

Using cosmic variance to constrain the dark matter halo mass of Lyman-alpha emitting galaxies at $z=3.1$

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ABSTRACT

We use cosmological N-body simulations to find the characteristic mass of dark matter halos hosting Lyman-Alpha Emitting (LAE) galaxies at a redshift of $z = 3.1$. The method is based on matching the statistics for the number density between mock and observed fields. The mock catalogs are constructed using a simple model where a dark matter halo can only host one LAE with a probability f_{occ} if its mass is found within a certain range mass range delimited by two threshold values, M_{min} and M_{max} . We find that the most of the models that are consistent the observed cosmic variance statistics are those with halo masses in the range $10.5 < \log_{10} M_{\text{min}}/h\text{M}_{\odot} < 11.5$ and $\log_{10} M_{\text{max}}/h^{-1}\text{M}_{\odot} < 13.5$ with an occupation fraction that scales as $f_{\text{occ}} =$. We explore three additional constraints to narrow down these range: the number of mocks consistent with observations, observational constraints on the occupation fraction and the angular correlation function. The first two conditions narrow down the space parameter to $M_{\text{min}} =$ and $M_{\text{max}}, f_{\text{occ}}$. The angular correlation function does not add a significant constraint due to the cosmic variance in the small angular fields where this statistics has been computed so far. We make available the mock data for the best models in a public repository. Implications for galaxy formation models?

Key words: galaxies: kinematics and dynamics, Local Group, methods:numerical

1 INTRODUCTION

Lyman- α emitting galaxies (LAEs) have become in the last decade a central topic in studies of structure formation in the Universe. They are helpful in a diverse range of fields. LAEs can be used as probes of reionization (Dijkstra et al. 2011), tracers of large scale structure (Koehler et al. 2007), signposts for low metallicity stellar populations and markers of the galaxy formation process through cosmic history (Forero-Romero et al. 2012).

At the same time, theoretical and observational developments have contributed to the emergence of a paradigm to describe structure formation in a cosmological context. In this context it is considered that dominant matter content of the Universe is to be found in dark matter, whereby each galaxy is hosted by larger dark matter structure known as a halo.

Most models of galaxy formation find that the mass of the halo can be used to predict properties of the galaxy such as its stellar mass and star formation rate (Behroozi et al. 2012). Processes that regulate the star formation cycle are also thought to be strongly dependent on its mass. Furthermore, the spatial clustering of galaxies on large scales is entirely dictated by the halo distribution. For the reasons

mentioned above, finding the typical dark matter halo mass hosting LAEs represents a significant step forward to understand the nature of this population in the context of Lambda Cold Dark Matter (Λ CDM) paradigm.

Some theoretical approaches to this problem have been based on a forward modeling. Starting from the DM halo population, the corresponding intrinsic star formation properties are inferred and statistics such as the luminosity function, the correlation function and the equivalent width distributions. Such modelling has been implemented from analytic considerations, semi-analytic models and full N-body hydrodynamical simulations (Dayal et al. 2009; Forero-Romero et al. 2011; Yajima et al. 2012; Forero-Romero et al. 2012).

Added to the uncertainties in the astrophysical processes describing star formation in galactic populations, a highly debated step in this approach is the calculation of the fraction of Lyman- α photons that escape the galaxy to the observer. Given the resonance nature of the line, the radiative transfer of Lyman- α is sensitive to the density, temperature, topology and kinematics of the neutral Hydrogen in the interstellar medium (ISM) (Neufeld 1991; Forero-Romero et al. 2011; Laursen et al. 2013).

This complexity makes the use of monte-carlo simulations for the radiative transfer a required tool to obtain physically sound results, although the degeneracy in the physical parameters involved in the problem makes it difficult to achieve a robust consensus on what is the theoretical expected value for the Lyman- α escape fraction in high redshift.

Throughout this paper we assume a Λ CDM cosmology with the following values for the cosmological parameters, $\Omega_m = 0.27$, $\Omega_\Lambda = 0.73$ and $h = 0.70$, corresponding to the matter density, vacuum density and the Hubble constant in units of $100 \text{ km s}^{-1} \text{ Mpc}^{-1}$.

2 METHODOLOGY

In this paper we constrain the typical mass of dark matter halos hosting LAEs at $z = 3.1$. Our model is based on the number density information obtained in the recent large scale survey presented by XXX where XXX LAEs are detected over 7 fields of $\sim 46 \times 35 \text{ Mpc}^2 h^{-2}$ in area comoving in area corresponding to observed fields of XXX deg^2 .

Both the spatial distribution and luminosities of the galaxies have, at least in a statistical sense, information to constrain theoretical model of LAEs. The most detailed theoretical models are also faced to diverse physical and astrophysical uncertainties in obtaining statistical prediction for the Lyman- α line. This uncertainties are largest impediment to construct an ab-initio model for LAEs.

In this paper we want to step back and reduce the complexity of our model, with the sole objective of reproducing the cosmic variance in the number density of LAEs. Afterwards we will interpret the implications of this result for physical models for Lyman- α emitting galaxies.

Our model is based on the predictions of a large volume high resolution N-body simulation describing the gravitational dynamics of dark matter. We do not have an strong bias towards the theoretical expectation of what the mass of the dark matter halo hosting the galaxy should be. Instead, we fully explore the parameter space of our simplified model. The only cut we impose is that observed LAEs do not reside in dark matter halos with masses less than $10^{10} h^{-1} \text{ M}_\odot$ [citation].

In the following subsections we describe the most relevant features of the observational data, the N-body simulation we use, our model and its parameters together with the method to compare its predictions against observations.

