

## Evolutionary models of accretor stars in massive binaries: the case of $\zeta$ Ophiuchi

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### ABSTRACT

Most massive stars are born in binaries close enough for mass-transfer to occur. This modifies the appearance, internal structure, and future evolution of both stars, however the effects on the accretor are often overlooked. We model the evolution of a 100-day period binary consisting initially of a  $25 M_{\odot}$  and a  $17 M_{\odot}$ , which experiences stable mass transfer. We focus on the impact of mass accretion on the surface abundances, rotation rate, and internal structure of the accretor. To anchor our models, we show that our accretor star broadly reproduces the properties of  $\zeta$  Ophiuchi, a runaway thought to be the product of the disruption of a former binary system. We compare our accretor to single rotating stars and find that the later and stronger spin-up provided by binary interactions produces significant differences. Specifically, the core of the accretor retains a much higher spin, and its inner density profile is significantly different because of the appearance of an off-center convective shell during the mass transfer, which remains for the rest of the main sequence. Moreover, in our accretor surface pollution can occur because of the accreted material from the donor (which is progressively mixed inwards) rather than outward mixing of core material. Our models demonstrate the impact of mass accretion on the secondary star in a binary that is not captured in single rotating stellar models, with possible implications for its further evolution (either in a binary or as single stars), the final collapse, and the resulting spins of the compact object formed.

*Keywords:* stars: individual:  $\zeta$  Ophiuchi – stars: massive – stars: binaries

### 1. INTRODUCTION

The overwhelming majority of massive stars is born in multiple systems (e.g., Mason et al. 2009; Almeida et al. 2017; Moe & Di Stefano 2017), and a large fraction will exchange mass or merge with a companion in their lifetime (e.g., Sana et al. 2012). The most common type of interaction is a stable mass transfer through Roche lobe overflow (RLOF) after the end of the donor’s main sequence (case B, Kippenhahn & Weigert 1967), when stars experience most of their radial expansion. Many studies have focused on the dramatic impact these interactions have on the donor star (e.g., Morton 1960; Yoon et al. 2017; Götberg et al. 2017, 2018; Laplace et al. 2020, 2021; Blagorodnova et al. 2021). Often the accreting companion is treated as a point mass. However, binary interactions have a crucial impact on the initially less massive star too.

#### 1.1. *The importance of accretor stars*

During mass transfer, the initially less massive star is expected to accrete mass, spin-up to critical rotation (e.g., Packet 1981), and possibly be polluted by nuclearly processed material from the inner layers of the

donor star (e.g., Blaauw 1993). The increase in mass leads to growth of the convective core and consequent “rejuvenation” of the accretor (e.g., Neo et al. 1977; Schneider et al. 2016). Understanding the evolution of accretors in massive binaries has wide and crucial implications for stellar populations, electromagnetic transient observations, and gravitational-wave progenitors.

Accretor stars (and merger products) can appear as blue stragglers (e.g., Chen & Han 2009, 2010; Rain et al. 2021) and thus impact cluster populations, their age estimates, and their main sequence (e.g., Pols & Marius 1994; Wang et al. 2020) and post main sequence morphology (e.g., Wei et al. 2021). The high spin of the accretor post-mass-transfer might be the dominant explanation for the origin of Oe and Be stars (i.e., stars showing emission lines, e.g., Pols et al. 1991; Bodensteiner et al. 2020; Vinciguerra et al. 2020; Dorigo Jones et al. 2020; Wang et al. 2021; Hastings et al. 2021). Rotationally-mixed accretor stars at low metallicity have also been suggested to be important for the ionizing flux of high-redshift galaxies (e.g., Eldridge & Stanway 2012; Stanway et al. 2016).

After mass transfer, the majority of massive binaries will be disrupted by the first supernova, and eject the accretor (“binary SN scenario”, Blaauw 1961; De Donder et al. 1997; Eldridge et al. 2011; Boubert & Evans 2018; Renzo et al. 2019; Evans et al. 2020). A small fraction of these would be sufficiently fast to become runaway stars, but the majority will be too slow to stand out in astrometric surveys. Assuming a constant star formation history, Renzo et al. (2019) estimated that  $10.1^{+4.6\%}_{-8.6\%}$  ( $0.5^{+2.1\%}_{-0.5\%}$ ) of O-type stars could be slow “walkaway” (runaway) accretors ejected after the companion’s core collapses. Therefore, presently single O-type stars that accreted mass earlier on contribute to the populations of field massive stars (e.g., Dorigo Jones et al. 2020).

From the transients perspective, massive accretor stars are also important: Zapartas et al. (2019) estimated that  $14^{+4\%}_{-11\%}$  of hydrogen (H) rich type II SNe might come from progenitors ejected from a binary after the explosion of their companion. The fact that they accreted mass before exploding can influence their helium (He) core mass and thus the explosion properties and the inferred progenitors (Zapartas et al. 2021). The high post-mass transfer rotation rate of accretor stars in binaries might have implications for the formation of long gamma-ray burst progenitors (e.g., Cantiello et al. 2007).

The majority of isolated binary evolution scenarios for gravitational-wave progenitors include a common-envelope phase. This is initiated by the originally less massive star, who accretes mass from its companion before the formation of the first compact object (e.g., Belczynski et al. 2016; Tauris et al. 2017; Broekgaarden et al. 2021). Therefore, it is possible that accretion of mass before the formation of the first compact object could modify the internal structure of the star that will initiate the common-envelope (e.g., Law-Smith et al. 2020; Klencki et al. 2021). Specifically, the rotation rate, chemical composition, and innermost structure of the envelope (because of rejuvenation) will differ from a single star.

Despite their importance, accretor stars in binaries have so far received much less attention than the donor stars, with the pioneering studies of Ulrich & Burger (1976); Hellings (1983, 1984), and Braun & Langer (1995) as notable exceptions. Large grids of accretor models are lacking, most of the studies focus on lower mass systems (e.g.,  $M_1 \lesssim 16 M_\odot$  in van Rensbergen et al. 2011) or neglect the crucial impact of rotation (e.g., Sravan et al. 2019, but see also Wang et al. 2020) and only few sparse massive models exist (e.g., Petrovic et al. 2005; Cantiello et al. 2007).

Modeling accretor stars requires following the coupled evolution of two rotating stars exchanging mass. Making robust predictions is challenging because of the large number of free parameters necessary to model each individual star and their interactions. Population synthesis calculations based on single star models cannot include the effects of binary mass transfer on the internal structure, and rely on the implicit assumption that the accretor is sufficiently well described by a (possibly fast-rotating) single star model. Here, we compute detailed evolutionary models of both stars in a binary and compare our accretor to rotating single star models to test these assumptions.

## 1.2. A prototypical example of accretor: $\zeta$ Ophiuchi

The nearest O-type star to Earth,  $\zeta$  Ophiuchi ( $\zeta$  Oph) provides a unique opportunity to constrain massive binary evolution models.  $\zeta$  Oph has a distance from Earth of  $107 \pm 4$  pc (Neuhäuser et al. 2020, and references therein), and a spectral type O9.5IVnn (Sota et al. 2014). It occasionally shows emission lines, making it an Oe star (Walker et al. 1979; Vink et al. 2009). Its surface rotation rate is extremely large, with most estimates of the projected rotational velocity from optical spectra exceeding  $v \sin(i) \gtrsim 400 \text{ km s}^{-1}$  (corresponding to the “nn” in the spectral type, Zehe et al. 2018 and references therein). By comparing the observed  $v \sin(i) = 432 \pm 16 \text{ km s}^{-1}$  to the theoretical breakup rotation, Zehe et al. (2018) constrained the inclination angle to  $i \gtrsim 56$  degrees. Using optical interferometry, Gordon et al. (2018) measured the centrifugal distortion of  $\zeta$  Oph, finding a polar radius of  $7.5 R_\odot$  and centrifugally increased equatorial radius of  $9.1 R_\odot$ , corresponding to a  $v \sin(i) = 348 \text{ km s}^{-1}$ .

$\zeta$  Oph was originally identified as a runaway because of its large proper motion by Blaauw (1952). Unfortunately, the *Gaia* data for this object are not of sufficient quality<sup>1</sup> to improve previous astrometric results, but estimates of the peculiar velocity range in  $20 - 50 \text{ km s}^{-1}$  (e.g., Zehe et al. 2018; Neuhäuser et al. 2020). The large velocity with respect the surrounding interstellar material is also confirmed by the presence of a prominent bow-shock (e.g., Bodensteiner et al. 2018).

Because of its young apparent age, extremely fast rotation, and nitrogen (N) and He rich surface (e.g., Herrero et al. 1992; Blaauw 1993; Villamariz & Herrero 2005; Marcolino et al. 2009),  $\zeta$  Oph is a prime example of runaway from the binary SN scenario (Blaauw 1993). Many studies have suggested  $\zeta$  Oph might have accreted

<sup>1</sup> The renormalized unit weighted error (RUWE) of this star in Gaia EDR3 is 4.48.

mass from a companion before acquiring its large velocity, both from spectroscopic and kinematic considerations (e.g., Blaauw 1993; Hoogerwerf et al. 2000, 2001; Tetzlaff et al. 2010; Neuhäuser et al. 2020) and using stellar modeling arguments (e.g., van Rensbergen et al. 1996). Recently, Neuhäuser et al. (2020) suggested that  $1.78 \pm 0.21$  Myr ago a SN in Upper-Centaurus-Lupus produced the pulsar PSR B1706-16, ejected  $\zeta$  Oph, and also injected the short-lived radioactive isotope  $^{60}\text{Fe}$  on Earth. This argues strongly for a successful SN explosion of the companion with a  $\sim 250 \text{ km s}^{-1}$  natal kick, sufficient in most cases to disrupt the binary (e.g., Tauris 2015; Renzo et al. 2019; Evans et al. 2020).

Although the nature of  $\zeta$  Oph as a binary product is well established, because of its observed large surface rotation rate, previous attempts to model it rely purely on rotational mixing to explain the N- and He-rich surface composition (e.g., Maeder & Meynet 2000). Even the binary models of van Rensbergen et al. (1996) assumed rotational mixing from the inside of the accreting star driven by the spin-up during mass transfer (see also Cantiello et al. 2007). However, Villamariz & Herrero (2005) (hereafter, VH05) were unable to obtain a satisfying fit for the stellar spectra using the single-star rotating models from Meynet & Maeder (2000, 2003): by the time rotational mixing enriches the surface, single massive stars have significantly spin down through wind mass loss.

This may not be surprising: parametrized models of rotational mixing predict lower efficiency for metal-rich and relatively low mass stars. The reason is the increased importance of mean molecular weight gradients and the longer thermal timescales compared to more massive stars (e.g., Yoon et al. 2006; Perna et al. 2014). The parent association of  $\zeta$  Oph identified by Neuhäuser et al. (2020) has a metallicity  $Z = 0.01 \simeq Z_\odot$  (based on asteroseismology from Murphy et al. 2021), and mass estimates for  $\zeta$  Oph range from  $13 - 25 M_\odot$ , at the lower end of the range where rotational mixing can efficiently bring He and CNO-processed material to the surface (chemically homogeneous evolution, Maeder & Meynet 2000).

Given the challenges in explaining the surface composition of  $\zeta$  Oph as a rotating single star and the evidence for its past as a member of a binary system, this star offers a unique opportunity to constrain the evolution of accretors in massive binaries. Here, we present self-consistent binary evolution models of the coupled evolution of both stars and their orbit. After describing our MESA setup in Sec. 2, we present a model which reproduces the majority of the salient features of  $\zeta$  Oph in Sec. 3 and compare it extensively to fast rotating single

star models. We discuss the sensitivity of our results to the many free parameters in Sec. 6, before concluding in Sec. 7. ■ [check structure] ■

## 2. MODELING MASSIVE BINARIES WITH MESA

We follow self-consistently the coupled evolution of two massive stars in a binary system using MESA (version 15140, Paxton et al. 2011, 2013, 2015, 2018, 2019). Our choice of input parameters and our numerical results are available at <https://doi.org/10.5281/zenodo.4701565>. We discuss here only the main relevant parameters and describe the effects of varying some of the assumed values in Sec. 6. Appendix A gives more details on our choice of input physics, and appendix B discuss the numerical resolution in space and time.

We adopt the Ledoux (1947) criterion to determine convective stability and a mixing length parameter of 1.5. We allow for time-dependent convection as in Renzo et al. (2020) based on Arnett (1969). We include semi-convection and thermohaline mixing following Langer et al. (1983) and Kippenhahn et al. (1980), respectively, each with efficiency 1.0. We use the exponential core overshooting from Herwig (2000) with free parameters  $(f, f_0) = (4.25 \times 10^{-2}, 10^{-3})$  (Claret & Torres 2017) which broadly reproduces the width of the main sequence from Brott et al. (2011) for a  $16 M_\odot$  single star. We do not use over/undershooting for off-center convective shells. We also use the local implicit enhancement of the convective flux in superadiabatic regions introduced in MESA 15140 (`use_superad_reduction = .true.`).

