

Cosmic web correlations with the shape, spin and peculiar velocities of dark matter halos

Jaime E. Forero-Romero¹, Sergio Contreras², Nelson Padilla²

¹ *Uni A* ² *Uni B*

4 February 2014

ABSTRACT

We study the alignment of DM halos with the cosmic web as described by the tidal and velocity shear fields. We focus on the alignment of halos shape, spin and peculiar velocities. We use a cosmological N-body simulation that allows to study dark matter halos spanning four orders of magnitude in mass ($(10^{10}\text{--}10^{14}) h^{-1}M_{\odot}$) and spatial scales of $(0.5\text{--}1.0)h^{-1}\text{Mpc}$ to define the cosmic web. After varying the numerical parameters in our experiments we find that the tidal field provides a robust description for the different kinds of alignments. We measure a strong here is a strong alignment of the peculiar velocity field and the halo shape with the cosmic web defined by the tidal field. The results for the spin alignment are less significant. The shape alignment is stronger for halos that are either massive ($M > 10^{12}h^{-1}M_{\odot}$), have large or are located in high density regions.

Key words: methods: N-body simulations, galaxies: haloes, cosmology: theory, dark matter, large-scale structure of Universe

1 INTRODUCTION

The formation of large scale structure in the Universe proceeds by gravitational instability. This is an intrinsic anisotropic process. [Pancake formation ...]

Numerical simulations allow the exploration of a deep nonlinear regime, following the formation of virialized dark matter halos, which host observable galaxies. Some characteristics of these halos have to be linked to the environment in which they formed. One way to study the connection halo-environment is to quantify the degree of correlation between different halo properties and their surrounding density and velocity fields.

Among the properties that are of interest we find the shape, spin and peculiar velocity. [Relevance of dark matter halo shapes ...] [Relevance of halo spin ...] [Relevance of peculiar velocities]

There is large tradition of theoretical precedents for the measurements of shape and spin alignment. In this paper we review most of these studies and offer our own study with complementary numerical techniques and simulations. [Brief summary of the appendix ...]

We also present the results for the alignment of peculiar velocities with large scale structure. Such a study has received less attention in the literature. However, we consider that it presents relevant insights for the study of cosmic flows.

The structure of this paper is the following. In §2 we

present the theoretical antecedents for the alignment studies we present in this paper. In §3 we present the N-body cosmological simulation and halo catalogs next to describe in §4 the two web-finding algorithms we use. In §6 we present our main results about the alignment of shape, spin and peculiar velocities with respect to the cosmic web. We continue with a discussion of these results in §7 paying special attention to resolution effects and the conditions that might drive the strong alignment signals observed in the precedent section. In §8 we present our conclusions.

2 THEORETICAL CONSIDERATIONS: NOTATION AND PRECEDENTS

There is abundant literature on the issue of shape and angular momentum alignment of dark matter haloes with respect to the cosmic web. In this paper we focus our attention on results published during the last decade that have made use of large N-body dark matter only cosmological simulations. There are many works that have addressed this problem using observational data from large surveys such as the Sloan Digital Sky Survey (SDSS), however we choose to narrow our discussion to simulation based studies which are prone to comparison.

Out of the three alignments that we study in this paper -shape, angular momentum and peculiar velocity- only the first two have received wide attention in the literature,

Author	Web Method	Halo Finder	Major Axis	Correlation
SHAPE				

being angular momentum the most popular with twice the number of studies for shape alignment.

These alignments are often measured from the distribution of the $\mu = |\cos \theta|$ where θ is the angle between the two axes of interest. This is often directly measured as the absolute value of the dot product between the two unit vectors along the directions being tested, for instance in the case of angular momentum one would compute $\mu = |\hat{j} \cdot \hat{n}|$. In the case of shape alignments the major axis is the chosen direction to compare against the cosmic web.

For an isotropic distribution of the vector around the direction defined by \hat{n} the μ distribution, ranging between 0 and 1, should be flat and its mean value should be $\langle \mu \rangle = 0.5$. If a distribution is biased towards 1 ($\langle \mu \rangle > 0.5$) we call this an statistical alignment along \hat{n} , while in the case of a bias towards 0 ($\langle \mu \rangle < 0.5$) we talk about an anti-alignment, meaning an perpendicular orientations with respect to the \hat{n} direction.

a Trowland et al. (2013) presented a parameterization for the μ distribution in the case of angular momentum alignment based on theoretical considerations by Lee et al. (2005) (Eq. 6 in the Appendix A). Under this parameterization it was found a unique correspondence between the full shape of the $|\mu|$ distribution and its average. In this paper we use that result to only present the results for the average $\langle |\mu| \rangle$.

Table 1 and Table 2 summarize recent results found in the literature for shape and angular momentum alignment. The Appendix A includes a detailed description of the definitions, algorithms and simulations used in each one of these studies. In these tables the first column describe the reference, the second column summarizes the web finding method with a single name, the third scale associates an spatial scale to the web finding methods, in most cases it corresponds to the grid size or smoothing scale used to interpolate the underlying matter density or velocity field; The fourth column indicates along which web element (filament or wall) was measured the alignment; the fifth column indicates the strenght of the alignment/anti-alignment, $++/--$ indicate a strong alignment/anti-alignment while $+/-$ indicate a weaker signal; the last columns indicates whether the described signal is present within a defined range of halo mass.

These results can be summarized in three important points:

- The halo mass of $1 - 5 \times 10^{12} h^{-1} M_{\odot}$ is a threshold mass between behaviours of no-alignment, alignment or anti-alignment.
- Halo shape provides a strong alignment signal along filaments and sheets, more so for massive haloes.
- Halo spin tends to be oriented perpendicular to filaments and parallel to sheets, but it is a weaker than shape alignment.

3 N-BODY SIMULATION AND HALO FINDING

In this paper we use the Bolshoi simulation that follows the non-linear evolution of a dark matter density field on cosmo-

logical scales. The volume is a cubic box with $250h^{-1} \text{Mpc}$ on a side, the matter density field is sampled with 2048^3 particles. The cosmological parameters in the simulation correspond to the results inferred from WMAP5 data (Dunkley et al. 2009), which are also consistent with the more recent results of WMAP9 (Hinshaw et al. 2013). These parameters are $\Omega_m = 0.27$, $\Omega_{\Lambda} = 0.73$, $\sigma_8 = 0.82$, $n_s = 0.95$ and $h = 0.70$ for the matter density, cosmological constant, normalization of the power spectrum, the slope in the spectrum of the primordial matter fluctuation and the dimensionless Hubble constant. With this conditions the mass of each dark matter particle in the simulation corresponds to $m_p = 1.4 \times 10^8 h^{-1} M_{\odot}$. A more detailed description of the simulation can be found in (Klypin et al. 2011).

In this paper we use groups found with a Friends-Of-Friends (FOF) halo finder using a linking length of $b = 0.17$ times the mean interparticle separation. This choice translates into halos with a density of 570 times the mean density at $z = 0$. The measurements for the shape, angular momentum and peculiar velocity are done using the set of particles in each dark matter halo. The definition we use in this paper for the shape comes from the diagonalization of the reduced inertia tensor.

$$\mathcal{I}_{lm} = \sum_i \frac{x_{i,l} x_{i,m}}{R_i^2}, \quad (1)$$

where i is the particle index in the halo and l, m run over the three spatial indexes and $R_i^2 = x_{i,1}^2 + x_{i,2}^2 + x_{i,3}^2$, where the positions are measured with respect to the center of mass.