2.1 The Observational Constraints

Our observational reference are the recently published results of a panoramic survey of LAEs at $z = 3.1$ by Yamada et al. (2012). This survey was conducted with the Subaru 8.2m telescope and the Subaru Prime Focus Camera, which has a field of view covering $34 \times 27 \text{ arcmin}$, corresponding to a comoving scale of $46 \times 35 \text{ Mpc } h^{-1}$ at $z = 3.09$. The narrow band filter is centered at 4977 \AA with a 77 \AA width, corresponding to the redshift range $z = 3.062 - 3.125$ and $41 \text{ Mpc } h^{-1}$ comoving scale for the detection of the Lyman- α line centered at $z = 3.09$.

The choice to have only one the data from Yamada et al 2012 as reference was made because their surveys is

the largest in area with a set of homogenous conditions that define the LAE sample. Other surveys by XXX and XXX that cover similar regions, but they use different criteria on the equivalent width (EW) cuts to construct the LAE samples. Different cuts in the EW can change the number of LAEs to be included in the catalog. This cuts have an impact on the fainter LAEs which are more abundante than brighter ones. Different definitions of the EW cuts can yield number densities different by a factor of two [REF, I think Yamada has some numbers].

The survey covered four independent fields. The first is the SSA22 field of 1.38 deg^2 with 1394 detected LAEs, this field has been known to harbor a region with a large density excess of galaxies. The second observed region is composed by the fields Subaru/XMM-Newton Deep Survey (SXDS)-North, -Center and -South, with a total of 0.58 deg^2 and 386 LAEs. The third and fourth fields are the Subaru Deep Field (SDF) with 0.22 deg^2 and 196 LAEs, and the field around the Great Observatory Optical Deep Survey North (GOODS-N) with 0.24 deg^2 and 185 LAEs. In Table 1 we summarize the values we use in throughout this paper for the each field, covered area, measured surface LAE number density and inferred number volume density.

2.2 The Simulation and Halo Catalogs

The Bolshoi simulation was performed in a cubic volume of $250 h^{-1} \text{ Mpc}$ on a side. It includes dark matter distribution is sampled using 2048^3 particles, which translates into a particle mass of $m_p = 1.35 \times 10^8 h^{-1} \text{ M}_\odot$. The cosmological parameters are consistent with a WMAP5 and WMAP7 data with a matter density $\Omega_m = 0.27$, cosmological constant $\Omega_\Lambda = 0.73$, dimensionless Hubble constant $h = 0.70$, slope of the power spectrum $n = 0.95$ and normalization of the power spectrum $\sigma_8 = 0.82$ [REF].

We use halo catalogs constructed with a Friend-of-Friends (FOF) algorithm with a linking length of 0.17 times the interparticle distance. We have verified that the main results we present in this paper also hold if instead we use halo catalogs constructed from a the Bound Density Maxima (BDM) algorithm (Klypin et al. 1999) that are defined to have a density of 200 times the critical density. The minimum halo mass in the models we construct in this paper correspond to groups of ~ 75 particles. The catalogs were obtained from the publicly available Multidark database¹ (Riebe et al. 2011).

2.3 Populating Halos with LAEs

The model that populates halos with LAEs is based on a one-to-one correspondence: each halo can only host a single LAE. There are three physical parameters in the model: the halo mass range $M_{\min} < M_{\text{halo}} < M_{\max}$ where LAEs reside and the fraction f_{occ} of such halos that effectively host a LAE. In what follows we will describe by the letter \mathcal{M} a model defined by these three parameters M_{\min} , M_{\max} y f_{occ} .

We stress that we do not intent to build a model for the luminosity of each LAE. Physically speaking we are primarily interested in constraining the halo mass above which

¹ <http://www.multidark.org/MultiDark/>

there are detectable LAEs. under the conditions defined by Yamada et al.

For each mode \mathcal{M} we create mock field from disjoint volumes in the simulation with the same geometry probed by Suprime-CAM and the narrow band filter, namely $46 \times 35 \times 41 h^{-3} \text{Mpc}^3$ where the last dimension goes in the redshift direction, corresponding to a total area of 880 arcmin^2 for each mock field. There is a total $5 \times 7 \times 6 = 210$ of such sub-volumes in a snapshot of the Bolshoi simulation.

Next we group these 210 mock fields in three different ways to construct the LAEs number density distribution. The first way (match method) we follow the observational setup and constructs 15 different mock surveys, each one composed of 12 mock fields, out of which 7 correspond to contiguous sub-boxes in the simulation to mimick the whole SSA22, 3 are also contiguous between them but not to the first 7 fields to mimick the SXDS fields and finally 2 non-contiguous fields that correspond to the SDF and GOODS-North fields. This will produce 15 different distributions for the number density for a given model M . The second (random method) is similar to the first one. There are 15 different mock surveys with 12 mock fields each, but this time each field corresponds to uncorrelated sub-boxes in the simulation. The third (full method) way in only has 1 mock survey containing all the 210 mock fields, in this setup there is only one predicted number density distribution for each model M .

The advantage of these three sampling ways is that they allow us to explore the effects of both cosmic variance and the correlation between fields. Comparing the results of the first and second method will help us to quantify the effect of field correlation, while comparing the first and the third method will serve us to gauge the impact of cosmic variance.