We treat rotation in the “shellular” approximation (e.g., Zahn 1992; Ekström et al. 2012), that is we assume constant rotational frequency  $\omega$  along isobaric surfaces. Furthermore, we assume tidal synchronization and rigid rotation at the beginning of our runs, that is zero age main sequence (ZAMS). With the initial configuration that we assume motivated by  $\zeta$  Oph –  $M_1 = 25 M_\odot$ ,  $M_2 = 17 M_\odot$  and period of  $P = 100$  days – this means initially slow rotation: the surface averaged rotational velocity is  $\lesssim 3 \text{ km s}^{-1}$  for both stars. Our models include a diffusive approximation for meridional currents (Eddington-Sweet circulations, Sweet 1950), which dominate the chemical mixing due to rotation. We also include the secular and dynamical shear instabilities, and the Goldreich-Schubert-Fricke (GSF) instability (see Heger et al. (2000) for a review of these processes). We assume a Spruit-Tayler dynamo for the transport of angular momentum (Spruit 2002), and chose the same free parameters as Heger et al. (2000). This also includes the rotational enhancement of the wind mass loss as in Langer (1998). Moreover, when a star approaches critical rotation, we re-calculate im-

plicitly the wind mass loss rate to keep its rotation sub-critical. Specifically, at each timestep we calculate a wind enhancement factor to reduce the ratio  $\omega/\omega_{\text{crit}} \lesssim 0.95$  where  $\omega$  is the angular frequency and  $\omega_{\text{crit}} = \sqrt{(1 - L/L_{\text{Edd}})GM/R^3}$  with  $L_{\text{Edd}}$  the Eddington luminosity computed using the stellar opacity down to optical depth  $\tau = 2/3$ ,  $L$  is the luminosity,  $R$  the radius of the equator, and  $G$  the gravitational constant. However, we allow a tolerance of  $\pm 0.05$  on this limit.

We assume a metallicity of  $Z = 0.01$  informed by the present-day  $Z$  of  $\zeta$  Oph's parent cluster (Murphy et al. 2021), and assume the relative element abundances to scale with solar values (Grevesse & Sauval 1998). We include wind mass loss following Vink et al. (2000, 2001) for effective temperature  $T_{\text{eff}} \gtrsim 10^{4.2}$  K, and de Jager et al. (1988) for  $T_{\text{eff}} \lesssim 10^{4.2}$  K, both with a scaling factor of 1. Compared to the measurements for  $\zeta$  Oph, we might be over-estimating the wind mass-loss rate by a factor of 100 (“weak wind problem”, see Marcolino et al. 2009). This suggests that spin-down due to wind mass-loss may also be overestimated compared to reality. However, Lucy (2012) and Lagae et al. (2021) proposed that the temperature structure of the winds of low-luminosity O-type stars might affect the spectral lines and cause an empirical underestimate of the mass-loss rate.

We follow Kolb & Ritter (1990) to calculate the mass transfer rate with an implicit scheme. Moreover, we assume that the specific angular momentum and entropy of the transferred mass match the surface of the accretor, while the chemical composition is set by the stratification of the donor star. Mass transfer is conservative until the accretor reaches critical rotation, after which rotationally-enhanced mass loss governs the mass transfer efficiency (e.g., Petrovic et al. 2005). Transferred matter which is not successfully accreted carries away the specific angular momentum corresponding to the accretor’s orbital motion (e.g., Soberman et al. 1997; van den Heuvel et al. 2017).

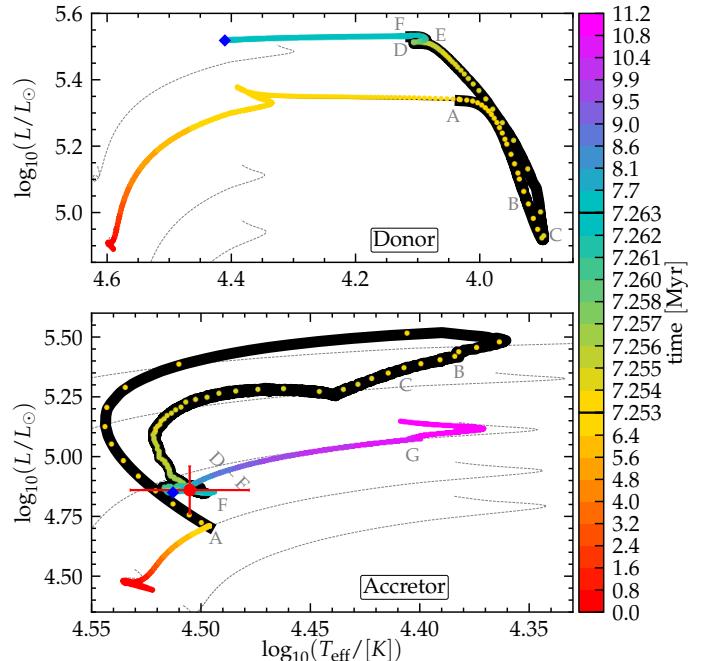
For simplicity, we evolve the accretor star as single after the mass transfer is completed. We define this based on the surface properties of the donor star, specifically, we require that it has lost most of its hydrogen-rich envelope – surface He mass fraction exceeds 0.35 – and its radius is smaller than both its Roche radius and the ZAMS radius. The stripped donor star is contracting at this point, and will likely develop strong wind mass loss, possibly appearing as a Wolf-Rayet star. Depending on the amount of leftover H-rich material and the wind mass-loss rate, it is possible that the donor would expand again later in the evolution, filling its Roche lobe anew (cf., Gilkis et al. 2019). However, this is not ex-

pected for the mass range we focus on here (Laplace et al. 2020, e.g.). We stop the evolution of the accretor star at terminal age main sequence (TAMS) defined as when the central mass fraction of H drops below  $10^{-4}$ .

To compare with our accretor model, we also compute the main sequence evolution of non-rotating single stars of 15, 17, 20, 25, and  $30 M_{\odot}$  and four more single  $20 M_{\odot}$  stars rigidly rotating with  $\omega/\omega_{\text{crit}} = 0.2, 0.3, 0.4$ , and 0.5 at birth. Apart from the initial rotation rate and the fact that they are single, the setup of these models is otherwise identical to our stars in the binary system.

### 3. EVOLUTION OF AN ACCRETOR STAR

We describe here the evolution of a binary system in which the accretor can broadly reproduce the observed features of  $\zeta$  Oph. We assume initial masses  $M_1 = 25 M_{\odot}$ ,  $M_2 = 17 M_{\odot}$ , and initial period  $P =$



**Figure 1.** HR diagram for the donor star (top) and accretor star (bottom) of the progenitor binary of  $\zeta$  Oph. Each point is separated by 50 years, and the black outline corresponds to the RLOF phase. The colors show the age on a non-uniform scale: we use smaller time-bins during RLOF ( $7.25 \text{ Myr} \lesssim t \lesssim 7.26 \text{ Myr}$ ). The red data point shows the position of  $\zeta$  Oph according to VH05, and the blue diamonds mark the end of the binary run. We continue the accretor evolution as a single star from there until core H depletion, hence the bottom panel shows a longer time. We emphasize the different scales on the two panels. The thin gray dashed line show the main sequence evolution of non-rotating single stars of  $15, 17, 20, 25$ , and  $30 M_{\odot}$  at  $Z = 0.01$  for comparison.

100 days (corresponding to a separation  $a \simeq 314 R_\odot$ ) with a metallicity of  $Z = 0.01$ .

Fig. 1 shows the Hertzsprung-Russell (HR) diagrams of both stars, the donor and accretor are shown separately on the top and bottom panel, respectively (see Fig. 8 for an HR diagram of both stars on the same scale). Each marker in the figure corresponds to an elapsed interval of 50 years in physical time.

After 7.24 Myr, the most massive star in the system evolves off the main sequence and 8400 years later, at point A in Fig. 1, it overfills its Roche lobe and starts to donate mass. This results in a stable case B RLOF on a thermal timescale from point A to F (black outline of the curves).

At the onset of RLOF (point A in Fig. 1), the accretor star is still on the main sequence with  $T_{\text{eff}} \simeq 10^{4.5}$  K and its central mass fraction of hydrogen is  $X(^1\text{H}) \simeq 0.42$ . The accretion of material drives the star out of thermal equilibrium and it quickly becomes over-luminous to radiate away the excess internal energy. The accretor reaches  $L \simeq 10^{5.5} L_\odot \gg L_{\text{nuc}} \simeq 10^{5.1} L_\odot$ , with  $L_{\text{nuc}}$  the total energy released per unit time by nuclear burning (integrated throughout the star). However, the duration of this phase is only  $\lesssim 2 \times 10^3$  years, corresponding roughly to the thermal timescale of the outer envelope of the accretor. During the same phase, the radius of the accretor increases dramatically from  $\sim 7.5 R_\odot$  to  $\sim 35 R_\odot$ .

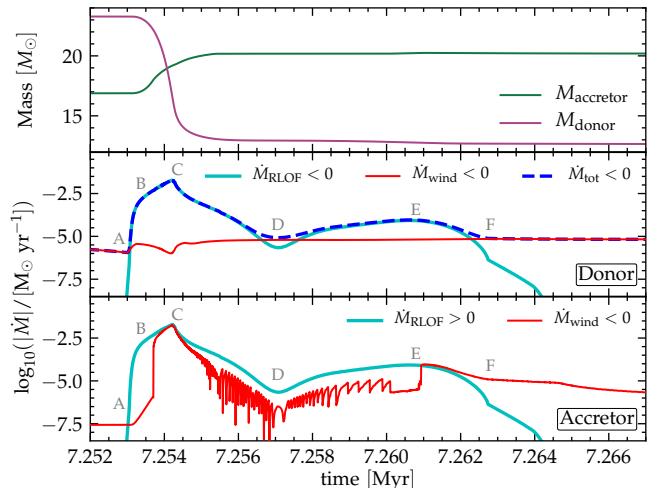
At point B – roughly corresponding to the lowest  $T_{\text{eff}}$  – the accretor reaches critical rotation, which briefly decreases the efficiency of mass transfer. This allows the star to contract, and increase its  $T_{\text{eff}}$ . The contraction also increases  $\omega_{\text{crit}}$ , allowing for further accretion to resume.

At  $T_{\text{eff}} \simeq 10^{4.43}$  K, slightly after point C, the inner layers of the donor's envelope are exposed. These layers were part of the convective core earlier on, before it receded in mass coordinate, therefore the transferred material becomes progressively more He-rich and CNO-processed. The difference in composition of the incoming material affects the opacity and mean molecular weight in the outer layers of the accretor and causes a kink in its evolutionary track. Specifically, material with a high mean molecular weight  $\mu$  is placed on top of the comparatively low- $\mu$  primordial envelope of the accretor. Because of the increasing gradient in mean molecular weight, thermohaline mixing starts in the outer layers of the accreting star, and, together with rotational mixing, it progressively dilutes the surface  ${}^4\text{He}$  and  ${}^{14}\text{N}$  mass fractions. Because of numerical reasons, this process causes noisy features from point D to F on the HR diagram of the accretor (e.g., Cantiello et al. 2007).

From D to E the donor star briefly expands again: by point D the surface is He-rich, and partial recombination of  ${}^4\text{He}$  drives a convective layer which is extremely thin in mass ( $\lesssim 10^{-4} M_\odot$ ) but can expand to significantly large radii<sup>2</sup>.

We emphasize that our adopted (standard) choices to model mixing and rotation are likely to impact the morphology of the accretor's evolutionary track during RLOF. The entire duration of RLOF from A to F is only about  $10^4$  years - of the order of the thermal timescale of the donor star. Moreover, the accretor spends most of this time close to the final, post-RLOF position (blue diamond in the bottom panel). Therefore, observation of a (population of) mass-transferring binary(ies) is unlikely to provide a direct probe of the accuracy of our mixing treatment. We expand on the mixing processes inside the accretor in Sec. 5.

### 3.1. Mass and mass transfer rate evolution



**Figure 2.** The top panel shows the total mass of each star as a function of time during RLOF. The middle and bottom panel show the contributions to the rate of mass change for the donor and accretor star, respectively. The cyan solid lines show the mass transfer rate due to RLOF, the red lines show the (mechanically enhanced) wind mass loss rates. In the middle panel, the dashed blue lines shows their sum, corresponding to the total mass loss rate of the donor. During RLOF the accretor reaches critical rotation, which leads to oscillations in the rotationally-enhanced wind mass loss.

Fig. 2 shows the mass evolution (top panel) and the rate of mass change (middle and bottom panel) for each individual star during the mass transfer phase.

<sup>2</sup> With previous MESA releases, we found it challenging to compute models beyond this phase: the large increase in radius impacts significantly the mass transfer rate.

The donor star (middle panel) loses mass via RLOF ( $\dot{M}_{\text{RLOF}} < 0$ , cyan line) and wind ( $\dot{M}_{\text{wind}} < 0$ , thin red line). The dashed blue line shows the sum of these two negative terms and represents the total rate of mass change of the donor. Conversely, the accretor (bottom panel) gains mass via RLOF (i.e.,  $\dot{M}_{\text{RLOF}} > 0$  from the accretor's point of view), but still loses mass to the wind ( $\dot{M}_{\text{wind}} < 0$ ). At peak (point C), the mass transfer rates reaches values above  $10^{-2} M_{\odot} \text{ yr}^{-1}$  and taps into the optically thick matter of the donor (i.e., the donor's Roche radius becomes smaller than its photosphere  $R_{\text{RL},1} < R_1$ ).