The spin is calculated as

$$\vec{J} = \sum_i m_p R_i \vec{v}_i, \quad (2)$$

where the velocities are also measured with respect to the center of mass velocity. Finally the peculiar velocity of a halo is computed as the center of mass velocity.

The halo data used in this paper is publicly available through the MultiDark database¹ and is thoroughly described in Riebe et al. (2013).

4 WEB FINDING ALGORITHMS

We use two algorithms to define the cosmic web in cosmological N-body simulations. Both are based on the same algorithmic principle, which determines locally a symmetric tensor which can be diagonalized to yield three real eigenvalues $\lambda_1 > \lambda_2 > \lambda_3$ and their corresponding eigenvectors \mathbf{e}_1 , \mathbf{e}_2 and \mathbf{e}_3 . This allows for a local classification into one of the following four web types: void, sheet, filament and peak depending on the number of eigenvalues larger than a given threshold λ_{th} is 3, 2, 1 or 0, respectively.

We use two different symmetric tensors. The first is the shear tensor, defined as the Hessian of the gravitational potential, normalized in such a way as to be dimensionless:

$$T_{\alpha\beta} = \frac{\partial^2 \phi}{\partial r_{\alpha} \partial r_{\beta}}, \quad (3)$$

¹ <http://www.multidark.org/MultiDark/>

Author	Web Method	Spatial Scale	Along	Alignment	Mass dependence
Forero-Romero et al. (2014)	T-Web	$0.5 - 1 h^{-1} \text{Mpc}$	filament wall	++ --	all masses $h^{-1} \text{M}_\odot$ all masses
Forero-Romero et al. (2014)	Vp-Web	$0.5 - 1 h^{-1} \text{Mpc}$	filament filament wall	? ? ?	$> 10^{12} h^{-1} \text{M}_\odot$ $< 10^{12} h^{-1} \text{M}_\odot$ all masses
Libeskind et al. (2013)	V-Web	$1 h^{-1} \text{Mpc}$	filament filament wall	++ + ++	$> 10^{12} h^{-1} \text{M}_\odot$ $< 10^{12} h^{-1} \text{M}_\odot$ all masses
Zhang et al. (2009)	Hessian density field	$2.1 h^{-1} \text{Mpc}$	filament filament	++ +	$> 10^{12} h^{-1} \text{M}_\odot$ $< 10^{12} h^{-1} \text{M}_\odot$
Aragón-Calvo et al. (2007)	Hessian density field	- - - -	wall wall filament filament	++ + ++ +	$> 10^{12} h^{-1} \text{M}_\odot$ $< 10^{12} h^{-1} \text{M}_\odot$ $> 10^{12} h^{-1} \text{M}_\odot$ $< 10^{12} h^{-1} \text{M}_\odot$

Table 1. Shape alignment with the cosmic web. Summary of theoretical results provided by methods comparable to ours.

Author	Web Method	Spatial Scale ($h^{-1} \text{Mpc}$)	Along	Alignment	Mass dependence
Forero-Romero et al. (2014)	T-Web	$0.5 - 1$	filament filament wall wall	none - none -	$< 10^{12} h^{-1} \text{M}_\odot$ $> 10^{12} h^{-1} \text{M}_\odot$ $< 10^{12} h^{-1} \text{M}_\odot$ $> 10^{12} h^{-1} \text{M}_\odot$
Forero-Romero et al. (2014)	Vp-Web	$0.5 - 1$	filament filament wall wall	none none none +	$< 10^{12} h^{-1} \text{M}_\odot$ $> 10^{12} h^{-1} \text{M}_\odot$ $< 10^{12} h^{-1} \text{M}_\odot$ $> 10^{12} h^{-1} \text{M}_\odot$
Libeskind et al. (2013)	V-Web	1	filament filament wall	- + ++	$> 10^{12} h^{-1} \text{M}_\odot$ $< 10^{12} h^{-1} \text{M}_\odot$ all masses
Trowland et al. (2013)	Hessian density	$2 - 5$	filament filament	- +	$> 5 \times 10^{12} h^{-1} \text{M}_\odot$ $< 5 \times 10^{12} h^{-1} \text{M}_\odot$
Codis et al. (2012)	Morse Theory & T-Web	$1 - 5$	filament filament wall	-- ++ ++	$> 10^{12} h^{-1} \text{M}_\odot$ $< 10^{12} h^{-1} \text{M}_\odot$ all masses
Zhang et al. (2009)	Hessian density	2.1	filament filament	++ --	if anticorrelated with shape if correlated with shape
Aragón-Calvo et al. (2007)	Hessian density	- - - -	wall wall filament filament	++ + - +	$> 10^{12} h^{-1} \text{M}_\odot$ $< 10^{12} h^{-1} \text{M}_\odot$ $> 10^{12} h^{-1} \text{M}_\odot$ $< 10^{12} h^{-1} \text{M}_\odot$
Hahn et al. (2007)	Tidal Web	2.1 2.1	filament wall	- ++	none stronger $> 10^{12} h^{-1} \text{M}_\odot$

Table 2. Spin alignment with the cosmic web. Summary of theoretical results provided by methods comparable to ours.

where ϕ is the gravitational potential rescaled by a factor $4\pi G \bar{\rho} = 3/2 \Omega_m H_0^2$ in such a way that the Poisson equation can be written as $\nabla^2 \phi = \delta$, where δ is the matter overdensity, $\bar{\rho}$ is the average matter density, H_0 is the Hubble constant at present time and Ω_m is the matter density parameter. A detailed presentation of this algorithm can be found in (Forero-Romero et al. 2009).

The second tensor is the velocity shear:

$$\Sigma_{\alpha\beta} = -\frac{1}{2H_0} \left(\frac{\partial v_\alpha}{\partial r_\beta} + \frac{\partial v_\beta}{\partial r_\alpha} \right), \quad (4)$$

where v_α correspond to the components of the peculiar co-moving velocities. With this definition the trace of the shear tensor is minus the divergence of the velocity field normal-

ized by the Hubble constant $-\nabla \cdot v/H_0$. A detailed description of this algorithm can be found in (Hoffman et al. 2012).

4.1 Numerical considerations

In this paper we compute the cosmic web on cubic grids of two different resolutions 256^3 and 512^3 . For the T-Web we interpolate first the matter density field using a Cloud-In-Cell (CIC) scheme. Then we smooth using a Gaussian kernel with a spatial variance equal to the size of one grid cell. This smoothed matter density field is transformed into Fourier space to solve the Poisson equation and find the gravitational potential ϕ . The Hessian is computed a finite differences method. Finally, the eigenvalues and eigenvectors are computed on each grid point.

For the V-Web we interpolate first the momentum density field over a grid using the CIC scheme and then apply a gaussian smoothing with a spatial variance of one grid cell. We use the matter density field, which is also CIC interpolated and gaussian smoothed, to normalize the momentum field. This ratio between the momentum and matter density field is what we consider as the velocity field to compute the shear tensor on each grid point. In this case we also compute the eigenvalues and eigenvectors on each grid point.

5 OUR NUMERICAL EXPERIMENTS

In this paper we use the data and the methods described above to perform three different kinds of measurements.

5.1 Preferential Alignment

The first measurement is a rough approximation to find out along which axes are halos aligned. We call this Preferential Alignment (PA).

