

# Minimal Scotogenic (DM) models with (Dirac) neutrino masses

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Oct 11, 2019 - USP [PDF: <http://bit.ly/darkusp>]

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**Focus on**

[arXiv:1811.11927 \[PRD\]](https://arxiv.org/abs/1811.11927), [arXiv:1906.09685 \[PRD\]](https://arxiv.org/abs/1906.09685) and [arXiv:1909.09574](https://arxiv.org/abs/1909.09574)

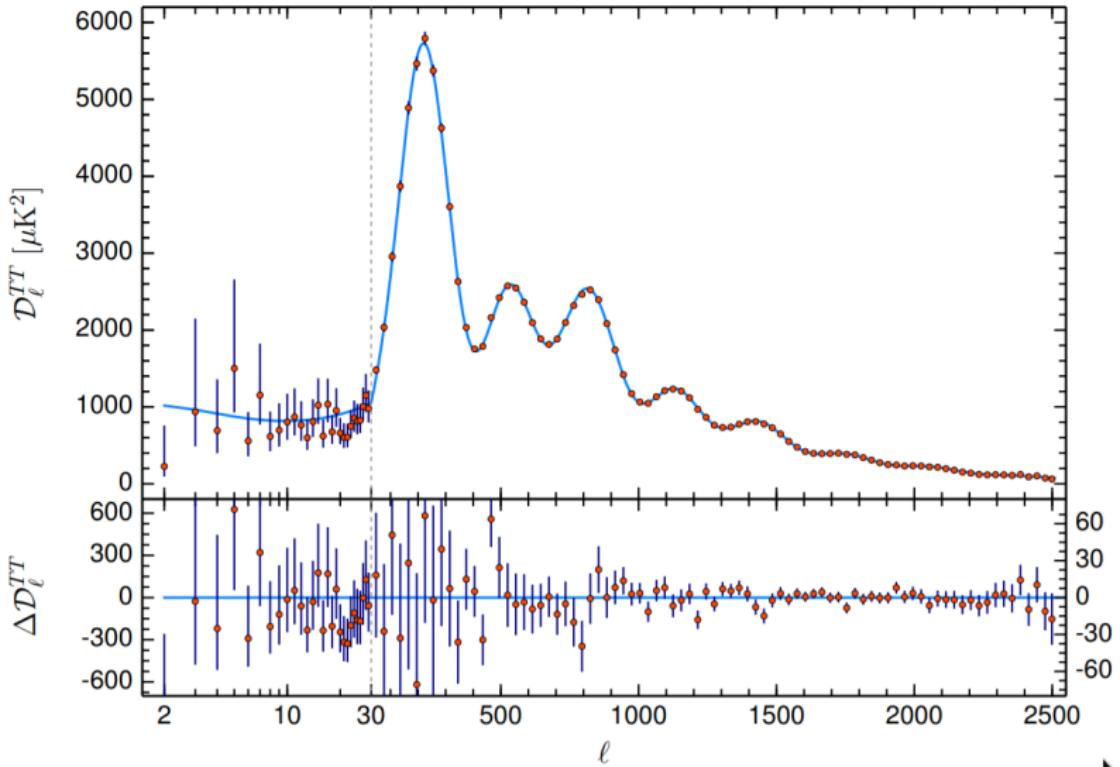
**In collaboration with**

Carlos Yaguna (UPTC), Julian Calle, Óscar Zapata, Andrés Rivera (UdeA),  
Walter Tangarife (Loyola University Chicago)

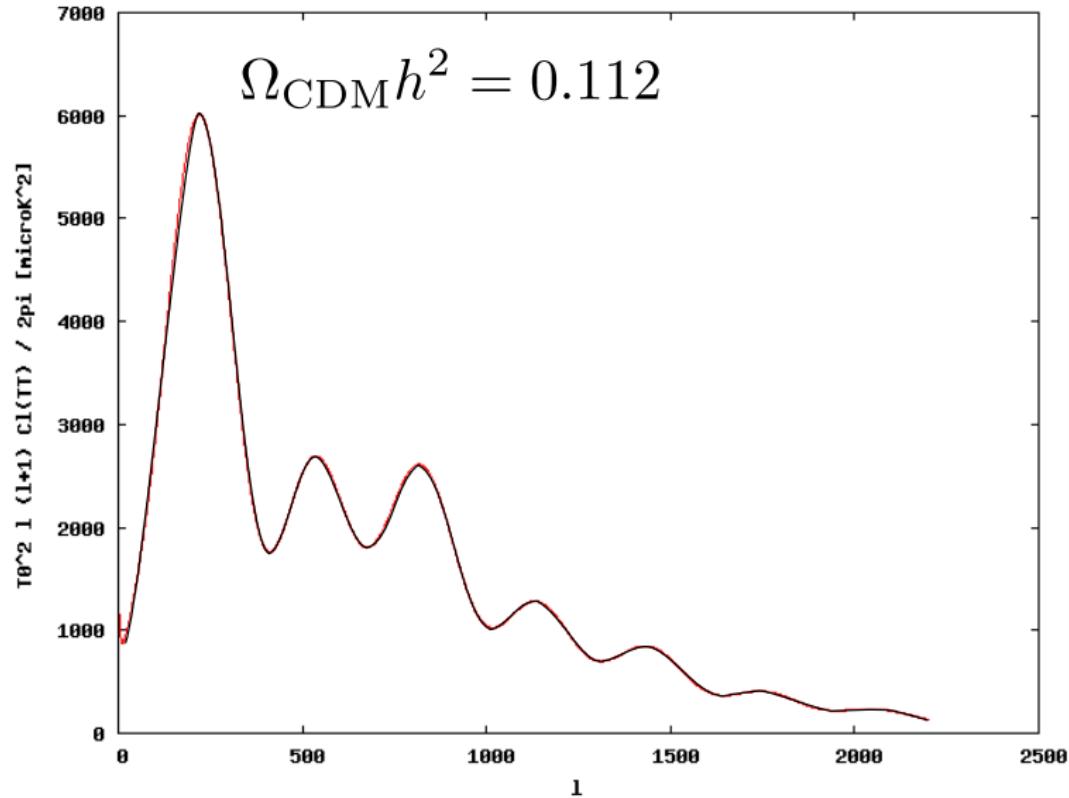


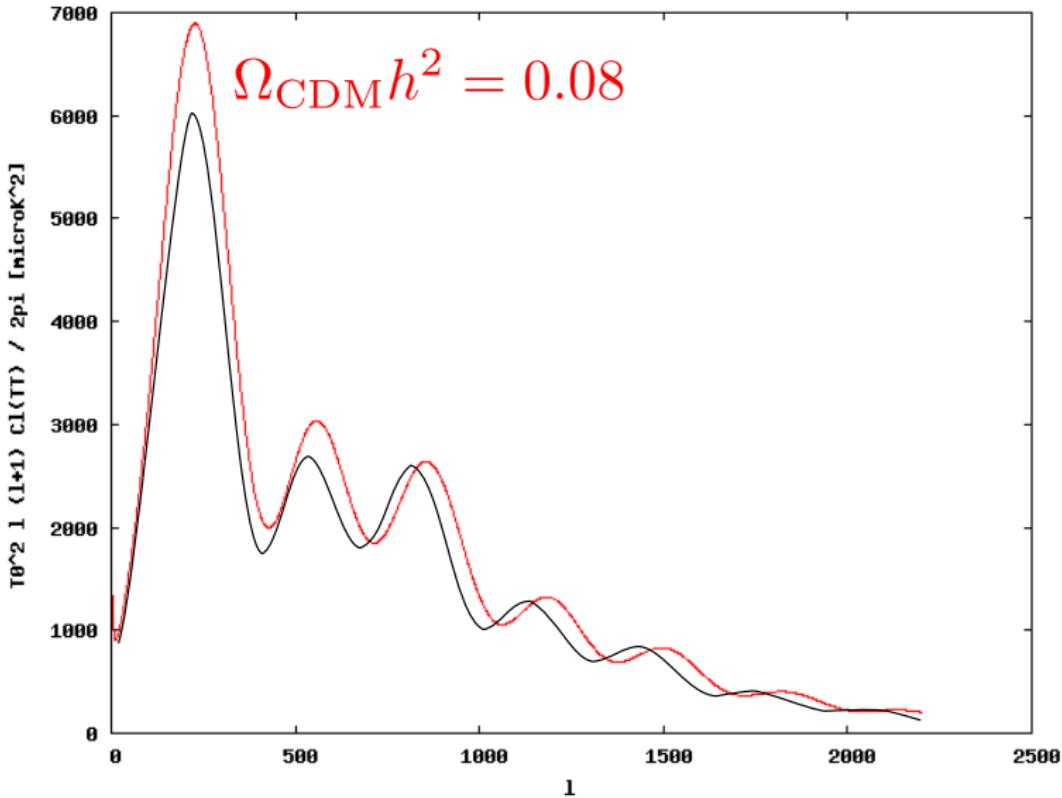
$\Lambda$ CDM paradigm (with baryonic effects)

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Credit: Planck 2018





# Cosmic Miso Soup

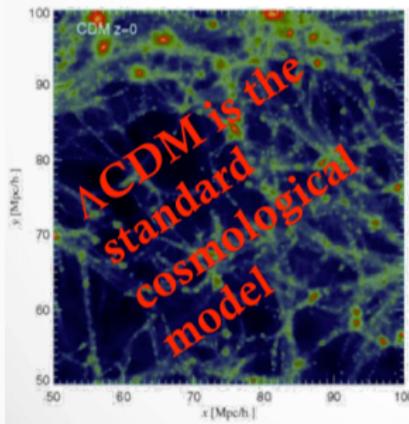
- When matter and radiation were hotter than 3000 K, matter was completely ionised. The Universe was filled with plasma, which behaves just like a soup
- Think about a Miso soup (if you know what it is). Imagine throwing Tofus into a Miso soup, while changing the density of Miso
- And imagine watching how ripples are created and propagate throughout the soup



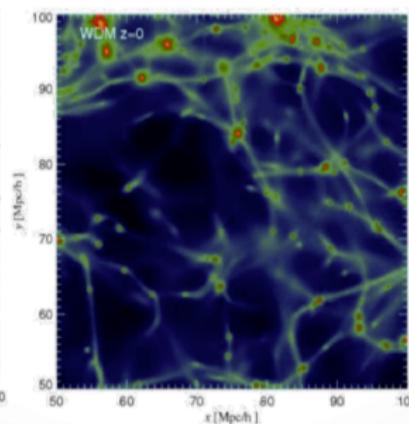
Nobu São Paulo version

# Dark matter simulations

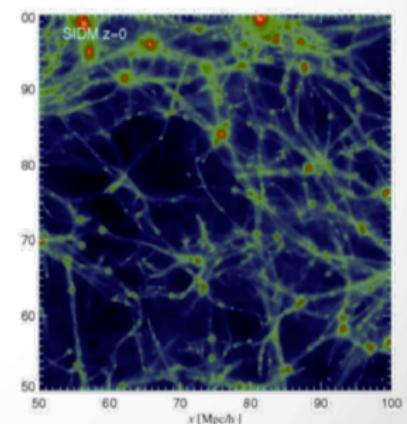
Cold Dark Matter  
(Slow moving)  
 $m \sim \text{GeV-TeV}$   
Small structures form  
first, then merge



Warm Dark Matter  
(Fast moving)  
 $m \sim \text{keV}$   
Small structures are  
erased

