

Stellar Dynamics

1. Basics & Nomenclature

Stellar dynamics is almost entirely collisionless, due to the low number density of stars relative to their radii. It is therefore governed by the *collisionless Boltzmann equation* (sometimes called the *Vlasov equation*) acting under gravity.

In the continuum limit, we can express the *distribution function* of stars in phase space as $f(\vec{x}, \vec{v}, t)$, in units of per length-cubed per unit velocity-cubed. If we define the phase space vector $\vec{w} = \{\vec{x}, \vec{v}\}$ then we can write $f(\vec{w}, t)$. The distribution in phase space can be arbitrarily complicated. It will not be thermalized as in a gas or fluid, or obey any particular equation of state.

The continuum limit will be violated most rapidly by two-body interactions. We find in the exercises that the time for an N -body system to relax due to this effect, the *two-body relaxation time*, is:

$$t_{\text{relax}} \sim \frac{0.1N}{\ln N} t_{\text{cross}} \quad (1)$$

where t_{cross} is the crossing time of the system. Globular clusters have relaxation times short compared to their ages. Galaxies, over most of their extent, have relaxation times high compared to their ages.

In the continuum limit, the system obeys the collisionless Boltzmann equation under just gravity:

$$\frac{\partial f}{\partial t} + \vec{v} \cdot \vec{\nabla} f - \vec{\nabla} \Phi \cdot \frac{\partial f}{\partial \vec{v}} = 0 \quad (2)$$

It can be shown that this equation obeys a special case of *Liouville's Theorem*:

$$\frac{df}{dt} = 0 \quad (3)$$

where in this case the substantive derivative is:

$$\frac{d}{dt} = \frac{\partial}{\partial t} + \sum_{\alpha} \frac{\partial w_{\alpha}}{\partial t} \frac{\partial}{\partial w_{\alpha}} \quad (4)$$

The meaning of Liouville's Theorem is that as a particle travels through phase space, the phase space density of particles around it remains constant.

2. Jeans Equations

The first few moments of the collisionless Boltzmann equation are instructive and can be useful. These equations are known as the *Jeans Equations*.

The zeroth moment integrated over velocity yields the equation of continuity:

$$\frac{\partial n}{\partial t} + \vec{\nabla} \cdot (n \langle \vec{v} \rangle) = 0 \quad (5)$$

where $n(\vec{x})$ is the mean density per unit volume and $\langle \rangle$ indicates a density weighted-mean over all velocities.

The first moment integrated over velocity yields something akin to Euler's equations:

$$n \frac{\partial \langle \vec{v} \rangle}{\partial t} + n \langle \vec{v} \rangle \cdot \vec{\nabla} \langle \vec{v} \rangle = -n \vec{\nabla} \Phi(\vec{x}, t) - \vec{\nabla} \cdot (n \sigma^2) \quad (6)$$

where σ^2 is the tensor second moment of the velocity field, analogous to pressure:

$$\sigma_{ij} = \langle (v_i - \langle v_i \rangle) (v_j - \langle v_j \rangle) \rangle \quad (7)$$

Each successively higher moment of the collisional Boltzman equation involves terms of higher order in this fashion; whereas in a collisional fluid the system would close with an equation of state, in a collisionless system the equations never close.

In steady state, where the density is not changing with time anywhere and therefore the mean velocity is zero everywhere, we find a relation analogous to the hydrostatic equation:

$$\vec{\nabla} \cdot (n \sigma^2) = -n \vec{\nabla} \Phi(\vec{x}, t) \quad (8)$$

Under spherical symmetry in configuration space, this can be rewritten:

$$\frac{\partial (n \sigma_{rr}^2)}{\partial r} + \frac{n}{r} [2\sigma_{rr}^2 - (\sigma_{\theta\theta}^2 + \sigma_{\phi\phi}^2)] = -n \frac{\partial \Phi}{\partial r} \quad (9)$$

and in addition $\sigma_{\theta\theta}^2 = \sigma_{\phi\phi}^2$. In this equation, σ_{rr} and $\sigma_{\theta\theta}$ may both be functions of r .