We use the fact that for a given vector under study \hat{a} and the three eigenvectors the following identity holds

$$(\hat{s} \cdot \hat{e}_1)^2 + (\hat{s} \cdot \hat{e}_2)^2 + (\hat{s} \cdot \hat{e}_3)^2 = 1. \quad (5)$$

Using this we know that all halos can be splitted into three groups:

- (i) Halos with $(\hat{s} \cdot \hat{e}_1)^2 > (\hat{s} \cdot \hat{e}_2)^2$ and $(\hat{s} \cdot \hat{e}_1)^2 > (\hat{s} \cdot \hat{e}_3)^2$, which can be considered to aligned mostly along \hat{e}_1 .
- (ii) Halos with $(\hat{s} \cdot \hat{e}_2)^2 > (\hat{s} \cdot \hat{e}_1)^2$ and $(\hat{s} \cdot \hat{e}_2)^2 > (\hat{s} \cdot \hat{e}_3)^2$, which can be considered to aligned mostly along \hat{e}_2 .
- (iii) Halos with $(\hat{s} \cdot \hat{e}_3)^2 > (\hat{s} \cdot \hat{e}_1)^2$ and $(\hat{s} \cdot \hat{e}_3)^2 > (\hat{s} \cdot \hat{e}_2)^2$, which can be considered to aligned mostly along \hat{e}_3 .

If the halo population does not show any preferential alignment, then all the halos must be evenly distributed along these three populations. On the contrary, if there is more than one third of the halo population in one of these sets, then we can talk about a preferential alignment along one of the axis. However, this statistic does not give a precise answer on the degree of the alignment

5.2 Average Alignment Angle

We emphasize that all the quantities that we compute in this paper are independent of any threshold on the eigenvalues that could be used to make a web classification into web types. Instead, we focus on the alignments with respect to the eigenvectors regardless of the web type. In this context we recall that the eigenvector \hat{e}_1 is vector perpendicular to the plane defining a sheet and also a filament and that \hat{e}_3 is the vector that marks the direction of a filament and lies on the plane of a sheet. Therefore we focus on quantifying the degree of alignment along these two eigenvectors.

This experiment complements the results obtained by the PA statistic by computing the average and standard deviation of $|\langle \hat{s} \cdot \hat{e}_1 \rangle|$ and $|\langle \hat{s} \cdot \hat{e}_3 \rangle|$. We perform this tests in different populations split into different mass bins logarithmically spaced between $1 \times 10^9 h^{-1} M_\odot$ and $1 \times 10^{14} h^{-1} M_\odot$.

In a separate test we make the same measurements but this time splitting the halo sample by other properties such as: circularity, concentration, local matter density, spin and triaxiality. In this case we take the upper and lower 30% of the halos according to each property and measure the strenght of the alignment by the average value of $|\langle \hat{s} \cdot \hat{e}_1 \rangle|$ and $|\langle \hat{s} \cdot \hat{e}_3 \rangle|$.

6 RESULTS

6.1 Preferential Alignment

Figure 1 presents all the results for the preferential alignment (PA).

6.2 Angular Momentum Alignment

We now take a look at the alignment of the halo spin with the cosmic web. This characterization is done in terms of the angle between the angular momentum vector and the first/third eigenvector.

Figure 2 shows the results as a function of the halo mass. In the case of the V-web the trends can be summarized as follows. Low mass halos $10^{12} h^{-1} M_\odot$ do not show any preferential alignment with the cosmic web. Halos more massive than this threshold have their spin slightly perpendicular to the direction defined by the first eigenvector (defining sheets) and uncorrelated to the third eigenvector (filament direction). This translates into a weak tendency for the spin to lie on the sheet-plane.

In the case of the T-web there is a similar trend for low mass halos $10^{12} h^{-1} M_\odot$ not to have any preferential direction with respect to the cosmic web. High mass halos present a weak alignment with the first eigen vector and anti-alignment with the third eigenvector. Summarizing, the spin of massive halos tends to be slightly perpendicular to filaments and sheets.

6.3 Shape Alignment

The shape alignment shows stronger trends. Figure 3 presents the main results as a function of halo mass. In this case we characterize the shape by the direction determined by the major axis.

In the case of the V-web we have a clear anti-alignment

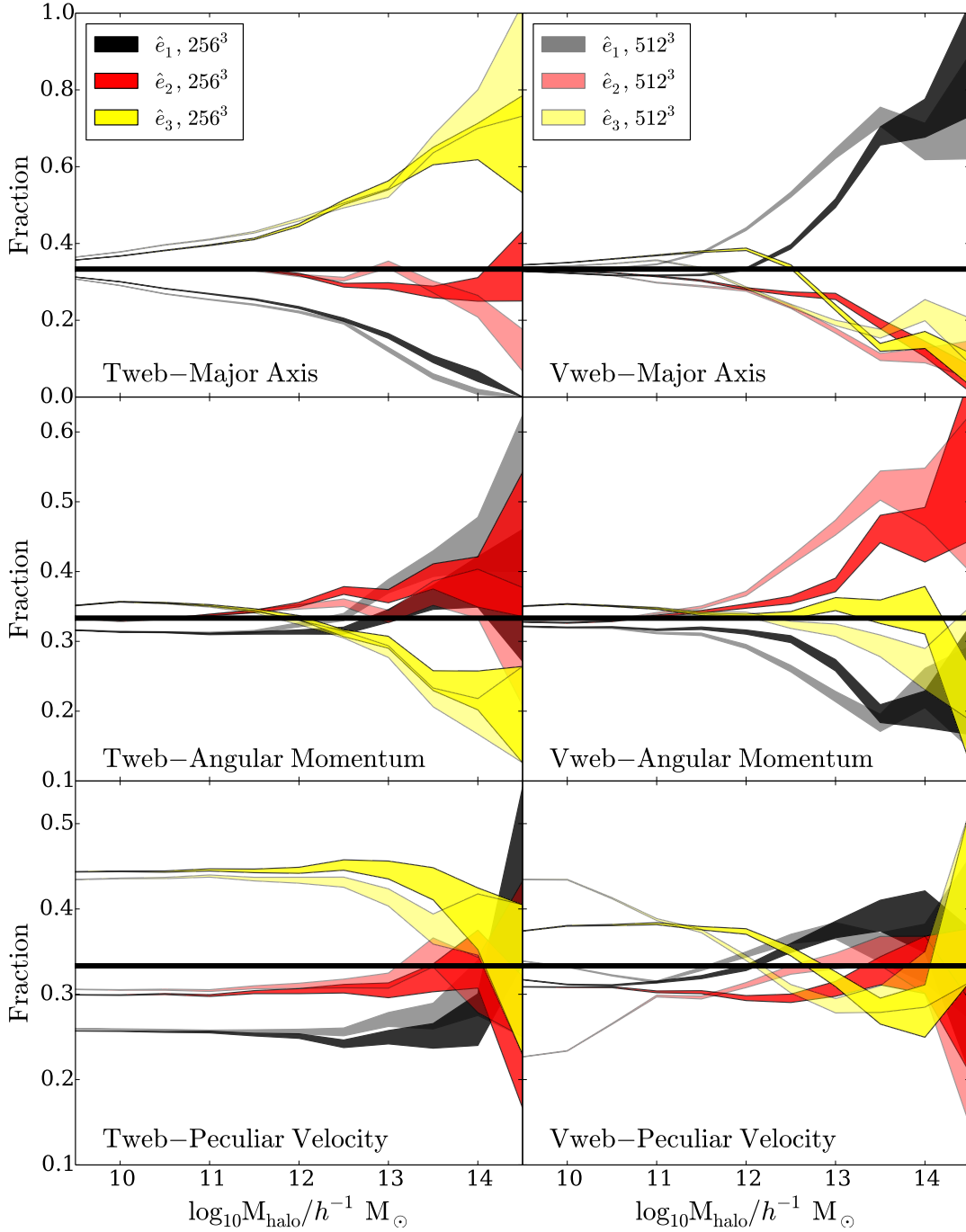


Figure 1. Fraction of halos in a mass bin that show a preferential alignment with respect to an eigenvector for the three properties studies in this paper: shape (major axis), angular momentum and peculiar velocity. Left column presents the results against the Tweb and the right the Vweb. Strong colors refer to 256^3 grid resolutions and lighter colors to a 512^3 grid.