Initially, between point A and B, the mass transfer rate equals the mass accretion rate (bottom panel of Fig. 2), that is initially the accretion is (by construction) fully conservative. The bulk of the mass that is accreted,  $\sim 2 M_{\odot}$  out of  $\sim 3.1 M_{\odot}$ , is accreted during this initial phase, which lasts about  $\sim 2 \times 10^3$  years. As the mass and surface rotation rate of the accretor increase, the assumed rotational-enhancement of the wind progressively increases the mass-loss rate by  $\sim 5$  orders of magnitude. The mechanically-enhanced wind controls the accretion efficiency, and at  $\sim 7.254$  Myr (from B to C, where the red solid line and the cyan line overlap) the mass transfer becomes briefly non-conservative. In our setup, the majority of the mass transferred during this phase is ejected as a fast wind from the accretor, carrying the same specific angular momentum as the orbit of the accretor star. The decreased accretion efficiency allows, the accretor to contract (point B to C, cf. Fig. 1). In the remaining evolution from C to F, the interplay between the wind mass loss rate, the spin-up due to accretion and the spin-down due to inward transport of angular momentum (see Sec. 4) cause large oscillations in the wind mass loss rate. Nevertheless, for most of the evolution, accretion still occurs, albeit not-fully conservatively. This allows for CNO-processed material from the donor to reach the surface of the accretor during late stages of mass transfer.

The minimum of the mass transfer rate in point D corresponds to a brief phase of contraction (see also Fig. 1). However, from D-E, the donor star briefly expands again. This is due to the partial recombination of the now He-rich outer layers, which causes a transient surface convection layer. The convective layers expand, leading to an increase in the mass-transfer rate, despite at this point the donor is less massive than the accretor and the binary orbit is widening. During this phase, the mass transfer becomes highly non-conservative again (in the bottom panel the wind and the mass accretion rate nearly cancel each other again at  $\sim 7.261$  Myr), until the

donor completely detaches from its Roche lobe at point F.

### 3.2. Massive binary evolution naturally explains $\zeta$ Ophiuchi's properties

The mass and orbital velocity of our accretor star model agrees well with the measured values for the presently single O-star  $\zeta$  Oph. At the end of our binary evolution, well after the donor detaches from its Roche lobe (blue diamonds in Fig. 1) the accretor is a H-rich fast-rotating star of  $\sim 20.1 M_{\odot}$ . This matches well with the estimates for  $\zeta$  Oph, which although highly uncertain, typically include  $20 M_{\odot}$  in their range (e.g., Hoogerwerf et al. 2001, VH05, Neuhauser et al. 2020). The post-RLOF orbital velocity of the accretor is  $v_2 \simeq 52 \text{ km s}^{-1}$ . In the subsequent evolution, wind mass loss from both stars will widen the binary and decrease the orbital velocity of the accretor star. As a test, we evolved one binary assuming the Nugis & Lamers (2000) wind mass loss rate for the stripped donor until its He core depletion. At that point the accretor star's orbital velocity has decreased to  $\sim 40 \text{ km s}^{-1}$ , and it is expected to decrease even further during the remaining evolution. The precise amount of the orbital widening and decrease of the accretor's orbital velocity depends on the very uncertain stripped star mass loss rate (e.g., Vink 2017; Sander & Vink 2020). Nevertheless, the value we obtain is in broad agreement with estimates of the observed runaway velocity of  $\zeta$  Oph.

Accounting for both wind mass loss and the amount of mass transferred, at the end of our binary run (blue diamonds in Fig. 1) the donor becomes a He star of  $\sim 9.4 M_{\odot}$ , likely to contract further. Depending on its wind mass-loss rate, the stripped donor's spectrum might show absorption lines, emission lines, or a mixture of both (e.g., Crowther 2007; Neugent et al. 2017; Götberg et al. 2018). Its surface H mass fraction is  $\lesssim 0.2$  and even this residual H-rich layer might possibly be removed by further wind mass loss (e.g., Yoon et al. 2017; Götberg et al. 2017; Laplace et al. 2020).

In the evolution beyond the blue diamond in Fig. 1 (computed as a single star) the accretor settles on a main-sequence track at higher luminosity compared to the original track because of the accretion of mass, and its slope is slightly steeper due to the close-to-critical rotation and the accretion of partially nuclear processed (He- and N-rich) material.

The effective temperature, bolometric luminosity and kinematic age of  $\zeta$  Oph are also reasonably well reproduced by our model. After detachment, the donor star has approximately 0.5 Myr left until core-collapse, which likely will disrupt the binary system and eject the ac-

accretor star (e.g., Renzo et al. 2019). As seen by the colors of the track in Fig. 1, our accretor model spends around 2 Myr within the error bars for  $\zeta$  Oph estimated by VH05. This means that after the donor explodes, the accretor star will look similar to  $\zeta$  Oph for around 1.5 Myr, in good agreement with the kinematic age of  $1.78 \pm 0.21$  Myr for  $\zeta$  Oph (Neuhäuser et al. 2020).

Moreover, according to our model, the present-day age of  $\zeta$  Oph is  $\sim 9.5$  Myr. The age of the parent association of  $\zeta$  Oph, Upper-Centaurus-Lupus, is relatively uncertain, with estimates from pre-main sequence isochrone fitting of about  $9 \pm 1$  Myr, but an average age of  $15 \pm 3$  Myr (Pecaut & Mamajek 2016). However, given the sensitivity of our model to rejuvenation (see Sec. 5.2), the large age scatter of the region, and the unknown systematics in age measurements, we consider our model broadly compatible with the existing constraints.

In the following sections, we discuss in detail the rotational profile (Sec. 4) and surface composition (Sec. 5) of our accretor model. As a short summary, we find that both the surface rotation and chemical abundances of  $\zeta$  Oph are more easily explained by an accretor model than by single rotating models.

#### 4. ANGULAR MOMENTUM TRANSPORT AND ROTATION

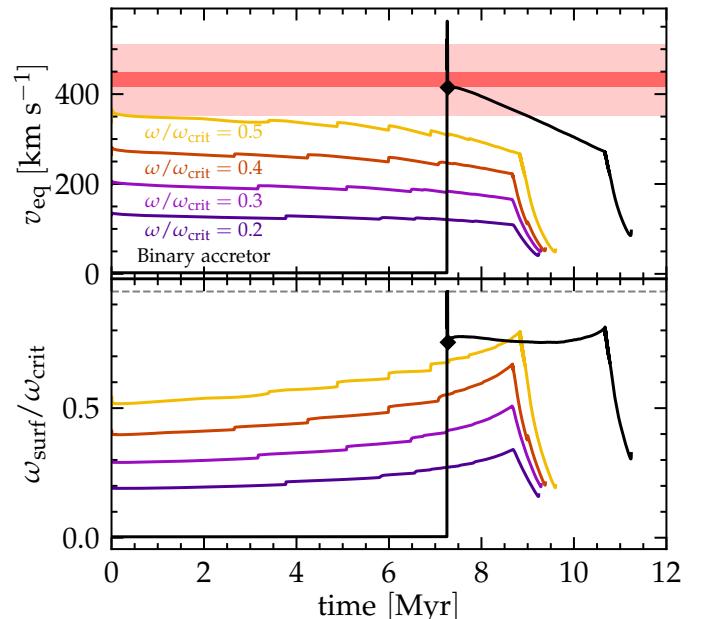
##### 4.1. Surface rotation

Rapid rotation is one of the main properties expected as a result of spin up due to mass accretion (Packet 1981). One of the main distinguishing features of  $\zeta$  Oph is its extremely high surface rotation rate. The black line in the top panel of Fig. 3 shows the evolution of the surface equatorial rotational velocity  $v_{\text{eq}}$  for our accretor model, not including any projection effect. More precisely,  $v_{\text{eq}}$  is a mass-weighted average of the rotational velocity of layers with opacity  $\tau \leq 100$ . The dark horizontal red band corresponds to the  $v \sin(i) = 432 \pm 16$  km s $^{-1}$  measured for  $\zeta$  Oph by Zehe et al. (2018), and the lighter band shows a range of 5 times their error bar, which roughly encloses the majority of the estimated  $v \sin(i)$  for  $\zeta$  Oph in the literature  $350 \text{ km s}^{-1} \lesssim v \sin(i) \lesssim 600 \text{ km s}^{-1}$  (e.g., Walker et al. 2005; Gordon et al. 2018). For comparison, the colored solid lines show also  $v_{\text{eq}}$  for models of rotating  $20 M_{\odot}$  single stars with birth spins of  $\omega/\omega_{\text{crit}} = 0.2, 0.3, 0.4$  and  $0.5$ . The bottom panel of Fig. 3 shows instead the ratio of the surface rotational frequency  $\omega_{\text{surf}}$  to the critical rotational frequency  $\omega_{\text{crit}}$ .

The initial binary is wide enough that assuming tidal synchronization at ZAMS implies a very low  $v_{\text{eq}} \lesssim 3$  km s $^{-1}$ . At 7.25 Myr, mass transfer rapidly spins up

the accretor to critical rotation,  $v_{\text{crit}} \sim 520$  km s $^{-1}$ . This corresponds to  $\omega/\omega_{\text{crit}} \simeq 0.95$  (dashed horizontal line in the bottom panel of Fig. 3), which is the upper-limit we impose in our models (see Sec. 2).

The star remains fast rotating throughout the mass transfer phase, and the remaining evolution in a binary which ends at the black diamond in Fig. 3 (corresponding to the blue diamonds in Fig. 1). In the remaining evolution, the star spins down progressively through wind mass loss, and within  $\sim 2$  Myr its averaged surface rotational velocity drops below  $\sim 350$  km s $^{-1}$ . Both the single star models and our accretor (after being spun up) evolve to higher  $\omega/\omega_{\text{crit}}$  because of the increase in stellar radii and corresponding decrease in  $\omega_{\text{crit}}$  (e.g., Langer 1998; Zhao & Fuller 2020). However, our accretor model remains at a higher  $\omega_{\text{surf}}/\omega_{\text{crit}} \simeq 0.75$  for a significantly longer time: the chance of observing a single rotating star at very high  $\omega_{\text{surf}}/\omega_{\text{crit}}$  is lower than for an accretor from a massive binary system. Moreover,  $\omega_{\text{surf}}/\omega_{\text{crit}}$  we find is in good agreement with the



**Figure 3.** Equatorial surface rotational velocity (top panel) and  $\omega/\omega_{\text{crit}}$  (bottom panel) for the accretor model (black) and single rotating  $20 M_{\odot}$  stars (colored lines). At  $\sim 7.25$  Myr the mass transfer quickly spins up the accretor to critical rotation: the dashed horizontal line in the bottom panel shows the upper-limit we impose. By the time the donor detaches from the RLOF the accretor is still spinning at  $\sim 400$  km s $^{-1}$ . From this point (diamonds) onwards, we continue the evolution as a single star, and the accretor spins down because of wind mass loss. Note however that we use the wind mass-loss rate from Vink et al. (2001), which might be  $\sim 2$  orders of magnitude too high for  $\zeta$  Oph (Marcolino et al. 2009).

observationally constrained values for typical Be stars (see [Rivinius et al. 2013](#), for a review).

Close to the end of the main sequence, the increase in wind mass loss rate strengthens the surface spin-down. This effect is also seen in the late main-sequence evolution of single stars rotating rapidly at birth.

[Fig. 3](#) shows that our model retains a significant surface rotation for a long period of time, comparable to the kinematic age of  $\zeta$  Oph. Since the spin up of the accretor happens roughly half-way through its main-sequence, the star is much faster rotating than single stars of the same (post-RLOF) mass initialized as fast rotators at ZAMS. Although [Fig. 3](#) does not account for the projection angle, [Zehe et al. \(2018\)](#) argued for  $i \geq 56$  degrees, corresponding to an upward shift of the red band in [Fig. 3](#) of  $\lesssim 20\%$ . This shift impacts the comparison of  $\zeta$  Oph to our accretor model and to single star models in the same way.

We emphasize that our model is computed using the [Vink et al. \(2000, 2001\)](#) wind mass-loss rate with full efficiency throughout its evolution. This is two orders of magnitude higher than the wind mass loss rate reported by [Marcolino et al. \(2009\)](#) (weak wind problem, however, see also [Lucy 2012](#); [Lagae et al. 2021](#)). While this may impact the evolution of the binary even before RLOF, it presumably increases the spin-down rate of our model compared to the observations. We expect that an accreting star modeled with lower wind-mass loss rate post-RLOF would retain an even higher surface rotation for longer (see also Sec. 6.1).

#### 4.2. Interior structure and rotation profile

Our model suggests that the internal rotational profile of accretor stars evolve differently compared to those of rotating single stars. To illustrate the angular momentum transport in our accretor stars, it is again helpful to compare with single rotating  $20 M_{\odot}$  models. Typically, single stars are assumed to be solid-body rotators at ZAMS (e.g., [Maeder & Meynet 2000](#)).

