



**Autonomous Vehicle Simulation (AVS) Laboratory,  
University of Colorado**

**Basilisk Technical Memorandum**  
**Document ID: Basilisk-SphericalHarmonics**  
**SPHERICAL HARMONICS C++ MODEL**

Prepared by	S. Carnahan
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<b>Status:</b> Tested
<b>Scope/Contents</b>
The gravity effector module is responsible for calculating the effects of gravity from a body on a spacecraft. A spherical harmonics model and implementation is developed and described. A unit test has been written and run which test basic input/output, single-body gravitational acceleration, and multi-body gravitational acceleration.

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1.0	First version - Mathematical formulation and implementation	M. Diaz Ramos
1.1	Added test documentation	S. Carnahan

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## 1 Introduction

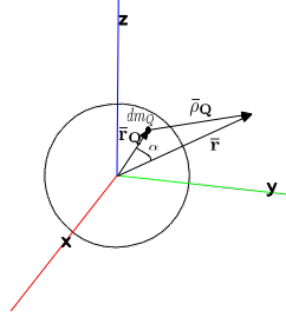
The gravity effector module is responsible for calculating the effects of gravity from a body on a spacecraft. A spherical harmonics model and implementation is developed and described below. The iterative methods used for the software algorithms are also described. Finally, the results of the code unit tests are presented and discussed.

## 2 Mathematical model

### 2.1 Gravity models

Gravity models are usually based on solutions of the Laplace equation ( $\nabla^2 U(\bar{\mathbf{r}}) = 0$ ). It is very important to state that this equation only models a gravity potential outside a body. For computing a potential inside a body the Poisson equation is used instead.

The spherical harmonic potential is a solution of the Laplace equation using orthogonal spherical harmonics. It can be derived solving the Laplace equation in spherical coordinates, using the separation of variables technique and solving a Sturm-Liouville problem. In this work, the solution will be found using another technique, which essentially follows Vallado's book.<sup>4</sup>



**Fig. 1:** Geometry of the Spherical Harmonics Representation.

For each element of mass  $dm_Q$  the potential can be written as

$$dU(\bar{\mathbf{r}}) = G \frac{dm_Q}{\rho_Q} \quad (1)$$

where  $\rho_Q$  is the distance between the element of mass and the position vector  $\bar{\mathbf{r}}$  where the potential is computed. This position vector is usually given in a body-fixed frame. The relation between the position vector  $\bar{\mathbf{r}}$ , the position of the element of mass  $\bar{\mathbf{r}}_Q$  and  $\rho_Q$  can be given using the cosine theorem and the angle  $\alpha$  between the two position vectors, as can be appreciated in Figure 1.

$$\rho_Q = \sqrt{r^2 + r_Q^2 - 2rr_Q \cos(\alpha)} = r \sqrt{1 - 2\frac{r_Q}{r} \cos(\alpha) + \left(\frac{r_Q}{r}\right)^2} = r \sqrt{1 - 2\gamma \cos(\alpha) + \gamma^2} \quad (2)$$

where  $\gamma = r_Q/r$ .

The potential can be obtained by integrating  $dU$  through the whole body.

$$U(\bar{\mathbf{r}}) = G \int_{body} \frac{dm_Q}{r \sqrt{1 - 2\gamma \cos(\alpha) + \gamma^2}} \quad (3)$$

If the potential is computed outside the body,  $\gamma$  will always be less than 1, and the inverse of the square root can be approximated using the Legendre polynomials  $P_l[\beta]$ .<sup>4</sup> Even though this derivation does not use the Laplace equation, it still assumes that the potential is computed outside the body.

