the AAVSO website for uploading new observations of variable stars and searching for previously uploaded observations. Note: solar observations are not included in this data set, and are found at <http://www.aavso.org/solar>.

[<http://www.aavso.org/webobs>]

Southern Gems: A part of the Citizen Sky Project (<http://www.aavso.org/citizensky>); ten naked-eye southern hemisphere variable stars. [<http://www.citizensky.org/teams/southern-gems/important-southern-gems-resources>]

Useful Links for ANOVA

StatSoft, Introduction to ANOVA/MANOVA: An overview of ANOVA techniques from the publisher of a statistics textbook. The page links to a great deal of other information on statistical analysis. [<http://www.statsoft.com/textbook/anova-manova/?button=1>]

Wikipedia - One-way ANOVA (analysis of variance): Standard Wikipedia entry summarizing the basics of the ANOVA technique. [http://en.wikipedia.org/wiki/One-way_analysis_of_variance]

Useful Links for Period and Time Series Analysis (TS and WWZ)

AAVSO Software Directory: A collection of software that was developed with variable star observations and data analysis in mind. [<http://www.aavso.org/software-directory>]

TS (Fortran): A direct link for downloading the time series statistical program.
[<http://www.aavso.org/sites/default/files/software/ts.tar.gz>]

WWZ (Fortran): A direct link for downloading a Fortran version of WWZ.
[<http://www.aavso.org/sites/default/files/software/wwz.tar.gz>]

Useful Links for Polynomial Fit

Lutus, P, Interactive Polynomial Regression Data Fit: An interactive example using JavaScript.
[<http://www.arachnoid.com/polysolve>]

Weisstein, E.W. Root-Mean-Square: Concise explanation from Wolfram MathWorld with useful links. [<http://mathworld.wolfram.com/Root-Mean-Square.html>]

Wikipedia - Polynomial Regression: Standard Wikipedia entry, including history and references.
[http://en.wikipedia.org/wiki/Polynomial_regression]

Wikipedia - Akaike Information Criteria (AIC): Standard Wikipedia entry, including history and references. [http://en.wikipedia.org/wiki/Akaike_information_criterion]

Useful Links for Regular Expressions (for Observation List Pattern Search)

Regular Expressions: A detailed website on the use of regular expressions in data searches, including numerous examples and tutorials. [<http://www.regular-expressions.info>]

Regular Expressions Quick Start: A quick tutorial on the basics of using regular expressions in data searches. [<http://www.regular-expressions.info/quickstart.html>]

Wikipedia - Regular Expressions: Standard Wikipedia entry, including history and references. [http://en.wikipedia.org/wiki/Regular_expression]

Regular Expressions - User Guide: An overview of regular expressions that explains the basics very thoroughly. [<http://www.zytrax.com/tech/web/regex.htm>]

Further Reading

Benn, D., 2012, *Algorithms + Observations = VStar*, JAAVSO, vol 40, p 852. An overview of the history and possible future of VStar.

[<http://www.aavso.org/sites/default/files/jaavso/v40n2/852.pdf>]

Benn, D. *VStar Blog*. An ongoing blog (with archive) that discusses the usage of VStar as well as updates to the program. [<http://dbenn.wordpress.com/category/astronomy-science/vstar>]

Ferraz-Mello, S. 1981, *Estimation of Periods From Unequally Spaced Observations*, Astron. J., vol 86, p 619. A technical article on the use of Fourier Transforms as applied to time series data with gaps in observations. [<http://adsabs.harvard.edu/full/1981AJ.....86..619F>]

Foster, G. 1995, *The CLEANest Fourier Spectrum*, Astron. J., vol 109, no 4, p 1889. A technical article describing the development and application of the CLEANest algorithm.

<http://adsabs.harvard.edu/full/1995AJ....109.1889F>

Foster, G. 1996, *Wavelets for Period Analysis of Unevenly Sampled Time Series*, Astron. J., vol 112, p 1709. A technical article describing the development and application of the Weighted Wavelet Z-transform (WWZ). [<http://adsabs.harvard.edu/full/1996AJ....112.1709F>]

Foster, G., *Analyzing Light Curves: A Practical Guide*, 2010, Lulu. Considered the standard introduction to variable star light curve analysis; details the application of various statistical approaches and algorithms to variable star data.

<http://www.lulu.com/shop/grant-foster/analyzing-light-curves-a-practical-guide/paperback/product-11037112.html>

Foster, G., *Understanding Statistics: Basic Theory and Practice*, 2013, Lulu. An introduction to statistical principles and techniques utilizing basic, everyday examples.

[<http://www.lulu.com/shop/grant-foster/understanding-statistics-basic-theory-and-practice/paperback/product-20680689.html>]

Templeton, M., 2004, *Time-Series Analysis of Variable Star Data*, JAAVSO vol 32, p 41. An introduction to the application of Time-series Analysis to variable star data; based on an AAVSO meeting workshop. [<http://www.aavso.org/files/jaavso/v32n1/41.pdf>]

Variable Star Astronomy Chapter 12: Variable Stars and Phase Diagrams: A very basic introduction to the statistics of variable star data written for the *Variable Star Astronomy*

(formerly *Hands-on Astrophysics*) curriculum.

<http://www.aavso.org/sites/default/files/Chapter12.pdf>

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Like most modern software, VStar is dependent upon libraries of code written by others for some of its functionality, most notably for plotting, database access, and some statistical operations.

See <https://sourceforge.net/p/vstar/code//trunk/extlib/ReadMe.txt> for the most up to date information about the libraries used by VStar.

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Revision History

Date	Primary Authors	Comments
Friday, 27 September 2013	David Benn, Kristine Larsen	Initial document release; for VStar 2.15.3.
Sunday, 29 September 2013	David Benn	Resized context diagram. Added AAVSO logo to title page. Removed 2.15.3 from title page.
Saturday, 15 March 2014	David Benn	Added note about Excel model equation format.
Friday, 9 May 2014	David Benn	Added Plugin-Manager material to Plug-ins section.
Sunday, 22 June 2014	David Benn, Mike Simonsen	Corrections and improvements from VStar CHOICE course, May 2014 course participants (Paul York and Brad Walter). Star naming section by Mike Simonsen.