2.4 Model Sampling and Selection

We generate a series of models \mathcal{M} with different input parameters $\{M_{\min}, M_{\max}, f_{\text{occ}}\}$ as follows. M_{\min} and M_{\max} are allowed to take 30 different values evenly spaced by 0.1 dex, M_{\min} ranges from $\log_{10} M_{\min} = 10.0$ up to $\log_{10} M_{\min} = 12.9$, while M_{\max} range from $\log_{10} M_{\min} = 10.1$ up to $\log_{10} M_{\min} = 13.0$. The occupation fraction f_{occ} takes 100 different values from 0.01 to 1.00 regularly spaced by 0.01. In total the number of different sets of input parameters to be explored is $30 \times 30 \times 100 = 9 \times 10^4$.

For each model \mathcal{M} we compute the LAE surface density distributions for the three different ways of grouping the mock fields, as described in the previous section. For each sub-volume we project the positions of the LAE hosting halos along the z direction and calculate its surface number density in units of sources per arcmin^2 . For each number density distribution we perform a Kolmogorov-Smirnov against the 12 surface density observational values. From this test we obtain the value $0 < P < 1$ to reject the null hypothesis, namely that the two data sets come from the same distribution. In this paper we use values of $P > 0.1$ to consider that the simulated and observed number densities come from the same distribution.

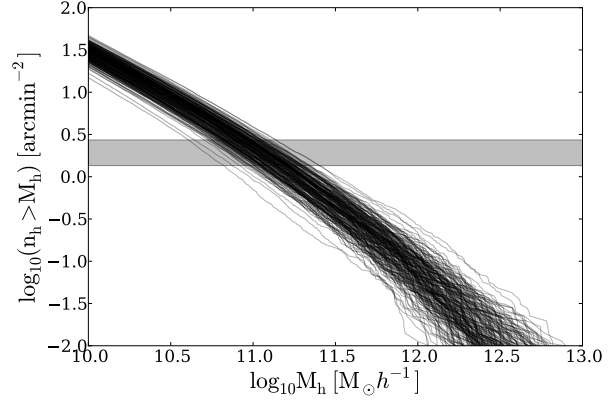


Figure 1. Cumulative mass function of dark matter haloes in the 210 sub-volumes of $46 \times 35 \times 41 h^{-3} \text{Mpc}^3$. The variation in the total number of dark matter halos per sub-volume evidences the effect of cosmic variance at such sub-volume scale. It is also appreciable the low population $\lesssim 10^{-3} h^2 \text{Mpc}^{-2}$ of halos with $\log(M/M_{\odot}) > 12.0$

3 RESULTS

3.1 Dark Matter Halo Number Density

In Figure ?? we present the results for the integrated dark matter halo surface density as a function of halos mass. Each line corresponds to one of the 210 sub-volumes in the Bolshoi simulation. The shadowed area indicates the surface density values for LAEs allowed by the observations.

This result allows us to better understand the expected trends for the LAEs' preferred mass and the occupation fraction. From this Figure we can read which models do not have a chance to reproduce the observations. Regions in the plot where the halo surface density values are below the observational constraint correspond to high masses halo masses. For a LAE model \mathcal{M} with a minimum mass $M_{\min} > 3 \times 10^{11} h^{-1} M_{\odot}$ located in that mass range, the surface density is too low compared with observations.

Conversely, there are regions in the plot where the halo surface density is always higher than the observational constraints correspond to models \mathcal{M} with a minimum mass below $M_{\min} < 3 \times 10^{10} h^{-1} M_{\odot}$. Models with this minimum mass have a chance for successfully reproducing observations if the occupation fraction $f_{\text{occ}} < 1$ is tuned as to lower the halo number density down to the observed value.

3.2 Kolmogorov-Smirnov Tests

Figure presents the regions in the parameter space M_{\min} - M_{\max} where the KS test yields values of $P > 0.05$. We consider that for those models it is not possible to rule out the null hypothesis, namely that the number density in simulated data and the observations come from the same parent distribution. Each panel corresponds to the three different ways of grouping the mock fields. In the case of the methods **Match** and **Random** the color code indicates the fraction of these 15 mock surveys with $P > 0.1$. The third panel shows the result for the method **Full**, in this case the color