The top row of [Fig. 4](#) shows the internal rotational velocity,  $v_{\text{rot}} = \omega \times r$ , while the bottom row shows the angular frequency profile  $\omega$  as a function of mass coordinate. The first four panels in each row show single rotating stars with increasing initial  $\omega/\omega_{\text{crit}}$ , the last panels shows the accretor model. The gray area in the right-most panels highlights matter accreted during RLOF. We use colors to show the age of the star at each profile shown. The accretor starts as a  $17 M_{\odot}$  and is rejuvenated by binary interactions (see Sec. 5), reaching TAMS at 11.2 Myr. For comparison, the lifetime of a non-rotating single star of  $17 M_{\odot}$  is  $\sim 11.1$  Myr, while the initially  $20 M_{\odot}$  rotating models have a main-sequence

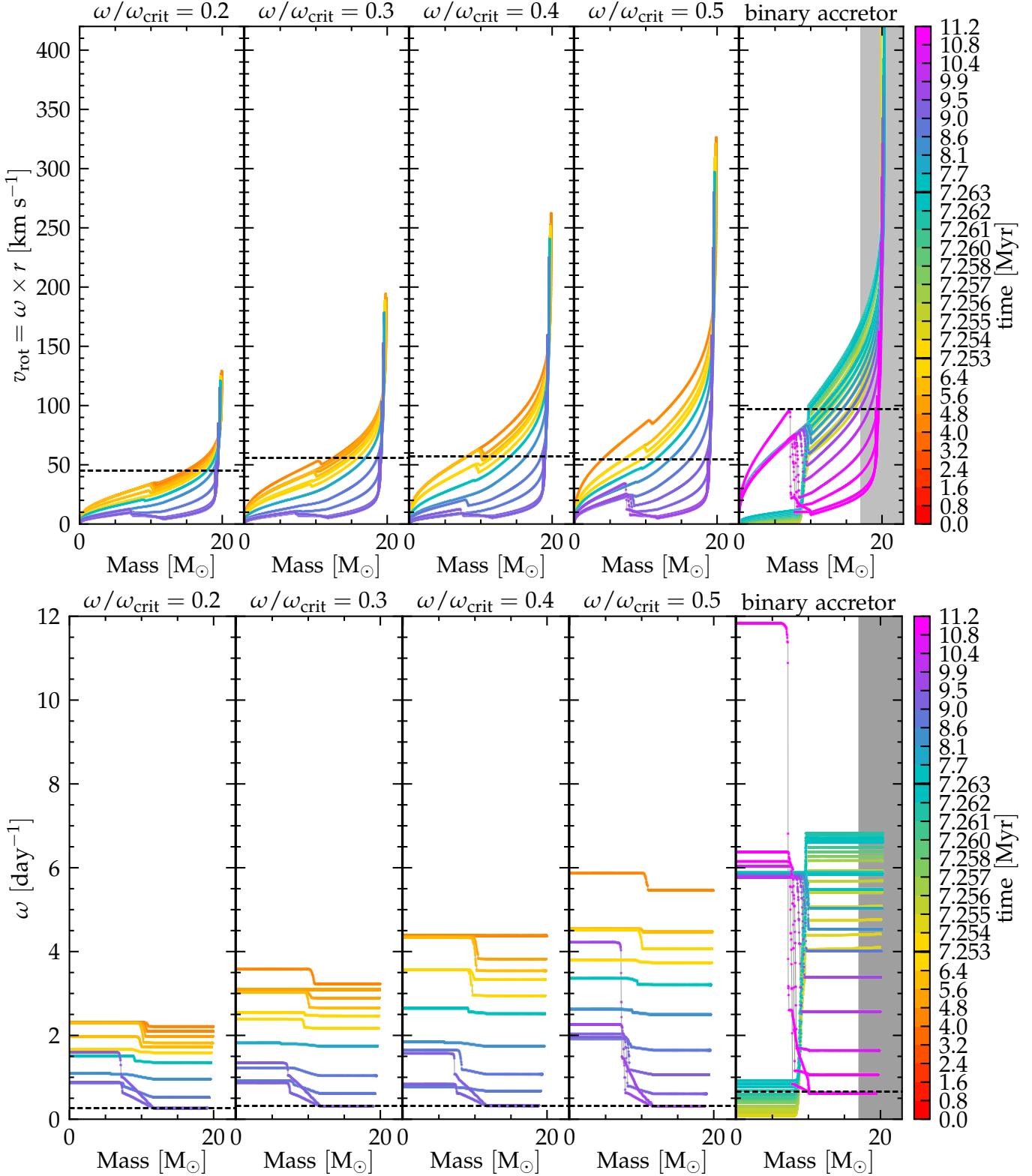
lifetime of  $\sim 9.2$ - $9.6$  Myr (longer for higher initial rotation rates), hence their TAMS profile are a shade of purple rather than pink in [Fig. 4](#).

The thin black horizontal dashed lines in each panel of the top row of [Fig. 4](#) mark the TAMS surface rotation rates: all our single star models reach a TAMS surface  $v_{\text{rot}} \simeq 50 - 60 \text{ km s}^{-1}$ . Initially faster rotating models spin down more in their outer layers, have slightly longer main sequence lifetimes (because of rotational mixing increasing the available fuel), and develop stronger differential rotation. As the core contracts and spins up, the single star profiles show the progressive development of a core-envelope interface.

Conversely, the entire interior of the accretor has a negligible rotational velocity until RLOF (starting at  $\sim 7.25$  Myr). Because of binary mass transfer, the accretor is spun up from the surface inwards, late in its evolution, and it reaches critical rotation  $\omega/\omega_{\text{crit}} \simeq 1$ . In our model, inward transport of angular momentum creates a  $v_{\text{rot}}$  profile monotonically decreasing from the surface to the center. After the end of mass transfer, roughly at 7.27 Myr, the accretor achieves rigid and close to critical rotation (flat profiles in the last panel on the bottom row of [Fig. 4](#),  $\omega \sim 6 \text{ day}^{-1}$ ). At the end of our binary run, the accretor is still rigidly rotating, and persists<sup>3</sup> for a total duration of  $\sim 10^4$  years. Afterwards, the accretor's envelope spins down because of winds and its evolutionary expansion. By the end of the accretor's main-sequence, the surface still spins with  $v_{\text{rot}} \simeq 100 \text{ km s}^{-1}$ , which is approximately twice as fast as the single star models.

Conversely, in the remaining evolution, the core contracts (decrease in  $r$ ). The weak core-envelope angular momentum coupling provided by the Spruit-Taylor dynamo leads to an approximately constant total angular momentum in the core, therefore, as it contracts and decreases its momentum of inertia, its rotation rate  $\omega$  increases dramatically. At the end of the main sequence, the outer-edge of the core of the accretor star has a similar rotational velocity as the surface, and much larger than for the single star models. Consequently, the TAMS core-envelope interface for the accretor is much more prominent than for in single rotating stars. It might be possible to distinguish accretors from initially single stars by using asteroseismology to measure the core rotation rate (e.g., [Cantiello et al. 2014](#)). However, this requires the detection of mixed modes which is presently challenging for massive stars.

<sup>3</sup> To calculate the duration, we consider rotation to be rigid if the difference between the minimum and maximum frequency throughout the star is  $\Delta\omega \lesssim 10^{-2} \text{ days}^{-1}$ .



**Figure 4.** Top: Internal rotational profile for  $20 M_{\odot}$  single star models with increasing  $\omega/\omega_{\text{crit}}$  at birth (first four panels), and for the accretor of our fiducial binary. Bottom: internal rotational frequency profile. As in Fig. 1, the colorbar is non-uniform. In the top (bottom) panel the thin dashed black lines mark the TAMS surface rotation rate (surface rotation frequency). In the rightmost panels, the gray areas indicate mass accreted during RLOF. The pink lines (TAMS) in the last panel show that the core of the accretor is rotating almost as fast as its surface despite its much smaller radius, and both are faster than the surface of single star models.

Moreover, the higher core spin of accretors may have important implications for their future explosions (e.g. MacFadyen & Woosley 1999; Cantiello et al. 2007), the spin of the resulting compact objects, and the analysis of gravitational-wave events (e.g., Zaldarriaga et al. 2018; Qin et al. 2018; Callister et al. 2020).

## 5. MIXING AND COMPOSITION

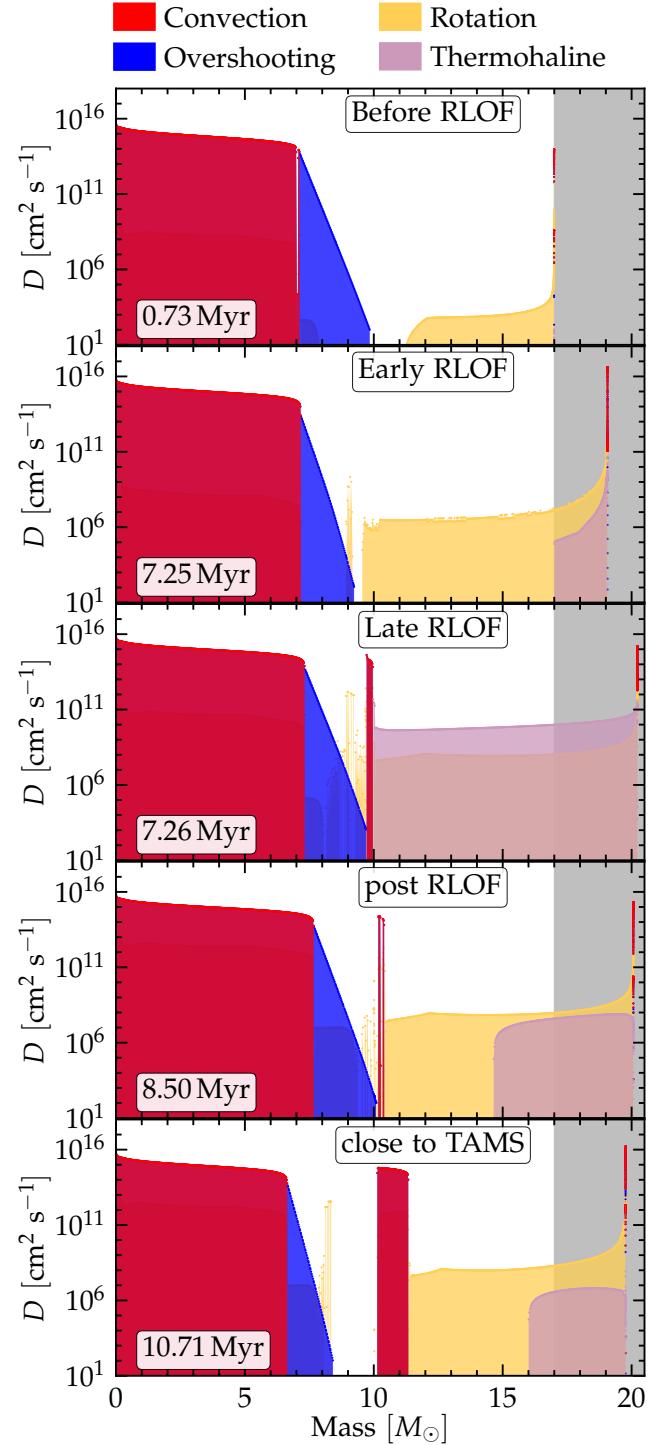
### 5.1. Mixing processes

Rotational mixing and thermohaline mixing induced by mass accretion affect significantly the composition profile of our model. They both act primarily in the envelope of the accretor, starting from the surface and growing inwards. MESA treats mixing using a diffusion approximation (Paxton et al. 2011), and to illustrate the dominant processes we show in Fig. 5 the diffusion coefficients for the elemental mass fractions as a function of Lagrangian mass coordinate at selected times. From top to bottom, we show the mixing contribution in the interior of our accretor star model at:

- main sequence before RLOF (i.e., before point A in Fig. 1 and Fig. 2);
- early during RLOF (between B and C);
- late during RLOF (between point C and D);
- post-interaction structure (within the red error-bars for  $\zeta$  Oph in Fig. 1);
- close to TAMS (point G in Fig. 1).

In each panel of Fig. 5, the gray background highlights mass coordinates exceeding the initial mass of the star. The remaining colors show convection (red), overshooting (blue), rotational mixing (yellow), and thermohaline mixing (pink). Rotational mixing includes all the rotational instabilities that we consider – meridional currents, secular and dynamical shear instabilities, and GSF instability. However, throughout the evolution rotational mixing is strongly dominated by the meridional currents (Eddington-Sweet circulations). The only exception is at the interface between core and envelope (i.e., at the outer edge of the overshooting region), where the post-RLOF contraction and spin-up of the core (see Sec. 4) can lead to significant dynamical shear (and GSF mixing to a lesser extent). For clarity, we do not show semiconvection which never dominates the mixing in our accretor model.

The top panel shows the typical structure of a main sequence massive star: the convective core initially reaching  $\sim 7 M_{\odot}$  with the overshooting extension to  $\sim 9 M_{\odot}$ . The slow rotation imparted by the assumed initial tidal



**Figure 5.** Mixing diffusion coefficients in the accretor star as a function of mass (the center corresponds abscissa 0 and the surface to the maximum abscissa for which a diffusion coefficient is plotted). The gray area on the right highlights accreted material. From top to bottom, each panel shows a profile during the main sequence (before point A in Fig. 1), early during RLOF (close to point C), mid-RLOF (close to point D), late during RLOF (between point D and E), and after RLOF (point G). A movie of the entire evolution is available at <https://doi.org/10.5281/zenodo.4701565>.