The Legendre polynomials can be written as

$$P_l[\beta] = \frac{1}{2^l l!} \frac{d^l}{d\beta^l} (\beta^2 - 1)^l \quad (4)$$

The potential is

$$U(\bar{\mathbf{r}}) = \frac{G}{r} \int_{body} \sum_{l=0}^{\infty} \gamma^l P_l[\cos(\alpha)] dm_Q \quad (5)$$

The angle  $\alpha$  must be integrated. However, the cosine of the angle  $\alpha$  can be decomposed using the geocentric latitude and the longitude associated to vectors  $\bar{\mathbf{r}}$  and  $\bar{\mathbf{r}}_Q$ . These angles will be called  $(\phi, \lambda)$  and  $(\phi_Q, \lambda_Q)$  respectively. Using the addition theorem it is possible to write.<sup>4</sup>

$$P_l[\cos(\alpha)] = P_l[\sin(\phi_Q)]P_l[\sin(\phi)] + 2 \sum_{m=1}^l \frac{(l-m)!}{(l+m)!} (a_{l,m}a'_{l,m} + b_{l,m}b'_{l,m}) \quad (6)$$

where

$$a_{l,m} = P_{l,m}[\sin(\phi_Q)] \cos(m\lambda_Q) \quad (7)$$

$$b_{l,m} = P_{l,m}[\sin(\phi_Q)] \sin(m\lambda_Q) \quad (8)$$

$$a'_{l,m} = P_{l,m}[\sin(\phi)] \cos(m\lambda) \quad (9)$$

$$b'_{l,m} = P_{l,m}[\sin(\phi)] \sin(m\lambda) \quad (10)$$

where  $P_{l,m}[x]$  are the associated Legendre functions. " $l$ " is called degree and " $m$ ", order. The polynomials can be computed as

$$P_{l,m}[\beta] = (1 - \beta^2)^{\frac{m}{2}} \frac{d^m}{d\beta^m} P_l[\beta] \quad (11)$$

As can be seen,  $a_{l,m}$  and  $b_{l,m}$  must be integrated, but  $a'_{l,m}$  and  $b'_{l,m}$  can be taken outside the integral. Therefore, it is possible to define

$$C'_{l,m} = \int_{body} (2 - \delta_m) r_Q^l \frac{(l-m)!}{(l+m)!} a_{l,m} dm_Q \quad (12)$$

$$S'_{l,m} = \int_{body} (2 - \delta_m) r_Q^l \frac{(l-m)!}{(l+m)!} b_{l,m} dm_Q \quad (13)$$

where  $\delta_m$  is the Kronecker delta.

Then

$$U(\bar{\mathbf{r}}) = \frac{G}{r} \sum_{l=0}^{\infty} C'_{l,0} \frac{P_l[\sin(\phi)]}{r^l} + \frac{G}{r} \sum_{l=0}^{\infty} \sum_{m=1}^l \frac{P_{l,m}[\sin(\phi)]}{r^l} [C'_{l,m} \cos(m\lambda) + S'_{l,m} \sin(m\lambda)] \quad (14)$$

Non-dimensional coefficients  $C_{l,m}$  and  $S_{l,m}$  are usually used

$$C'_{l,m} = C_{l,m} R_{\text{ref}}^l m_Q \quad (15)$$

$$S'_{l,m} = S_{l,m} R_{\text{ref}}^l m_Q \quad (16)$$

where  $m_Q$  is the total mass of the body and  $R_{\text{ref}}$  is a reference radius. If the coefficients  $C_{l,m}$  and  $S_{l,m}$  are given, the reference radius must be specified. Usually, the reference is chosen as the maximum radius or the mean radius.<sup>3</sup>

The potential is then

$$U(\bar{\mathbf{r}}) = \frac{\mu}{r} \sum_{l=0}^{\infty} C_{l,0} \left( \frac{R_{\text{ref}}}{r} \right)^l P_l[\sin(\phi)] + \frac{\mu}{r} \sum_{l=0}^{\infty} \sum_{m=1}^l \left( \frac{R_{\text{ref}}}{r} \right)^l P_{l,m}[\sin(\phi)] [C_{l,m} \cos(m\lambda) + S_{l,m} \sin(m\lambda)] \quad (17)$$