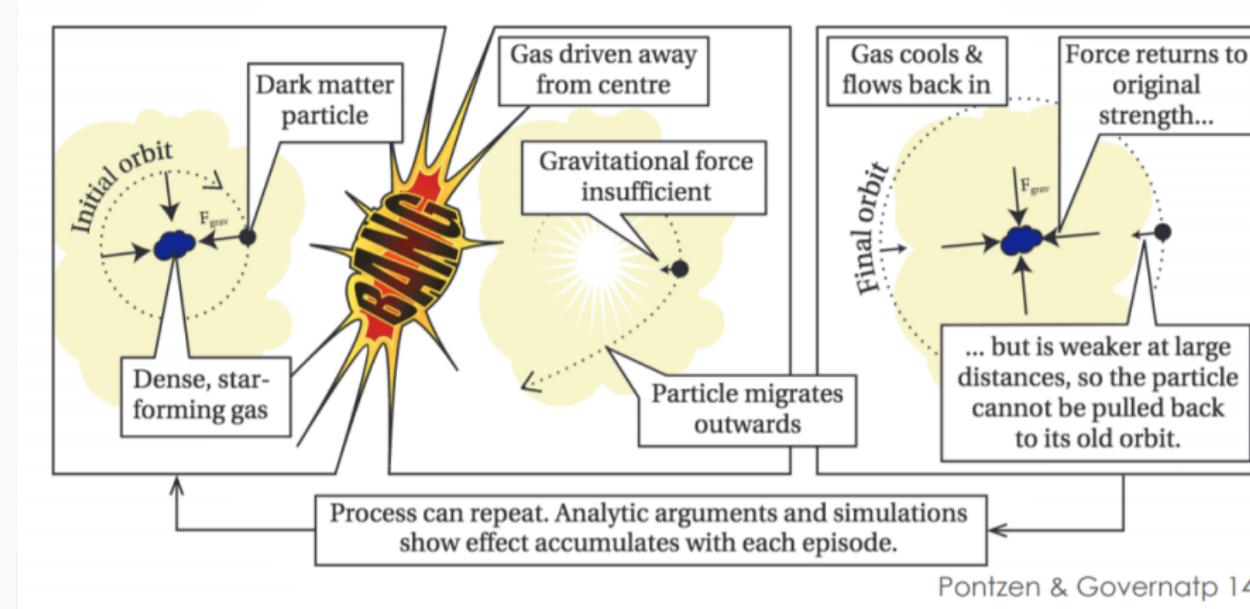


Self-Interacting Dark Matter  
Strongly interact with itself  
Large scale similar to CDM,  
Small galaxies are different



Credit: Arianna Di Cintio (Conference on Shedding Light on the Dark Universe with Extremely Large Telescopes, ICTP - 2018)

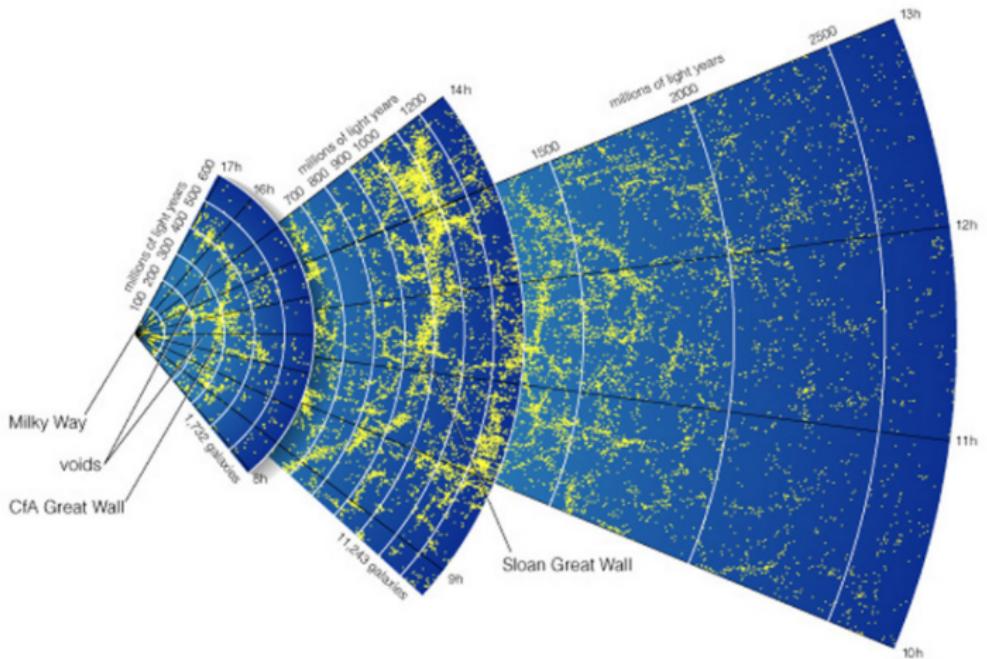
# Baryonic effects



Pontzen & Governato 14

Once the effect of baryonic physics is included, it is hard to distinguish between WDM/SIDM/CDM

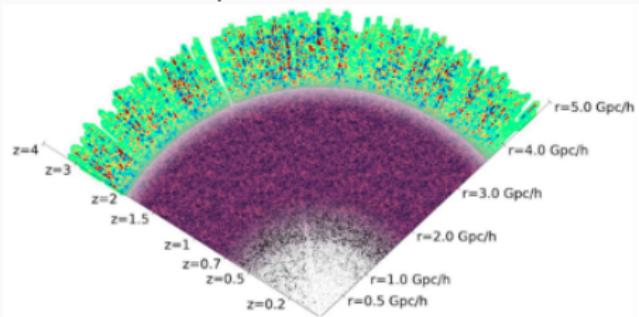
# Goal



Maps of galaxy positions reveal extremely large structures: ***superclusters*** and ***voids***

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## The DESI experiment



Credits: J. Forero

<http://cosmology.univalle.edu.co/>

Dark matter in the universe evolves through gravity to form a complex network of halos, filaments, sheets and voids, that is known as the cosmic web

A.C Rodriguez *et al* arXiv:1801.09070 [CAC]

Cosmological simulations of structure formation predict that the majority of gas in the intergalactic medium (IGM) is distributed in a cosmic web of sheets and filaments as a consequence of gravitational collapse. The intersections of these structures become the locations at which galaxies and their supermassive black holes (SMBHs) form and evolve.

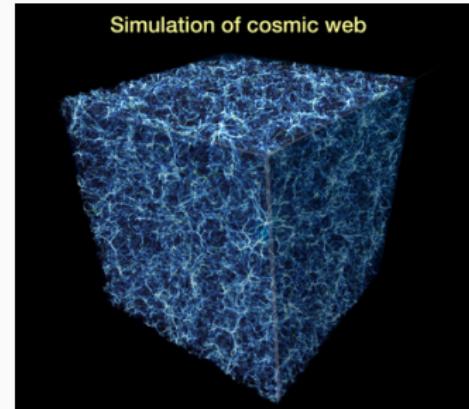
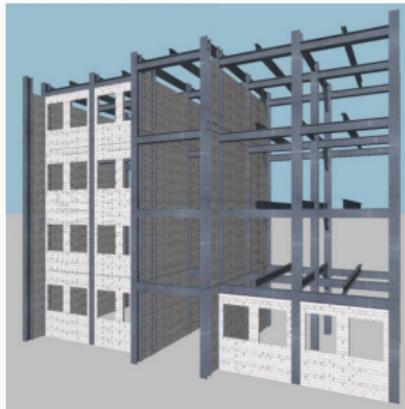
[...] at a  $z$  of  $\sim 3$ ,  $> 60\%$  of all gas in the Universe resides in filaments

H. Umehata *et al*, Science 366, 97, 4 Oct 2019

# Cooking the soup: Cosmic web

*Dark matter connects clusters of galaxies with massive tendrils, forming a cosmic web that serves as an unseen skeleton for the universe.*

<https://phys.org/news/2018-06-years-scientists-account-universe.html>



These great filaments are made largely of **dark matter** located in the space between galaxies  
and filled with 60% of the **primordial gas**!

[<https://hubblesite.org>]

An excess of a gas ( $20\sigma$ ) is observed between Milky Way and Andromeda (M31): arXiv:1403.7528 [MNRAS]<sup>1</sup>

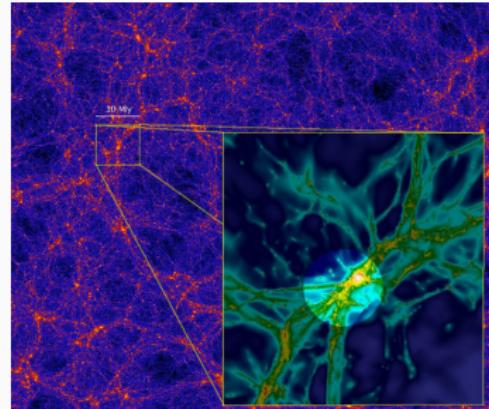
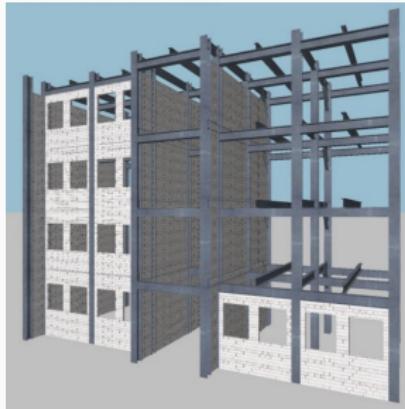
Clouds of HI likely embedded in a filament between M31 and M33: arXiv:1305.1631 [nature]

<sup>1</sup> See also: arXiv:1603.05400 [A&A]

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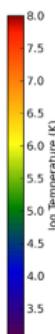
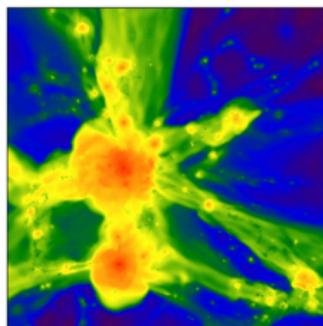
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# Direct observations of filaments

## Where are the Baryons? (Cen, Ostriker, astro-ph/9806281 [AJ])

*Thus, not only is the universe dominated by dark matter, but more than one half of the normal matter is yet to be detected. (the muscles)*



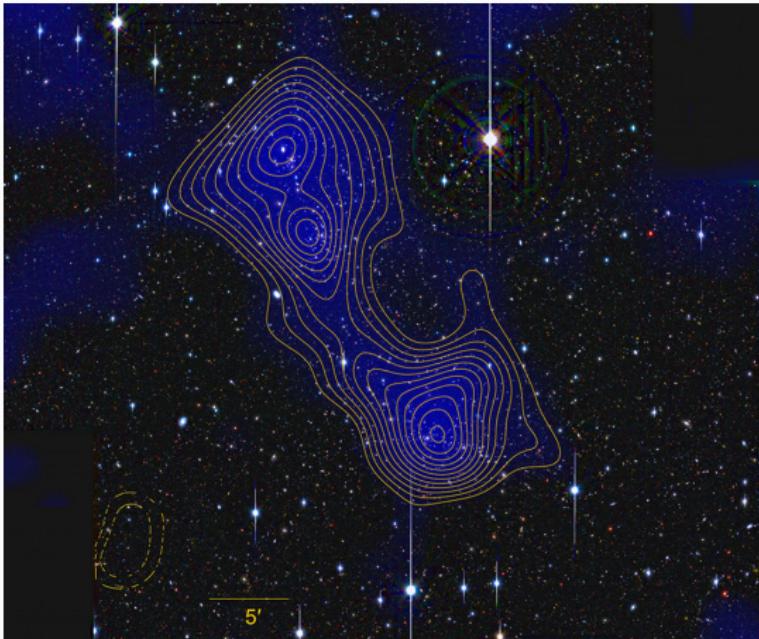
Warm-hot intergalactic medium (WHIM)  
Density-weighted temperature projection of a portion of the refinement box of the C run of size  $(18 h^{-1}\text{Mpc})^3$ .  
Low temperature WHIM confirmed by O VI line that peak at  $T \sim 3 \times 10^5 \text{ K}$

Credit: Cen, arXiv:1112.4527 [AJ]



Hotter phases of the WHIM: Observations of the missing baryons in the warm-hot intergalactic medium (Nicastro, et al. arXiv:1806.08395 [Nature]).

