
Creating an Automated Classifier to Approximate the By-eye Selection of HAT Exoplanet Candidates

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Abstract

A proposal to use machine learning techniques to aid in the binary classification of data from the Hungarian Automated Telescope (HAT) arrays. The HAT arrays are designed to detect periodic dimmings of stars, potentially caused by orbiting planets around those stars. There are many signals, both physical and spurious, that can approximate planet signals, so it is important to weed out as many false positives as possible. After an automated pipeline removes probable false planets, a final manual by-eye selection is made to create the final set of planet candidates. We propose to train classifier models to approximate this by-eye selection with the hope of finding a sufficiently well-performing model that can stand as a proxy for or even entirely replace the manual portion of the selection pipeline.

1 Introduction

The field of exoplanets (“extra-solar planets”, or planets around other stars) has established itself as the hip new field in astrophysics. Since the discovery of the first exoplanets in the early 90’s [9], nearly 3000 exoplanets have been discovered [6], with potentially thousands more exoplanets remaining unconfirmed or thus far undiscovered in existent data. This explosion in exoplanet discovery has been fueled by the funding and construction of many exoplanet detection surveys. The majority of exoplanets have been found by the space-based *Kepler* telescope and its primary and “K2” surveys [3], [8]. Ground-based surveys, while not as sensitive to smaller planets as *Kepler* (due to atmospheric distortion of image quality and other problems associated with observing from the ground), are sufficiently sensitive to discover Jupiter-sized planets on short-period orbits. Such planets are called “hot Jupiters” because they are the same size as Jupiter but are sufficiently close in to the stars they orbit (much closer than Mercury is to our sun) that the star heats them up to several thousands of degrees Celsius.

Since ground-based surveys are much cheaper to run per area of sky monitored than space-based surveys, the majority of known hot Jupiters have been discovered by large-scale ground-based surveys. Of the few hundred hot Jupiters that have been found to date, the plurality (~100) have been found by Princeton’s own Hungarian Automated Telescope (HAT) collaboration, headed by Prof. Gaspar Bakos of the Department of Astrophysical Sciences. This collaboration runs two surveys. The longest-running of the two surveys is HATNet, a collection of five telescopes in Arizona and two telescopes in Hawaii [7]. The other survey is HATSouth, which consists of three locations with eight telescopes each: Chile, Namibia, and Australia [2]. A third survey, HATPI, which will consist of 63 telescopes in Chile, is currently being constructed [4]. The telescopes themselves are small as telescopes go: only about 10 cm across, they are more similar to camera lenses than the “normal” telescopes usually associated with astronomy. Figure 1 shows a picture of four of the HAT-South telescopes.

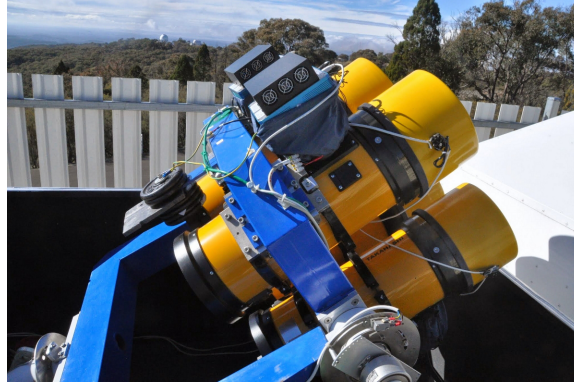


Figure 1: The four HATSouth telescopes at Siding Springs Observatory, Australia. Photo credit Gaspar Bakos. From [5].