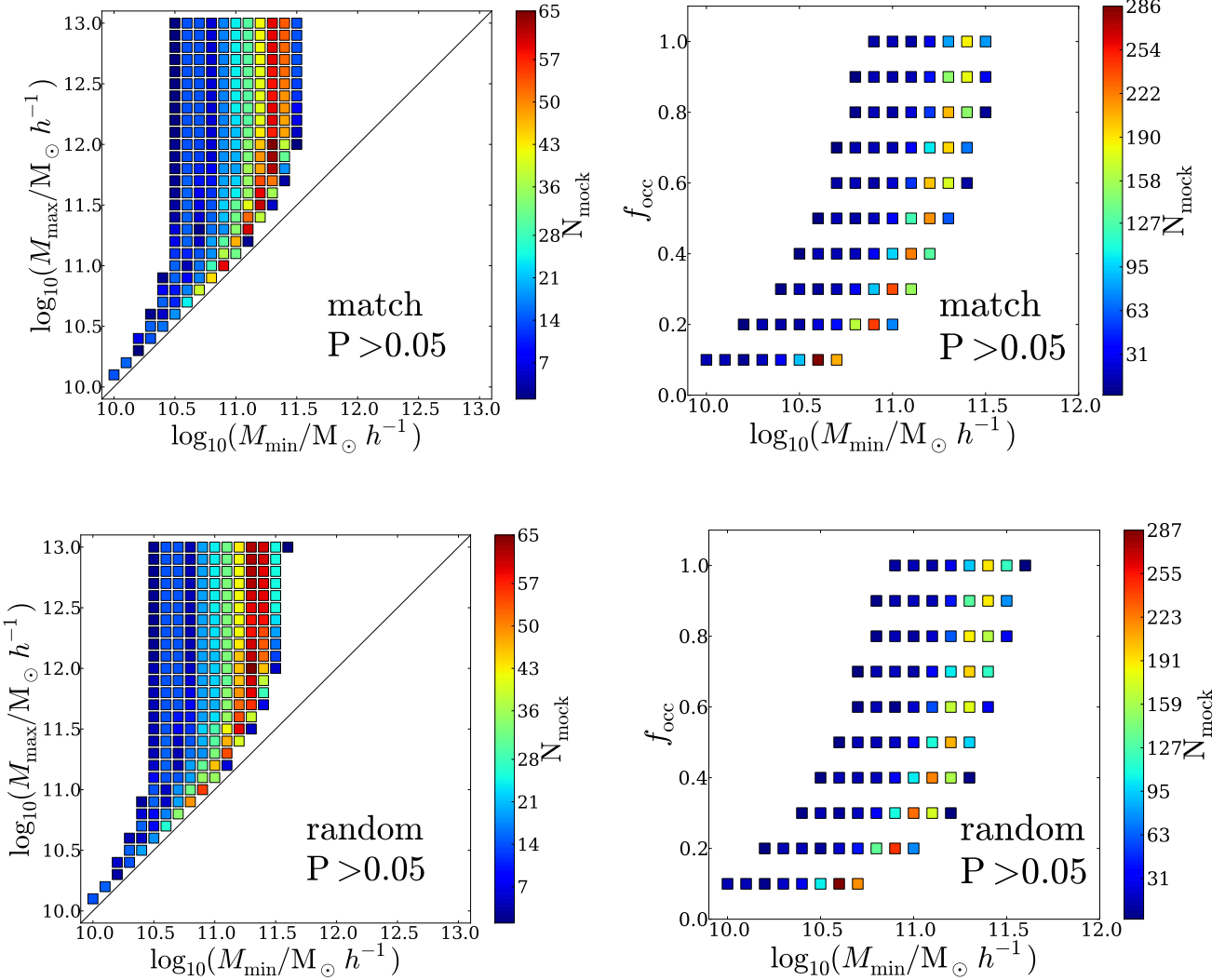


Figure 2. M_{\min} - M_{\max} plane for all models with $P > 0.05$ in the three different of grouping the mock fields. In the case of the **match** and **random** methods the color code corresponds to the number of mock surveys that are found to be compatible with observations. For the **full** method the colour code corresponds to the results of 100 times the maximum P value resulting from the KS test. .

code correspond to the value maximum value of $100 \times P$ for a model with those mass ranges.

These results clearly distinguish three mass regimes. In the first regime, at high mass values, we find that LAE models with minimum mass of $M_{\min} > 10^{11.5} h^{-1} M_{\odot}$ are not compatible with observations. There is a second regime for masses below $M_{\min} = 3 \times 10^{10}$ any values for M_{\min} and M_{\max} can be made compatible with observations, provided that f_{occ} is fine tuned to do it. In an intermediate mass regime, for minimum mass values $3 \times 10^{10} h^{-1} M_{\odot} < M_{\min} < 3 \times 10^{11} h^{-1} M_{\odot}$ only a limited range of models with M_{\max} with occupation fraction $f_{\text{occ}} \sim 1$ is able to reproduce observations.

In these three different mass regimes the occupation monotonically decreases as a function of the minimum halo mass M_{\min} . In Figure XXX we show this trend in three panels following the same correspondence as Figure XXX. From these results we interpret that the different mass

regimes that were identified correspond to best fit models with $f_{\text{occ}} \sim 1$, $0 < f_{\text{occ}} < 1$ and $f_{\text{occ}} \sim 0$, respectively.

The two **match** and **random** methods present the highest number of matching mock surveys in the medium mass regime. However, it is important to keep in mind that not all the mock surveys for a successful model M present a high value $P > 0.1$, only a modest fraction seems to be consistent with observations. This shows that the cosmic variance is still present on the physical scales probed by observations. This will be considered in more detail in the discussion section.

To illustrate this point, in Figure ?? we present the results for two mock fields for the **Match** method for a model with the same parameters, but two extreme values for the KS test.

Conversely there are different models where the KS test yield values of $P \sim 1$. To illustrate the kind of success represented by these models, we have selected these best ones in the case of the method **MatchObs**. Figure XXX shows in

the main panel the spatial distribution of the mock surveys, the smaller panel shows the corresponding surface density distribution and the observational constraint.

In what follows we will focus our discussion on the mocks constructed with the **Match** and **Random** methods.

3.3 Additional Constraints

The raw results of our experiments yield a preferred range for halo masses hosting LAEs bounded by a minimum mass $10^{10} h^{-1} M_{\odot} < M_{\min} < 10^{11.5} h^{-1} M_{\odot}$ and without any limitation on the maximum mass. In this regime the average occupation fraction varies with an approximate dependence of the kind $f_{\text{occ}} \approx 0.1(M_{\min}/10^{10})^{X}$. These results discard models that were nevertheless disfavored from the very beginning based on the results of the mass functions (see Figure 1.).

In this section we consider three different ways to impose tighter constraints on this mass range by making use of the results we have derived so far together with additional statistical and astrophysical constraints. In the first constraint we select the models where all the mock surveys present KS-test values consistent with observations. The second constraints uses recent observational results on the average occupation fraction for LAEs at high- z . The third exploits the information in the Angular Correlation Function (ACF).