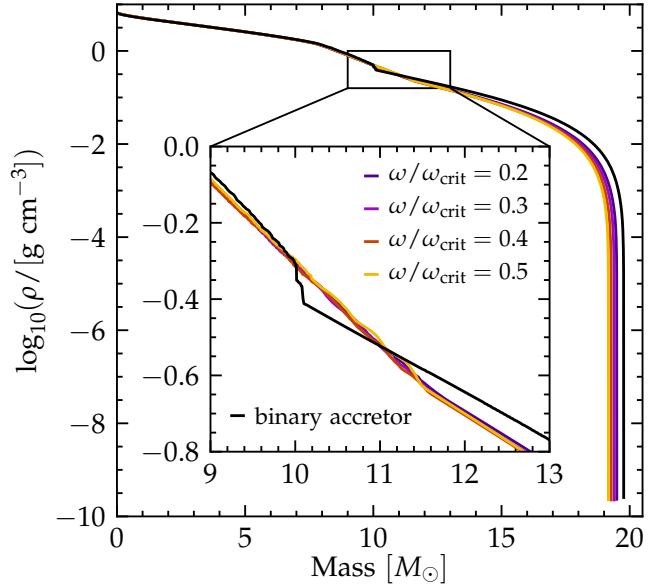
synchronization causes a meridional circulation in the envelope. Meridional circulations are also present in the core throughout the evolution, but with a diffusivity more than nine orders of magnitude lower than core convection. A small sub-surface convective zone is also appreciable at the very surface (see, e.g., Cantiello et al. 2021).

In the second panel from the top, the star has already accreted  $\sim 2 M_{\odot}$  (extending in the gray region), including some CNO-processed material from the inner layers of the core, and its surface is already spun up to  $\sim 330 \text{ km s}^{-1}$ . Thermohaline mixing has started in the newly accreted layers, but it is subdominant compared to rotational mixing due to meridional circulations in the envelope. Angular momentum transport (by the Spruit-Tayler dynamo) has already imparted some rotation to the inner layers of the envelope. This leads to the disconnected spike in rotational mixing at the outer edge of the core (at mass coordinate  $\sim 9 M_{\odot}$ ), dominated by dynamical shears. We note that between the top panel and the onset of RLOF, the convective core recedes in mass coordinate, but, by the time shown in the second panel, the accretion of mass has caused the core to grow back to its initial size.

In the third panel, the star has already accreted all the mass that it will, including CNO-enriched layers from the donor’s interior. Thermohaline mixing takes over the dominant role in the envelope (although meridional circulations persist behind it, with a mixing coefficient two orders of magnitude smaller). The core is still growing in mass, meaning that rejuvenation is still on-going. The mixing at the core-envelope boundary at  $\sim 9 - 10 M_{\odot}$  is partially still due to dynamical shear, but a thin, off-center, convective region has also appeared. Because we adopt an exponentially decreasing overshooting, the diffusion coefficient at the outer edge of the overshooting region is small, and therefore the mixing between the core and the off-center convective region is weak. We do not include over/undershooting for off-center convective layers, which could increase the coupling between these layers.

After detachment (the fourth panel in Fig. 5, roughly corresponding to  $\zeta$  Oph’s structure today), thermohaline mixing is progressively shutting down from inside out, while meridional circulations remain on.

Finally, the last panel displaying the interior mixing shortly before TAMS shows that, as the post-RLOF evolution proceeds, the previously thin, off-center convective layer grows thicker, including almost  $1.5 M_{\odot}$  of material between the mass coordinates  $10 - 11.5 M_{\odot}$ . In addition, the hydrogen-burning core has receded, leaving a  $\sim 1.5 M_{\odot}$  thick layer that is un-affected by mixing,



**Figure 6.** Comparison of the density profiles of the accretor model (black) and single  $20 M_{\odot}$  rotating stars (colors). The models are compared when they reach the same central H mass fraction of  $X_c = 0.085$  (point G in Fig. 1 for the accretor). The inset magnifies the region above the core, where the outcome of common envelope events is decided. Because of the growth of the core, the density profile of the accretor in this region is significantly different.

meaning that the core and the convective layer have disconnected.

To illustrate the effect of such off-center convective layer, Fig. 6 shows a comparison of the density profiles between the accretor model (point G in Fig. 1) and our four initially  $20 M_{\odot}$  single rotating stars when they reach the same central H mass fraction. Convection significantly alters the density profile above the core of the accretor compared to a single star. The sharper inner density drop and shallower density profile of accretors may be possible to explore with asteroseismology:  $\zeta$  Oph itself has been observed to show non-radial pulsations (Walker et al. 2005), which may be involved in the transient appearance of emission lines and a decretion disk. If compared to a star of comparable mass which evolved as single, the pulsations could shed light on the structural differences between single stars and binary products.

Moreover, if the binary system remains bound after the explosion of the donor and the evolutionary expansion of the accretor leads to a common envelope (e.g., Paczynski 1976), the layer above the He core is crucial to determine the success or failure of the common envelope ejection. Although this is not the future fate of  $\zeta$  Oph, this might be crucial for our understanding of gravitational wave progenitors. Common envelope sim-

ulations have so far neglected the impact of previous RLOF phase(s) on the density structure of the stars initiating the dynamically unstable mass transfer.

■ [from here to the next sec. feels like discussion] ■ YG: A convective layer above the hydrogen burning core that develops during main-sequence evolution is uncommon for  $20 M_{\odot}$  single star models at  $Z=0.01$  (see however REFs). Apart from affecting the composition profile and therefore also the density profile of the star if it develops in a region that previously belonged to the main-sequence core, it is possible that it affects the future evolution of the star. Convective layers in this region of a star have been associated with late expansion and helium core burning as blue supergiants (Schootemeijer+19, Klencki+20a), which in turn can hamper future binary interaction or cause nuclear timescale mass transfer and long lasting X-ray binary phases (Marchant+21). Considering the role of accretor stars similar to our model as progenitors of merging binary neutron stars (Tauris+17, Vigna-Gomez+20), it will be important to account for these structural changes when investigating the evolutionary pathways resulting in gravitational wave events. Post-main sequence evolution of the accretor star is a next step and the topic of our future study. YG: Note: we probably want to cite Groh+19, just haven't read it carefully enough yet.

The core growth caused by the accretion of mass, the consequent ingestion in the burning region of fresh H-rich material, and the resulting rejuvenation of the accretor are strongly dependent on the presence of convective boundary mixing. YG: I don't understand what is the convective boundary mixing - do you mean the convective layer above the core? In our model, this is initially provided by overshooting and to a lesser extent dynamical shear (between the top and second panel), at later times, the development of a convective layer above the overshooting region dominates this mixing. YG: Are you absolutely sure about the origin of this convective layer? I have had many discussions about convective layers similar to this one and I have not yet been convinced by any explanation for its origin. This could be different. What proves that it was provoked by overshooting and shear? Wait, I think I may be misunderstanding what we are talking about?

Instead, in the envelope, Eddington-Sweet circulations first dominate, then thermohaline mixing takes over – roughly from the kink after point C in Fig. 1, corresponding to the accretion of CNO-processed material, until the end of RLOF. The combination of these two processes mixes CNO-processed material from the donor star from the surface inward, and CNO-processed material from the core of the accretor outward. YG: Do we

want to show a composition profile? This sounds like the composition of the accretor star is not as the initial composition in any part of the star - is that correct? It may show interesting gradients too...

YG: In summary, the star accretes about  $3.4 M_{\odot}$  material during mass transfer (out of  $\sim 10.6 M_{\odot}$  total transferred). This causes rejuvenation and comparing to the single stellar tracks in the HR diagram (Fig. 1) the star appears to be X Myr after detachment, instead of the actual 7.3 Myr. After detachment, the core and the envelope are disconnected and the core rigidly rotates rapidly at a roughly constant spin rate until TAMS, while the envelope spins down but also rotates rigidly. YG: Here, I am thinking we can check what is the age of the  $20 M_{\odot}$  single star model that the track overlaps with at detachment.

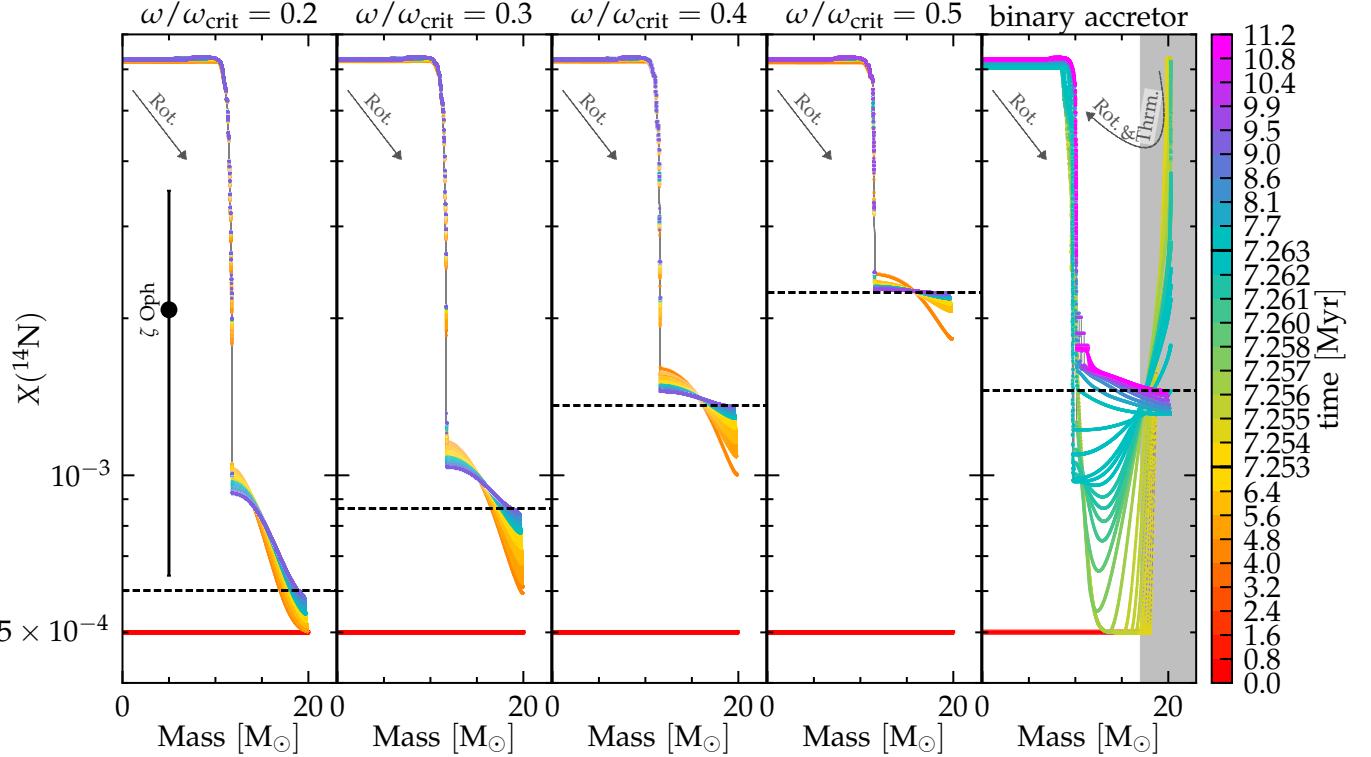
## 5.2. Chemical Composition

We now compare the composition profiles in our accretor and in single rotating massive star models. The first four panels of Fig. 7 show the mass fraction of  $^{14}\text{N}$  as a function of mass coordinate along the evolution of four  $20 M_{\odot}$  stars initialized with  $\omega/\omega_{\text{crit}} = 0.2, 0.3, 0.4, 0.5$ . The last panel of Fig. 7 shows our accretor model. While we focus here on  $^{14}\text{N}$ , Fig. 9 shows a similar plot containing also  $^{12}\text{C}$  and  $^{16}\text{O}$ .

The first four panels of Fig. 7 show the typical rotational mixing profiles:  $^{14}\text{N}$  rapidly rises by about an order of magnitude in the core because of the CNO burning, and it is then mixed outwards (as indicated by the arrows in the top left corner of each panel). At any time the  $^{14}\text{N}$  profile is monotonically decreasing with mass coordinate, and the higher the initial rotation, the higher the surface  $^{14}\text{N}$  mass fraction reached at TAMS because of the more efficient rotational mixing.

Conversely, the  $^{14}\text{N}$  mass fraction profile of the accretor is *not* monotonic throughout the evolution. The profile is shaped by two main processes: (i) accretion of CNO-processed material from the donor star mixed inwards by meridional circulations and thermohaline mixing (as indicated by the top right arrow in the last panel), (ii) outward mixing of the CNO-processed material from the accretor's core caused by rotational mixing – as in the single stars – and rejuvenation.

Initially, the tidally synchronized accretor star rotates too slowly ( $\lesssim 3 \text{ km s}^{-1}$ ) for significant outward rotational mixing out of the core, and until the onset of RLOF (roughly at 7.25 Myr) no appreciable variation of the envelope  $^{14}\text{N}$  mass fraction occurs. During late RLOF after the “kink” feature between C and D in Fig. 1, N-enriched material from the donor's core piles onto the accretor's surface – inside the gray area.



**Figure 7.**  $^{14}\text{N}$  mass fraction as a function of mass coordinate for  $20 M_{\odot}$  single star models with increasing  $\omega/\omega_{\text{crit}}$  at birth (first four panels), and for the accretor of our fiducial binary (last panel). The colorbar is non-uniform and allows for more color variation during RLOF. In each panel, the red flat line marks the primordial value for  $Z = 0.01$ , the thin dashed black line marks the surface value at TAMS. In the last panel, the gray area highlights mass accreted during RLOF. The black errorbar in the first panel shows the surface  $^{14}\text{N}$  of  $\zeta$  Oph estimated by VH05 assuming the surface mass fraction of H from Tab. 1. The abundance of  $^{14}\text{N}$  alone is not strongly constraining.