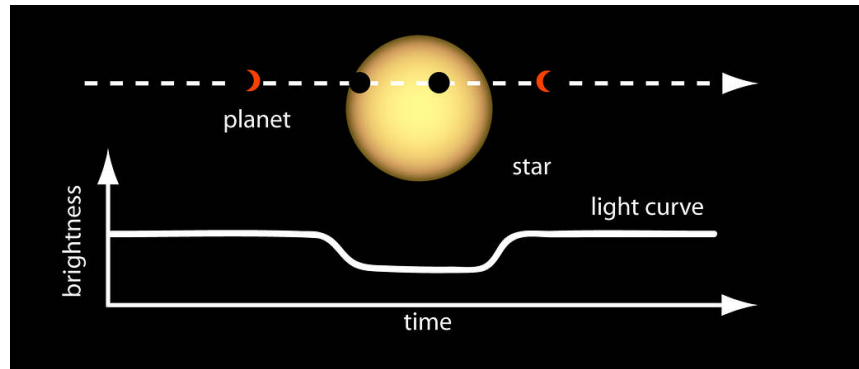


Figure 2: An example “light curve” (brightness as a function of time) from a star that hosts a transiting exoplanet. As the planet crosses in front of the star, the measured brightness from the star is less than the out-of-transit value. Partial eclipses as the planet is transitioning from being out-of-transit to fully in transit put a slope in the light curve between in-transit and out-of-transit magnitudes. Public domain image from [1].

The detection technique used by HAT*, *Kepler*, and many other exoplanet surveys is the “transit” technique. Planets are far too faint compared to their host stars to be able to detect directly around most stars. Instead, in the transit technique, the brightness of stars is monitored as a function of time (usually by taking periodic long-exposure photographs of large areas of the sky). When a planet crosses in front of the star it orbits, the amount of light detected from that star is decreased. This is shown in Figure 2. Transits are found in data using matched filtering (with either a box- or trapezoid-shaped filter, matching the signal from a transiting exoplanet), folding the data on a variety of periods and phases. When a star is found to have periodic dips in brightness, it is labeled as a potentially planet-hosting star. However, transiting planets are not the only source of periodic signals like that seen in Figure 2. Starspots rotating in and out of view also provide periodic dips in brightness and binary stars that eclipse each other in a grazing way (so the amount of light that goes missing during an eclipse is about the same as a planet) are two examples of astrophysical false positive signals. Various cuts are made to help ensure that a prospective planet is neither one of these false positives nor a spurious signal. The various parameters of the fit to the transit signal (depth, time of ingress, length of transit) are often good indicators to separate between planet and not. These cuts are automated. However, after the cuts, there are still false positives that are not obvious to the pipeline but are obvious to a trained human eye. Thus, the final step in the HAT pipeline is a manual, by-eye examination of the signal. Planet candidates that pass this step then are sent for follow up (and hopefully confirmation!) at bigger telescopes.

Other than this manual step, the vetting of possible planet candidates is completely automatic. This single manual step is quite labor intensive, and prevents a fully automatic characterization of planet detection efficiency, which is necessary for the calculation of statistics related to the occurrence rate of these planets. Thus, it would be nice to have an automated approximation to the by-eye work that has occurred. This is the purpose of this project.

2 Related Work

3 Data

The data consists of ~ 1000 examples of potential planet candidates that were decided, by eye, to be included as planet candidates, and $\sim 30,000$ examples of potential planet candidates that were decided to not be regarded as planet candidates. There are ~ 90 features, all continuous, for each potential planet candidate. These features include parameters and statistics related to the fitting of the brightness dips, parameters related to the measured planet size and orbital period, and properties of the host star (how bright it is in different colors, etc.). The data were extracted from the proprietary HAT database using a script provided by Joel Hartman, Research Scientist at Princeton University. The database itself has not reference currently.

4 Methods

Our first approach to this problem will be to investigate the simple binary classifiers that we have covered in class so far. These include Naive Bayes's, random forests, support vector machine (SVM), and logistic regression. With logistic regression will investigate two types of regularization using l_1 and l_2 penalties. We will also explore a set of feature selection methods such as mutual information.

Whilst our main goal is to predict which objects would be manually selected, we are also interested in running a set of unsupervised learning algorithms to view any hidden structure in the data set. These methods will include K-means and gaussian mixture models. This will be useful for potentially identifying new approaches to automating this process as well as potentially probing some new relationships.

5 Results

6 Discussion and Conclusion

Acknowledgments

We are grateful to Joel Hartman for providing a script to extract the necessary data from the HAT database. We are also grateful to the entire HAT team (PI: Gaspar Bakos) for their many dedicated years of constant observations and the huge pile of astronomical data they've collected.

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