with respect to both first and third eigenvectors for massive halos $> 10^{12} h^{-1} M_{\odot}$. This implies that a halo is typically perpendicular to filaments and lies along the plane defined by a sheet. For low mass halos $< 10^{12} h^{-1} M_{\odot}$ we do not find any correlation with the cosmic web.

The T-web shows alignment trends starting at masses of $10^{10} h^{-1} M_{\odot}$, two orders of magnitude less than the V-web. In this case we see an alignment along the first and third eigenvectors. It means that halos in filamentary envi-

ronments tend to be aligned with the filament, while halos in pancake environments are perpendicular to the sheet.

The most massive halos showing the strongest alignment trends are not usually located in regions classified as peaks where all the eigenvectors are positive, meaning that there is collapse/compressive stress in all directions.

In this situation we can interpret alignment both along the first and third eigenvector as a signature that most massive halos tend to be aligned midway between the first and

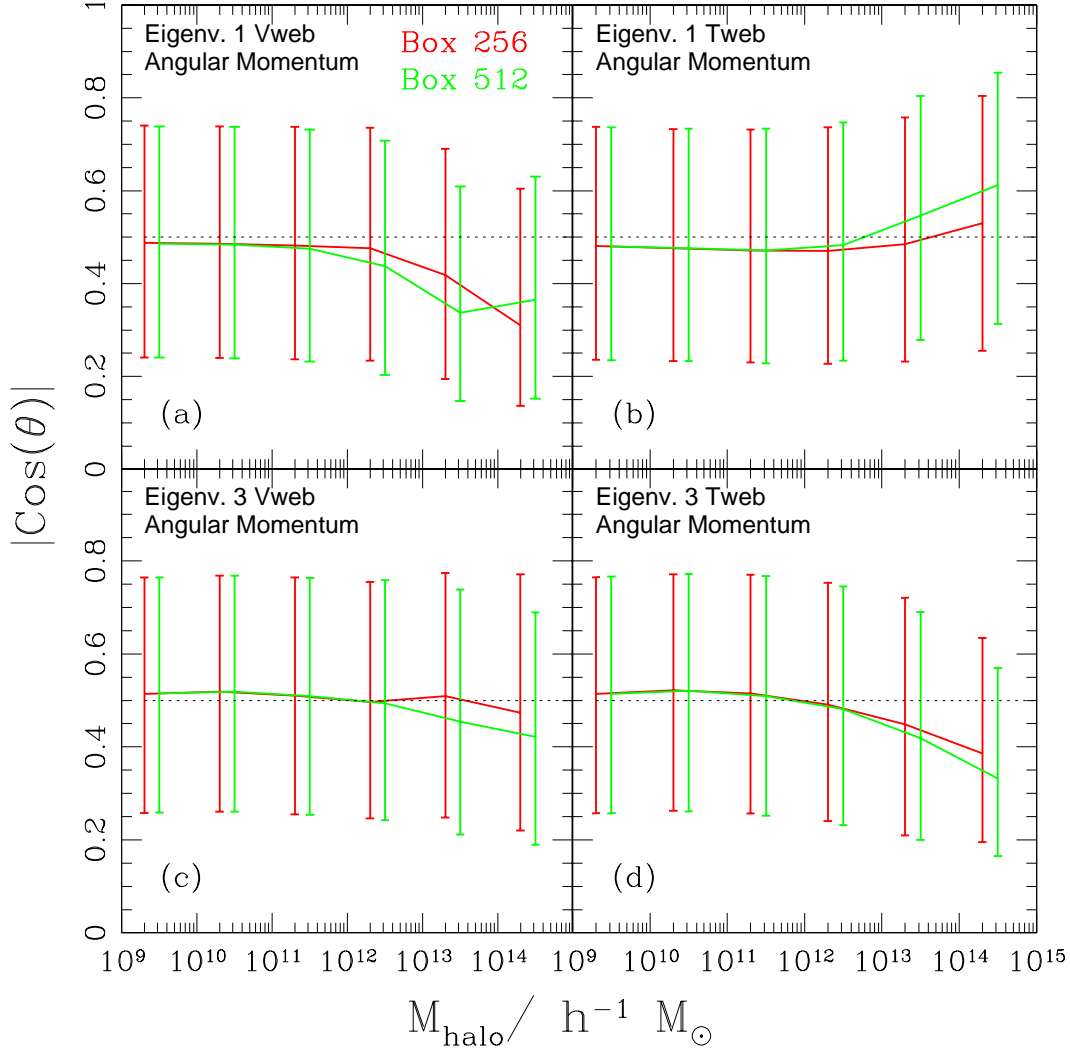


Figure 2. Median of $\cos \theta$ quantifying the angular momentum alignment for the Vweb (left) and the Tweb (right) at two different grid resolutions. In the upper (lower) panels the angle θ is measured between the first (third) eigenvector and the angular momentum vector.

second eigenvector, lying along the plane defined by the second and third eigenvectors.

6.4 Peculiar velocity Alignment

Possible conclusion. Prolateness follows concentration. In the Tweb Angular momentum alignment is influenced by spin at higher masses. In the Vweb angular momentum alignment is influenced by all factors. In the peculiar velocity alignment only spin seems to have an effect.

6.5 What drives the alignment

We wish to understand what other selection criteria on halo properties can produce a stronger local alignment for the shape, spin and peculiar velocities. We split the halo population into low and high mass halos imposing a cut at $M_{\text{halo}} = 10^{11} h^{-1} M_{\odot}$. This allows us to have robust statistics on the high mass end. We have also computed these results

for a cut at $M_{\text{halo}} = 10^{12} h^{-1} M_{\odot}$ and checked that the results we report below are not affected by this change.

For each mass interval we perform cuts in the following properties: spin, concentration, triaxiality, circularity and density. We measure the web alignments in two sets, each one including the 30% of halos in the lower/higher end of the corresponding property.

7 DISCUSSION

We do not find a strong signal for the alignment of spin and the cosmic web. A significant correlation signal has been recently measured between the spin and the cosmic vorticity (Laigle et al. 2013). This two facts are consistent with spin being produced by the non-symmetric part of the shear and tidal tensors.