The close-to-critical rotation (through Eddington-Sweet circulations) and the inversion in the mean molecular weight  $\mu$  (through thermohaline mixing) drive inward mixing of the N-rich material and dilute it in the envelope (see also Fig. 5).

Simultaneously, the mere growth in mass causes the steepening of the core-temperature gradient and increase in the convective core mass (rejuvenation, e.g., Schneider et al. 2016), driving some outward convective mixing of N-rich material. The evolution of the structure modified by accretion also causes the formation of an off-center convective region (cf. Fig. 5) which persists at least until TAMS, when we stop our model. Because the convective turnover timescale is much shorter than evolutionary timescales, convection homogenizes the composition of this region and produces “steps” at the outer edge of the core (slightly outside mass coordinate  $10 M_{\odot}$ ).

### 5.3. Comparison to $\zeta$ Oph

■ [this section is now floating] ■

In Fig. 7, the black errorbar in the first panel shows  $\zeta$  Oph’s surface  $^{14}\text{N}$  from VH05 (assuming the surface H mass fraction from our model listed in Tab. 1). The mass fraction of  $^{14}\text{N}$  alone is not sufficient to distinguish between these models, and already a moderate  $\omega/\omega_{\text{crit}} \geq 0.3$  is sufficient for models to reach the lower-limit of the error bar.

Tab. 1 summarizes the surface properties of the accretor star at 8.5 Myr (fourth panel of Fig. 5), roughly corresponding to  $\zeta$  Oph’s position on the HR diagram today. Both the mass and radius agree reasonably well with the estimates from VH05 and previous studies, that is  $20 M_{\odot}$  and  $8.3 \pm 1.5 R_{\odot}$ , respectively. Our radius of  $9.8 R_{\odot}$  is larger by  $\sim 0.6 R_{\odot}$  than the equatorial radius recently measured by Gordon et al. (2018), and our model has a  $T_{\text{eff}} \simeq 31\,300$  K, on the lower end of the range considered by VH05. On the other hand, the age of 8.5 Myr is maybe slightly lower than the estimate for  $\zeta$  Oph’s age (the kinematic age plus the entire donor’s lifetime). The surface rotational velocity in excess of  $350 \text{ km s}^{-1}$  is also in the correct ballpark albeit possibly on the low end.

$M [M_{\odot}]$	$R [R_{\odot}]$	$\omega [\text{days}^{-1}]$	$v_{\text{rot}} [\text{km s}^{-1}]$	$X(^1\text{H})$	$X(^4\text{He})$	$X(^{12}\text{C})$	$X(^{14}\text{N})$	$X(^{16}\text{O})$
20.1	9.8	4.575	361.4	0.678044	0.312058	0.001339	0.001346	0.004149

**Table 1.** Surface properties of the accretor close to the present-day position of  $\zeta$  Oph on the HR diagram, corresponding to a physical age of 8.50 Myr, shortly after the end of RLOF (between the blue diamond and the lower estimate of  $\zeta$  Oph’s  $T_{\text{eff}}$  in Fig. 1, fourth panel of Fig. 5).

We report the surface H mass fraction<sup>4</sup>, lower than primordial because of the accretion of nuclearly processed material, and the surface mass fraction of the most prominent species  $^4\text{He}$ ,  $^{12}\text{C}$ ,  $^{14}\text{N}$ ,  $^{16}\text{O}$ . Assuming our surface H mass fraction  $X(^1\text{H})$ , the corresponding mass fractions of  $^4\text{He}$ ,  $^{12}\text{C}$ ,  $^{14}\text{N}$ ,  $^{16}\text{O}$  obtained by VH05 are  $0.34^{+0.14}_{-0.05}$ ,  $0.0006 \pm 0.0004$ ,  $0.002 \pm 0.001$ , and  $0.005 \pm 0.004$ .

By the accretor’s TAMS, rotational mixing (in the form of Eddington-Sweet circulations) and thermohaline mixing nearly homogenize the composition of the envelope of our accretor’s model. The surface mass fractions we obtain are sensitive to the interplay between several poorly understood processes treated in one dimension: mass accretion efficiency, rotationally enhanced wind mass loss, thermohaline, and inward rotational mixing. These also impact the composition of the envelope, and thus its radius and  $T_{\text{eff}}$ . Therefore, although not perfect, we consider the match with the mass fractions reported by VH05 surprisingly satisfactory.

## 6. ROBUSTNESS OF THE MODELS AND DISCUSSION

Models of the interior evolution of stars require the use of several poorly constrained parameters, most arising from the one-dimensional representation of multi-dimensional phenomena (convection, mixing, rotation, etc.). This remains true when modeling two stars in a binary, with the added caveat that an even larger number of parameters enters in the treatment of binary interactions (and in particular mass transfer). This emphasizes the need for observational constraints and motivated us to compare our models to the observationally well characterized  $\zeta$  Oph.

Accretor stars are expected in most populations of (massive) stars, both in clusters (e.g., Chen & Han 2009; Wang et al. 2020) and in the field (e.g., de Mink et al. 2011, 2013). However, these might not obviously stand out as binary products in kinematics surveys (e.g., Renzo et al. 2019). Therefore, to inform the search for accretor stars in observed samples, it is also necessary to

characterize the robustness of model predictions against numerical, physical, and algorithmic choices. In Sec. 6.1 we report on exploration of parameter variations for each individual star in the binary. We discuss parameters governing the mass transfer phase in Sec. 6.2, the initial binary architecture in Sec. 6.3.1, and the consequences of the assumed SN explosion of the companion in Sec. 6.3.2.

### 6.1. Uncertainties in the single-star physics

*Rotation*—Rotation is a critical ingredient of our models: it governs the equatorial radius and thus  $\omega/\omega_{\text{crit}}$ . Therefore, by assumption, it also controls the mass transfer efficiency through mechanical enhancement of the accretor’s wind (see Sec. 6.2). Through Eddington-Sweet meridional circulations, rotation affects outward mixing from the core, and more importantly inward mixing from the surface. We emphasize that the shellular approximation used in one-dimensional stellar evolution codes might not be appropriate for  $\omega/\omega_{\text{crit}} \simeq 1$ , which our accretor star reaches during RLOF. Decreasing the diffusion coefficient for Eddington-Sweet meridional circulations by a factor of 10 has a very small effect on the evolution of the accretor on the HR diagram. However, the noisiness during the late RLOF phase (beyond point B in Fig. 1) increases in amplitude, confirming that the details of this part of the evolution are sensitive to the treatment of rotational mixing (and its interplay with other processes).

*Angular momentum transport*—Our models assume a Spruit-Tayler dynamo (Spruit 2002) for the transport of angular momentum throughout the evolution. Adopting the stronger angular momentum transport from Fuller et al. (2019) might result in a more efficient spin-down of the surface during RLOF, possibly allowing for more accretion of mass.

The accretor is spun up late, from the surface inwards. It also reaches critical rotation in contrary to the single star models, which are initialized at ZAMS with rigid rotation. Late during the mass transfer, even the weak core-envelope coupling of Spruit (2002) is sufficient for the accretor to achieve rigid rotation. Subsequently, the envelope spins down because of wind mass loss and more importantly the core spins significantly up because of its evolutionary contraction. We expect that following Fuller et al. (2019) the core would lose more angular momentum to the envelope, limiting its ability to spin up as

<sup>4</sup> We obtain the mass fractions of individual elements inverting the definition  $\varepsilon(X) = 12 + \log_{10}(N_X/N_H)$ , where  $N_X$  and  $N_H$  are the number fractions of species  $X$  and H, respectively (e.g., Lodders 2019).

it contracts and decreasing its rotational velocity. Nevertheless, we expect that the difference with single star rotational profiles would remain, albeit possibly smaller, because of the shorter evolutionary time left.

*Thermohaline mixing*—During RLOF, thermohaline mixing in the envelope becomes the dominant mixing process. In MESA, each mixing process is represented by its own diffusion coefficient, and they are then summed together (e.g., Paxton et al. 2011), under the implicit assumptions that mixing processes are independent from each other. This is typically reasonable since locally one process dominates the mixing. In the envelope of our accretor model, initially Eddington-Sweet circulations are dominant, however thermohaline mixing reaches and exceeds their diffusivity late during the mass transfer because of accretion of chemically enriched material from the donor. If fast rotation can physically modify thermohaline mixing processes, this could impact our accretor models. We also computed models with enhanced efficiency of thermohaline mixing (a factor of 100 higher, Schootemeijer et al. 2019), but these proved numerically unstable when accreting CNO-processed material.

*Convective overshooting*—On the basis of nucleosynthesis arguments (e.g., Herwig 2000) and asteroseismology (e.g., Moravveji et al. 2016), an exponentially decreasing overshooting is generally considered preferable. Nevertheless, we have also explored models with a step-function overshooting from Brott et al. (2011). Our fiducial exponential overshooting was chosen to reproduce the width of the main-sequence of the models of Brott et al. (2011), and, not-surprisingly, the qualitative evolution of our fiducial model and models with step overshooting is similar. However, adopting a step overshooting provides a higher diffusivity at the outer edge of the core (cf. exponential decrease), which ultimately impacts the details of the chemical profile at the outer edge of the core, and the morphology of the evolutionary tracks during late RLOF.

*Stellar winds*— $\zeta$  Oph is one of the low-luminosity O-type stars for which Marcolino et al. (2009) found a lower-than-predicted wind mass loss rate. To address the “weak wind problem”, we also attempted running models with artificially decreased wind mass loss rate (e.g., Renzo et al. 2017), but these resulted in supercritically rotating ( $\omega/\omega_{\text{crit}} > 1$ ) post-RLOF accretor stars with untrustworthy numerical results. The solution to the “weak wind problem” is not currently known, but Lucy (2012) and Lagae et al. (2021) suggest that observed mass loss rates might be underestimated, suggesting that the theoretically motivated hot star wind

mass loss rate might still be appropriate to model low luminosity O-type stars.

*Metallicity*—Throughout this study, we assumed an initial metallicity  $Z = 0.01$  informed by the asteroseismology of low mass stars in Upper-Centaurus-Lupus (e.g., Murphy et al. 2021), identified as the parent association for  $\zeta$  Oph by Neuhäuser et al. (2020). Moreover, we have assumed that mass fractions of each element scale with the Solar values (Grevesse & Sauval 1998), which might not be appropriate especially for massive stars (e.g., Grasha et al. 2021). With these assumptions, the initial mass fraction of  $^{12}\text{C}$  and  $^{14}\text{N}$  are lower than the values reported by VH05 for  $\zeta$  Oph. Even though both values increase during mass transfer, our model still slightly under-predicts them. Improved agreement could be obtained changing the ratio of abundances to non-solar values, or by changing the efficiency of downward rotational and thermohaline mixing which dilutes the accreted material into the secondary’s envelope.

We also run a model identical to the one described in Sec. 3, except with  $Z = Z_{\odot} = 0.0142$  (Asplund et al. 2009) with the same composition scaling from Grevesse & Sauval (1998). Qualitatively, the binary evolution remains similar, with the higher metallicity stars having slightly larger radii and cooler  $T_{\text{eff}}$  at a given luminosity. This still produces a stable case B RLOF, however, matching the high present-day  $T_{\text{eff}} = 32\,000 \pm 2\,000$  of  $\zeta$  Oph (e.g., VH05) requires more massive and hotter accretors at higher  $Z$  (see also Sec. 6.3.1).

## 6.2. Uncertainties in the treatment of mass transfer

*Mass transfer efficiency,  $\beta_{\text{RLOF}}$* —We regulate the accretion efficiency through the rotational enhancement of mass loss (e.g., Langer 1998). However, whether critical rotation can effectively stop the accretion of matter is unclear. Popham & Narayan (1991) and Paczynski (1991) argued that accretion of mass (but not angular momentum) might be possible even at or beyond critical rotation.

During RLOF, the total amount of mass lost by the donor is  $\Delta M_{\text{donor}} \simeq 10.6 M_{\odot}$ , of which only  $\Delta M_{\text{accretor}} \simeq 3.4 M_{\odot}$  are successfully accreted by the companion. This corresponds to an overall mass transfer efficiency  $\beta_{\text{RLOF}} \equiv |\Delta M_{\text{accretor}}|/|\Delta M_{\text{donor}}| \simeq 0.32$ , although the accretion efficiency is *not* constant throughout the mass transfer (e.g., van Rensbergen et al. 2006). In our models, the mass transfer efficiency depends on the radial and rotational evolution of the accreting star. During RLOF, the accretor is out of gravothermal equilibrium with significant impact on its radius and ultimately on the amount of mass transferred and its angular momentum. In reality, the gas stream between

the two stars, the hot-spot due to the RLOF stream hitting the accretor's surface (see below), and the geometric distortion of the outer layers because of the centrifugal forces would not follow the spherical symmetry imposed by 1D codes such as MESA.

While the mass transfer efficiency,  $\beta_{\text{RLOF}}$ , and importantly its time-evolution need further attention, it is also likely that this parameter and its evolution depend on the details of the system (masses, mass ratio, period, etc.). For instance, to explain the lower mass sdO+Be binaries found by Wang et al. (2021) it is likely that a larger mass transfer efficiency would be required. Conversely, Petrovic et al. (2005) argued for  $\beta_{\text{RLOF}} \simeq 0.1$  to reproduce WR+O star binaries.