Since  $P_l[x] = P_{l,0}[x]$  the potential can be written in a more compact way

$$U(\bar{\mathbf{r}}) = \frac{\mu}{r} \sum_{l=0}^{\infty} \sum_{m=0}^l \left( \frac{R_{\text{ref}}}{r} \right)^l P_{l,m}[\sin(\phi)] [C_{l,m} \cos(m\lambda) + S_{l,m} \sin(m\lambda)] \quad (18)$$

Some coefficients have a very interesting interpretation.

$$C_{0,0} = 1 \quad (19)$$

$$S_{l,0} = 0 \quad \forall l \geq 0 \quad (20)$$

$$C_{1,0} = \frac{Z_{\text{CoM}}}{R_{\text{ref}}} \quad (21)$$

$$C_{1,1} = \frac{X_{\text{CoM}}}{R_{\text{ref}}} \quad (22)$$

$$S_{1,1} = \frac{Y_{\text{CoM}}}{R_{\text{ref}}} \quad (23)$$

where  $[X_{\text{CoM}}, Y_{\text{CoM}}, Z_{\text{CoM}}]$  represents the center of mass of the celestial body. Therefore, if the origin of the coordinate system coincides with the center of mass, all these coefficients are identically zero. Similarly, the second order coefficients are related to the second order moments (moments of inertia).

Finally, the coefficients and Legendre polynomials are usually normalized to avoid computational issues. The factor  $N_{l,m}$  is called the normalization factor

$$N_{l,m} = \sqrt{\frac{(l-m)!(2-\delta_m)(2l+1)}{(l+m)!}} \quad (24)$$

The normalized coefficients are

$$\bar{C}_{l,m} = \frac{C_{l,m}}{N_{l,m}} \quad (25)$$

$$\bar{S}_{l,m} = \frac{S_{l,m}}{N_{l,m}} \quad (26)$$

The normalized associated Legendre functions are

$$\bar{P}_{l,m}[x] = P_{l,m}[x]N_{l,m} \quad (27)$$

The potential may be written as

$$U(\bar{\mathbf{r}}) = \frac{\mu}{r} \sum_{l=0}^{\infty} \sum_{m=0}^l \left( \frac{R_{\text{ref}}}{r} \right)^l \bar{P}_{l,m}[\sin(\phi)] [\bar{C}_{l,m} \cos(m\lambda) + \bar{S}_{l,m} \sin(m\lambda)] \quad (28)$$

## 2.2 Pines' Representation of Spherical Harmonics Gravity

There are many ways to algorithmically compute the potential and its first and secondary derivatives. One of such algorithms is the one proposed by Pines.<sup>2</sup>

The spherical harmonics representation as it was presented has a singularity at the poles for the gravity field. The Pines' formulation avoids this problem and is more numerically stable for high degree and high order terms.

Unfortunately, this formulation does not contain the normalization factor which is necessary if the coefficients are normalized. In a paper written by Lundberg and Schutz,<sup>1</sup> a normalized representation of the Pines' formulation is given, but it contains an approximation.

For this work, and in order to code the spherical harmonics formulation, a formulation similar to Pines' using the Lundberg-Schutz paper will be derived. However, no approximations will be used. Therefore, the algorithm will be developed here without using the exact formulations given in those

papers. For the sake of brevity, not every single derivation will be carried out, but it is possible to get the results following the expressions obtained in this section.

In the Pines' formulation the radius and the director cosines are used as coordinates. The potential will be given as  $U[r, s, t, u]$ , where

$$r = \sqrt{x^2 + y^2 + z^2} \quad (29)$$

$$s = \frac{x}{r} \quad (30)$$

$$t = \frac{y}{r} \quad (31)$$

$$u = \frac{z}{r} \quad (32)$$

For a function of these coordinates, the dependance will be given using square brackets (e.g.  $f[r, s, t, u]$ ).