Although the spherical symmetry in configuration space means that f does not depend on θ or ϕ , it can clearly depend on v_θ and v_ϕ . Thus, σ_{rr}^2 does not have to equal $\sigma_{\theta\theta}^2$. The orbit distribution can be anisotropic, and the degree anisotropy affects the radial distribution $n(r)$. This anisotropy is usually quantified by

$$\beta(r) = 1 - \frac{\sigma_{\theta\theta}^2}{\sigma_{rr}^2} \quad (10)$$

We can then write:

$$M(< r) = \frac{r v_c^2}{G} = -\frac{r \sigma_{rr}^2}{G} \left[\frac{d \ln n}{d \ln r} + \frac{d \ln \sigma_{rr}^2}{d \ln r} + 2\beta \right], \quad (11)$$

where we have defined v_c as the circular velocity of a stable circular orbit. In this equation, the right hand side consists of in-principle observables. Particularly, $n(r)$ is the tracer density; the potential could be set by other unobserved masses. However, in practice β proves hard to constrain for most systems, for which only line-of-sight velocities and projected densities are known.

If we take σ_{rr} to be constant and $\beta = 0$, and we assume that the particles are *self-gravitating* — meaning that they are the source of the potential — then a solution to the equations is given by the singular isothermal sphere:

$$M(< r) = \frac{rv_c^2}{G} \propto \frac{r^{-2}}{2r\sigma_{rr}^2} \quad (12)$$

yielding the relation $v_c^2 = 2\sigma_{rr}^2$, which is a useful order-of-magnitude relationship between the circular velocity and the one-dimensional (for example, line-of-sight) velocity dispersion in gravitating systems.

3. Virial Theorem

The virial equations establish the relationship between kinetic and potential energy in collisionless gravitating systems. They are obtained as a further moment of the Jeans Equation. Specifically, one takes the first moment of position over the analog of Euler’s equation. For a time-independent system, in the center-of-mass frame, we can establish the *tensor virial theorem*:

$$2K_{jk} + W_{jk} = 0 \quad (13)$$

where the internal *kinetic energy tensor* is:

$$K_{jk} = \frac{1}{2} \int d^3\vec{x} \rho \sigma_{jk}^2 \quad (14)$$

(where ρ is the mass density, so for particles of equal mass m , $\rho = nm$). The *potential energy tensor* is:

$$W_{jk} = -\frac{G}{2} \int d^3\vec{x}' d^3\vec{x} \rho(\vec{x}') \rho(\vec{x}) \frac{(x'_j - x_j)(x'_k - x_k)}{|\vec{x}' - \vec{x}|^3} \quad (15)$$

The trace of the tensor virial theorem yields the *scalar virial theorem*:

$$2K + W = 0 \quad (16)$$

where K is the total kinetic energy in the center of mass frame, and W is the total potential energy.

4. Jeans Theorem

Jeans Theorem yields an important tool for modeling equilibrium self-gravitating systems. These systems can be described by the set of orbits of the particles comprising them. The density field resulting from this distribution of orbits generates a potential. To remain in equilibrium, the orbit distribution needs to be stable in that potential. Jeans Theorem yields a way of generating orbit distributions that are self-consistent in this sense that they are in equilibrium.

We showed earlier that f is conserved along orbits in phase space:

$$\frac{df}{dt} = 0 \quad (17)$$

Further, if $\Phi(\vec{x})$ is time-independent, there are six *constants of motion* $C(\vec{x}, \vec{v}, t)$ conserved along the orbit (there must be six because the orbit is fully defined by $\vec{w}(t=0)$).

The *integrals of motion* are related. These are functions only of phase space position that are conserved along orbits:

$$\frac{dI(\vec{x}, \vec{v})}{dt} = 0 \quad (18)$$

Each is also a constant of the motion; so there are at most six of them. The existence of these integrals of motion implies:

$$f(\vec{x}, \vec{v}) = f(I_1(\vec{x}, \vec{v}), I_2(\vec{x}, \vec{v}), \dots, I_6(\vec{x}, \vec{v})) \quad (19)$$

If this were not the case, then f would be an independent integral of the motion itself!

An integral of motion that always exists is total energy. It is conserved for each particle along its orbit. Under specific symmetries, other useful integrals of motion exist. For example, in spherical symmetry, the angular momentum \vec{J} is conserved; note that only its amplitude is physically significant in spherical symmetry however. Therefore under spherical symmetry all equilibrium distribution functions can be written as $f(E, J)$.

A specific case of interest is the *isothermal sphere*. This distribution results from the choice $f \propto \exp(-E/\sigma^2)$. The resulting f can be shown to have a velocity distribution that has a Gaussian width σ in each dimension. Generically, at large radius $\rho \propto r^{-2}$ for an isothermal sphere. At these radii the circular velocity $v_c = \sqrt{2}\sigma$. The case in which r^{-2} at all radii is known as the *singular isothermal sphere*, because of the infinite value of the density at the center.