# A filament of dark matter between two clusters of galaxies



Supercluster system of three galaxy clusters

- Abell 222 (south) detected at  $\sim 8\sigma$
- Abell 223 (north) double galaxy cluster seen at  $\sim 7\sigma$

reconstructed surface mass density (blue)

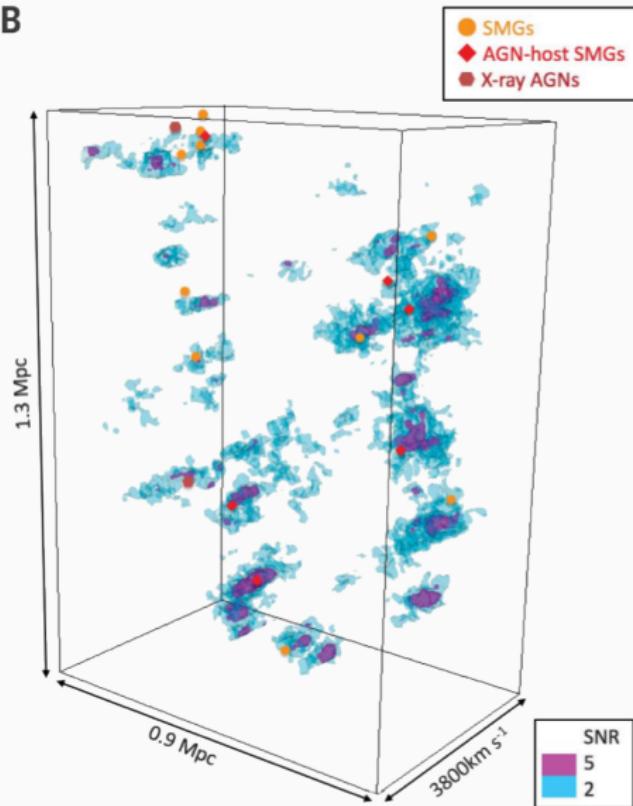
significance contours from  $0.5\sigma$  to  $2.5\sigma$

J.P. Dietrich *et al*, arXiv:1207.0809 [Nature]

For a recent review see: arXiv:1905.08991

# Three-dimensional pictures of Ly $\alpha$ filaments

B



The 3D distribution of Ly $\alpha$  filaments shown with

signal-to-noise ratio (SNR) > 5

signal-to-noise ratio (SNR) > 2

H. Umehata *et al*, Science 366, 97, 4 Oct 2019

Agrawal, et al, arXiv:1610.04611 [JCAP]

*Apart from its manifold gravitational influences, dark matter has so far eluded detection, prompting model builders to think more broadly about what dark matter can be and in the process consider other and more subtle ways to search for it.*

## Dark sectors

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In the following discussion we use the following doublets

$$H = \begin{pmatrix} H^+ \\ H^0 \end{pmatrix}, \quad L_i = \begin{pmatrix} \nu_{Li} \\ e_{Li}^- \end{pmatrix}. \quad (1)$$

corresponding to the Higgs doublet and the lepton doublets (in Weyl Notation) respectively, such that

$$L_i \cdot H = \epsilon_{ab} L_i^a H^b, \quad a, b = 1, 2$$



SM





SM

$$m_{\text{Majorana}}^\nu = \frac{h_\nu}{\Lambda} L \cdot H L \cdot H \quad (\text{three-level})$$

Type-I arXiv:1808.03352, II arXiv:1607.04029, III arXiv:1908.04308



$$\mathcal{L} = y(N_R)^\dagger L \cdot H + M_N N_R N_R + \text{h.c}$$



$$m_{\text{Majorana}}^\nu = \frac{h_\nu}{\Lambda} L \cdot H L \cdot H$$

Type-I arXiv:1808.03352, with N. Bernal, C. Yaguna, and Ó. Zapata [PRD]

$$U(1)_X \rightarrow Z_7$$

$$\mathcal{L} = y(N_R)^\dagger L \cdot \langle H \rangle + y' \langle S \rangle N_R N_R + \text{h.c}$$



$$m_{\text{Majorana}}^\nu = \frac{h_\nu^+}{\Lambda} L \cdot H L \cdot H \frac{S}{\Lambda}$$

Type-I arXiv:1808.03352

: Also new terms arise  
from spontaneous  
breakdown of a new  
gauge symmetry

Local  $U(1)_X \rightarrow Z_7$

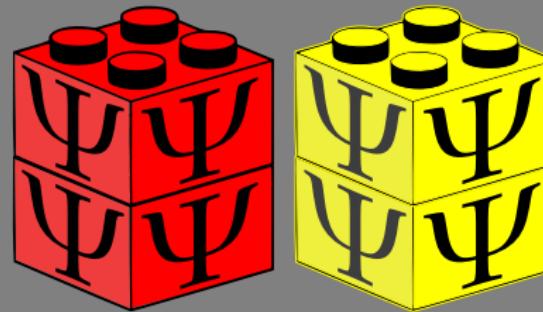
$$\mathcal{L} = y(N_R)^\dagger L \cdot \langle H \rangle + y' \langle S \rangle N_R N_R + \text{h.c}$$



$$m_{\text{Majorana}}^\nu = \frac{h_\nu}{\Lambda} L \cdot H L \cdot H \frac{S}{\Lambda}$$

Type-I arXiv:1808.03352

: Also new terms arise  
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## Standard model extended with $U(1)_X$ gauge symmetry

Fields	$SU(2)_L$	$U(1)_Y$	$U(1)_X$
$L$	2	-1/2	$l$
$Q$	2	-1/6	$q$
$d_R$	1	-1/2	$d$
$u_R$	1	+2/3	$u$
$e_R$	1	-1	$e$
$H$	2	-1/2	$h$
$\psi$	1	0	$n$

Table 1: The new and fermions with their respective charges.

$$[\mathrm{SU}(3)_c]^2 \mathrm{U}(1)_X :$$

$$[3u + 3d] - [3 \cdot 2q] = 0$$

$$[\mathrm{SU}(2)_L]^2 \mathrm{U}(1)_X :$$

$$-[2\textcolor{blue}{l} + 3 \cdot 2q] = 0$$

$$[\mathrm{U}(1)_Y]^2 \mathrm{U}(1)_X :$$

$$\left[ (-2)^2 e + 3 \left(\frac{4}{3}\right)^2 u + 3 \left(-\frac{2}{3}\right)^2 d \right] - \left[ 2(-1)^2 \textcolor{blue}{l} + 3 \cdot 2 \left(\frac{1}{3}\right)^2 q \right] = 0 \quad (2)$$

with solution

$$u = -e + \frac{2l}{3}, \quad d = e - \frac{4l}{3}, \quad q = -\frac{l}{3}, \quad (2)$$

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which satisfy

$$\mathrm{U}(1)_Y [\mathrm{U}(1)_X]^2 : \quad [(-2)e^2 + 3\left(\frac{4}{3}\right)u^2 + 3\left(-\frac{2}{3}\right)d^2] - [2(-1)\underline{l}^2 + 3 \cdot 2\left(\frac{1}{3}\right)q^2] = 0 \quad (3)$$

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For  $N$  extra quiral fields  $\psi_\alpha$  ( $\alpha = 1, \dots, N$ ) with  $X$ -charges  $n_\alpha$ :

$$[SO(1, 3)]^2 U(1)_X : \quad \sum_{\alpha} n_{\alpha} + 3(e - 2l) = 0,$$

$$[U(1)_X]^3, \quad \sum_{\alpha} n_{\alpha}^3 + 3(e - 2l)^3 = 0$$

with solution

$$u = -\textcolor{violet}{r} - \frac{4\textcolor{blue}{l}}{3}, \quad d = \textcolor{violet}{r} + \frac{2\textcolor{blue}{l}}{3}, \quad q = -\frac{\textcolor{blue}{l}}{3}, \quad e = \textcolor{violet}{r} + 2\textcolor{blue}{l}, \quad (2)$$

which satisfy

$$\mathrm{U}(1)_Y [\mathrm{U}(1)_X]^2 : \quad [(-2)e^2 + 3\left(\frac{4}{3}\right)u^2 + 3\left(-\frac{2}{3}\right)d^2] - [2(-1)\textcolor{blue}{l}^2 + 3 \cdot 2\left(\frac{1}{3}\right)q^2] = 0 \quad (3)$$

For  $N$  extra quiral fields  $\psi_\alpha$  ( $\alpha = 1, \dots, N$ ) with  $X$ -charges  $n_\alpha$ :  $\textcolor{violet}{r} \equiv e - 2\textcolor{blue}{l}$

$$\begin{aligned} [\mathrm{SO}(1,3)]^2 \mathrm{U}(1)_X : & \quad \sum_{\alpha} n_{\alpha} + 3\textcolor{violet}{r} = 0, \\ [\mathrm{U}(1)_X]^3, & \quad \sum_{\alpha} n_{\alpha}^3 + 3\textcolor{violet}{r}^3 = 0 \end{aligned}$$

Then the general anomaly free two-parameter solution can be written as

$$X(\textcolor{violet}{r}, \textcolor{blue}{l}) = \textcolor{violet}{r}R + \textcolor{blue}{l}Y.$$

with solution

$$u = -1 - \frac{4l}{3}, \quad d = 1 + \frac{2l}{3}, \quad q = -\frac{l}{3}, \quad e = 1 + 2l, \quad (2)$$

which satisfy

$$\mathrm{U}(1)_Y [\mathrm{U}(1)_X]^2 : \quad [(-2)e^2 + 3\left(\frac{4}{3}\right)u^2 + 3\left(-\frac{2}{3}\right)d^2] - [2(-1)l^2 + 3 \cdot 2\left(\frac{1}{3}\right)q^2] = 0 \quad (3)$$

For  $N$  extra quiral fields  $\psi_\alpha$  ( $\alpha = 1, \dots, N$ ) with  $X$ -charges  $n_\alpha$ :  $r \equiv e - 2l = 1$

$$\begin{aligned} [\mathrm{SO}(1,3)]^2 \mathrm{U}(1)_X : & \quad \sum_{\alpha} n_{\alpha} + 3 = 0, \\ [\mathrm{U}(1)_X]^3 : & \quad \sum_{\alpha} n_{\alpha}^3 + 3 = 0 \end{aligned}$$

Since  $f \rightarrow f' \rightarrow f/r$ , without lost of generality:  $r \rightarrow 1$

$$X(l) = R + lY.$$

We impose  $\nu_{R1} = \psi_N$ ,  $\nu_{R2} = \psi_{N-1}$ , to have at most one massless neutrino.