Since  $u = \sin(\phi) = \cos(90^\circ - \phi)$ , it is possible to write

$$P_{l,m}[\sin(\phi)] = P_{l,m}[u] \quad (33)$$

The derived Legendre functions  $A_{l,m}[u]$  are defined such that

$$P_{l,m}[u] = (1 - u^2)^{\frac{m}{2}} A_{l,m}[u] \quad (34)$$

From the definition of  $P_{l,m}$  (Equation 11), it is possible to write

$$A_{l,m}[u] = \frac{d^m}{du^m} P_l[u] = \frac{1}{2^l l!} \frac{d^{l+m}}{du^{l+m}} (u^2 - 1)^l \quad (35)$$

The term  $(1 - u^2)^{\frac{m}{2}}$  can be written as  $(1 - \sin^2(\phi))^{\frac{m}{2}} = |\cos(\phi)|^m = \cos^m(\phi)$ .

If the complex number  $\xi$  is defined such that ( $j$  is the imaginary unit)

$$\xi = \cos(\phi) \cos(\lambda) + j \cos(\phi) \sin(\lambda) = \frac{x}{r} + j \frac{y}{r} = s + jt \quad (36)$$

it is possible to write

$$\xi^m = \cos^m(\phi) e^{jm\lambda} = (s + jt)^m \quad (37)$$

The following sequences may be defined

$$R_m[s, t] = \text{Re}\{\xi^m\} \quad (38)$$

$$I_m[s, t] = \text{Im}\{\xi^m\} \quad (39)$$

Putting all together, it is possible to write

$$U(\bar{\mathbf{r}}) = \frac{\mu}{r} \sum_{l=0}^{\infty} \sum_{m=0}^l \left( \frac{R_{\text{ref}}}{r} \right)^l A_{l,m}[u] \{C_{l,m} R_m[s, t] + S_{l,m} I_m[s, t]\} \quad (40)$$

In order to normalize the coefficients ( $\bar{C}_{l,m}$  and  $\bar{S}_{l,m}$ ) and the derived Legendre functions ( $\bar{A}_{l,m} = N_{l,m} A_{l,m}$ ), each term is divided and multiplied by the normalization factor  $N_{l,m}$ . Then

$$U(\bar{\mathbf{r}}) = \frac{\mu}{r} \sum_{l=0}^{\infty} \sum_{m=0}^l \left( \frac{R_{\text{ref}}}{r} \right)^l \bar{A}_{l,m}[u] \{\bar{C}_{l,m} R_m[s, t] + \bar{S}_{l,m} I_m[s, t]\} \quad (41)$$

The sets  $D_{l,m}[s, t]$ ,  $E_{l,m}[s, t]$ , and  $F_{l,m}[s, t]$ , are defined as

$$D_{l,m}[s, t] = \bar{C}_{l,m}R_m[s, t] + \bar{S}_{l,m}I_m[s, t] \quad (42)$$

$$E_{l,m}[s, t] = \bar{C}_{l,m}R_{m-1}[s, t] + \bar{S}_{l,m}I_{m-1}[s, t] \quad (43)$$

$$F_{l,m}[s, t] = \bar{S}_{l,m}R_{m-1}[s, t] - \bar{C}_{l,m}I_{m-1}[s, t] \quad (44)$$

The value  $\rho_l[r]$  is also defined as

$$\rho_l[r] = \frac{\mu}{r} \left( \frac{R_{\text{ref}}}{r} \right)^l \quad (45)$$

The gravity potential may be finally computed as

$$U(\bar{\mathbf{r}}) = \sum_{l=0}^{\infty} \sum_{m=0}^l \rho_l[r] \bar{A}_{l,m}[u] D_{l,m}[s, t] \quad (46)$$

This is the final expression that will be used to compute the gravity potential.

## 2.3 Recursion Formulas

Several recursion formulas are needed in order to algorithmically implement the Pines' formulation. They will be given without proof, but they are easily derived using the definitions above.

### 2.3.1 Recursion formula for $\rho_l[r]$

Initial condition:  $\rho_0[r] = \frac{\mu}{r}$

$$\rho_l[r] = \rho \cdot \rho_{l-1}[r] \quad (47)$$

where  $\rho = R_{\text{ref}}/r$ .