5. Chandrasekhar Dynamical Friction

Collisionless, gravitating, dynamical systems exhibit an effect known as *dynamical friction* that converts “bulk” kinetic energy into “internal” kinetic energy, even in systems with long two-body relaxation times. This effect is calculated in the exercises below, where it is shown that for a mass M moving through a system with density ρ with an isothermal distribution function of velocity distribution σ there is a drag force:

$$\frac{d\vec{v}_M}{dt} = -\frac{4\pi \ln \Lambda G^2 M \rho}{v_M^3} \left[\operatorname{erf} X - \frac{2X}{\sqrt{\pi}} \exp(-X^2) \right] \vec{v}_M \quad (20)$$

where $X = v_M/\sqrt{2}\sigma$ and $\Lambda \sim M_{\text{total}}/M$. For an initial circular orbit of radius r_i , this drag leads to a dynamical friction time scale:

$$t_f = \frac{2.6 \times 10^{11} \text{ yr}}{\ln \Lambda} \left[\frac{r_i}{2 \text{ kpc}} \right]^2 \left[\frac{v_c}{250 \text{ km s}^{-1}} \right] \left[\frac{10^6 M_\odot}{M} \right] \quad (21)$$

6. Tidal radius

Two point masses M and m separated by distance D will create a potential with a saddle point that separates zones dominated by one of the two potential wells or the other. If these two point masses are orbiting each other with frequency Ω , a fixed potential can only be found in the frame rotating about the center of mass at Ω (and this potential is only relevant for telling you how stationary objects in that frame will start to move in that frame).

The co-rotating frame by design has the relative $1/r^2$ forces from the two masses roughly balancing each other, meaning the position of the saddle point is defined by the derivatives of the gravitational forces, that is, the tidal forces. The dimensional scaling for the tidal forces is $1/r^3$, so:

$$r = \left(\frac{m}{3M} \right)^{1/3} D \quad (22)$$

This tidal radius limits the size of any system of mass m orbiting a body of mass M . Particles or gas more distant than the tidal radius become unbound from mass m . They tend to form *tidal tails*. Particles closer to mass M enter orbits that lead mass m , and particles farther from mass M enter orbits that trail mass m . These tidal features are observable for numerous systems near the Milky Way.

7. Commentary

The fact that collisionless systems have a non-trivial phase space is of enormous significance. It provides another way that each object's history may be encoded in its dynamics. It also means that the properties of a system have a full six-dimensional structure to their description. This complicates accurate predictions of N -body gravitating systems when N cannot be achieved computationally.

The virial theorem is often spoken of in the casual terms that $v^2 \sim GM/r$ for the characteristic, v , M and r for the system. While this relation follows from dimensional analysis alone, the virial theorem goes further and is a precise relationship. It is clearest to think of the virial theorem establishing the relationship between K and U for a bound, equilibrium system. However, as described in the problems, one can create definitions of “characteristic” for v , M , and r for which the equation $v^2 = GM/r$ holds strictly, and will scale with a constant coefficient in homologous systems.

8. Important numbers

9. Key References

- *Galactic Dynamics*, Binney et al. (2009)

10. Order-of-magnitude Exercises

1. Assuming the Milky Way halo is isothermal and spherical, given the circular velocity at the radius of the Sun is 220 km s^{-1} , what do you expect the one-dimensional velocity dispersion of dark matter particles is? What is the mass interior to the Sun?
2. What is the typical relaxation time for globular clusters? Galaxies? Clusters of galaxies?
3. What are the tidal radii for:
 - (a) A Milky Way mass galaxy ($10^{12} M_{\odot}$ total) inside a Coma-mass cluster of galaxies ($10^{15} M_{\odot}$), at 500 kpc from the cluster center.
 - (b) An LMC-mass galaxy ($10^{11} M_{\odot}$ total) inside a Milky Way-mass galaxy, at 50 kpc from the galactic center.
 - (c) A relatively massive globular cluster ($10^5 M_{\odot}$) inside a Milky Way mass galaxy, at 10 kpc from the galactic center.
4. What is the dynamical friction time scale for:
 - (a) A Milky Way mass galaxy ($10^{12} M_{\odot}$ total) inside a Coma-mass cluster of galaxies ($10^{15} M_{\odot}$ total) at 500 kpc from the cluster center.
 - (b) An LMC-mass galaxy ($10^{11} M_{\odot}$ total) inside a Milky Way-mass galaxy at 50 kpc from the galactic center.
 - (c) A relatively massive globular cluster ($10^5 M_{\odot}$) inside a Milky Way mass galaxy, at 10 kpc from the galactic center.