## One parameter $U(1)_X$ SM extension

Fields	$SU(2)_L$	$U(1)_Y$	$U(1)_X$	$U(1)_{B-L}$	$U(1)_R$	$U(1)_D$	$U(1)_G$	$U(1)_D^*$
$L$	2	-1/2	$\textcolor{blue}{l}$	-1	0	-3/2	-1/2	0
$Q$	2	-1/6	$-l/3$	$1/3$	0	1/2	$1/6$	0
$d_R$	1	-1/2	$1 + 2l/3$	$1/3$	1	0	$2/3$	0
$u_R$	1	+2/3	$-1 - 4l/3$	$1/3$	-1	1	$-1/3$	0
$e_R$	1	-1	$1 + 2l$	-1	1	-2	0	0
$H$	2	1/2	$-1 - l$	0	-1	1/2	$-1/2$	0
$\sum_\alpha n_\alpha$	1	0	-3	-3	-3	-3	-3	0
$\sum_\alpha n_\alpha^3$	1	0	-3	-3	-3	-3	-3	0

\*  $r = \textcolor{blue}{l} = 0$  ( $n_\alpha \neq 0$ )

Solutions in terms of a parameter: arXiv:1811.11927, N. Okada, et al [PRD];

and some specific examples from: arXiv:1705.05388, Farinaldo Queiroz, et al [JHEP]

All known  $U(1)_{B-L}$  (radiative) neutrino solutions apply for  $U(1)_X$ :

Known solutions with  $\sum n_\alpha = -3$  and  $\sum n_\alpha^3 = -3$

$(\nu_{R1}, \nu_{R2}, \psi_{N-2}, \dots)$	Ref
$(-1, -1, -1)$	hep-ph/0611205, S. Khalil [JPG]
$(-4, -4, +5)$	 arXiv:0706.0473, Montero, V. Pleitez [PLB]
$\left(-\frac{2}{3}, -\frac{2}{3}, -\frac{4}{3}, -\frac{1}{3}\right)$	 arXiv:1607.04029, S. Patra , W. Rodejohann, C. Yaguna [JHEP]
$\left(-\frac{8}{5}, -\frac{8}{5}, -\frac{2}{5}, -\frac{7}{5}, +2\right)$	 arXiv:1812.05523, with J. Calle, C. Yaguna, Ó. Zapata [PRD]
$\left(-1, -1, -\frac{10}{7}, -\frac{4}{7}, -\frac{2}{7}, \frac{9}{7}\right)$	 1808.03352, with N. Bernal, C. Yaguna, Ó. Zapata [PRD]

**Table 2:** The possible solutions of the Dirac neutrino mass models with at least two repeated charges and until six chiral fermions.

Known solutions with  $\sum n_\alpha = -3$  and  $\sum n_\alpha^3 = -3$

$(\nu_{R1}, \nu_{R2}, \psi_{N-2}, \dots)$	Ref
$(-1, -1, -1)$	hep-ph/0611205, S. Khalil [JPG]
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[https://en.wikipedia.org/wiki/Sums\\_of\\_three\\_cubes](https://en.wikipedia.org/wiki/Sums_of_three_cubes)

Only known integer solutions for  $-3$  (1953)

September 2019:

$$42 = (-80538738812075974)^3 + 80435758145817515^3 + 12602123297335631^3$$

**Table 2:** The possible solutions of the Dirac neutrino mass models with at least two repeated charges and until six chiral fermions.

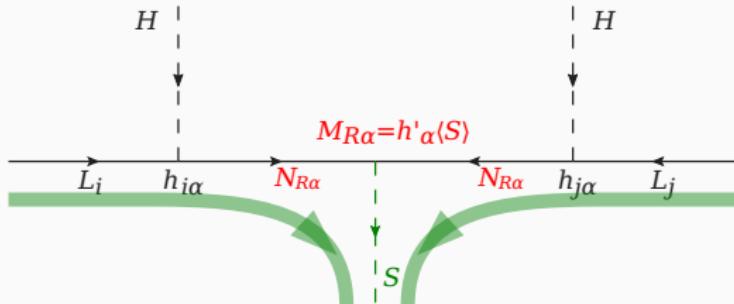
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Not known solution for  
one-loop neutrino Majorana masses  
with local  $U(1)_X$ .

**Table 2:** The possible solutions of the Dirac neutrino mass models with at least two repeated charges and until six chiral fermions.

Fields	$U(1)_{B-L}$	$Z_2^1$	$Z_2^1$
$L$	-1	+	+
$Q$	1/3	+	+
$d_R$	1/3	+	+
$u_R$	1/3	+	+
$e_R$	-1	+	+
$H$	0	+	+
$S$	-2	+	+
$N_{R1}$	-1	+	+
$N_{R2}$	-1	+	+
$\psi_1 \rightarrow (\xi_L)^\dagger$	-10/7	-	+
$\psi_2 \rightarrow \eta_R X$	-4/7	-	+
$\psi_3 \rightarrow \zeta_R$	-2/7	+	-
$\psi_4 \rightarrow (\chi_L)^\dagger$	+9/7	+	-
$S'$	1	+	+

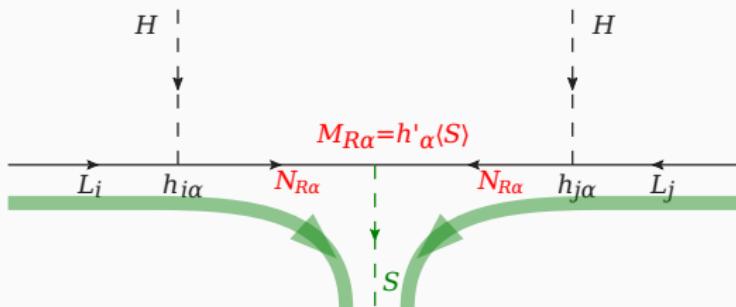


After integrating out heavy fermions, we obtain  
light neutrino masses

$$\mathcal{M}_\nu^{ij} = \sum_{\alpha=1}^2 (h^{i\alpha} v) \frac{1}{M_R^\alpha} (h^{j\alpha} v)$$

With only two heavy fermions, one massless neutrino is left

Fields	$U(1)_{B-L}$	$Z_2^1$	$Z_2^1$
$L$	-1	+	+
$Q$	1/3	+	+
$d_R$	1/3	+	+
$u_R$	1/3	+	+
$e_R$	-1	+	+
$H$	0	+	+
$S$	-2	+	+
$N_{R1}$	-1	+	+
$N_{R2}$	-1	+	+
$\psi_1 \rightarrow (\xi_L)^\dagger$	-10/7	-	+
$\psi_2 \rightarrow \eta_R \chi$	-4/7	-	+
$\psi_3 \rightarrow \zeta_R$	-2/7	+	-
$\psi_4 \rightarrow (\chi_L)^\dagger$	+9/7	+	-
$S'$	1	+	+



Two component Dirac fermion dark matter



$$\chi_1 = \begin{pmatrix} \xi_L \\ \eta_R \end{pmatrix},$$

$$\mathcal{L} = M_1 \overline{\chi}_1 \chi_1$$

$$\chi_2 = \begin{pmatrix} \chi_L \\ \zeta_R \end{pmatrix}$$

$$+ M_2 \overline{\chi}_2 \chi_2$$

## (One-loop) Dirac neutrino masses

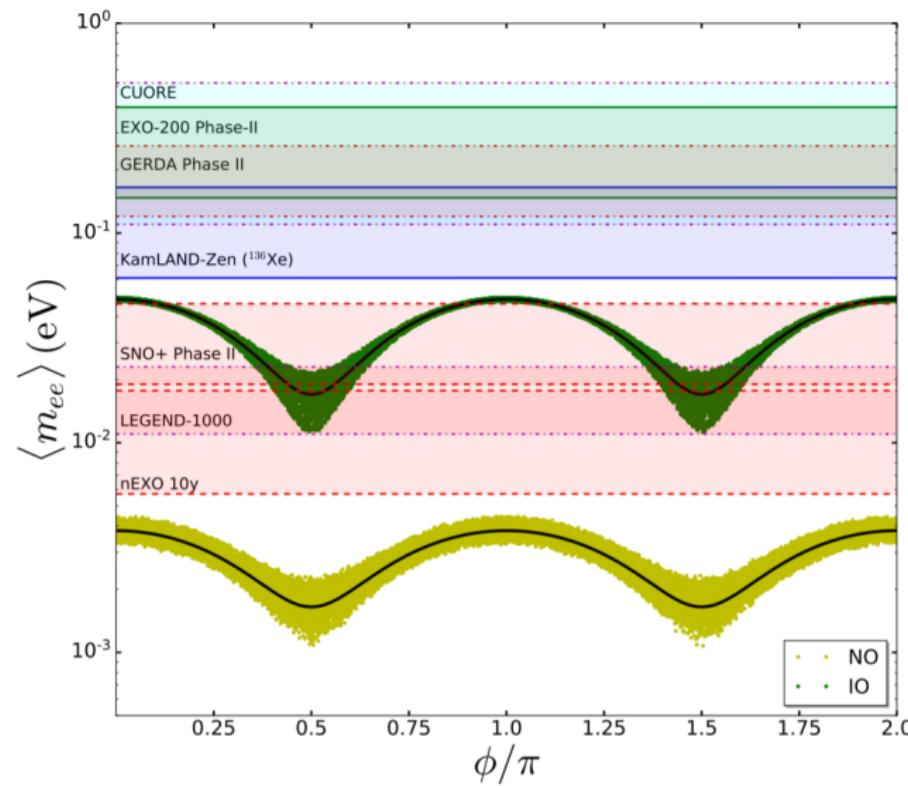
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## Lepton number

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- Lepton number ( $L$ ) is an accidental discrete or Abelian symmetry of the standard model (SM).
- Without neutrino masses  $L_e, L_\mu, L_\tau$  are also conserved.
- The processes which violate individual  $L$  are called Lepton flavor violation (LFV) processes.
- All the neutrino mass models predict, to some extent, LFV processes
- Only models with Majorana neutrinos predict processes with total  $L = L_e + L_\mu + L_\tau$  violation, like **neutrino less doublet beta decay** (NLDBD).
- NLDBD is experimentally challenging, specially if there is a massless neutrino in the spectrum.

# NLDBD prospects for a model with a massless neutrino (arXiv:1806.09977 [PLB] with Reig, Valle and Zapata)



Total lepton number:  $L = L_e + L_\mu + L_\tau$

### Majorana $\cancel{U(1)_L}$

Field	$Z_2$ ( $\omega^2 = 1$ )
SM	1
$L$	$\omega$
$(e_R)^\dagger$	$\omega$
$(\nu_R)^\dagger$	$\omega$

### Dirac $U(1)_L$

Field	$Z_3$ ( $\omega^3 = 1$ )
SM	1
$L$	$\omega$
$(e_R)^\dagger$	$\omega^2$
$(\nu_R)^\dagger$	$\omega^2$

$$\mathcal{L}_\nu = h_D (\nu_R)^\dagger L \cdot H + M_R \nu_R \nu_R + \text{h.c.}$$

$$\mathcal{L}_\nu = h_D (\nu_R)^\dagger L \cdot H + \text{h.c.}$$

$$h_D \sim \mathcal{O}(1)$$

$$h_D \sim \mathcal{O}(10^{-11})$$

Total lepton number:  $L = L_e + L_\mu + L_\tau$

Majorana  $\cancel{U(1)_L}$

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SM	1
$L$	$\omega$
$(e_R)^\dagger$	$\omega$
$(\nu_R)^\dagger$	$\omega$

Dirac  $U(1)_{B-L}$

Field	$Z_3$ ( $\omega^3 = 1$ )
SM	1
$L$	$\omega$
$(e_R)^\dagger$	$\omega^2$
$(\nu_R)^\dagger$	$\omega^2$

$$\mathcal{L}_\nu = h_D (\nu_R)^\dagger L \cdot H + M_R \nu_R \nu_R + \text{h.c.}$$

$$\mathcal{L}_\nu = h_D (\nu_R)^\dagger L \cdot H + \text{h.c.}$$

$$h_D \sim \mathcal{O}(1)$$

$$h_D \sim \mathcal{O}(10^{-11})$$

Explain smallness ala Peccei-Quinn:

$$U(1)_{B-L} \xrightarrow{\langle S \rangle} Z_N, \quad N \geq 3.$$

## Small Dirac neutrino masses

---

To explain the **smallness** of Dirac neutrino masses choose  $U(1)_{B-L}$  which:

- Forbids tree-level mass (TL) term ( $Y(H) = +1/2$ )

$$\begin{aligned}\mathcal{L}_{T,L} &= h_D \epsilon_{ab} (\nu_R)^\dagger L^a H^b + \text{h.c} \\ &= h_D (\nu_R)^\dagger L \cdot H + \text{h.c}\end{aligned}$$

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## Small Dirac neutrino masses

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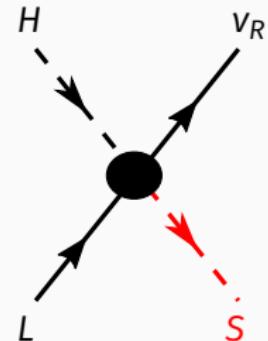
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$$U(1)_{B-L} \xrightarrow{\langle S \rangle} Z_N$$