3.3.1 Models with the highest success rates

For each model there are 15 different mock survey realizations. In the previous section we presented the models that had at least one (1) mock survey realization with $P > 0.05$.

In Figure 3 we show the number of models that have at least n realizations with $P > 0.05$ for the **match** and **random** methods. This shows that there are around 550 to 600 different models that have at least one mock survey realization consistent with observations. At the other extreme, there are 80 to 100 models with all the 15 realizations with $P > 0.05$. Here we focus on the latter models. The best models represent $\sim 15\%$ of the number of initially considered good models.

In Figure XX we present the loci of these models in the parameter space $M_{\min} - M_{\max}$ and $M_{\min} - f_{\text{occ}}$. In Figure XX we show the spatial distribution for two mock surveys corresponding to one of such models.

3.3.2 Observational constraints on the occupation fraction

We now impose a different restriction using the observational results by Hayes et al. (2010). These authors constrained the value of f_{occ} at $z = 2.2$ to be $f_{\text{occ}} = 0.10$. This estimation was based on blind surveys of the $\text{H}\alpha$ and Lyman alpha line with the European Southern Observatory (ESO) Very Large Telescope (VLT). Using corrections by extinction to obtain an estimate for the intrinsic $\text{H}\alpha$ luminosity, and using values for the theoretical expectation of the ratio $\text{Lyman}\alpha/\text{H}\alpha$ they derive a bulk escape fraction for the Lyman α radiation of $f_{\text{esc}} = (5.3 \pm 3.8)\%$ or $f_{\text{esc}} = (10.7 \pm 2.8)\%$ if a different dust correction is used. The authors show that the luminosity function for LAEs at $z = 2.2$ is consistent

with the escape fraction being constant for every galaxy regardless of its luminosity. From this results they derive that almost 90% of the star forming galaxies emit insufficient Lyman-alpha to be detected, effectively setting the occupation fraction to be $f_{\text{occ}} = 0.10$.

For the cosmological parameters used in this paper the age of the universe between $z = 3.1$ and $z = 2.2$ has changed by ~ 1 Gyr. We assume that the physical conditions that determine the escape fraction f_{esc} and the occupation fraction f_{occ} remain constant over that time scale. This assumption allows us to further pick models that have an occupation fraction of $f_{\text{occ}} = 0.10$. Under this selection only 18 models can be selected. Considering an occupation fraction $f_{\text{occ}} = 0.20$ another 39 models can be considered. This constraint help us to select $\sim 10\%$ of the original models that were considered as consistent with observations in the previous section.

Figure XX shows the preferred models in the planes $M_{\min} - M_{\max}$ and $M_{\min} - f_{\text{occ}}$ for the **match** and **random** methods. The list for the model parameters is found in the appendix in Table XX.

In terms of the constraints done in the previous subsection, we find that the models in this region of parameter space have an average number of XX mock catalogs consistent with observations.

3.3.3 Constraints from the Angular Correlation Function

We calculate the mean angular correlation function (ACF) for all models with $P > 0.05$ using the **match** and **random** methods. These correlation functions is calculated over the densest subfield in all the mock surveys corresponding to the SSA22 region. These results are compared against the observations reported by Hayashino et al in 2004 over the same region, which were also performed on the densest field.

Figure 5 (**match**) and Figure 6 (**random**) present such comparison. The error bars in these figures represent the standard deviation of the ACF over all the sub-fields. In general, we observe that the standard deviation of the computed ACF in the subfields increases with M_{\min} following the same trend as in Figure 1, as a direct consequence of cosmic variance.

The comparison between the simulated and observed ACFs is also done using a R^2 statistic which includes the information on measurement uncertainties

$$R^2 = \sum_{\theta_i} \frac{(\xi_{\text{obs}}(\theta_i) - \xi_{\text{sim}}(\theta_i))^2}{\sigma_{\text{obs}}^2(\theta_i) + \sigma_{\text{sim}}^2(\theta_i)}, \quad (1)$$

where the sum is done over all the angle values θ_i where the ACF has been computed. In Figure XX we plot the integrated distributions for this R^2 statistics

Given that the ACF reported by Hayashino et al in 2004 is taken over the densest field observed in the SSA22 region by Yamada et al in 2012 it is expected that the predicted ACF in the SSA22 region should reproduce this observation. In the left panel of figure 5 we can see the predicted ACFs and their corresponding standard deviation over the seven fields that mock the SSA22 region. It can be seen that the model with $M_{\min} = 10.6$ seems to better reproduce the Hayashino's ACF and that the corresponding field is in fact