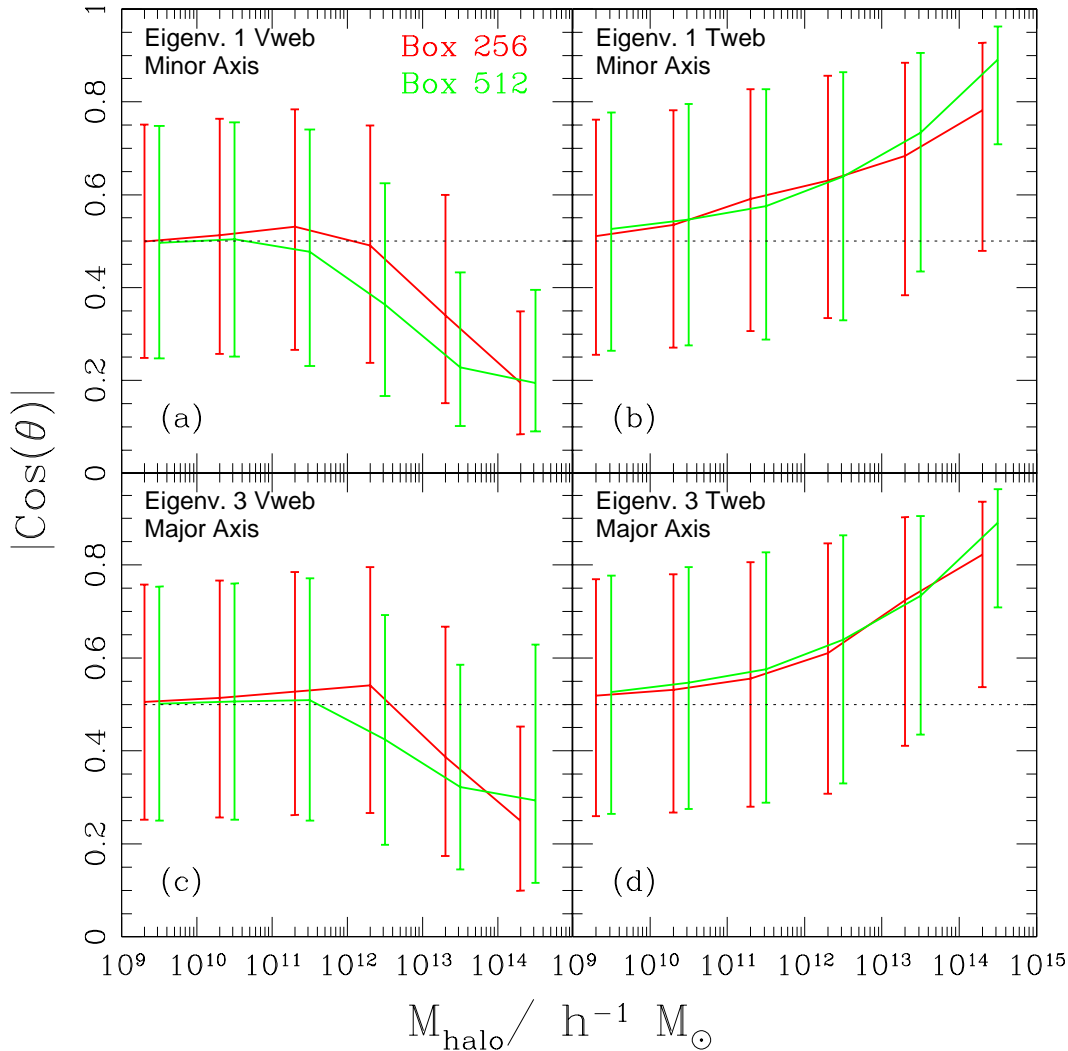


Figure 3. Median of $\cos \theta$ quantifying the shape alignment for the Vweb (left) and the Tweb (right) at two different grid resolutions. In the upper (lower) panels the angle θ is measured between the halo minor (major) axis and the first (third) eigenvector.

7.1 Interweb Alignment

In the first place we wish to directly compare the results of the two algorithms. For the two algorithms we have the information for the eigenvectors and the eigenvalues on exactly the same positions defined by the grids. This allows us to compute the pair-wise alignment between the eigenvectors in the two web finders.

We are mainly interested in the alignment of the web with the halos. Therefore we restrict our analysis to the grid cells that are occupied by halos. If we had decided to perform this kind of analysis on all the grid cells, the statistics would be dominated by the void regions, because they dominate in number the fraction of cells in the simulation.

Figure 6 shows the results of the inter-web alignment as a function of halo mass. It clearly stands out that the eigenvectors are aligned for low mass halos but are close perpendicular to each other at the location of the most massive halos. There is a transition scale around $(10^{11.5} - 10^{12.5})h^{-1}M_{\odot}$ depending on the grid resolution. The coarse

grid (256^3) shows the transition at higher masses than the fine grid (512^3).

For both resolutions, the halos more massive than $10^{13}h^{-1}M_{\odot}$ are located in regions where the fastest momentum-weighted collapse direction (defined by the Vweb) is perpendicular to the direction where the tidal compression is the highest.

The fact that this trend is less evident at low resolution and low masses, hints that the flip is actually related to the overdensity values, i.e. to the degree of non-linearity, in the corresponding spatial scales.

8 CONCLUSIONS

ACKNOWLEDGMENTS

APPENDIX A. DETAILED DESCRIPTION OF PREVIOUS THEORETICAL RESULTS

The tidal connection between voids in simulations (Platen et al. 2008), the alignment of observed galaxies with the in-