- Forbids Majorana term:  $\nu_R \nu_R$
- Realizes of the 5-dimension operator which conserves lepton number in  $SU(3)_c \times SU(2)_L \times U(1)_Y \times U(1)_{B-L}$ :

$$\mathcal{L}_{5-D} = \frac{h_\nu}{\Lambda} (\nu_R)^\dagger L \cdot H \cancel{S} + \text{h.c}$$



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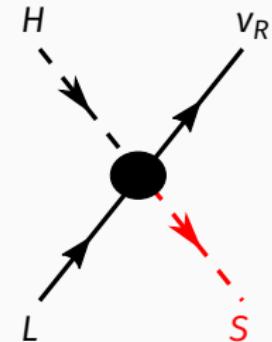
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$$\mathcal{L}_{5-D} = \frac{h_\nu}{\Lambda} (\nu_R)^\dagger L \cdot H \bar{S} + \text{h.c}$$



- Enhancement to the *effective number of degrees of freedom in the early Universe*  $\Delta N_{\text{eff}} = N_{\text{eff}} - N_{\text{eff}}^{\text{SM}}$  (see arXiv:1211.0186)

See E. Ma, Rahul Srivastava: arXiv:1411.5042 [PLB] for tree-level realization

From 1210.6350 and 1805.02025:  $\Delta N_{\text{eff}} = 3 (T_{\nu_R}/T_{\nu_L})^4$

$$\begin{aligned}\Gamma_{\nu_R}(T) &= n_{\nu_R}(T) \sum_f \langle \sigma_f (\nu_R \bar{\nu}_R \rightarrow f\bar{f}) v \rangle \\ &= \sum_f \frac{g_{\nu_R}^2}{n_{\nu_R}} \int \frac{d^3 p}{(2\pi)^3} \frac{d^3 q}{(2\pi)^3} f_{\nu_R}(p) f_{\nu_R}(q) \sigma_f(s) (1 - \cos \theta),\end{aligned}$$

$$s = 2pq(1 - \cos \theta), \quad f_{\nu_R}(k) = 1/(e^{k/T} + 1)$$

$$n_{\nu_R}(T) = g_{\nu_R} \int \frac{d^3 k}{(2\pi)^3} f_{\nu_R}(k), \quad \text{with } g_{\nu_R} = 2$$

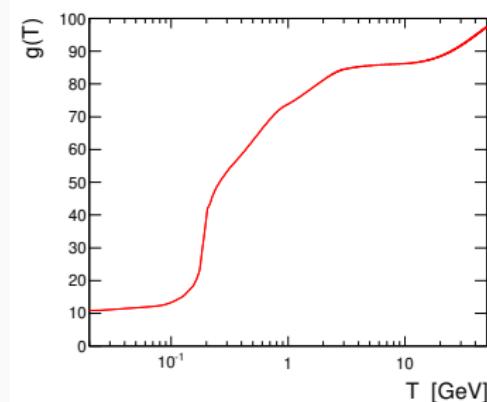
$$\sigma_f(s) \simeq \frac{N_C^f (Q_{BL}^f)^2 Q^2 s}{12\pi} \left( \frac{g'}{M_{Z'}} \right)^4, \quad \text{In the limit } M_{Z'}^2 \gg s.$$

with three right-handed neutrinos, the Hubble parameter is

$$H(T) = \sqrt{\frac{4\pi^3 G_N [g(T) + 21/4]}{45}} T^2.$$

The right-handed neutrinos decouple when

$$\Gamma_{\nu_R}(T_{\text{dec}}^{\nu_R}) = H(T_{\text{dec}}^{\nu_R}).$$



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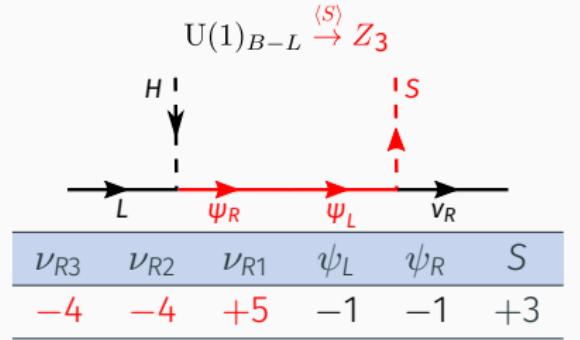
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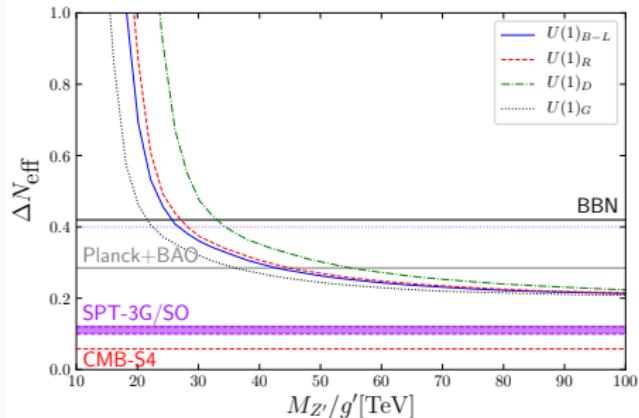
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with J. Calle and Ó. Zapata, arXiv:1909.09574

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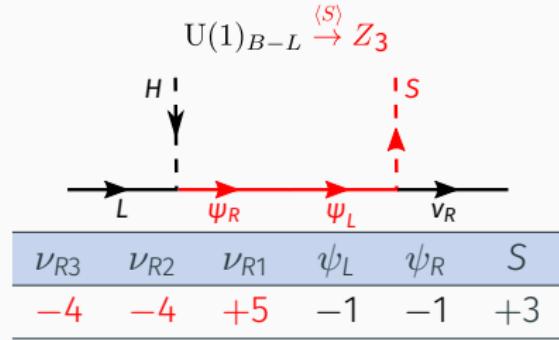
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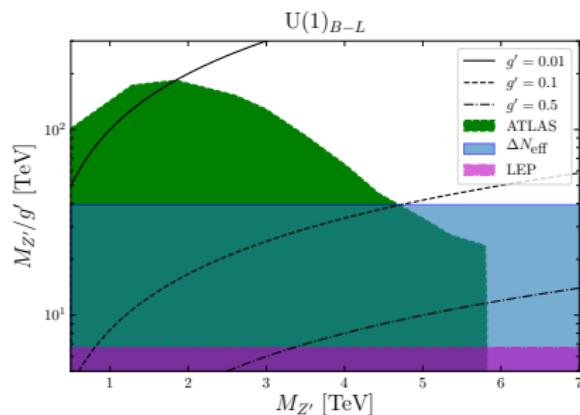
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One-loop realization of  $\mathcal{L}_{5-D}$  with  
total  $L$

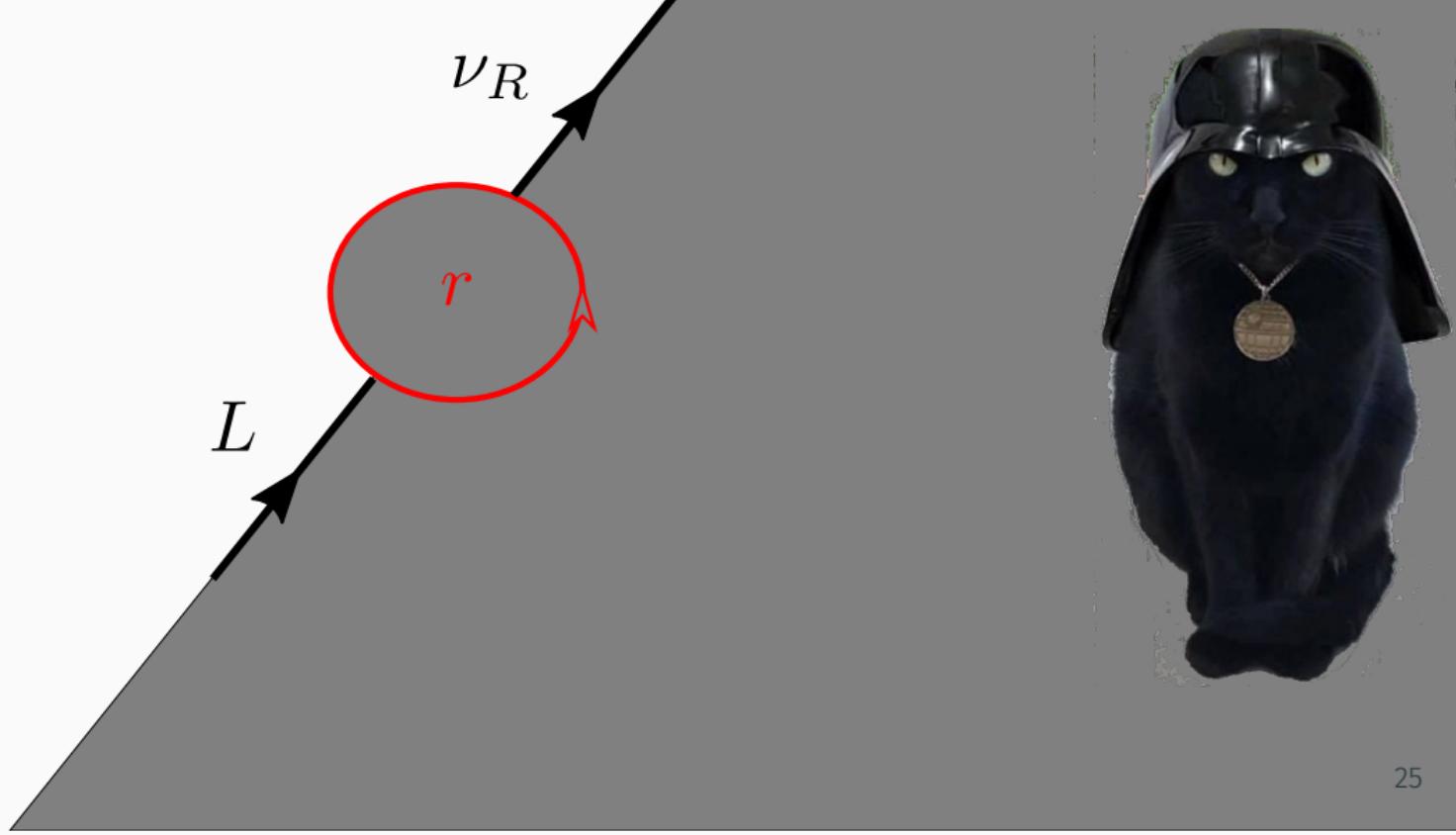
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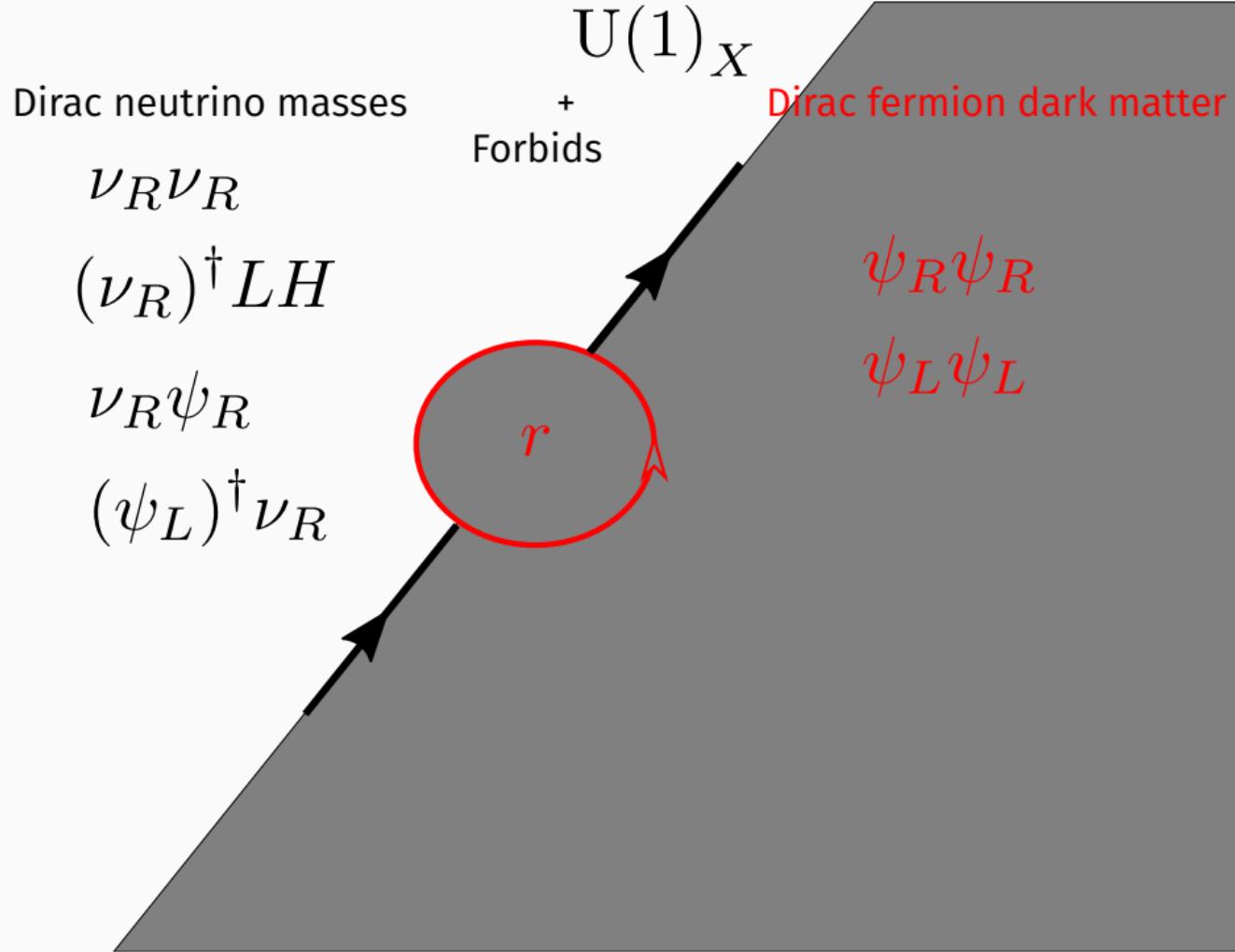
$U(1)_X$ 