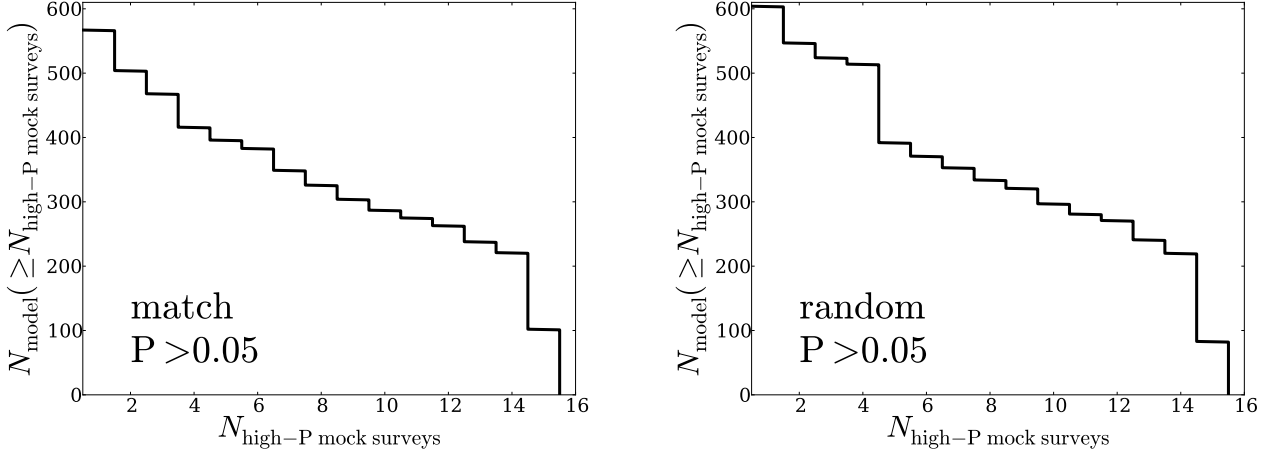


Figure 3. Number of models with a minimum number of mock survey realizations that are consistent with observations. .

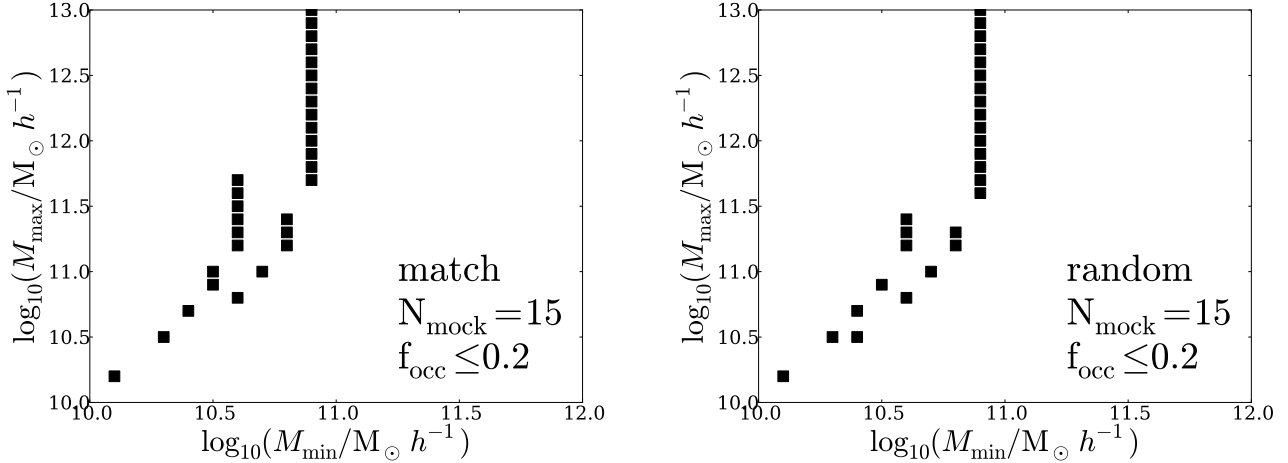


Figure 4. Favored models when the constraints on the maximal number of consistent mocks and the observational constraint on the occupation fraction are included.

an overdense field in the SSA22 region covered by Yamada et al.

From these tests we conclude that the ACF on small fields does not provide additional constraints to further select models for halos hosting LAEs. The reason is that cosmic variance is large and the statistical uncertainties on the ACF render almost any model compatible with the observational constraints.

We also present the ACF in the case of the full SSA22 region which has been homogeneously observed by (Yamada et al. 2012). To this date the observational ACF has not been reported in the literature, therefore our calculations can be considered as predictions.

In Figure X we present the results for the models. The full list of these correlation functions can be found in the data repository for this paper in [github](#).

4 DISCUSSION

When we include the tightest constraints on the mock catalogs, we find that there are 30 set of parameters of our model, out of the original 90000 initial models, that are consistent with the observational constraints at redshift 3.1: the distribution of the number density, the inferred values for the average occupation fraction and the angular correlation function.

These 30 models can be classified into two families of the same size. The first, where the range $M_{\min} - M_{\max}$ is narrow, typically of less than < 1.0 dex. While in the second family the extent > 1.0 dex. In the first case the minimum halo mass is found to be in a wide range $10^{10} h^{-1} M_{\odot} < M_{\min} < 10^{11.5} h^{-1} M_{\odot}$ while in the second case, only models with $M_{\min} \sim 10^{10.9} h^{-1} M_{\odot}$ are compatible with the observational constraints. In what follows we discuss the implications of the existence of these two families of models.

