

# Dark matter halo concentrations with a bayesian approach

Christian Poveda<sup>1</sup> & Jaime E. Forero-Romero<sup>1</sup>

<sup>1</sup>*Departamento de Física, Universidad de los Andes, Cra. 1 No. 18A-10, Edificio Ip, Bogotá, Colombia*

26 May 2014

## ABSTRACT

asd

**Key words:** methods: numerical, galaxies: haloes, cosmology: theory, dark matter

## 1 INTRODUCTION

(Navarro et al. 1997)

## 2 BASIC PROPERTIES OF THE NFW DENSITY PROFILE

The Navarro-Frenk-White density profile can be written as

$$\rho(r) = \frac{\rho_c \delta_c}{r/r_s (1 + r/r_s)^2}, \quad (1)$$

where  $\rho_c \equiv 3H^2/8\pi G$  is the Universe critical density,  $\delta_c$  is the halo dimensionless characteristic density and  $r_s$  is known as the scale radius, the radius that marks the transition between the two power law behaviour in the  $\rho \propto r^{-1}$  for  $r < r_s$  and  $\rho \propto r^{-2}$  for  $r > r_s$ .

We define the virial radius of a halo,  $r_v$ , as the boundary of the spherical volume that encloses an average density of  $\Delta_h$  times the average density of the Universe. The corresponding mass  $M_v$ , the virial mass, can be thus expressed as  $M_v = \frac{4\pi}{3} \bar{\rho} \Delta_h r_v^3$ .

The total mass enclosed within a radius  $r$  can be computed to be:

$$M(< r) = 4\pi \rho_c \delta_c r_s^3 \left[ \ln \left( \frac{r_s + r}{r} \right) - \frac{r}{r_s + r} \right]. \quad (2)$$

We can now define the concentration of the halo as  $c = r_v/r_s$ , the dimensionless variable  $x \equiv r/r_v$  and  $m \equiv M(< r)/M_v$ , which allows us to express the total enclosed mass within a dimensionless radius  $x$  as:

$$m(< x) = \frac{1}{A} \left[ \ln(1 + xc) - \left( \frac{xc}{xc + 1} \right) \right], \quad (3)$$

where

$$A = \left[ \ln(1 + c) - \left( \frac{c}{c + 1} \right) \right], \quad (4)$$

meaning that the concentration is the only free parameter to determine the density profile of the halo. In Figure

**Figure 1.** FIXME: Mass profile of a halo

?? we show the family of normalized integrated mass profiles for different values of the concentration in the range  $1 \leq c \leq 20$ .

## 3 A BAYESIAN APPROACH TO HALO FITTING

We proceed to find the value of the concentration parameter that best describes the simulation data, following the model in Eq. (3). We use the integrated mass profile because it allows us to use the data directly with the simulation without binning the particle positions and estimating a density.

We construct the integrated mass profile by ranking the particles by their increasing distance to the center of the halo. Once they are ranked, the total mass at a radius  $r_i$ , increases by  $m_p$ , where  $r_i$  is the position of the  $i$ -th particle and  $m_p$  is the mass of the computational particle. We define the center of the halo to be at the position of the particle with the lowest gravitational potential. In the process of building the mass profile we discard the particle at the center.

We stop the construction of the integrated mass profile once we arrive at an average density of  $\Delta_h \bar{\rho}$ . This radius marks the virial radius and the virial mass. We divide the total mass enclosed mass  $M_i$  and the radii  $r_i$  by these values to obtain the dimensionless variables  $x_i$  and  $m_i$ .

Using these positions and masses we define the following  $\chi^2$  function

$$\chi^2(c) = \sum_i [\log m_i - \log m(< x_i; c)]^2, \quad (5)$$

where  $m(< x_i; c)$  corresponds to the values in Eq.(3) at  $x = x_i$  and a given value of the concentration parameter  $c$ .

Finally we use a Metropolis-Hastings algorithm to sample the likelihood function distribution defined by  $\mathcal{L}(c) =$

[H]

**Figure 2.** FIXME: Test results

$\exp(-\chi^2(c)/2)$  to find the optimum value of  $c$  and its associated uncertainty  $\sigma_c$ .

## 4 NUMERICAL EXPERIMENTS

### 4.1 Mock Halos

We generate 100 mock halos with concentration values randomly placed in the range  $1 < c < 10$ . Each one of the halos is generated with four different total particle numbers: 20, 200, 2000 and 20000.

The method we use to generate the halos is based on the integrated mass profile. **Our algorithm works as follows...**

### 4.2 Simulation Data

We use data from the MultiDark cosmological volume. This simulation follows the non-linear evolution of a dark matter density field sampled with  $2048^3$  particles over a cubic box of  $1000 h^{-1}\text{Mpc}$  on a side. The data is publicly available, more details about the structure of the database and the simulations can be found in (Riebe et al. 2013).

We select a sample of halos in a cubic sub-volume of  $100 h^{-1}\text{Mpc}$  on a side centered on the most massive halo in the simulation at  $z = 0^1$ . We select first all the halos at  $z = 0$  detected with a Friends-of-Friends (FoF) algorithm with masses in the interval  $10^{11} \leq M_{\text{FoF}}/h^{-1}\text{M}_\odot \leq 10^{15}$ . The FoF algorithm ran with a linking length of 0.17 times the average interparticle distance. This choice translates into an overdensity  $\Delta_h \sim 400 - 700$  that is dependent on the halo concentration (More et al. 2011).

For each selected halo with the previous procedure we select from the database all the particles that belong to it. From the particles we follow the procedure spelled out in Section 3 with  $\Delta_h = 740$  (corresponding to 200 times the critical density) to find the halo concentration. Finally, we store the values obtained for the virial radius, virial mass and concentration.

Finally we compare this results with the values obtained by BDM, since the number of halos obtained by FOF and BDM is different, each BDM halo gets related with the nearest halo in FOF. We compute two concentrations, one where the overdensity limit is  $360\rho_{\text{back}}$  (like in BDMV) an other where the overdensity limit is  $740\rho_{\text{back}}$  (like in BDMW).

## 5 RESULTS

### 5.1 Mock Halos

### 5.2 Simulation Data

## 6 DISCUSSION

### 6.1 Comparison against other methods

### 6.2 Concentration as a function of halo mass

### 6.3 Implication for comparisons against observations

## 7 CONCLUSIONS

## REFERENCES

- More S., Kravtsov A. V., Dalal N., Gottlöber S., 2011, ApJS, 195, 4  
 Navarro J. F., Frenk C. S., White S. D. M., 1997, ApJ, 490, 493  
 Riebe K., Partl A. M., Enke H., Forero-Romero J., Gottlöber S., Klypin A., Lemson G., Prada F., Primack J. R., Steinmetz M., Turchaninov V., 2013, Astronomische Nachrichten, 334, 691

<sup>1</sup> This corresponds to the **miniMDR1** database in the MultiDark webpage