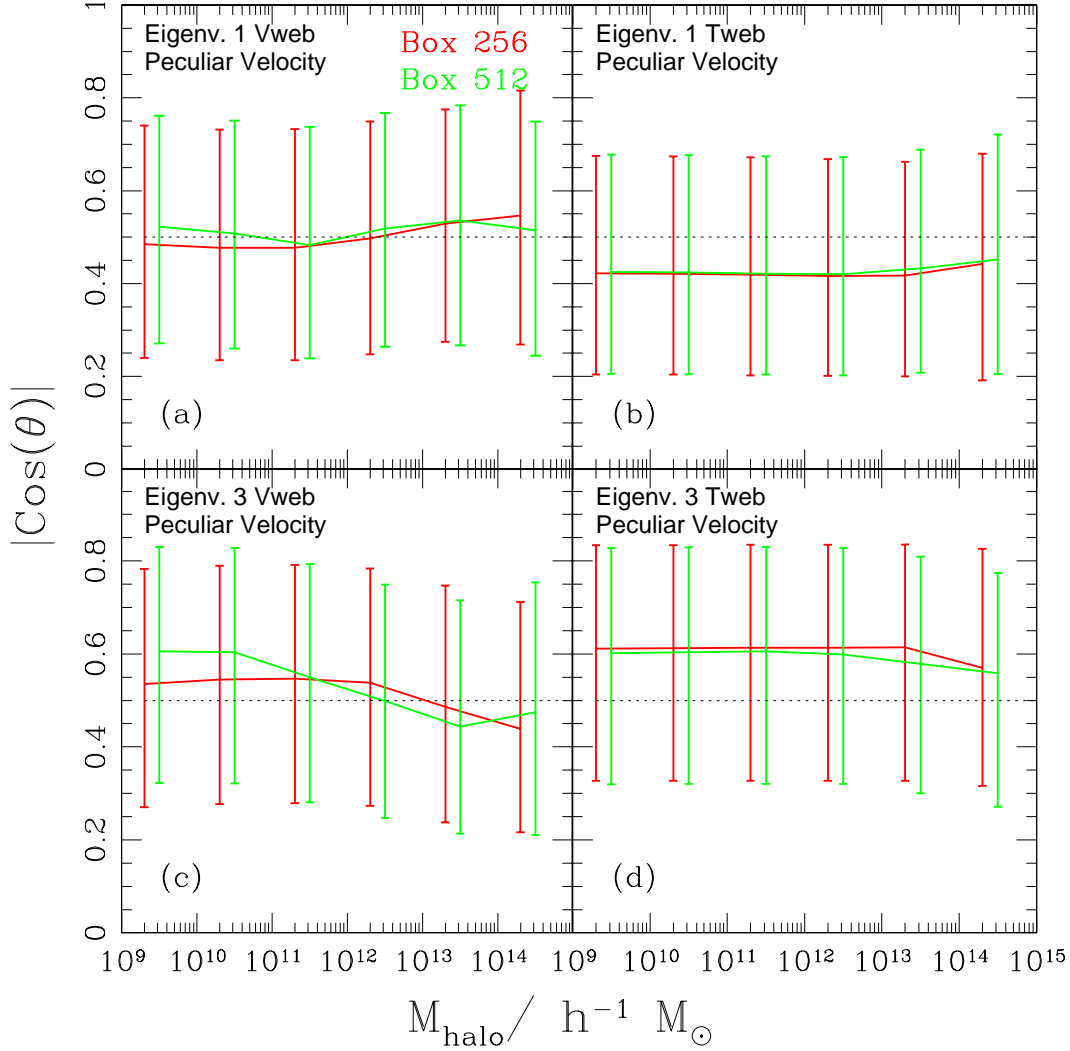


Figure 4. Median of $\cos \theta$ quantifying the peculiar velocity alignment with the Vweb (left) and the Tweb (right) at two different grid resolutions. In the upper (lower) panels the angle θ is measured between the first (third) eigenvalue and a halo's peculiar velocity.

ferred tidal field (Lee & Erdogdu 2007; Jones et al. 2010). There are other kind of alignment statistics based on modifications of the correlation (Paz et al. 2008; Faltenbacher et al. 2009) that go beyond a local computation and therefore are not reviewed in detail in this paper.

- (Libeskind et al. 2013) Studies shape and spin alignment with the cosmic web defined by the velocity shear tensor method described in this paper.

The simulation has 2048^3 particles in a box of $250h^{-1}\text{Mpc}$, corresponding to a particles mass of $1.3 \times 10^8 h^{-1} M_\odot$. Halos are found using a FOF halo finder with $b = 0.17$. The catalog only include halos more massiven than $3 \times 10^9 h^{-1} M_\odot$.

The shape is defined from the reduced inertia tensor. Results are reported for three mass bins $M_{\text{vir}} < 10^{11.5} h^{-1} M_\odot$, $10^{11.5} < M_{\text{vir}} < 10^{12.5} h^{-1} M_\odot$ and $M_{\text{vir}} > 10^{12.5} h^{-1} M_\odot$.

The identification of the cosmic web is done on a grid of 256^3 with a gaussian smoothing of $\sim 1h^{-1}\text{Mpc}$ over the velocity field. We define this velocity field as the ratio of a momentum density to the matter density field. The details of this approach are described in the Appendix .

The alignmet signal of the spin is very weak while the shape alignment signal is very strong. The shape alignment is such that the eigenvector corresponding to the smallest eigenvalue is aligned with the major axis. This effect is stronger for more massive halos. In other words the major axis of a halo is aligned with a filament, and lies on the plane that define a sheet. However the spin is antialigned with the filament for massive halos and weakly aligned for low mass halos.

- (Trowland et al. 2013)

The simulation is the millennium run, which has 2160^3 particles in a volume of $500h^{-1}\text{Mpc}$ on a side. This corresponds to a particle mass of $8.6 \times 10^8 h^{-1} M_\odot$. The catalog uses both halos and subhalos which were identified with SUBFIND. Only halos with more than 500 particles were kept to get a robust computation for the spin. For each halo the spin is defined as the sum of the angular momentum of each particle with respect to the center of mass.

The method to define the filamentary structure is based on the eigenvalues of the hessian of the density. However the analysis are reported on a box of $300h^{-1}\text{Mpc}$ on a side. Four

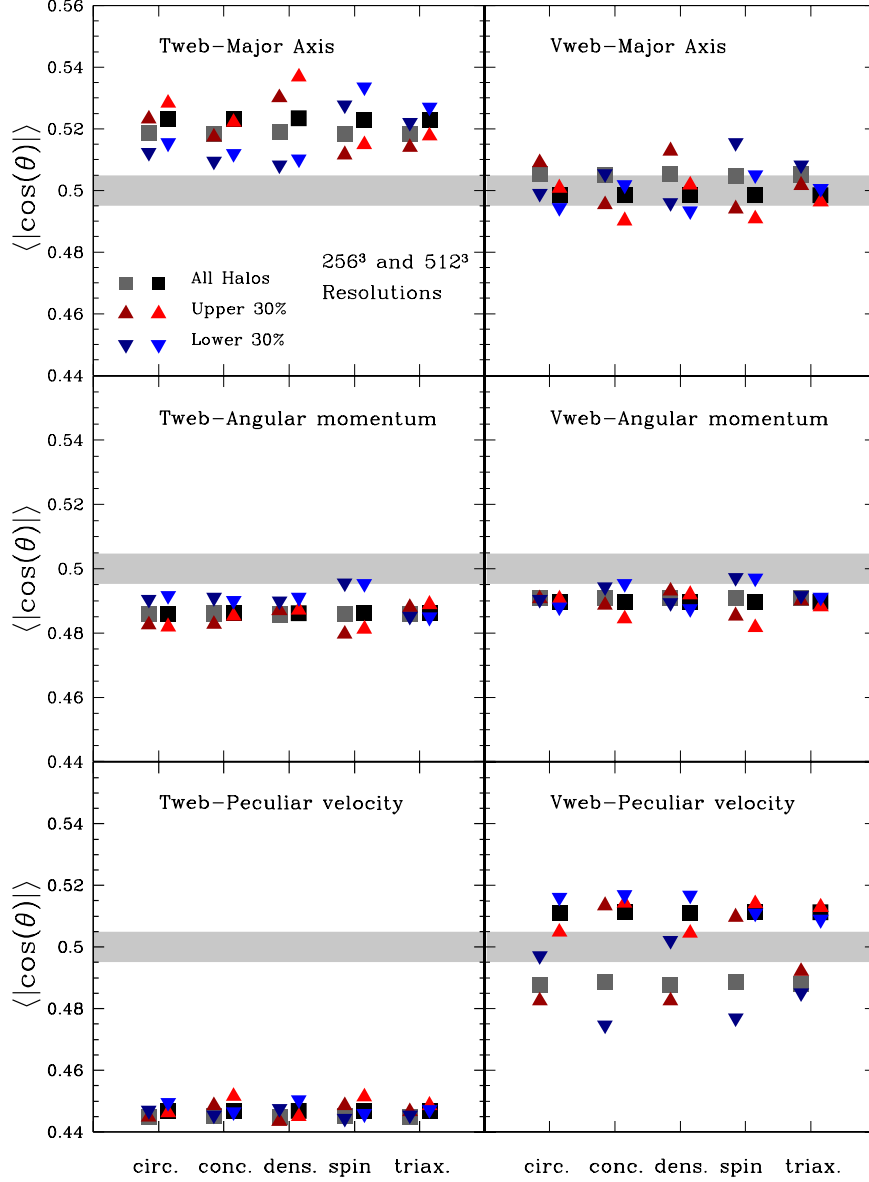


Figure 5. The alignment is always taken with respect to the first eigenvector.

different gaussian smoothing scales are used: 2.0, 3.0 and $5.0h^{-1}\text{Mpc}$.