### 2.3.2 Recursion formula for $R_m[s, t]$

Initial condition:  $R_0[s, t] = 1$

$$R_m[s, t] = sR_{m-1}[s, t] - tI_{m-1}[s, t] \quad (48)$$

### 2.3.3 Recursion formula for $I_m[s, t]$

Initial condition:  $I_0[s, t] = 0$

$$I_m[s, t] = sI_{m-1}[s, t] + tR_{m-1}[s, t] \quad (49)$$

### 2.3.4 Recursion formula for $\bar{A}_{l,m}[u]$

From Equation (35), it is possible to see that

$$A_{l,l}[u] = (2l - 1)A_{l-1,l-1}[u] \quad (50)$$

$$A_{l,l-1}[u] = uA_{l,l}[u] \quad (51)$$

There are several recursion formulas for computing Legendre polynomials  $A_{l,m}[u]$ , for  $m < l - 1$ . The following formula, which is stable for high degrees,<sup>1</sup> will be used:

$$A_{l,m}[u] = \frac{1}{l - m} ((2l - 1)uA_{l-1,m}[u] - (l + m - 1)A_{l-2,m}[u]) \quad (52)$$

Using Equations (50), (51), and (52), and the definition  $\bar{A}_{l,m}[u] = N_{l,m}A_{l,m}[u]$ , the following recursion formulas can be derived.

Initial condition:  $\bar{A}_{0,0}[u] = 1$

The diagonal terms are computed as

$$\bar{A}_{l,l}[u] = \sqrt{\frac{(2l-1)(2-\delta_l)}{(2l)(2-\delta_{l-1})}} \bar{A}_{l-1,l-1}[u] \quad (53)$$

The low diagonal terms are then calculated as

$$\bar{A}_{l,l-1}[u] = u \sqrt{\frac{(2l)(2-\delta_{l-1})}{2-\delta_l}} \bar{A}_{l,l}[u] \quad (54)$$

Finally, for  $l \geq (m+2)$ ,  $N1_{l,m}$  and  $N2_{l,m}$  are defined such that

$$N1_{l,m} = \sqrt{\frac{(2l+1)(2l-1)}{(l-m)(l+m)}} \quad (55)$$

$$N2_{l,m} = \sqrt{\frac{(l+m-1)(2l+1)(l-m-1)}{(l-m)(l+m)(2l-3)}} \quad (56)$$

and  $\bar{A}_{l,m}[u]$  computed using

$$\bar{A}_{l,m}[u] = u N1_{l,m} \bar{A}_{l-1,m}[u] - N2_{l,m} \bar{A}_{l-2,m}[u] \quad (57)$$

## 2.4 Derivatives

The first order derivatives of many of the values given are necessary to compute the gravity field (second order derivatives are needed if the Hessian is to be computed).

It is easy to show that

$$\frac{\partial D_{l,m}}{\partial s}[s, t] = m E_{l,m}[s, t] \quad (58)$$

$$\frac{\partial D_{l,m}}{\partial t}[s, t] = m F_{l,m}[s, t] \quad (59)$$

$$\frac{d\rho_l}{dr}[r] = -\frac{(l+1)}{R_{\text{ref}}} \rho_{l+1}[r] \quad (60)$$

$$\frac{\partial R_m}{\partial s}[s, t] = m R_{m-1}[s, t] \quad (61)$$

$$\frac{\partial R_m}{\partial t}[s, t] = -m I_{m-1}[s, t] \quad (62)$$

$$\frac{\partial I_m}{\partial s}[s, t] = m I_{m-1}[s, t] \quad (63)$$

$$\frac{\partial I_m}{\partial t}[s, t] = m R_{m-1}[s, t] \quad (64)$$

$$\frac{d\bar{A}_{l,m}}{du}[u] = \frac{N_{l,m}}{N_{l,m+1}} \bar{A}_{l,m+1}[u] \quad (65)$$

The gravity field can be computed using all the equations given. However, the gradient of the potential is needed. As a change of variables was realized, the chain rule must be applied. In order



to avoid filling up pages with math derivations, the results will be given. With patience, the following results can be obtained applying the chain rule and using all the derivatives given.