11. Analytic Exercises

1. In this exercise, we derive the equation for the relaxation time given in Equation 1. Relaxation refers to the effect of granularity in the potential due to the fact that there are a finite number of particles in the system, due primarily to two-body interactions. A distribution function that forms an equilibrium system, if sampled by a finite number of particles, will slowly (or quickly) wander away from that equilibrium, on a time scale associated with the relaxation time. We will calculate this for a system of mass M , consisting of N particles with mass m , within some system size R . For this exercise we define the crossing time $t_c = R/v$, where v is the typical velocity of a particle.
 - (a) What does the virial theorem tell us about the crossing time t_c ?
 - (b) Imagine two particles of mass m passing by each other with speed v with an impact parameter b , defined as their separation at infinity normal to their relative velocity. Argue from a heuristic point of view why the velocity perturbation normal to the original velocity will scale $\propto Gm/bv$ (in detail it is $\Delta v_{\perp} \approx 2Gm/bv$).

- (c) Consider interactions in some range of impact parameters, between b and $b + db$. During one crossing time of a particle, what is the mean $\langle \delta v_{\perp} \rangle$ and mean-squared $\langle \delta v_{\perp}^2 \rangle$ perturbation that these interactions cause on the velocity of the particle perpendicular to its motion?
 - (d) Below $b_{\min} = Gm/v^2$ our assumptions break down. We will take the perhaps questionable route of ignoring these close encounters. Considering only interactions with impact parameters between b_{\min} and R to express the total $\langle \delta v_{\perp} \rangle$ over all encounters in a crossing time. Express the result in terms of $\Lambda = R/b_{\min}$.
 - (e) Define the relaxation time as the time it takes for the total fractional perturbation in velocity to reach unity. How many crossing times does it take?
 - (f) Approximate the answer one step further, using the virial theorem to show $\Lambda \sim N$, and thus expressing the number of crossings just in terms of N .
2. Show that Liouville's Theorem follows from the collisionless Boltzmann equation.
 3. Verify the expressions for the first-order Jeans Equation.
 4. Show that Equation 9 follows from Equation 8 under spherical symmetry in configuration space.
 5. Using the spherically symmetric first-order Jeans Equation, Equation (9), show that Equation (11) holds.
 6. For a singular isothermal sphere with $n \propto r^{-2}$ and σ_{rr}^2 a constant, but $\beta > 0$, what choice of β will make the radial velocity dispersion equal to the velocity of a stable circular orbit in the potential?
 7. Starting with the first-order Jeans Equation, Equation (6), take another moment with respect to position. Rearrange in terms of the kinetic and potential energy tensors to obtain the tensor virial theorem.
 8. We will derive the Plummer model for a spherical equilibrium set of orbits. Under Jeans Theorem, all equilibrium models will have $f(E, J)$. We can take:

$$f = \begin{cases} k_1(-E)^p & E < 0 \\ 0 & E > 0 \end{cases} \quad (23)$$

Under this form, we can find self-consistent combinations of ρ and the gravitational potential ϕ .

- (a) Show that under this form, the density becomes:

$$\rho = k_2 (-\phi)^n \quad (24)$$

for $n = p + 3/2$. You may find the following integral useful:

$$\int_0^a dx x^m (a^n - x^n)^p = \frac{a^{m+1+np} \Gamma((m+1)/n) \Gamma(p+1)}{n \Gamma[(m+1)/n + p + 1]} \quad (25)$$

(b) Demonstrate that the potential:

$$\phi = -\frac{GM}{R} \frac{1}{(1 + r^2/R^2)^{1/2}} \quad (26)$$

satisfies this relation for $n = 5$.

9. We will derive the isothermal sphere model for a spherical equilibrium set of orbits. Under Jeans Theorem, all equilibrium models will have $f(E, J)$. We can take:

$$f = k \exp(-E/2\sigma^2) \quad (27)$$

There is something very unusual about this assumption, because it includes positive as well as negative energies. As it turns out, this feature has to do with the result we find later than the total size and mass of the system is infinite.