The authors report a slight alignment signal of spin against the principal filament axis. By fitting the following functional form to the $\cos(\theta)$ distribution

$$P(\cos \theta) = (1 - c) \sqrt{1 + \frac{c}{2}} \left[1 - c \left(1 - \frac{3}{2} \cos^2 \theta \right) \right]^{-3/2}, \quad (6)$$

they are able to quantify the degree of alignment ($c < 0$) or antialignment ($c > 0$). This parameterization is based on theoretical expectations of Tidal Torque Theory (TTT) (Lee et al. 2005). At $z = 0$, the reported value is $c = 0.035 \pm 0.004$, where the uncertainty was calculated using bootstrapping and resampling.

When the halo sample is divided between low mass and high mass halos with a transition scale $M_* = 5.9 \times 10^{12} M_\odot$,

there is an anti-alignment above this mass and an alignment below it.

- (Codis et al. 2012) Studies the alignment of the spin of dark matter halos relative to the surrounding large scale structure and to the tidal tensor eigenvalues.

They use a dark matter simulation with 4096^3 DM particles in a cubic periodic box of $2000h^{-1}\text{Mpc}$ on a side, which corresponds to a particle mass of $7.7 \times 10^9 M_\odot$. Halos are identified using a FoF algorithm with a linking length of 0.2 keeping all halos with more than 40 particles, which sets the minimum halo mass to be $3 \times 10^{11} M_\odot$. In their work the particles were sampled on a 2048^3 grid and the density field was smoothed with a gaussian filter over a scale of $5h^{-1}\text{Mpc}$ corresponding to a mass of 1.9×10^{14} . The skeleton was computed over 6^3 overlapping subcubes and then reconnected.

The filament finder algorithm is based on Morse theory

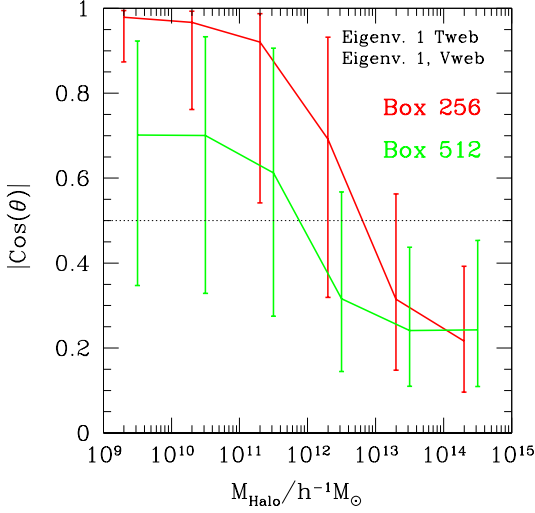


Figure 6. Median of the interweb alignment for the two grid resolutions as a function of the dark matter mass corresponding to the halos where the measurement was made. The error bars indicate the lower and upper quartiles (Figure in color in the web version).

and defines a Skeleton to be the set of critical lines joining the maxima of the density field through saddle points following the gradient. They also compute the hessian of the potential over the smoothed density field to get their eigenvectors.

The spin of the halo is defined as $m_p \sum_i (r_i - \bar{r}) \times (v_i - \bar{v})$ where \bar{r} is the center of mass of the halo and \bar{v} is the average velocity.

They measure the spin alignment with each one of the eigenvectors. With respect to the minor eigenvector e_1 there is anti-alignment for masses $M > 5 \times 10^{12} M_\odot$ and alignment for masses $< 5 \times 10^{12} M_\odot$. With respect to the intermediate eigenvector e_2 there is a strong alignment at high masses and no alignment for low masses, with respect to the major eigenvector e_3 there is an anti-alignment signal at all masses. The results from the Skeleton algorithm are in perfect agreement with the results from the Tidal web. The transitional mass is weakly dependent on the smoothing scale, varying between $1 - 5 \times 10^{12} h^{-1} M_\odot$ for smoothing scales between $1.0 - 5.0 h^{-1} \text{Mpc}$.

- (Zhang et al. 2009) Study the spin and shape alignment against filaments.

They use a dark matter simulation with 1024^3 DM particles in a periodic box of $100 h^{-1} \text{Mpc}$ on a side. The particle mass is $6.92 \times 10^7 h^{-1} M_\odot$. Dark matter haloes are found using a FOF algorithm with a linking length of 0.2 times the interparticle distance. Only halos with more than 500 particles are retained for further analysis.

The angular momentum is measured with positions respect to the center of mass and the shape is determined using the non-normalized moment of inertia tensor.

The environment is found using the hessian of the density. The density field was interpolated over a 1024^3 grid and then smoothed with a Gaussian filter of scale $R_s = 2.1 h^{-1} \text{Mpc}$. There are two methods to define the direction of a filament. The first method uses the eigenvalues of the hessian den-

sity, however they define the filament direction with the eigenvector corresponding the single positive eigenvalue of the hessian. The second method used a line that connects the two terminal halos in a filament segment.

The characterization of the alignment with the $\cos(\theta)$ statistic. For the method that uses the eigenvectors, they find that the strength of the spin alignment decreases with halo mass. For the shape they study the alignment of the major axis with the filament. They find an alignment signal in all mass bins, with a stronger effect for more massive halos.

In a final experiment they measure the spin alignment in four different samples that separated by the strength of the shape alignment. They find that halos anticorrelated in shape, show a strong spin correlation going to the extreme where there is a strong spin anticorrelation for halos with a strong shape correlation. This means that the halos with strong spin alignment are not the same halos showing strong shape alignment.

- (Aragón-Calvo et al. 2007) The method is the Multi-scale Morphology Filter which is based on the Hessian matrix of the density field, where the density field is computed from the particle distribution using a Delaunay tessellation field estimator (DTFE), which is self-adaptive. This allows them to identify clusters, filaments and walls.

The simulation has 512^3 particles in a cubic box of $150 h^{-1} \text{Mpc}$. The mass per particle is $2 \times 10^9 h^{-1} M_\odot$. Halo identification is done with the HOP algorithm. They keep halos with more than 50 particles and less than 5000, defining a mass range of $1 - 100 \times 10^{11} h^{-1} M_\odot$.

The principal axes of each halo are computed from the non-normalized inertia tensor. The inertia tensor and the angular momentum are computed with respect to the center of mass of the halo.

They compute two angles, one with respect to the direction defining the filaments and the other the walls. Their results make a distinction between halos of more massive and less massive than $10^{12} h^{-1} M_\odot$.

The halo spins tend to lie on the plane of the wall. This is stronger for massive halos. The effect for filaments is weaker: low mass halos tend to lie parallel to their host filament, while high mass halos tend to be perpendicular.

There is a very strong effect for the principal axes of halos in filaments to be strongly correlated with the direction of the filaments. The minor axis tend to be perpendicular to the filament. This effect is stronger for larger halos.

The effect in walls is less strong, but still the minor axis tend to lie perpendicular to the wall, while the other axis then to lie over the wall. The effect is stronger for massive halos.

They find that spins and shapes of dark matter halos are significantly correlated with each other and with the orientation of their host structures.