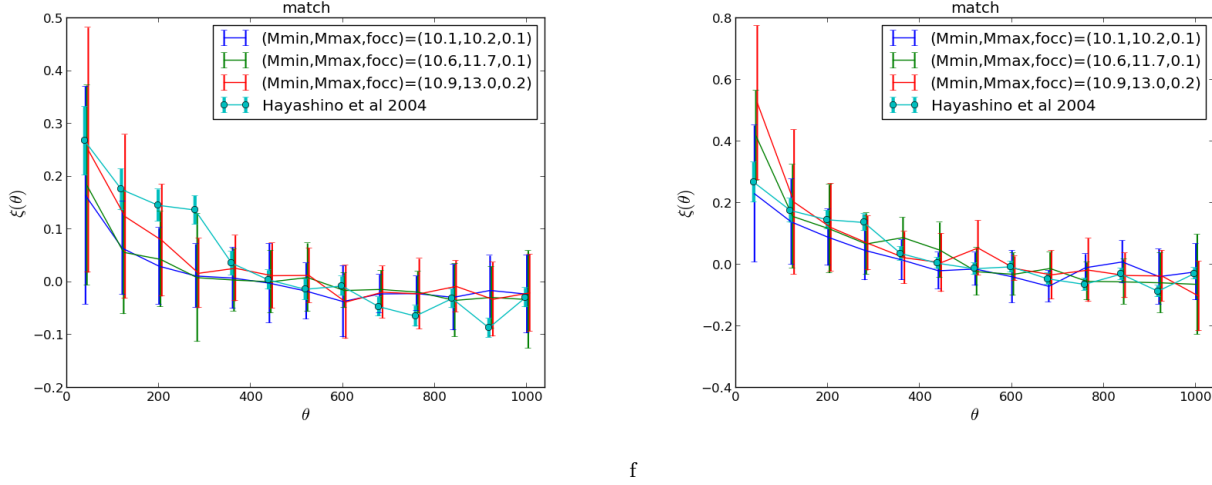


Figure 5. mean ACFs and their corresponding standard deviation (error bars) of some selected models in different mass ranges over the 7 subfields f of the SSA22 field (left) and the entire 12 field sample (right) using the match configuration.

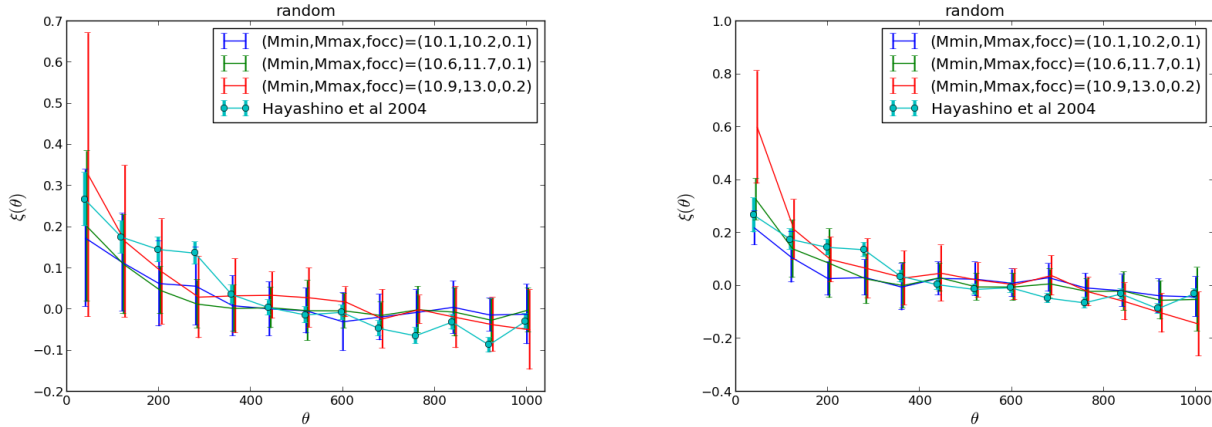


Figure 6. mean ACFs and their corresponding standard deviation (error bars) of some selected models in different mass ranges over the 7 subfields of the SSA22 field (left) and the entire 12 field sample (right) using the random configuration.

4.1 Implications for galaxy formation models

In the case of a narrow of masses to host LAEs the upper masses are bound to be $M_{\max} < 10^{11.5} h^{-1} M_{\odot}$ as it is show in Figure 4. For halos more massive than this bound it is naturally expected that the galaxies can be observed as Lyman Break Galaxies (LBGs). This would imply that not all the bright LBGs can be detected as a LAE.

We have the opposite situation in the second family of models. If we have a wide range in halo masses, where the upper end of the halo masses can be considered as observed LAEs, one can expect that bright LBGs will have a correspondence as observed LAEs. The most interesting aspect is that there is a clear cut in the minimal mass that can be attained by observed LAEs $M_{\min} > 10^{11} h^{-1} M_{\odot}$. This puts a tight constraint on the relationship between the minimum star formation rate required to be observed as a LAE and this minimal halo mass.

- ... Intrinsic emission and escape fraction.
- ... Star formation rate efficiency at this redshift.

... Mass dependence of the escape fraction.

4.2 Implications for large LAEs surveys

- ... The bias for the preferred halo mass.
- ... The scale at which cosmic variance drops.
- ... This can be observationally tested with HETDEX.

4.3 On the reproducibility of our results

- ... All the software to produce the results in this paper is publicly available.
- ... The raw catalogs can be obtained from the MultiDark database but can also be obtained in the repository of this paper on github.