(a) Show that:

$$\rho = mK (4\pi\sigma^2)^{3/2} \exp(-\phi/2\sigma^2) \quad (28)$$

(b) Show that the density must satisfy the equation:

$$\frac{\partial}{\partial r} \left(r^2 \frac{\partial \ln \rho}{\partial r} \right) = -\frac{2\pi G}{\sigma^2} r^2 \rho \quad (29)$$

- (c) Using the hydrostatic equation and the ideal gas law, and assuming isothermal gas, show that σ^2 is the equivalent of $kT/2m$ — i.e. $m\sigma^2$ is like the energy per degree of freedom.
- (d) Find the singular solution with $\rho \propto r^{-2}$ everywhere, and express $\rho(r)$, $M(< r)$, and $v_c(r)$ as functions of σ^2 and r .

We will show in a numerical exercise that we can numerically integrate the equations to a better solution by imposing a zero slope at $r = 0$ (but it is still infinite mass).

10. Tidal radius

11. We will here derive Equation (20) for the Chandrasekhar Dynamical Friction. We will calculate the drag on a particle of mass M moving at velocity \vec{v}_M through a system of density ρ with an isotropic Gaussian velocity dispersion σ in each Cartesian dimension.

(a) We use the same approach for calculating the relaxation time. Consider a two-body interaction in the frame of the particle with mass M , with a particle in the field with velocity \vec{v} . Using the results of Exercise 1b, estimate the change of the velocity of the second particle along the original direction \hat{v} .

- (b) What is the velocity change of the particle of mass M due to this interaction?
- (c) If the particle of mass M travels at velocity \vec{v}_M through a uniform sea of particles of mass m with a velocity distribution $f(\vec{v}_m)$, what is its rate of encounters with particles within some differential $d^3\vec{v}_m$ around a particular velocity \vec{v}_m , and within a differential db around impact parameter b ?
- (d) In the case of the perpendicular deflection, the mean velocity difference averages to zero. However, the mean velocity difference along the parallel direction is always in the same direction. Assume some b_{\max} is set by the size of the system in question, and define:

$$\Lambda = \frac{b_{\max} v_0^2}{GM} \sim \frac{M_{\text{total}}}{M} \quad (30)$$

where the approximation is from the virial relation. As in the case of two-body relaxation, we need to set a reasonable lower limit of $b_{\min} \approx GM/v_0^2$. Derive the change in \vec{v}_M per unit time from particles around velocity \vec{v}_m .

- (e) Now we have to integrate this over all velocities. This is easier than it appears, because the velocity distribution is isotropic around $\vec{v}_m = 0$; i.e. it has spherical symmetry in velocity space. Meanwhile the integral is over a kernel of the form $1/|\vec{v}_m - \vec{v}_M|^2$. This inverse square form has the property that we only need to consider the velocities $v_m < v_M$, as Newton showed a long time ago in the case of inverse-square gravity. Perform this integral to get the total drag force.
12. Imagine two masses, M_1 and M_2 , orbiting in a circular orbit around each other with a constant separation D . Transform into a frame rotating with the orbit, and find the effective potential Φ_{eff} such that in the rotating frame:

$$d\vec{v}/dt = -\vec{\nabla}\Phi_{\text{eff}}. \quad (31)$$

Plot contours of this effective potential in the orbital plane; you should see a number of stationary points (the Lagrange points), one of which is between the two masses. Then, for $M_2 \ll M_1$, show that the stationary point between the two masses is a distance from the small mass of:

$$r = \left(\frac{M_2}{3M_1} \right)^{1/3} D \quad (32)$$

That distance is approximately the tidal radius.

12. Numerics and Data Exercises

1. Using the inner boundary condition that $\partial\rho/\partial r = 0$, integrate Equation 29 outwards to derive the shape of an isothermal sphere with a core.

2. The spherically symmetric isothermal model has an issue that it has finite phase space density for positive energies. The “lowered isothermal model” deals with this issue by setting:

$$f(E) = \rho_1 (2\pi\sigma^2)^{-3/2} \left(e^{-E/\sigma^2} - 1 \right) \quad (33)$$

for $E < 0$ and $f(E) = 0$ otherwise. Analytically, integrate this phase space density to get the density ρ and the radial equation for the potential $\Phi(r)$. The *King models* are models which obey the condition $d\Phi/dr = 0$. Setting the initial value $\Phi(r = 0)$ to $-3\sigma^2$, $-6\sigma^2$, and $-12\sigma^2$, numerically integrate the equation outward until $\Phi(r)$ becomes positive and plot the resulting density profiles. Consult Binney & Tremaine §4.3.3(c) for assistance and to check your answer!

3. Φ , n and β
4. Dynamics estimates from Gaia
5. Globular cluster radial profiles.
6. Tidal potential

REFERENCES

Binney, J., Tremaine, S., & Freeman, K. 2009, *Physics Today*, 62, 56