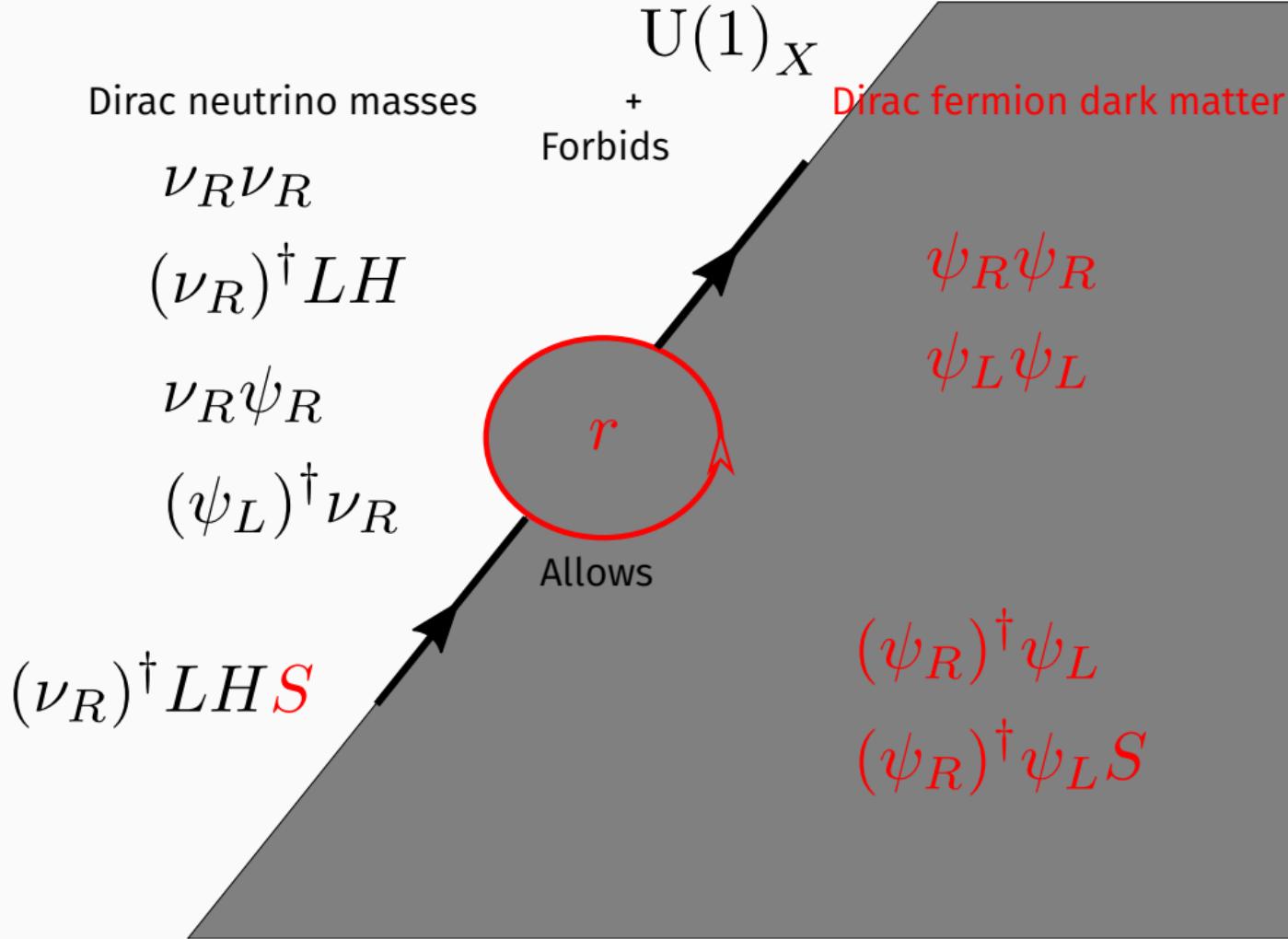
Dirac neutrino masses

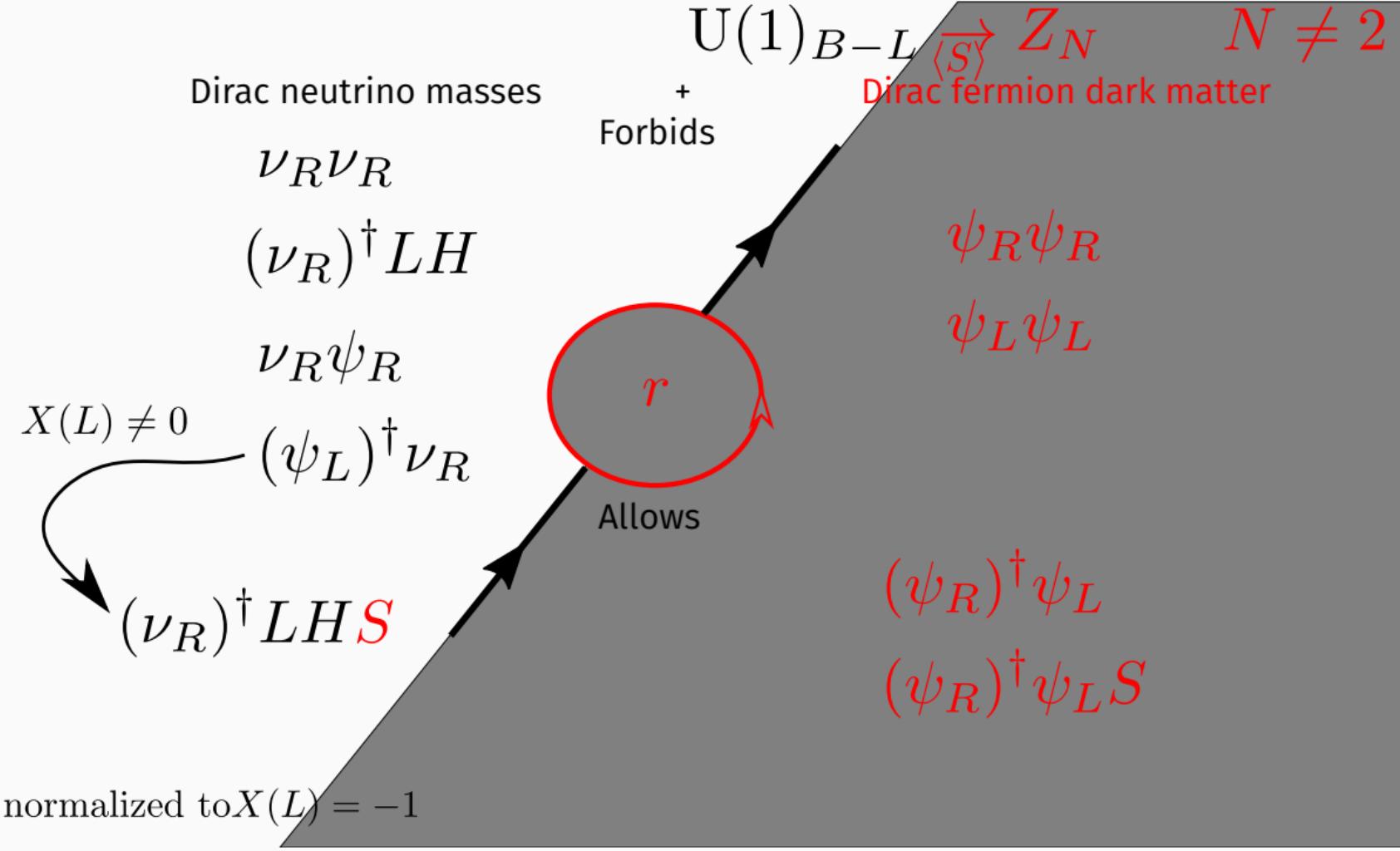
+

Dirac fermion dark matter

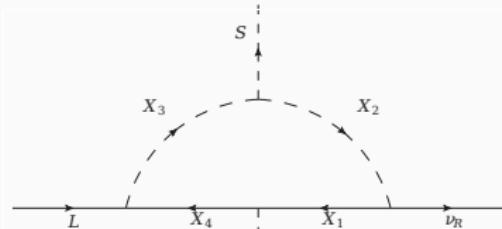




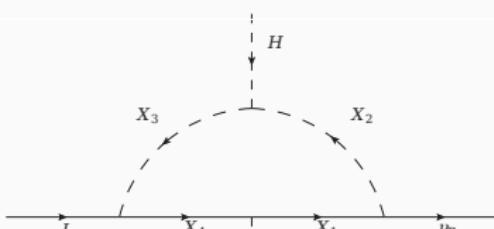




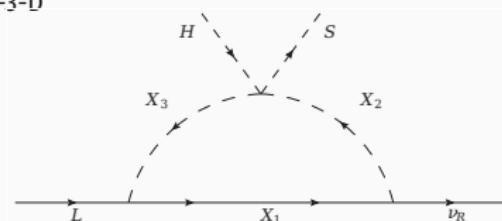
# One loop topologies $U(1)_{B-L} \oplus \mathbb{Z}_2 \oplus \mathbb{Z}_2$