- (Hahn et al. 2007) The method is the Tweb. They use three simulations each of 512^3 particles, with sizes $L_1 = 45 h^{-1} \text{Mpc}$, $L_2 = 90 h^{-1} \text{Mpc}$ and $L_3 = 180 h^{-1} \text{Mpc}$, this corresponds to particle masses of 4.7, 38.0, $300 \times 10^7 h^{-1} M_\odot$. The normalization is $\sigma_8 = 0.9$. Halo identification was done with a FOF algorithm with 0.2 times the interparticle distance. They consider halos of at least 300 particles.

The web is obtained for a grid of 1024^3 cells, the density field is obtained with a CIC interpolation and smoothed

using a Gaussian Kernel. In the rest of the paper all the results correspond to a smoothing scale of $R_s = 2.1h^{-1}\text{Mpc}$.

They Report on the angle between the halo angular momentum vector and the eigenvector corresponding to perpendicular directions to the sheets and the direction of the filaments. This is divided in two halo populations: $5 \times 10^{10} - 1.0 \times 10^{12}$ and $> 10^{12}$. There is a weak anti-alignment in the case of the filaments and a stronger anti-alignment in the case of the sheets. For the sheets the effect is stronger for the massive bin. In the filaments the alignment is weak regardless of the mass. They do not report any other significant statistic, but recognize that they suffer from small-number statistics in voids).

They do not see any strong dependance of the environment in the shape. They do not measure the shape alignment.

- (Brunino et al. 2007) Uses a DM only N-body simulation (the Millennium Run) to study the shape and angular momentum alignment with respect to large cosmological voids.

The voids are found with a spherical void finder. This algorithm first looks for potential void centers in the underdense regions of the simulation. These positions define the centers of spheres that are empty of halos above a given mass threshold. Only voids with radius larger than $10h^{-1}\text{Mpc}$ were considered for further analysis.

Halos were identified by linking particles with density exceeding 900 times the background density in order to keep the analysis of shape and angular momentum on the core of the halo. Only halos with masses in the range $8.6 \times 10^{11}h^{-1}\text{M}_\odot < M_h < 8.6 \times 10^{12}h^{-1}\text{M}_\odot$ were included analyzed to measure their angular momentum and the inertia tensor, both in a non-normalized way.

For the purpose of measuring alignments a second cut is imposed whereby only halos within a shell of radius $R_{\text{void}} < r < 1.05R_{\text{void}}$. The alignment is then measured between the vector connecting the center of the void to the halo in question and the vector of interest.

The results show that there is a strong alignment of halo shape in such a way that the major axis is parallel to the shell surrounding the void. In the case of angular momentum there is not a significant signal for alignment.

- (Basilakos et al. 2006) Use a cosmological SPH+N-body simulation to measure the alignment of cluster halos with their parent supercluster. For both the cluster halos and parent super-cluster they define the shape via the non-normalized inertia tensor. They find that strenght of the alignments increases with the degree of filamentarity of the supercluster.

The simulation has 2×512^3 particles in a box of side $500h^{-1}\text{Mpc}$. The dark matter particle mass is $6.6 \times 10^{10}h^{-1}\text{M}_\odot$ while for SPH particles is $1.2 \times 10^{10}h^{-1}\text{M}_\odot$. The halo finding is done with a FOF algorithm with a linking length of 0.17 and keep objects with more than 100 particles.

- (Lee & Pen 2002) Observational measurement for the alignment of galaxy spin axes with the local tidal shear field. For the measurement of shear, we have used the Point Source Catalog Redshift (PSCz) survey (a complete redshift survey from the IRAS Point Source Catalog) data. This was done down to a radial comoving distance of $\sim 150h^{-1}\text{Mpc}$.
- (Hatton & Ninin 2001) DM matter only simulation.

256^3 particles, $100h^{-1}\text{Mpc}$, particle mass. Look for alignment of angular momentum vectors with neighboring structures or other massive structures. They match halos using close pair statistics. They do not find any evidence for a statistically significant mutual alignment of halos on any scale.

REFERENCES

- Aragón-Calvo M. A., van de Weygaert R., Jones B. J. T., van der Hulst J. M., 2007, *ApJL*, 655, L5
- Basilakos S., Plionis M., Yepes G., Gottlöber S., Turchaninov V., 2006, *MNRAS*, 365, 539
- Brunino R., Trujillo I., Pearce F. R., Thomas P. A., 2007, *MNRAS*, 375, 184
- Codis S., Pichon C., Devriendt J., Slyz A., Pogosyan D., Dubois Y., Sousbie T., 2012, *MNRAS*, 427, 3320
- Dunkley J., Komatsu E., Nolte M. R., Spergel D. N., Larson D., Hinshaw G., Page L., Bennett C. L., Gold B., Jarosik N., Weiland J. L., Halpern M., Hill R. S., Kogut A., Limon M., Meyer S. S., Tucker G. S., Wollack E., Wright E. L., 2009, *ApJS*, 180, 306
- Faltenbacher A., Li C., White S. D. M., Jing Y.-P., Shu-DeMao Wang J., 2009, *Research in Astronomy and Astrophysics*, 9, 41
- Forero-Romero J. E., Hoffman Y., Gottlöber S., Klypin A., Yepes G., 2009, *MNRAS*, 396, 1815
- Hahn O., Carollo C. M., Porciani C., Dekel A., 2007, *MNRAS*, 381, 41
- Hatton S., Ninin S., 2001, *MNRAS*, 322, 576
- Hinshaw G., Larson D., Komatsu E., Spergel D. N., Bennett C. L., Dunkley J., Nolte M. R., Halpern M., Hill R. S., Odegard N., Page L., Smith K. M., 2013, *ApJS*, 208, 19
- Hoffman Y., Metuki O., Yepes G., Gottlöber S., Forero-Romero J. E., Libeskind N. I., Knebe A., 2012, *MNRAS*, 425, 2049
- Jones B. J. T., van de Weygaert R., Aragón-Calvo M. A., 2010, *MNRAS*, 408, 897
- Klypin A. A., Trujillo-Gomez S., Primack J., 2011, *ApJ*, 740, 102
- Laigle C., Pichon C., Codis S., Dubois Y., le Borgne D., Pogosyan D., Devriendt J., Peirani S., Prunet S., Rouberol S., Slyz A., Sousbie T., 2013, *ArXiv e-prints*
- Lee J., Erdogdu P., 2007, *ApJ*, 671, 1248
- Lee J., Kang X., Jing Y. P., 2005, *ApJL*, 629, L5
- Lee J., Pen U.-L., 2002, *ApJL*, 567, L111
- Libeskind N. I., Hoffman Y., Forero-Romero J., Gottlöber S., Knebe A., Steinmetz M., Klypin A., 2013, *MNRAS*, 428, 2489
- Paz D. J., Stasyszyn F., Padilla N. D., 2008, *MNRAS*, 389, 1127
- Platen E., van de Weygaert R., Jones B. J. T., 2008, *MNRAS*, 387, 128
- Riebe K., Partl A. M., Enke H., Forero-Romero J., Gottlöber S., Klypin A., Lemson G., Prada F., Primack J. R., Steinmetz M., Turchaninov V., 2013, *Astronomische Nachrichten*, 334, 691
- Trowland H. E., Lewis G. F., Bland-Hawthorn J., 2013, *ApJ*, 762, 72

Zhang Y., Yang X., Faltenbacher A., Springel V., Lin W.,
Wang H., 2009, ApJ, 706, 747