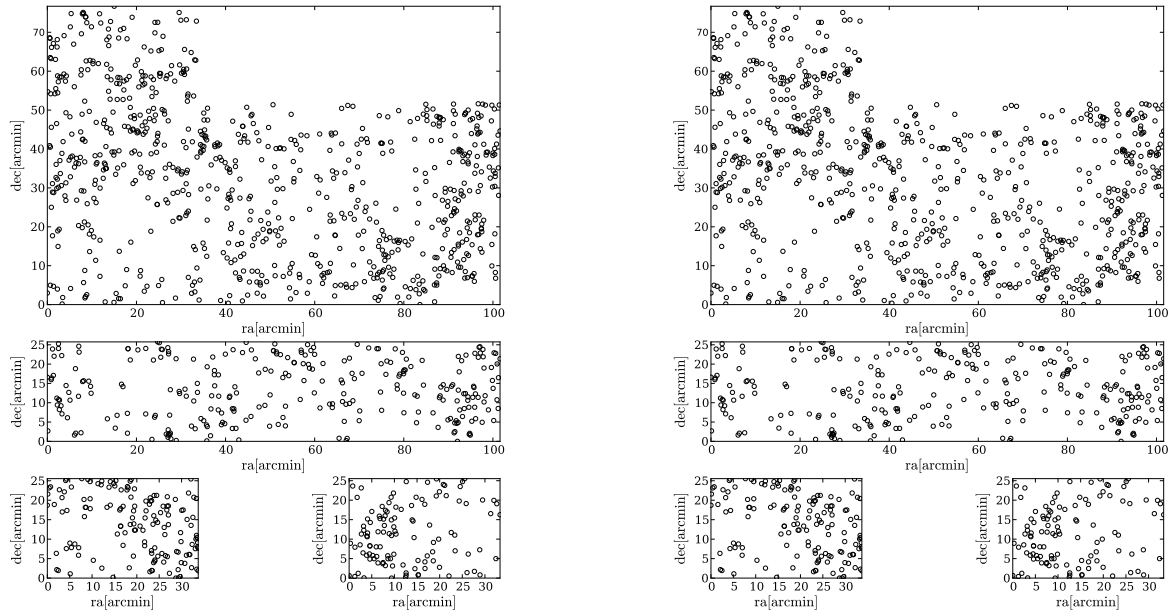


Figure 7. Spatial distribution for two mocks corresponding to the model $M_{\min} =$, $M_{\max} =$ and $f_{\text{occ}} =$. All the 15 different mock surveys for this model in the `match` configuration are consistent with observations at the $P > 0.05$ level. The full data for all the mocks can be found in the github repository for this paper. .

5 CONCLUSIONS

In this paper we constrain the preferred mass for dark matter halos hosting Lyman Alpha Emitters at a redshift $z = 3.1$. We use a method that matches the cosmic variance in the surface density number of LAEs between mock and real observations. The mock catalogs are based on a simplified model with three basic parameters: the halo mass range where LAEs can be found, $M_{\min} < M_h < M_{\max}$, and the fraction of the halos in this range that are actually occupied, f_{occ} . After a thorough exploration of the parameter space we are able to constrain the mass range of dark matter halos hosting LAEs to be in the range $< M_h <$ and a corresponding occupation fraction that scales as $f_{\text{occ}} = M_{\min}$.

We use three additional constraints to reduce the allowed range of models. The first imposes a tighter criterion to consider a model successful, namely that all the mock surveys for a given model must be consistent with observations. This restriction narrows down the allowed range of models to be.

The second constraint is based on the observational results that high redshift LAEs have a bulk Lyman alpha escape fraction of XX which can be also interpreted as an average occupation fraction of XX .

Including additional observational constraints on the occupation fraction allows us to reduce the range of allowed halo masses to be in a narrower range of $< M_h <$. Including the information from the angular correlation function (ACF) does not allow us to impose further constraints. This is due to the scatter in the ACF due to the cosmic variance on the field observed by XXX .

Our simulation allows us to extract 210 sub-boxes each of which has a comparable volume to the individual fields of view observed by Yamada et al. (2012). The comparison of the observed number density distribution against the

results from our model is based on three different ways of constructing mock surveys. The first reproduces the spatial correlation between the 12 observational fields (`match`), the second breaks this spatial correlation while keeping the number of fields (`random`) and the third one simply includes all the 210 sub-boxes (`full`). We find that the methods `match` and `random` allow a larger set of models than the `full` method. We do not find a significant difference between the two first methods.

ACKNOWLEDGMENTS

REFERENCES

- Behroozi P. S., Wechsler R. H., Conroy C., 2012, ArXiv e-prints
- Dayal P., Ferrara A., Saro A., Salvaterra R., Borgani S., Tornatore L., 2009, MNRAS, 400, 2000
- Dijkstra M., Mesinger A., Wyithe J. S. B., 2011, MNRAS, 414, 2139
- Forero-Romero J. E., Yepes G., Gottlöber S., Knollmann S. R., Cuesta A. J., Prada F., 2011, MNRAS, 415, 3666
- Forero-Romero J. E., Yepes G., Gottlöber S., Prada F., 2012, MNRAS, 419, 952
- Hayes M., Östlin G., Schaerer D., Mas-Hesse J. M., Leitherer C., Atek H., Kunth D., Verhamme A., de Barros S., Melinder J., 2010, Nature, 464, 562
- Klypin A., Gottlöber S., Kravtsov A. V., Khokhlov A. M., 1999, ApJ, 516, 530
- Koehler R. S., Schuecker P., Gebhardt K., 2007, A&A, 462, 7
- Laursen P., Duval F., Östlin G., 2013, ApJ, 766, 124
- Neufeld D. A., 1991, ApJL, 370, L85

10.1	10.2	0.1
10.3	10.5	0.1
10.4	10.7	0.1
10.5	10.9	0.1
10.6	11.2	0.1
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10.6	11.4	0.1
10.1	10.2	0.1
10.3	10.5	0.1
10.4	10.7	0.1
10.5	10.9	0.1
10.5	11.0	0.1
10.6	11.2	0.1
10.6	11.3	0.1
10.6	11.4	0.1
10.6	11.5	0.1
10.6	11.6	0.1
10.6	11.7	0.1
10.4	10.5	0.2
10.6	10.8	0.2
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10.9	11.8	0.2
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10.8	11.3	0.2
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10.9	12.2	0.2
10.9	12.3	0.2
10.9	12.4	0.2
10.9	12.5	0.2
10.9	12.6	0.2
10.9	12.7	0.2
10.9	12.8	0.2
10.9	12.9	0.2
10.9	13.0	0.2

mauchi R., Morimoto N., Kousai K., Umemura M., 2012, AJ, 143, 79