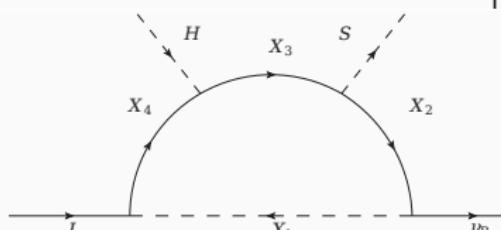
T1-3-D



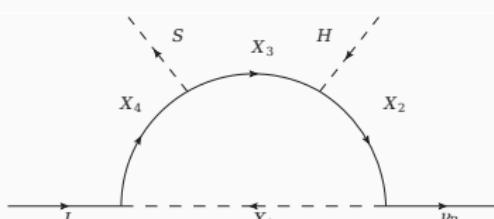
T1-3-E



T3-1-A



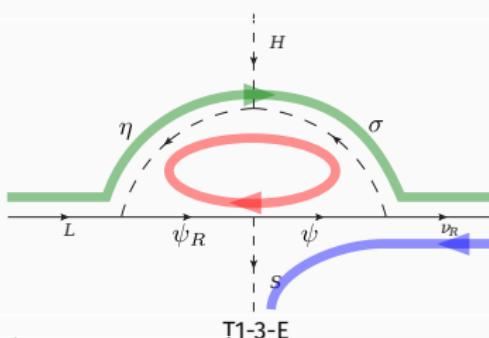
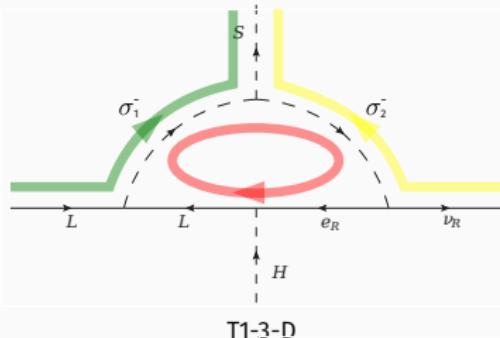
T1-2-A



T1-2-B

Chang-Yuan Yao and Gui-Jun Ding, arXiv:1802.05231 [PRD]

# One loop topologies $U(1)_{B-L}$ only!



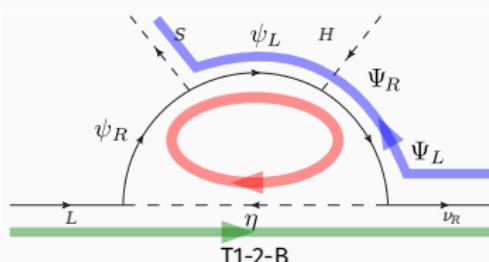
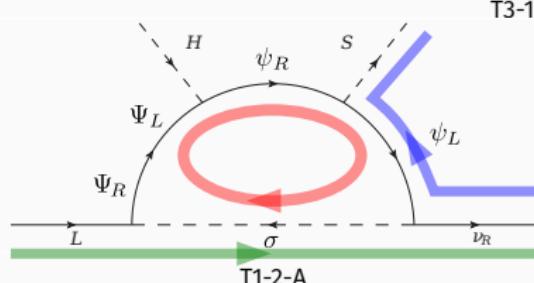
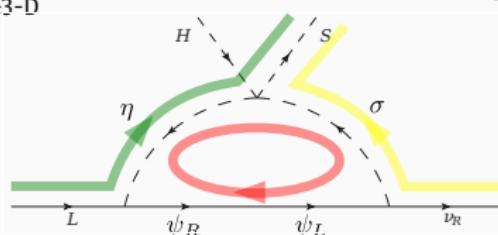
$\psi_{L,R} \rightarrow$  Singlet fermions

$\Psi_{L,R} \rightarrow$  Vector-like doublet fermions

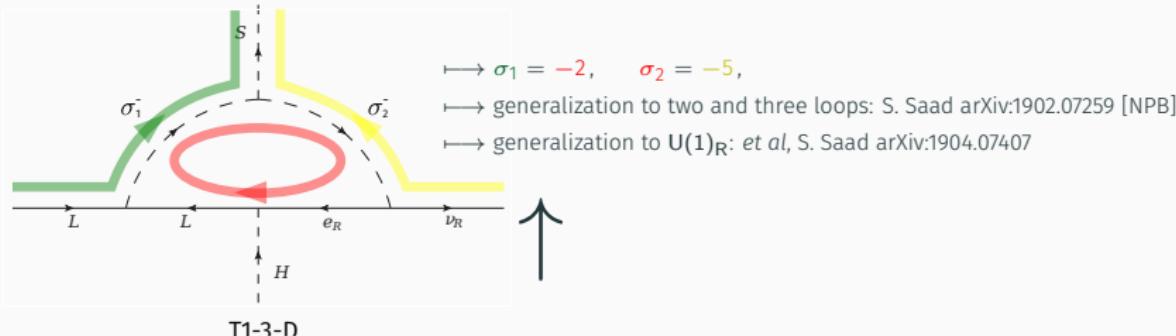
$\sigma \rightarrow$  Singlet scalar

$\eta \rightarrow$  Doublet scalar

with J. Calle, C. Yaguna, and O. Zapata, arXiv:1812.05523 [PRD]



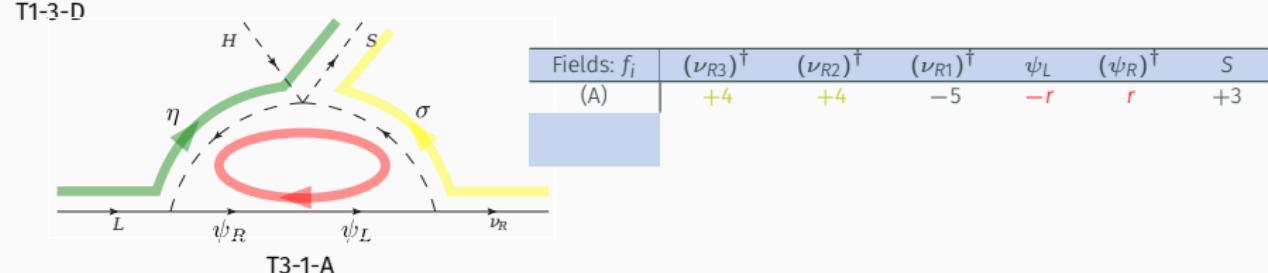
# One loop topologies $U(1)_{B-L}$ only! with J. Calle, C. Yaguna, and O. Zapata, arXiv:1812.05523 [PRD]



$\psi_{L,R} \rightarrow$  Singlet fermions (vector-like)

$\sigma \rightarrow$  Singlet scalar

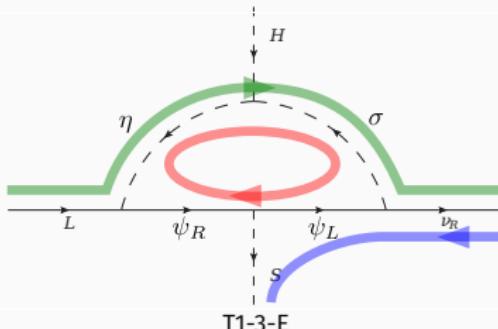
$\eta \rightarrow$  Doublet scalar



Anomaly cancellation conditions

$$\sum_i f_i = 3$$

$$\sum_i f_i^3 = 3$$



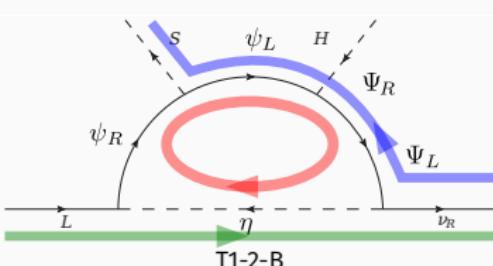
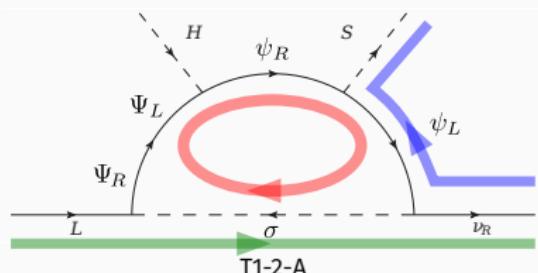
$\psi_{L,R} \rightarrow$ Singlet fermions (quiral)

$\Psi_{L,R} \rightarrow$ Vector-like doublet fermions

$\sigma \rightarrow$ Singlet scalar

$\eta \rightarrow$ Doublet scalar

Fields: $f_i$	$(\nu_{R3})^\dagger$	$(\nu_{R2})^\dagger$	$(\nu_{R1})^\dagger$	$\psi_L$	$(\psi_R)^\dagger$	$S$
(A)	+4	+4	-5	-r	r	+3
(B)	$+\frac{8}{5}$	$+\frac{8}{5}$	$+\frac{2}{5}$	$\frac{7}{5}$	$-\frac{10}{5}$	$+\frac{3}{5}$



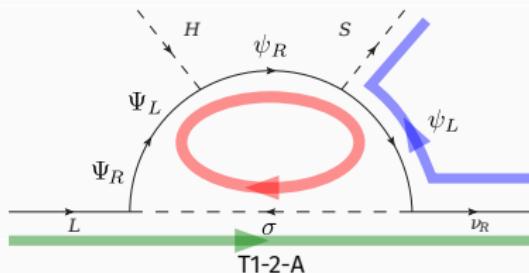
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$\psi_{L,R} \rightarrow$  Singlet fermions (quiral)  
 $\Psi_{L,R} \rightarrow$  Vector-like doublet fermions : 10/5  
 $\sigma \rightarrow$  Singlet scalar : 15/5

Fields: $f_i$	$(\nu_{R3})^\dagger$	$(\nu_{R2})^\dagger$	$(\nu_{R1})^\dagger$	$\psi_L$	$(\psi_R)^\dagger$	$S$
(A)	+4	+4	-5	-r	r	+3
(B)	$+\frac{8}{5}$	$+\frac{8}{5}$	$+\frac{2}{5}$	$\frac{7}{5}$	$-\frac{10}{5}$	$+\frac{3}{5}$



Anomaly cancellation conditions

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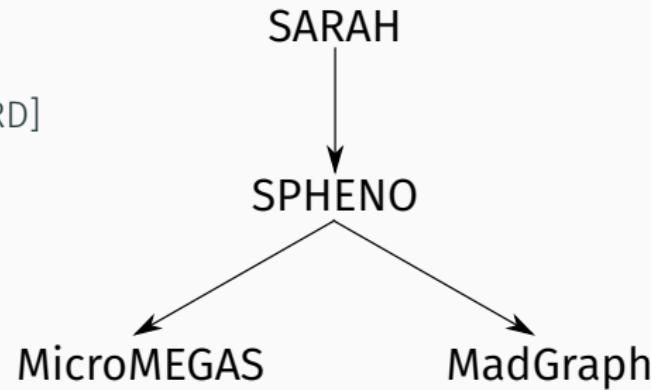
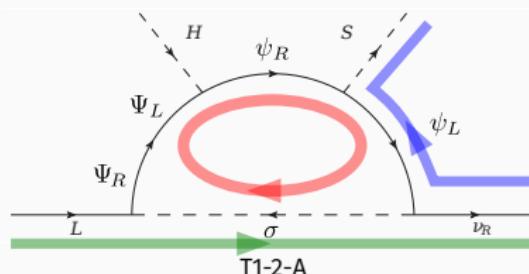
$$\sum_i f_i^3 = 3$$

