

	QSO	XRB	TDE	References
M_{BH}	10^{6-9}	$10^{0-1.8}$	$< 10^{6-7} (??)$	
\dot{M}	$\sim 1 M_{\odot}/\text{yr}$	$\sim 1 M_{\odot}$	$1 - 10 M_{\odot}/\text{yr}$	
\dot{M} (\Rightarrow LC shape?)	?	?	?	
a (BH spin)	low, mode, high	generally high		
$\log L/L_{\text{Edd}}$	-2 - 0	0.01-1		
preferential L_{Edd} ??	<i>maybe</i> for CLQs			
ang. momen (accn disk)				
$\frac{d}{dt}$ ang. momen	?	?	?	
fuel source	accn disk	accn disk	star	
opacity				
accn disk wind??				
host galaxy	\sim whole population*	—	post-starburst preference	
	*not local AGN	—		
evolution with z	peaks at $z \sim 2-3$	—	?	
binary BHs?	\checkmark		\times (probably)	
BLR?	\checkmark	No	\checkmark (but weird?)	
CL-BLR?	\checkmark	No (but...)		
BLR in polarimatory?	Yes	n/a	?	
He II ?	rare		\checkmark	
Coronal Lines	Sometimes	?	Sometimes	
Fe opacity important?	\checkmark	?	?	
[O III] /H β	defines Type #	(no NLR/BLR)	?	
Like a NLS1?	“no”	Yes	?	

Table 1: github.com/d80b2t

	QSO	XRB	TDE
PSD in opt.	changes with \dot{M}		
PSD in X-ray	no evolution		
PSD in IR	?		n/a?
X-rays	yes	By definition	No (except when there are)
Hard state?		Yes	
X-ray variability? (soft)	Yes	Yes	
X-ray variability? (hard)			
corona?	Yes	✓/× (Big debate)	?
Radio variability	✓		
Infrared variability	✓		✓ (probably)
Is x important?			
Viscous timescale	Incredibly		
X-ray Reprocessing	Yes		
IR Reprocessing	Yes		
Atomic Physics			
Challenges SS73?	AGN disk (x4) too big		

From Nadia Blagorodnova::

Using the last $M - \sigma$ relations for TDE hosts, they have a figure showing that preferentially they are close to L_{Edd} , but the range is 0.01-1 of L_{Edd} : <http://adsabs.harvard.edu/abs/2017arXiv170608965W>

From Ohad Shemmer::

Going back to my (and others) "NLS1 philosophy", in a nutshell: NLS1s have been identified back in 1986 as a "strange new class" of broad-line Seyferts. Many things happened since then, and 1999 should have pretty much marked the end of the "NLS1" terminology. Unfortunately, many folks are still having a hard time disengaging from this exotic "NLS1 class".

These sources are simply understood as type 1 AGN lying at some extreme corner of parameter space, driven mainly by high L/L_{Edd} . So their BELR lines are relatively narrow with respect to their luminosity, indicating high L/L_{Edd} and relatively low M_{BH} . This also dictates extremely low [O III]/H β ratios, strong Fe II lines, weak C IV lines, etc. etc.

So, for the Table, I'd simply remove (safely) the last two lines, i.e., "[O III]/H β " and "like NLS1", since these two lines are implicit in the $\log L/L_{\text{Edd}}$ line above.

Also, I think you can safely change to a "Yes" the XRB evolution with redshift; see, e.g., Lehmer+16, ApJ, 825, 7, and refs. therein.