Most studies, especially using rapid population synthesis tools, typically assume a constant  $\beta_{\text{RLOF}}$  and neglect the out-of-equilibrium phase of the accretor and how this can impact the binary and orbital evolution. Alternatively, rapid population synthesis can limit the accretion rate based on the thermal timescale of the accretor (calculated from models in gravothermal equilibrium). Based on this approach, Schneider et al. (2015) found a higher  $\beta_{\text{RLOF}} \simeq 0.7$  for a binary comparable to ours (initially  $M_1 = 20 M_\odot$ ,  $M_2 = 0.7M_1$  with separation  $a \simeq 300 R_\odot$ ), although their  $\beta_{\text{RLOF}}$  is very sensitive to the initial mass ratio and period in this regime.

*Specific angular momentum of accreted material*—This is an uncertain quantity and likely depends on the geometry of the accretion process, and in particular, whether the accretion stream through the first Lagrange point (L1) hits the accretor star directly, or if instead an accretion disk is formed (e.g., de Mink et al. 2013).

We calculate the minimum distance  $R_{\min}$  between the stream coming from L1 and the accretor using the fit from Ulrich & Burger (1976) to the numerical results of Lubow & Shu (1975). We find  $R_{\min} \simeq 1.5 R_\odot < R_{\text{accretor}}$ : this suggests that the stream should hit the accretor directly, without forming an accretion disk. Nevertheless, for the sake of numerical stability, we assume the incoming material and the stellar surface to have the same specific angular momentum. This provides a slow angular momentum accretion and consequent spin-up of the surface.

For a more physically motivated approach, we also attempted calculations using for specific angular momentum of the accreted material  $j = \sqrt{1.7GM_{\text{accretor}}}R_{\min}$ , representative for direct impact of the incoming stream with the accretor (Lubow & Shu 1975). This is typically much larger than the specific angular momentum of the accretor's surface. However, these models proved numerically more unstable and providing less trustworthy results after the accretor is spun up significantly. In

general, allowing for a faster accretion of angular momentum results in a faster spin-up, and a lower overall mass transfer efficiency  $\beta_{\text{RLOF}}$ .

*Specific entropy of the accreted material*—In our models, the composition of the transferred material is determined by the structure of the donor and the mass transfer rate calculated following Kolb & Ritter (1990), but we need to specify its specific entropy when it reaches the accretor surface. We follow the common practice of assuming the specific entropy of the incoming material to be same as the accreting surface. The scenario justifying this hypothesis is that during RLOF the matter is sufficiently optically thin so that radiative processes can rapidly equalize the entropy between the RLOF stream and the accreting surface. However, the very large mass-transfer rates we find (cf. Fig. 2) might result in optically thick flows for which this approximation might not be appropriate.

*Rejuvenation and core growth*—Because of the increase in mass, our accretor star is rejuvenated: its total main-sequence lifetime is longer than the lifetime of a single star born with the final post-RLOF mass of the accretor<sup>5</sup> (e.g., Schneider et al. 2016). The rejuvenation is due to the increase - in mass - of the core region, which brings fresh nuclear fuel inwards. Our results are in agreement with Hellings (1983), while Braun & Langer (1995) did not find any rejuvenation in their accretor models. We attribute this difference to the lack of convective boundary mixing (e.g., overshooting, efficient semiconvection, shear) in their models, which impedes the growth of the core. In our models, the growth of the core is initially driven by convection and overshooting, and to a lesser extent to dynamical shear, while the off-center convective layer of Fig. 5 and Fig. 6 does not contribute to the inward mixing of H-rich material and the rejuvenation itself. We cannot exclude that in the presence of a strong shell undershooting that convective layer would also mix efficiently with the core, enhancing further the rejuvenation effect

### 6.3. Comparing to $\zeta$ Oph

#### 6.3.1. Variations of initial binary parameters

The initial donor mass  $M_1$ , mass ratio  $q \equiv M_2/M_1$ , and the period of the progenitor binary of  $\zeta$  Oph cannot be directly constrained from observations. We have explored variation in these values, and the qualitative behavior of the models is similar. Shorter initial periods

<sup>5</sup> But not significantly longer than the lifetime of a single star of its initial, pre-RLOF mass.

results in larger post-RLOF orbital velocities, and thus larger runaway velocities if the binary is disrupted at the first SN (see Sec. 6.3.2). For example, taking  $P=75$  days (cf. 100 days in our fiducial model), the binary still experiences stable case B mass transfer, but the post-RLOF orbital velocity of the accretor is about  $60 \text{ km s}^{-1}$ , that is  $\sim 10 \text{ km s}^{-1}$  higher than in our fiducial model, because of the larger orbital velocity in the tighter binary system.

Increasing the donor mass also has a similar effect on the post-RLOF orbital velocity of the accretor. Using  $M_1 = 30 M_\odot$  (cf.  $25 M_\odot$  in Sec. 3),  $M_2 = 17 M_\odot$ , and  $P=100$  days, we obtain a post-RLOF velocity of  $65 \text{ km s}^{-1}$ . However, this produces a stripped donor of  $\sim 16 M_\odot$  at RLOF detachment, with stronger wind mass loss rate. Therefore this binary is expected to widen relatively more than our fiducial model of Sec. 3, slowing down the accretor. The increased mass of the stripped donor star could also imply a lower chance of exploding for the donor, which might instead collapse to a black-hole (however, see Sec. 6.3.2).

The higher  $M_1$  does not significantly change the post-RLOF total mass of the accretor, with  $M_2$  remaining about  $\sim 20.5 M_\odot$ , since in our models accretion is regulated mostly by the spin up of the accretor, and we do not couple the specific angular momentum of the transferred material to the orbit or the donor's spin.

However, changing the initial mass ratio also changes the difference between the main-sequence lifetime of the two stars, and thus how far along the main sequence the accretor is at the onset of RLOF. The observed position of  $\zeta$  Oph on the HR diagram, particularly its relatively high  $T_{\text{eff}}$  are difficult to reproduce assuming initially less massive accretors (which would remain too cool even after accreting mass), or more equal initial mass ratio (which would produce an accretor that is too evolved and cool at the onset of mass transfer).

### 6.3.2. The explosion of the donor star

Throughout this study, we have assumed the “binary SN scenario” to explain the runaway nature of  $\zeta$  Oph: after the mass transfer phase, the explosion of the donor disrupts the binary and ejects the accretor at roughly its pre-explosion orbital velocity (e.g., Blaauw 1961; Renzo et al. 2019). This fate occurs to the majority of massive binary systems, and  $\zeta$  Oph might be the best example of it (e.g., Blaauw 1952, 1961; Hoogerwerf et al. 2000). Neuhauser et al. (2020) suggested not only the companion successfully exploded producing the pulsar PSR B1706-16 and ejecting  $\zeta$  Oph, but also that the explosion produced radioactive  $^{60}\text{Fe}$  which polluted Earth.

From kinematic and orbital considerations they estimated the pulsar received a natal kick of  $253 \pm 54 \text{ km s}^{-1}$ , which would be sufficiently large to unbind the binary which has  $v_{\text{orb}} = \sqrt{G(M_1 + M_2)/a} \simeq 135 \text{ km s}^{-1}$  at the end of our binary simulation (blue diamond in Fig. 1) and this will decrease further in the remaining time to the donor’s core-collapse (Kalogera 1996; Tauris 2015).

The SN ejecta mass would depend on the post-RLOF wind mass loss of our donor star, which is very uncertain (e.g., Renzo et al. 2017; Vink 2017; Gilkis et al. 2019). YG: Vink 2017 and Gilkis+19 don't say that the wind mass loss rate for stripped stars are uncertain, but they create/use a new prescription for wind mass loss rate. I would suggest a simple change to "for a comparison of different wind mass loss rates, see Gilkis+19".

At the end of our binary evolution simulation, our stripped donor is  $\sim 9.4 M_\odot$ , with a surface H fraction of  $X \lesssim 0.2$  for a layer of  $\Delta M \simeq 2.5 M_\odot$ . Its wind mass-loss rate is  $\sim 10^{-5} M_\odot \text{yr}^{-1}$ , calculated assuming the empirical Wolf-Rayet wind mass loss prescription from Nugis & Lamers (2000, see also Fig. 2). We expect that the donor will explode in a H-free type Ib supernova. Although our stripped donor is rather massive, recent studies hints at a higher “explodability” of donor stars in binary systems (e.g., Schneider et al. 2021; Laplace et al. 2021; Vartanyan et al. 2021).

We have neglected the impact of the explosion on the structure of the accretor star. At the time of the explosion, the accretor subtends a solid angle  $\sim R^2/a^2 \simeq 2 \times 10^{-3}$  steradians with  $R$  the accretor radius and  $a$  the binary separation. We neglect the post-RLOF wind-driven orbital widening for this estimate. The blast wave will hit the accretor causing mass loss – directly via ablation and by injecting energy in the envelope, inflating it and enhancing its wind (Wheeler et al. 1975; Tauris & Takens 1998; Podsiadlowski 2003; Hirai et al. 2018; Ogata et al. 2021). Because of the SN shock, the just ejected new runaway star might appear bloated and redder (long before it overtakes the slowing SN remnant). The impact of this brief out of thermal equilibrium phase on the stellar spin should be investigated further.

Using 2D hydrodynamic simulations of the star-SN ejecta interactions in close binaries ( $a \lesssim 60 R_\odot$ , cf.  $a \gtrsim 343 R_\odot$  in our fiducial binary model), Hirai et al. (2018) and Ogata et al. (2021) found that the companion star recovers its pre-explosion luminosity and effective temperature within a few years to decades, and the amount of mass removed by the SN shock is  $\lesssim 10^{-2} M_\odot$ . The SN ejecta might also pollute the surface of the runaway depositing processed nuclear material (e.g., Przybilla et al. 2008; Suda et al. 2021). However, for the

large final separation of our model, little pollution is expected and enhanced mass loss and inward mixing might quickly dilute any signature below detectable levels.

## 7. SUMMARY & CONCLUSIONS

The impact of mass transfer on the structure and evolution of accretor stars in massive binaries has received relatively little attention in the literature. To investigate this, we have performed MESA calculations of massive binaries evolving two coupled stars simultaneously.

As a first application, we focused on finding a model in which the accretor properties are in qualitative agreement with observations of the nearest O-type star to Earth. This is the runaway star  $\zeta$  Oph, which has long been suggested to be a former accretor star ejected from a binary at the core-collapse of the donor star (binary SN scenario, Blaauw 1961). However, our models are also informative for the generic population of massive stars accreting in binaries.

### 7.1. Reproducing $\zeta$ Oph

We found that the main features of  $\zeta$  Oph can be reasonably well reproduced using standard stellar physics assumptions for the treatment of mass transfer, chemical mixing, and rotation. Our choices are described in Sec. 2 and Appendix A.

Our fiducial model is a binary starting with  $M_1 = 25 M_\odot$ ,  $M_2 = 17 M_\odot$ , and  $P = 100$  days at metallicity  $Z = 0.01$  (see Sec. 3). This binary experiences stable thermal-timescale Roche lobe overflow after the end of the main sequence of the donor (case B).

The accretor compares well with observations of  $\zeta$  Oph about  $1.5 - 2$  Myrs after the end of mass transfer, corresponding to the remaining donor's lifetime at the end of our simulations plus the kinematic age of  $\zeta$  Oph. Specifically, the position on the HR diagram (cf. Fig. 1), the runaway space velocity (estimated based on the accretor's orbital velocity), the surface composition and rotational velocity (cf. Tab. 1) are in the right ballpark.

Our model of  $\zeta$  Oph differs significantly from previous studies: in contrast with the accretor models of van Rensbergen et al. (1996), in our model the  $^{14}\text{N}$ - and  $^4\text{He}$ -rich surface composition is not the result of pure outward rotational mixing. Instead, this material is transferred from the receding core of the donor star and mixed from the surface inwards into the accretor by meridional circulations and more importantly thermohaline mixing. Thus, the present day surface mass fractions of  $\zeta$  Oph place a coupled constraint on the mass transfer efficiency and mixing in the accretor. Our results suggest  $\zeta$  Oph should not be used to calibrate models of rotational mixing in single star models.

We emphasize that the surface composition alone would not be a smoking-gun of a past in a binary, especially given the large uncertainties in the treatment of rotation and mixing in stellar evolution models. In our models, the accretion of  $^{14}\text{N}$  from the donor star allows the accretor star to be simultaneously fast rotating and  $^{14}\text{N}$ -rich, while single rotating stars can only mix upward  $^{14}\text{N}$  from their core, and have significantly spun down by the time they become  $^{14}\text{N}$ -rich. **YG: I am not convinced by this argument - looking at Fig. 7, the fast rotating single stars seem to have enriched surfaces already after 4Myr or so (second line in the diagram, orange color), then their surfaces become slightly more enriched with time, but not much. From the veq-t diagram, it doesn't seem to me that they are spun down at that time.** Moreover, alternative scenarios where  $\zeta$  Oph evolved as a single, fast-rotating star require ad-hoc explanations for the runaway velocity, and have been shown by VH05 to struggle in reproducing surface mass fractions, apparent age, mass, and rotation rate simultaneously.

The surface rotation rate of the accretor post-mass-transfer is always higher than the rotation rate of single stars initialized with half-critical rotation, but might still be on the low side compared to  $\zeta$  Oph. However, the wind spin down might be overestimated in our models (weak wind problem, cf. Marcolino et al. 2009; Lucy 2012; Lagae et al. 2021).