# $\text{SD}^3\text{M+SSDM}$ : $\sigma_a$ ( $a = 1, 2$ )

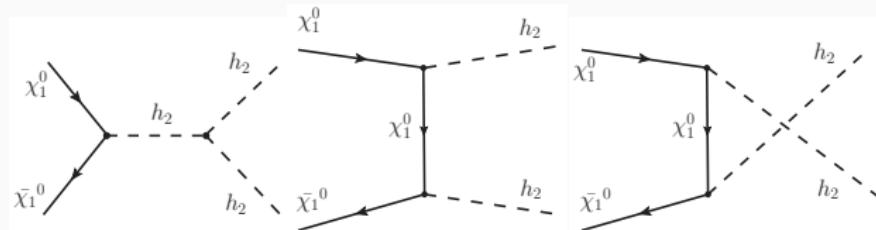
$M_\psi = h_1 \langle S \rangle$ ,  $y_2 = 0$ :

$$\mathcal{L} = \mathcal{L}_{\text{SD}^3\text{M}} + h_3^{ia} \widetilde{(\Psi_R)} \cdot L_i \sigma_a + h_2^{\beta a} (\nu_{R\beta})^\dagger \psi_L \sigma_a^* - V(\sigma_a, S, H).$$

with A.F Rivera, W. Tangarife, arXiv:1906.09685 [PRD]

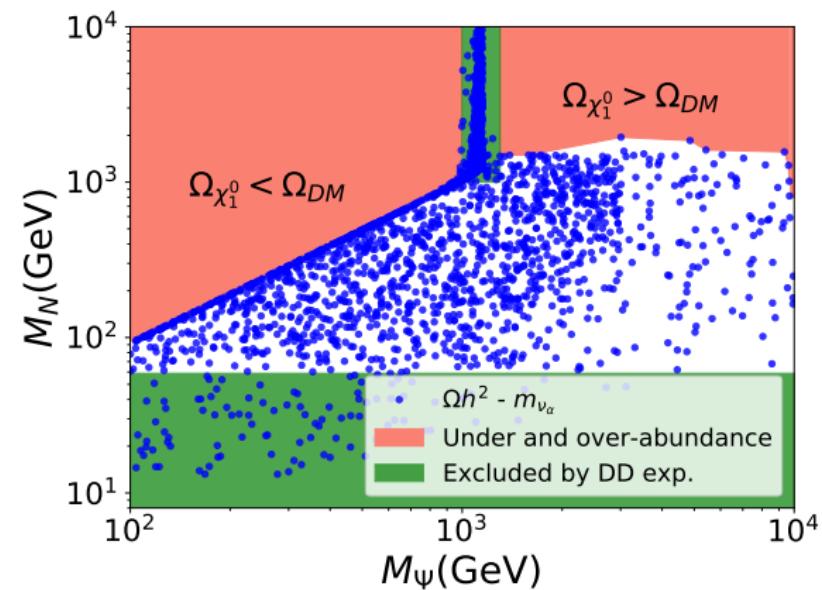


# Dark matter relic density

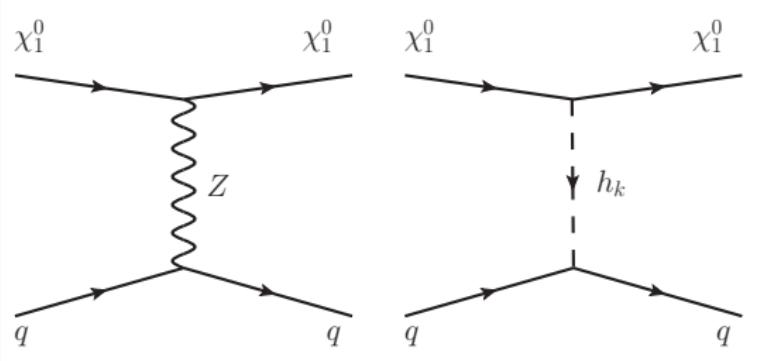


Decoupled  $Z'$  limit

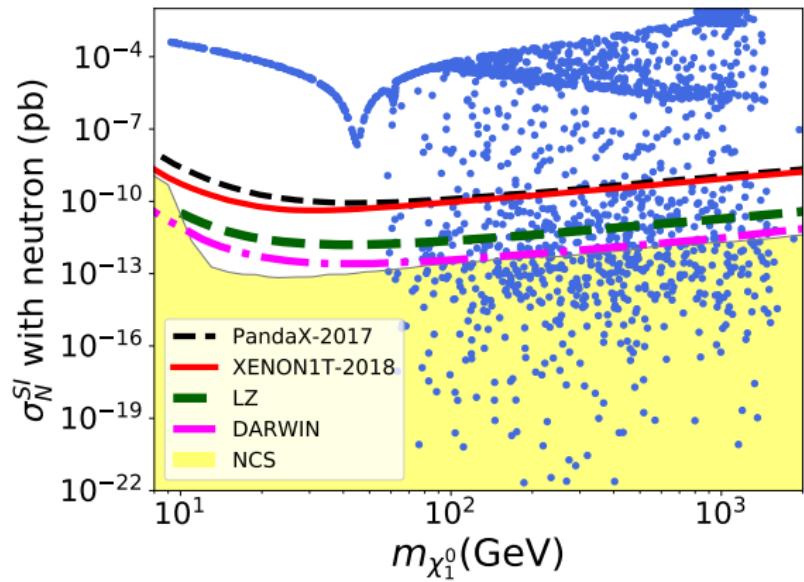
$$\begin{pmatrix} h \\ \text{Re}(S) \end{pmatrix} = \begin{pmatrix} \cos \alpha & \sin \alpha \\ -\sin \alpha & \cos \alpha \end{pmatrix} \begin{pmatrix} h_1 \\ h_2 \end{pmatrix}.$$



# Spin independent (SI) direct detection cross section



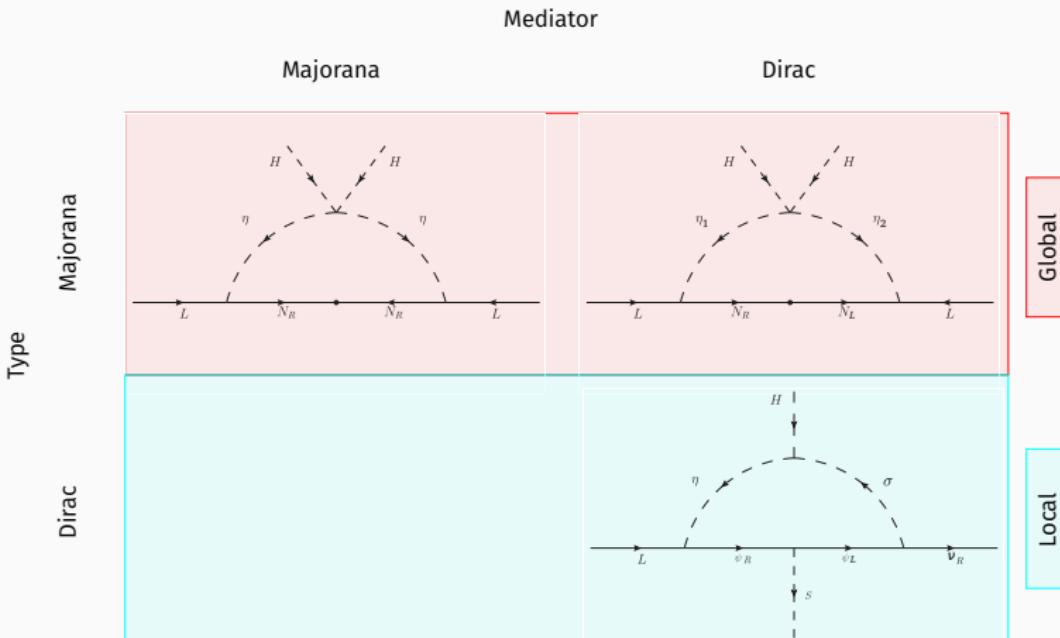
Decoupled  $Z'$  limit

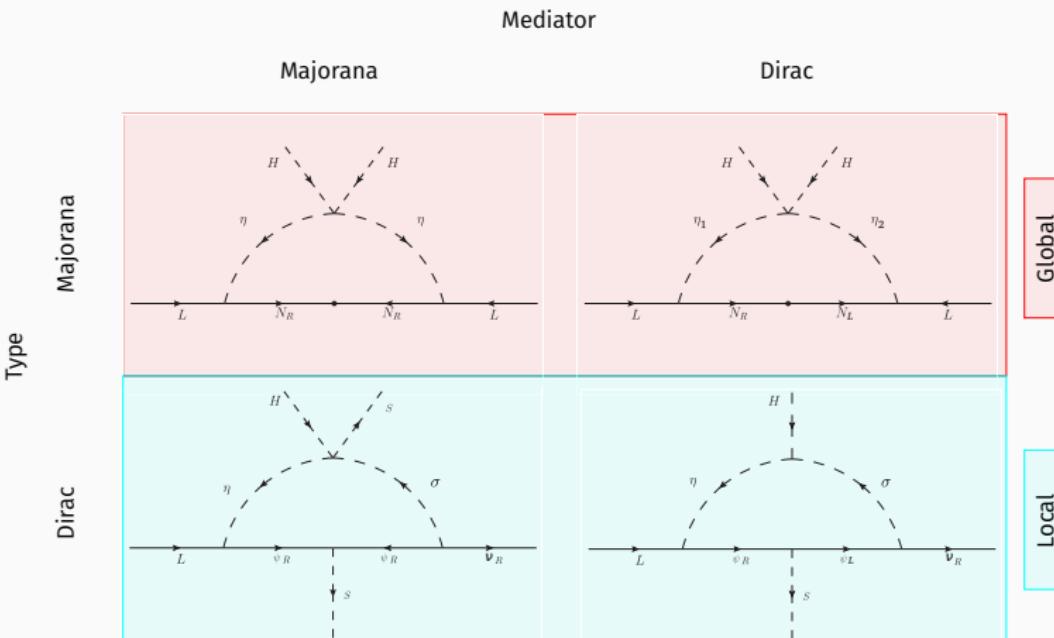


## Dark symmetry

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# Radiative Type-I seesaw





Fields	$SU(2)_L$	$U(1)_Y$	$U(1)_D^*$
$L$	2	-1/2	0
$Q$	2	-1/6	0
$d_R$	1	-1/2	0
$u_R$	1	+2/3	0
$e_R$	1	-1	0
$H$	2	1/2	0
$\eta$	2	1/2	1
$S$	1	0	2
$\sigma$	1	0	3
$\nu_{R1}$	1	0	4
$\nu_{R2}$	1	0	4
$\nu_{R3}$	1	0	-5
$\psi_{R1}$	1	0	-1
$\psi_{R2}$	1	0	-1
$\psi_{R3}$	1	0	-1
TOTAL			0

# Conclusions

It makes sense to focus our attention on models that can account for neutrino masses and dark matter (DM) **without adhoc symmetries**

## Tree-level type-I-II-III seesaws

Type-I seesaw: **Multicomponent dark matter**: In this extension of the SM by an  $U(1)_{B-L}$  gauge symmetry anomalies are canceled partially by two right-handed neutrinos and partially by two component DM Dirac fermions, providing a connection between neutrinos and DM analogous to that one between leptons and quarks in the SM.

The model predicts the existence of three scalar fields beyond the SM Higgs:  $H_1$ ,  $H_2$ ,  $A$

Model implemented in LanHEP. Implemented also in

SARAH <https://github.com/restrepo/BSM-Submodules/tree/B-L+DM/BSM/SARAH/Models/B-L/DM> (Tested with SARAH-4.14.1) to analyse perturbativity and stability conditions and higher scales with two-loop RGEs.

# Conclusions

It makes sense to focus our attention on models that can account for neutrino masses and dark matter (DM) **without adhoc symmetries**

## One-loop Dirac neutrino masses

A single  $U(1)_X$  gauge symmetry to explain both the smallness of Dirac neutrino masses and the stability of Dirac fermion dark matter

- Spontaneously broken  $U(1)_X$  generates a radiative Dirac neutrino masses
- A remnant symmetry makes the lightest field circulating the loop stable and good dark matter candidate.
- For T1-2-A: Either Singlet Doublet Dirac Dark Matter or Singlet Scalar Dark Matter with extra scalar and vector portal
- Dark symmetry for Majorana mediators

# Conclusions

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Thanks!