The gravity field can be computed as

$$\bar{\mathbf{g}} = (a_1[r, s, t, u] + s \cdot a_4[r, s, t, u])\hat{\mathbf{i}} + (a_2[r, s, t, u] + t \cdot a_4[r, s, t, u])\hat{\mathbf{j}} + (a_3[r, s, t, u] + u \cdot a_4[r, s, t, u])\hat{\mathbf{k}} \quad (66)$$

where

$$a_1[r, s, t, u] = \sum_{l=0}^{\infty} \sum_{m=0}^l \frac{\rho_{l+1}[r]}{R_{\text{ref}}} m \bar{A}_{l,m}[u] E_{l,m}[s, t] \quad (67)$$

$$a_2[r, s, t, u] = \sum_{l=0}^{\infty} \sum_{m=0}^l \frac{\rho_{l+1}[r]}{R_{\text{ref}}} m \bar{A}_{l,m}[u] F_{l,m}[s, t] \quad (68)$$

$$a_3[r, s, t, u] = \sum_{l=0}^{\infty} \sum_{m=0}^l \frac{\rho_{l+1}[r]}{R_{\text{ref}}} m \frac{N_{l,m}}{N_{l,m+1}} \bar{A}_{l,m+1}[u] D_{l,m}[s, t] \quad (69)$$

$$a_4[r, s, t, u] = \sum_{l=0}^{\infty} \sum_{m=0}^l \frac{\rho_{l+1}[r]}{R_{\text{ref}}} m \frac{N_{l,m}}{N_{l+1,m+1}} \bar{A}_{l+1,m+1}[u] D_{l,m}[s, t] \quad (70)$$

In order to avoid computing factorials, it is easy to see that

$$\frac{N_{l,m}}{N_{l,m+1}} = \sqrt{\frac{(l-m)(2-\delta_m)(l+m+1)}{2-\delta_{m+1}}} \quad (71)$$

$$\frac{N_{l,m}}{N_{l+1,m+1}} = \sqrt{\frac{(l+m+2)(l+m+1)(2l+1)(2-\delta_m)}{(2l+3)(2-\delta_{m+1})}} \quad (72)$$

Using all these expressions, the potential and the gravity field can be computed.

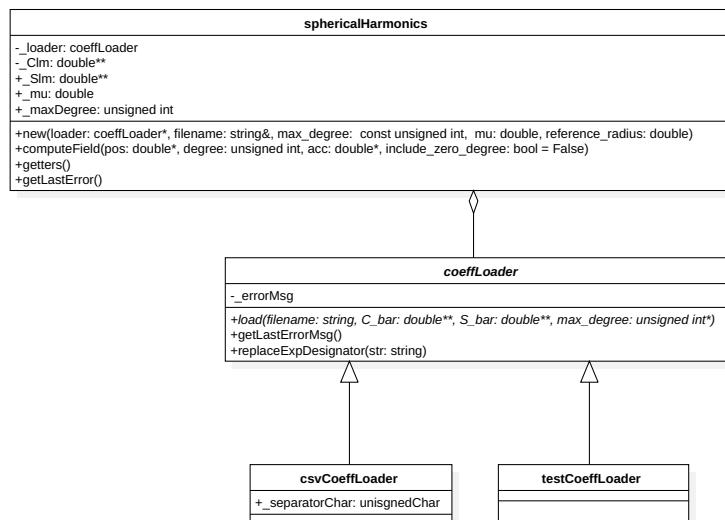


Fig. 2: UML class diagram.

### 3 Library

The library is structured around the class *sphericalHarmonics*, which computes the gravity field, and a family of classes which inherit from *coeffLoader* that load the coefficients from any source. The abstract class *coeffLoader* can be inherited if coefficient files with new formats are found.