We also tested the robustness of our fiducial model against variations in the initial parameters and algorithmic representation of physical phenomena, discussed in Sec. 6. Less massive accretors remain too cool throughout the evolution to be compatible with  $\zeta$  Oph, and initial mass ratios closer to unity lead to a more evolved accretor at the onset of mass-transfer, again resulting in too cool temperatures. Increasing the donor's initial mass might result in stripped stars unlikely to form a neutron star in their final SN explosion.

### 7.2. Accretors are not single rotating stars

Our models also highlight some general differences between accretors in massive binaries and stars evolving as single throughout their life. These might be important for several sub-fields of astrophysics, including asteroseismology, stellar populations, and time-domain and gravitational waves observations.

The first notable difference we find is the internal rotation profile. Single rotating stars are usually initialized as rigid rotators at birth, and throughout their evolution they spin down due to wind mass loss. Conversely, accretors are spun up later during their main-sequence evolution and from the surface inward. Moreover, for single stars, the maximum rotation rate, that

is the one assumed at the beginning of the evolution, is poorly understood theoretically and observationally (e.g., Ramírez-Agudelo et al. 2013; Ramírez-Agudelo et al. 2015). Conversely, accretors in binaries reach critical rotation  $\omega/\omega_{\text{crit}} \simeq 1$  (e.g., Packet 1981). The later spin-up and higher achieved rotation rate allow the accretor star to remain a fast rotator until the end of its main sequence.

The angular momentum accreted at the surface of the accretor is transported into the core (by the Tayler-Spruit dynamo in our simulations). This results in a much faster rotating helium core at the end of the main sequence compared to single stars. Such a fast spinning helium core has potential implications for the final explosion and the resulting compact object born from the accretor star in an interacting binary system.

Finally, in our models, the accretion of mass leads to rejuvenation and also the formation of a off-center convective layer above the main-sequence core (cf. Fig. 5). The latter ultimately results in a sharper density drop at the core edge (cf. Fig. 6), and a flatter density profile close to the end of the main sequence. If physical, the presence of such a convective layer could in principle be probed using asteroseismology. Depending on how the accretor (and the binary) evolves in the future, this difference could be crucial in determining the outcome of common envelope events between massive stars and

compact objects.

Improving our understanding of the evolution of the initially less massive stars in massive binary systems is crucial for the upcoming large surveys, stellar kinematics, and for the understanding of the evolution of gravitational-wave progenitors in isolated binaries. Although presently single, the nearest O-type star to Earth,  $\zeta$  Oph, can be used as an anchor point for the modeling of accretors. Our models demonstrate that a broad agreement with observations can be achieved with standard stellar evolution assumptions. Future efforts should extend these models to a wider mass, period, mass ratio, and metallicity range to investigate the impact of binary evolution on the life, explosion, and afterlife of the secondary stars in massive binary systems.

*Software:* MESA (Paxton et al. 2011, 2013, 2015, 2018, 2019), mesaSDK (Townsend 2018), ipython/jupyter (Pérez & Granger 2007), matplotlib (Hunter 2007), mesaPlot (Farmer 2018), NumPy (van der Walt et al. 2011).

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## APPENDIX

### A. MESA SETUP

We use MESA version 15140 to compute our models. The MESA equation of state (EOS) is a blend of the OPAL Rogers & Nayfonov (2002), SCVH Saumon et al. (1995), PTEH Pols et al. (1995), HELM Timmes & Swesty (2000), and PC Potekhin & Chabrier (2010) EOSes.

OPAL (Iglesias & Rogers 1993, 1996) provides the main radiative opacities, with low-temperature data from Ferguson et al. (2005) and the high-temperature from Buchler & Yueh (1976). Electron conduction opacities are from Cassisi et al. (2007).

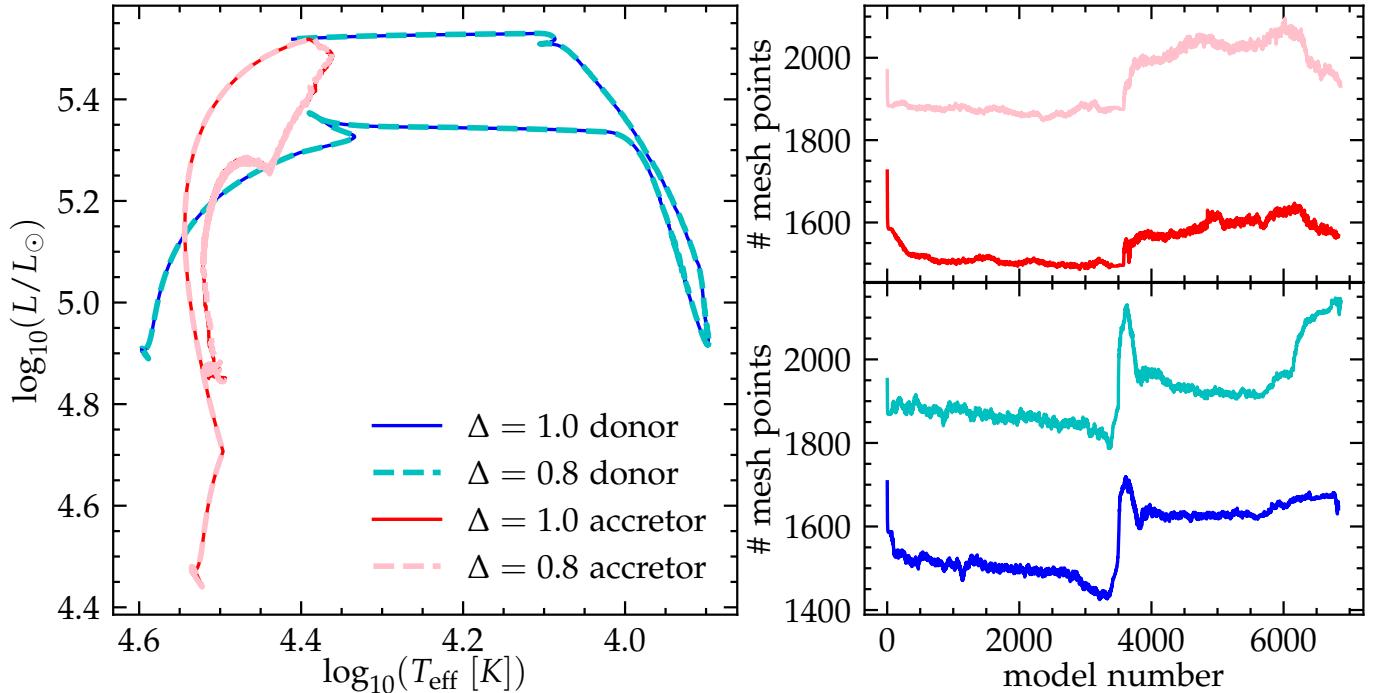
Nuclear reaction rates are a combination of rates from NACRE (Angulo et al. 1999), JINA REACLIB (Cyburt et al. 2010), plus additional tabulated weak reaction rates Fuller et al. (1985); Oda et al. (1994); Langanke & Martínez-Pinedo (2000). Screening is included via the prescription of Chugunov et al. (2007). Thermal neutrino loss rates are from Itoh et al. (1996). We use a 22-isotope nuclear network (approx\_21\_plus\_cr56).

The inlists, processing scripts, and model output are available at <https://doi.org/10.5281/zenodo.4701565>.

### B. RESOLUTION TESTS

We extensively checked the numerical convergence of our stellar evolution calculations with increasing number of mesh points. Fig. 8 shows that all the main features described here do not vary when increasing the spatial resolution by increasing the number of mesh points (i.e., decreasing `mesh_delta_coeff`). The right panel shows the number of mesh points for the accretor (top) and donor (bottom) as a function of the model number (akin to an arbitrary time coordinate). About  $\sim$ 7000 MESA models are used to compute the binary evolution. The higher resolution run has  $\sim 20\%$  more mesh points. The left panel shows the evolution on the HR diagram until the detachment of the binary for the two accretor models (pink/red) and the two donor models (blue/cyan).

Similarly, we tested the numerical convergence with decreasing timestep size. This can be done decreasing the parameter `mesh_time_coeff`. However, we were unable to successfully compute models at higher temporal



**Figure 8.** Left: HR diagram comparison for our fiducial binary model varying the number of mesh points. We only show the evolution until our definition of RLOF detachment. Right: number of mesh points as a function of timestep number. In both panels, the blue/cyan tracks show the donor stars, the red/pink tracks show the accretor. Thicker dashed lines correspond to the models at higher resolution (i.e., lower  $\Delta$  which indicates the value of `mesh_delta_coeff`).

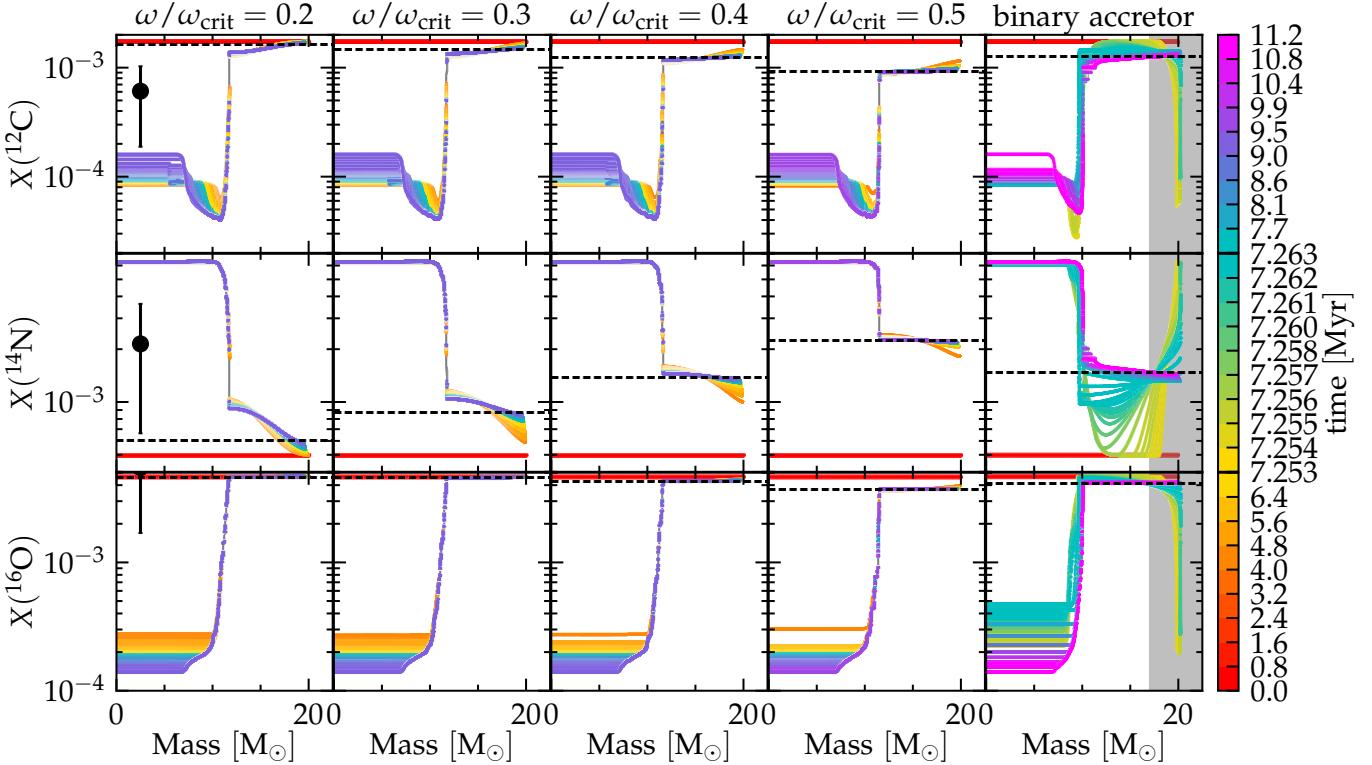
resolution. Partial results show a good agreement with our fiducial model until MESA becomes unable to find a satisfying numerical solution to the stellar structure equations (typically during RLOF). Lower temporal resolution models showed a similar qualitative agreement but increased noisiness during the late RLOF phase. For our fiducial model the adaptive timestep size never exceeds  $10^{3.8}$  years with typical pre-RLOF timesteps of the order of  $10^{3.2}$  years and sub-decade (occasionally sub-year) during RLOF. The main factor limiting the timestep sizes is the change of surface angular momentum in both stars during the mass transfer.

### C. INTERNAL COMPOSITION PROFILE EVOLUTION

Fig. 9 compares the internal evolution of the composition profile of single rotating stars with our accretor model. We show mass fractions of  $^{12}\text{C}$  and  $^{16}\text{O}$  to complement the mass fraction of  $^{14}\text{N}$  shown in Fig. 7, and reproduced also in the middle panel of Fig. 9.

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**Figure 9.** Same as Fig. 7, but for  $^{12}\text{C}$  (top panel), and  $^{16}\text{O}$  (bottom panel). The first four panels show single rotating stars of initially  $20 M_{\odot}$ , the rightmost panel shows the accretor in our fiducial binary. The middle panel is exactly the same as Fig. 7. The red lines mark the initial mass fractions, black dashed lines indicate the TAMS surface mass fraction, and the black error bars in the first column indicate the surface composition of  $\zeta$  Oph inferred by VH05 using the surface H mass fraction from our model.

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