A class diagram of the library can be seen in Figure 2. Figure 3 shows a use case diagram. In the typical use case, the user creates a *coeffLoader* object, and assigns it to the *sphericalHarmonics* object when its constructor is called. The class automatically loads the coefficients using *coeffLoader*'s *load()* interface and pre-computes all those values that are independent of the coordinates. The user then calls *computeField()* method to calculate the gravity field at a given planet-fixed position.

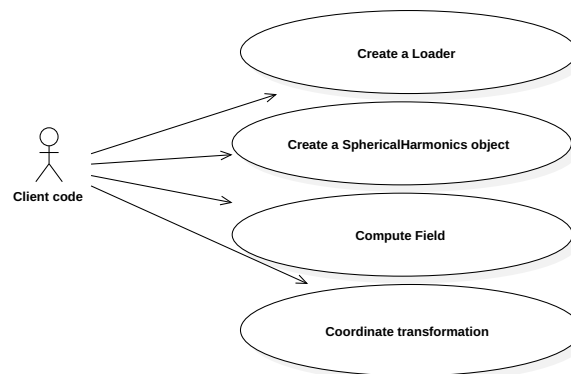


Fig. 3: UML use case diagram.

## 4 Unit Test

The unit test, `test_gravityDynEffector.py` validates the internal aspects of the Basilisk spherical harmonics gravity effector module by comparing module output to expected output. It utilizes spherical harmonics for calculation given the gravitation parameter for the massive body, a reference radius, and the maximum degree of spherical harmonics to be used. The unit test verifies basic set-up, single-body gravitational acceleration, and multi-body gravitational acceleration.

### 4.1 Test Description

This test is located in `SimCode/dynamics/gravityEffector/_UnitTest/test_gravityDynEffector.py`. In order to get good coverage of all the aspects of the module, the test is broken up into three sub-tests:

#### 4.1.1 Model Set-Up Verification

This test verifies, via three checks, that the model is appropriately initialized when called.

- 1.1 The first check verifies that the normalized coefficient matrix for the spherical harmonics calculations is initialized appropriately as a  $3 \times 3$  identity matrix.
- 1.2 The second check verifies that the magnitude of the gravity being calculated is reasonable (i.e. between  $9.7$  and  $9.9 \text{ m/s}^2$ ).
- 1.3 The final check ensures that the maximum degrees value is truly acting as a ceiling on the maximum number of degrees being used in the spherical harmonics algorithms. For example, if the maximum degrees value is set to 20, an attempt is made to call the spherical harmonics algorithms with a degrees value of 100 and another attempt is made with a degrees value of 20. The results are compared and should be equal due to the enforcement of the maximum degrees value.

#### 4.1.2 Single-Body Gravity Calculations

This test compares calculated gravity values around the Earth with ephemeris data from the Hubble telescope. The simulation begins shortly after 0200 UTC May 1, 2012 and carries on for two hours, evaluating the gravitational acceleration at two second intervals.

#### 4.1.3 Multi-Body Gravity Calculations

This test includes gravity effects from Earth, Mars, Jupiter, and the Sun. Calculations are verified by comparing against ephemeris data from the DAWN mission. The simulation begins at midnight, October 24th, 2008 and carries on for just under 22 minutes, evaluating the gravitational acceleration at two second intervals.

### 4.2 Test Parameters

This section summarizes the test parameters for each of the checks.

- Error Tolerance

There are specific error tolerances for each test. Error tolerances are determined based on whether the test results comparison should be exact or approximate due to integration or other reasons. Error tolerances for each test are summarized in table 2.

**Table 2:** Error tolerance for each test.

Test	Tolerated Error
Gravity Magnitude (check 1.2)	1.0e-01
Single-Body Gravity	1.0e-04
Multi-Body Gravity	5.0e-04

### 4.3 Test Results

All checks within `test_gravityDynEffector.py` passed as expected. Table 3 shows the test results.

**Table 3:** Test results.

Test	Pass/Fail	Notes
Setup Test	PASSED	
Single-Body Gravity	PASSED	
Multi-Body Gravity	PASSED	

### 4.4 Test Coverage

This test covers 2000% of the code.

## 5 Conclusion

What a great piece of code.

## REFERENCES

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